

# Neutrino measurements and abundances

F. L. Villante

<sup>1</sup>University of L' Aquila and LNGS-INFN

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(\*) partly based on work done in collaboration with **Aldo Serenelli (ICE and IEEC, BCN)**

# The Standard Solar Model (SSM)

Our comprehension of the Sun is based on the **Standard Solar Model (SSM)**.

This implies:

- ✓ Stellar structure equations;  
( $\alpha$  = mixing length)
- ✓ Chemical evolution paradigm:  
ZAMS homogenous model ( $Y_{\text{ini}}, Z_{\text{ini}}$ )  
Nuclear reactions + elemental diffusion
- ✓ Knowledge of the properties of solar plasma  
(i.e. opacity, equation of state, nuc. cross sections);

## No free parameters

The unknown quantities

-  $\alpha, Y_{\text{ini}}, Z_{\text{ini}},$

are fixed by requiring that the present Sun ( $t_{\text{sun}}=4.57$  Gyr) reproduces its observational properties

-  $R_{\text{sun}}, L_{\text{sun}}, (Z/X)_{\text{surf}}$

# The Standard Solar Model (SSM)

The predictions of SSMs can be **falsified** by other observations. e.g.:

## - *Solar neutrinos:*

Hydrogen fusion in the solar core produce a huge amount of neutrinos that can be measured in suitable detectors (Davis 1964, Bahcall 1964)



Solar Neutrino Problem  
Nuclear energy generation (cross sections, etc.)

## - *Helioseismology:*

Solar oscillations originally discovered by Leighton et al. 1962 and interpreted as standing acoustic waves

Elemental Diffusion  
Opacity, EoS, ...

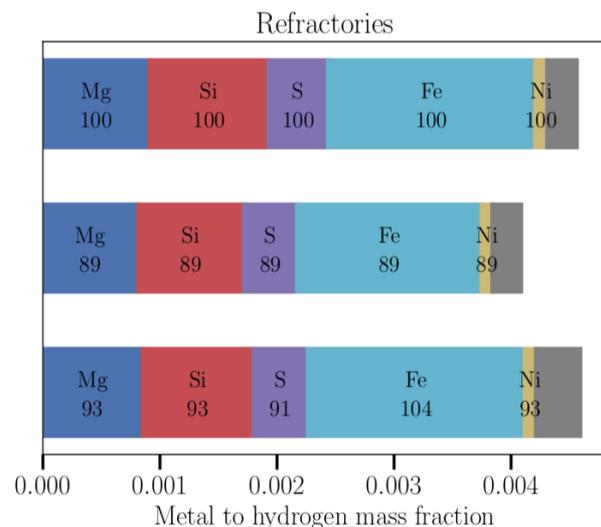
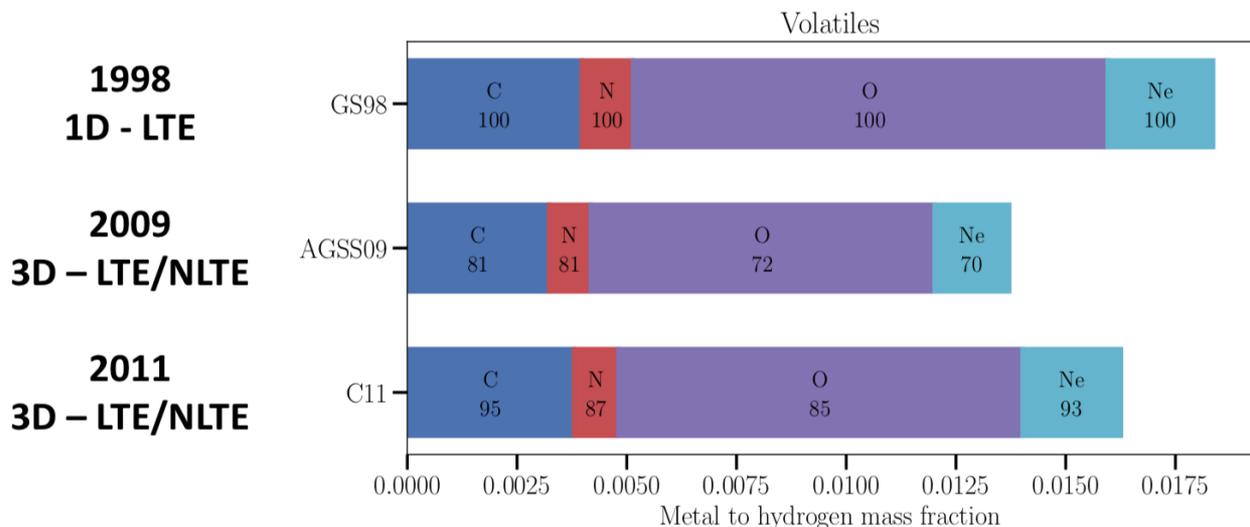
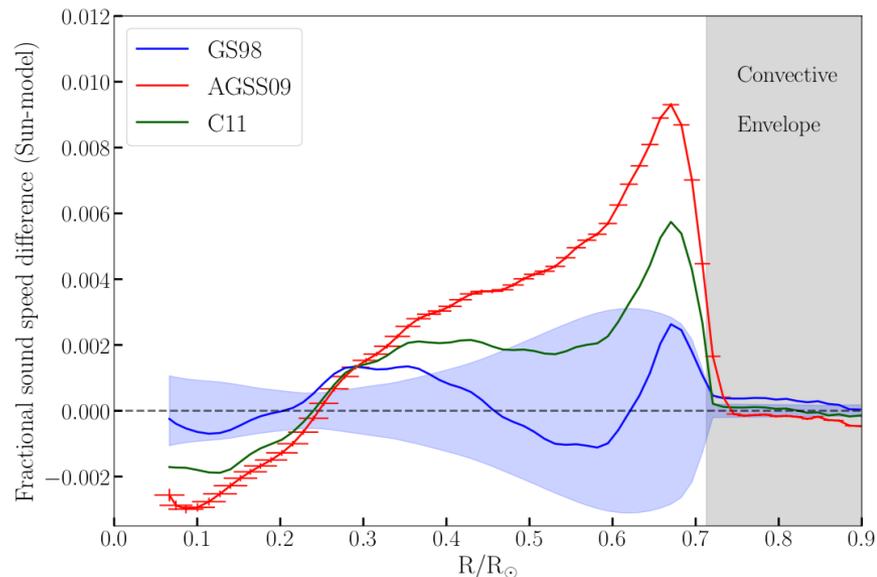
Constant improvement in SSM constitutive physics was triggered during last decades by solar neutrino and helioseismic data.

Impressive agreement with helioseismic data .... till few years ago

... till few years ago

**Downward revision of solar surface abundances**

**Solar surface composition** is a fundamental input for SSMs → determined with spectroscopic techniques (3D models of solar atmosphere, NLTE corrections, ...)



Orebi Gann et al. 2021/2022

GS98: Grevesse & Sauval 1998 – AGSS09: Asplund, Grevesse et al. 2009 – C11: Caffau et al. 2011

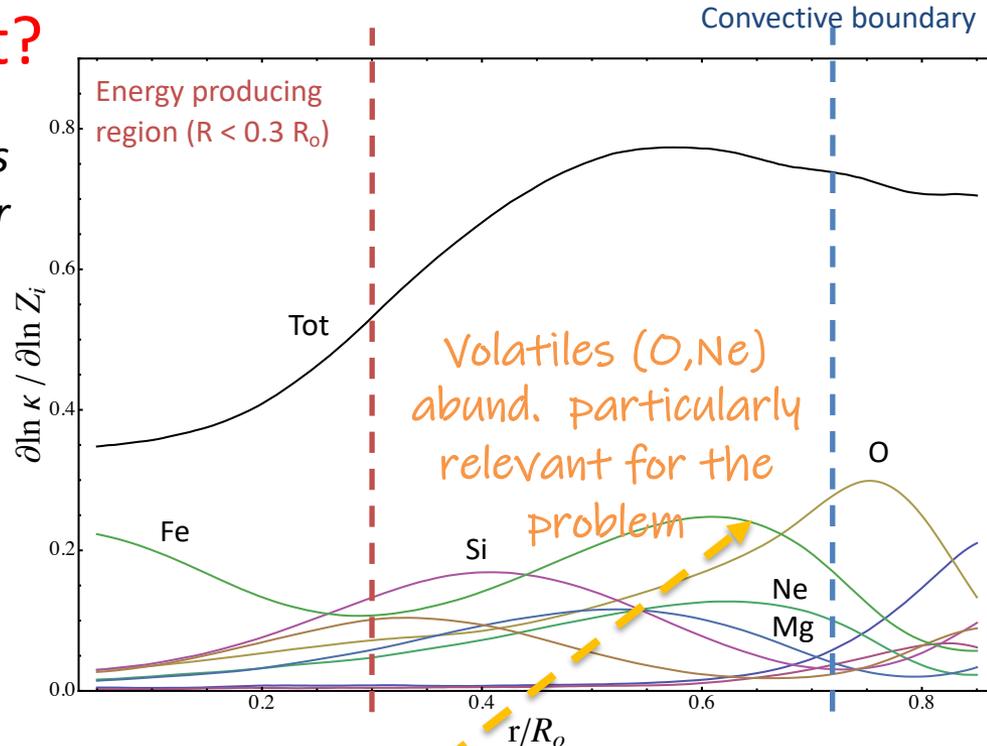
# Why metals are so important?

A change of the solar composition affects the efficiency of radiative energy transfer in the core of the Sun

Composition opacity change:

$$\delta\kappa(r) = \sum_j \frac{\partial \ln \kappa(r)}{\partial \ln Z_j} \delta z_j$$

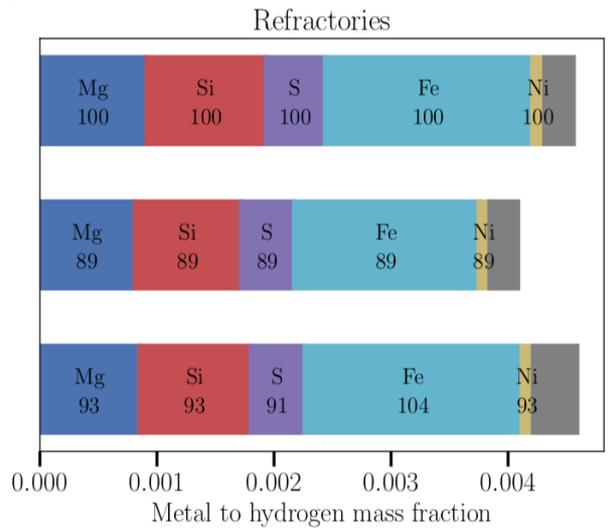
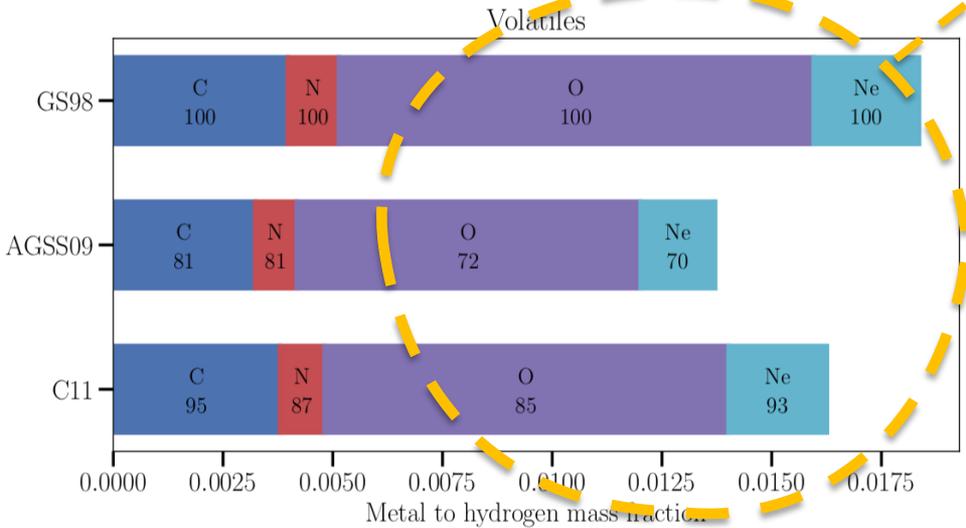
Different temperature stratification



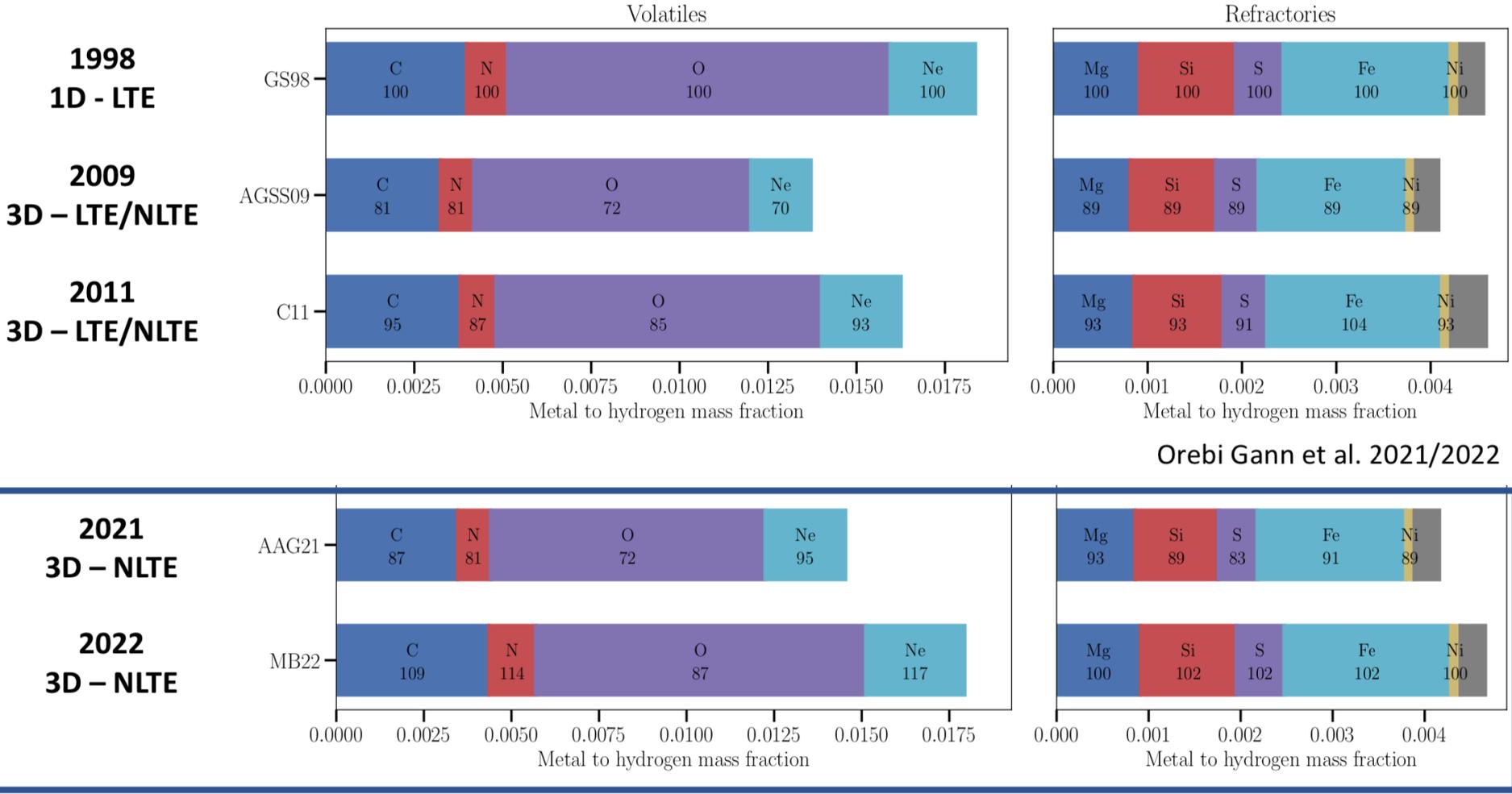
**1998**  
1D - LTE

**2009**  
3D - LTE/NLTE

**2011**  
3D - LTE/NLTE



# Updates in solar abundances



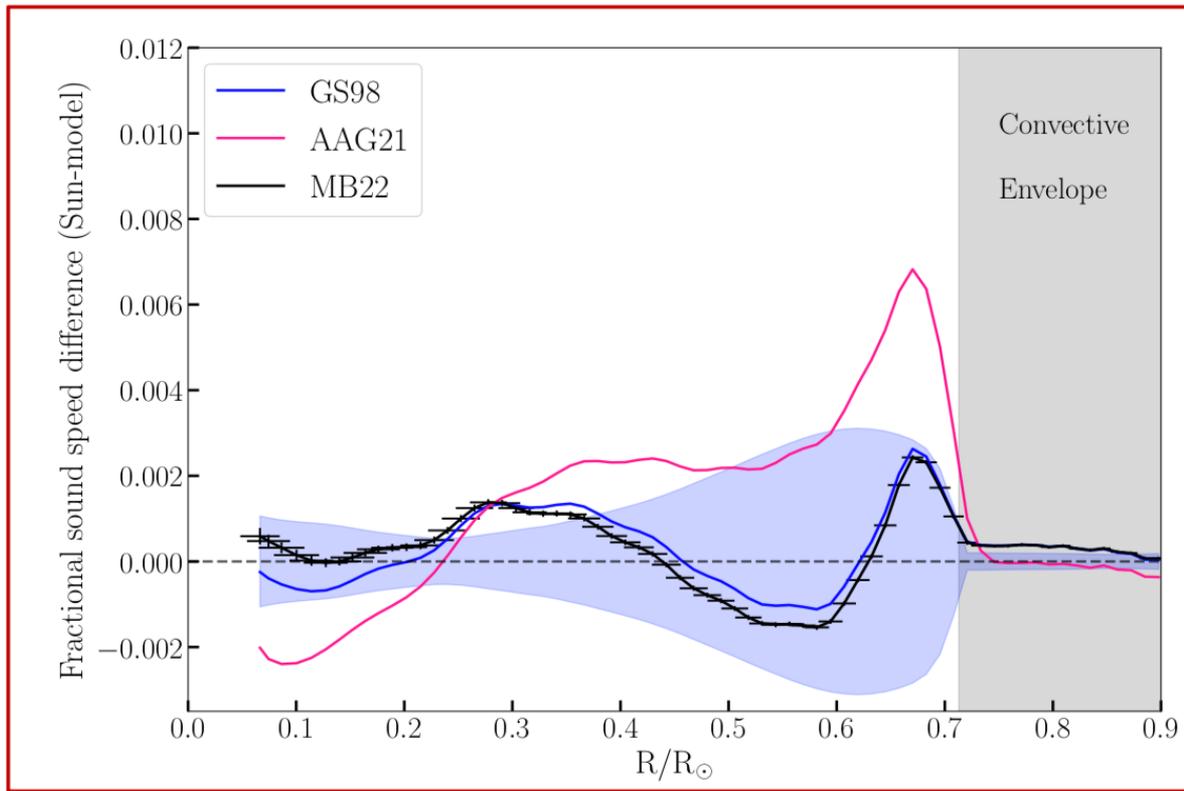
**AAG21: Asplund, Amarsi & Grevesse 2021 – MB22: Magg, Bergemann et al. 2022**

Solar composition “dichotomy” still persists but now based on 3D NLTE abundances

# Helioseismic results

| Model      | $R_{CZ}/R_{\odot}$ | $Y_S$  |
|------------|--------------------|--------|
| MB22-phot  | 0.7123             | 0.2439 |
| MB22-met   | 0.7120             | 0.2442 |
| AAG21      | 0.7197             | 0.2343 |
| AGSS09-met | 0.7231             | 0.2316 |
| GS98       | 0.7122             | 0.2425 |
| C11        | 0.7162             | 0.2366 |

Situation in 2022



*Helioseismic determinations*

$$R_b/R_{\odot} = 0.713 \pm 0.001$$

$$Y_b = 0.2485 \pm 0.0035$$

Magg et al. 2022

**HZ surface composition** provide a better description of helioseismic data

# Can we conclude that LZ abundances are wrong?

The interpretation is complicated by the **opacity-composition degeneracy**.

$$\delta\kappa(r) = \delta\kappa_I(r) + \sum_j \frac{\partial \ln \kappa(r)}{\partial \ln Z_j} \delta z_j$$

*Intrinsic opacity change (e.g. opacity table "errors")* ←

→ *Composition opacity change*

Q. Is opacity of the solar plasma sufficiently well calculated?

Note that the **shape of opacity variation** may be as **important as its overall magnitude**

$$\delta Q = \int dx K_Q(x) \delta\kappa(x)$$

*Fractional variation of observable quantity Q* ←

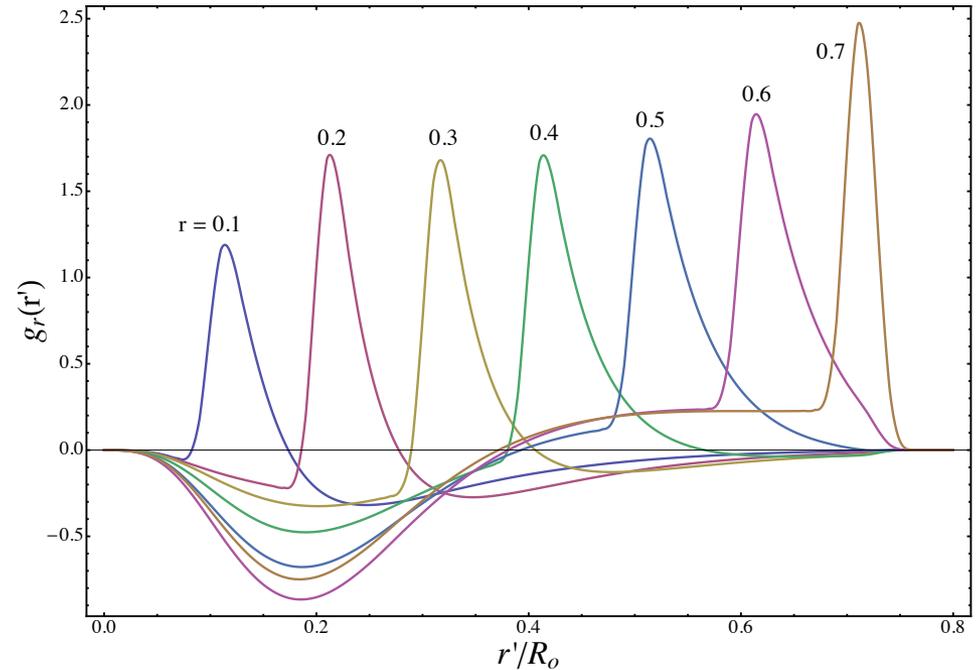
→ *Fractional variation of opacity at a given point x*

$$\left\{ \begin{array}{l} Q = \text{generic observable (e.g. surf. helium, conv. radius, sound speed, } \nu \text{ fluxes)} \\ x = \frac{r}{R_o}, \ln\left(\frac{T}{T_c}\right), \dots \\ K_Q(x) = \text{Opacity Kernel} \end{array} \right.$$

Tripathy & Christensen-Dalsgaard, 1998 (static or w/o diffusion)  
Villante et al. 2010 (Linear Solar Models)  
Villante, 2010 (LSM, diffusion)  
Vinyoles et al, 2017 (Full evolut. Models)

# The sound speed kernels

$$\delta u(r) = \int dr' K_u(r, r') \delta \kappa(r')$$



The kernels are not positive definite  $\rightarrow$  compensating effects can occur ...

$$\delta u_0(r) = \int dr' K_u(r, r') \simeq 0$$

The sound speed is *insensitive to a global rescaling of opacity*

$$\frac{GMm_u}{R} \sim \frac{k_B T}{\mu} = \frac{P}{\rho} = u$$

# The convective radius and the surface helium abundance

## Convective radius:

$$\delta R_b = \int dr K_R(r) \delta \kappa(r)$$

$$\begin{aligned} \delta R_b &= 0.12 A_{\text{in}} - 0.14 A_{\text{out}} \\ &\simeq 0.13 (A_{\text{in}} - A_{\text{out}}) \end{aligned}$$

$$\delta R_b = -0.02 A_0 - 0.10 A_1$$

## Surface helium:

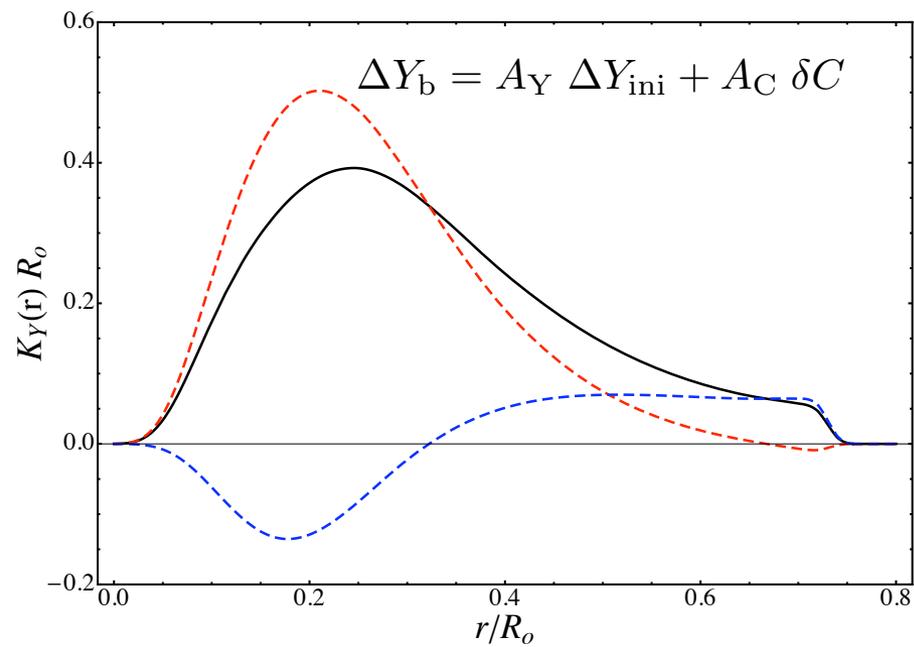
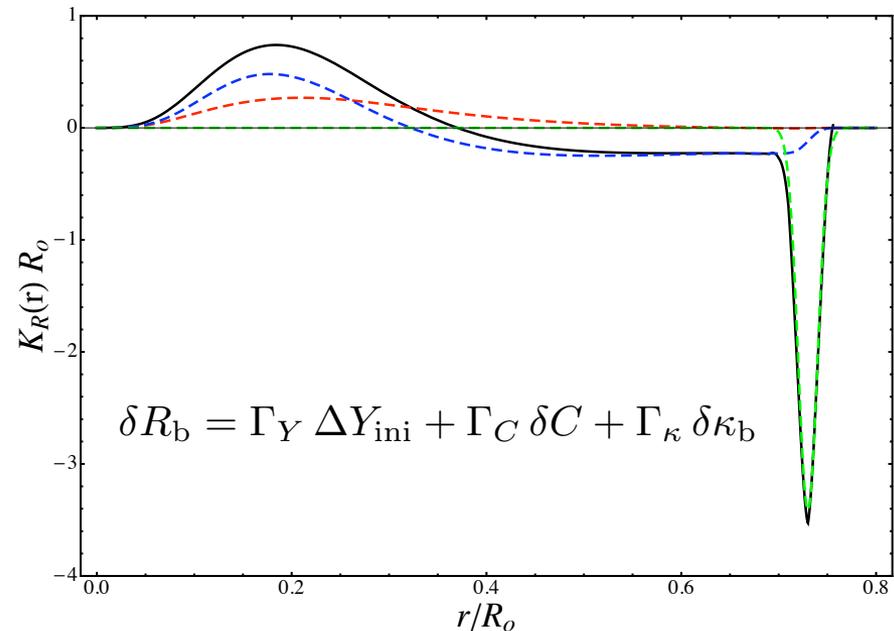
$$\Delta Y_b = \int dr K_Y(r) \delta \kappa(r)$$

$$\begin{aligned} \Delta Y_b &= 0.073 A_{\text{in}} + 0.069 A_{\text{out}} \\ &\simeq 0.07 (A_{\text{in}} + A_{\text{out}}) \end{aligned}$$

$$\Delta Y_b = 0.142 A_0 + 0.062 A_1$$

To reproduce helioseismic results:

$$A_{\text{in}} = 0.07 \pm 0.04 \quad A_{\text{out}} = 0.21 \pm 0.04$$



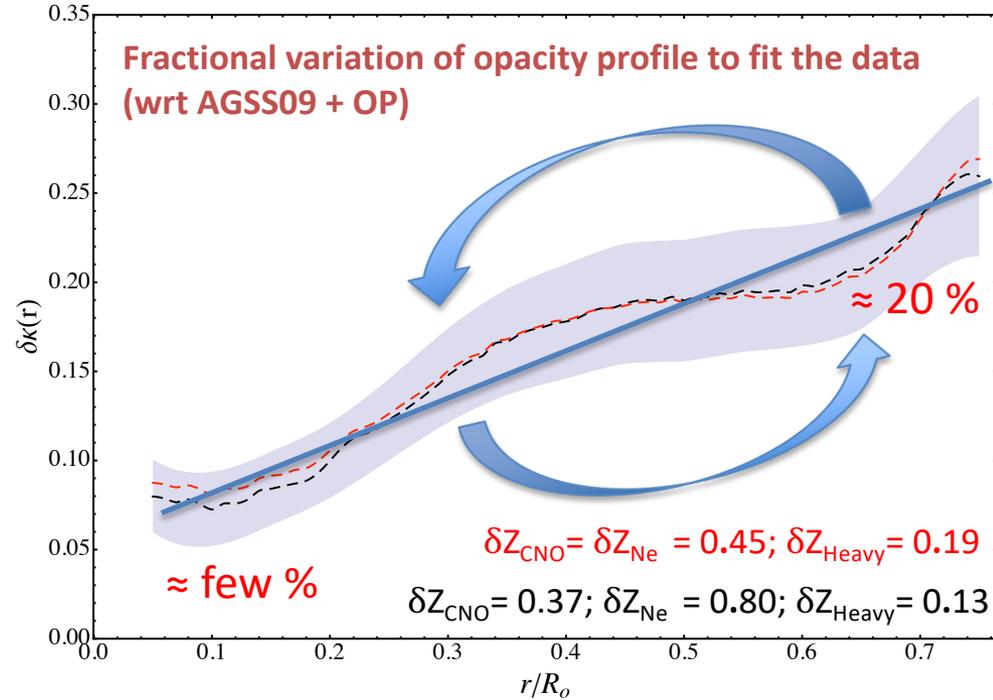
# The solar opacity profile

F.L. Villante and B. Ricci - *Astrophys.J.*714:944-959,2010  
 F.L. Villante – *Astrophys.J.*724:98-110,2010  
 F.L. Villante, A. Serenelli et al., *Astrophys.J.* 787 (2014) 13

The “optimal” composition/opacity profile of the Sun can be determined from obs. data

Note that:

- The sound speed and the convective radius determine **the tilt** of  $\delta\kappa(r)$  (but not **the scale**)
- The surface helium and the neutrino fluxes determine **the scale** for  $\delta\kappa(r)$



The **opacity at the bottom of the convective envelope** can be directly inferred from **helioseismic observables**:

$$\delta\kappa_b = C_Y \Delta Y_b + C_R \delta R_b + C_\rho \delta\rho_b = 0.24 \pm 0.03$$

(wrt AGS05 + OP)

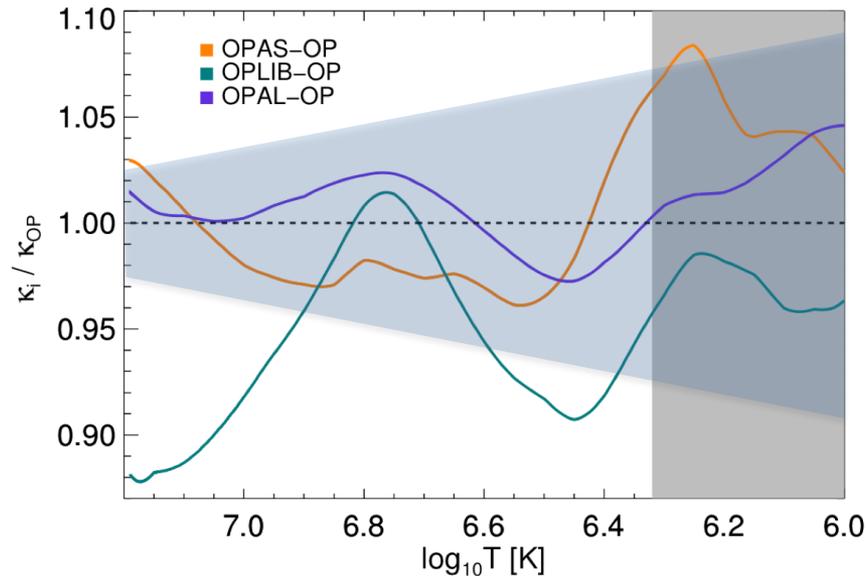
$$\begin{cases} C_Y = 6.27 \\ C_R = -11.71 \\ C_\rho = -1.58 \end{cases}$$

# Parameterizing uncertainty in opacity calculations ...

Opacity uncertainty in B16-SSMs is parameterized as:

$$\delta\kappa(T) = \kappa_a + (\kappa_b/\Delta) \ln(T/T_C)$$

$\kappa_a, \kappa_b$  = random variables  
(means equal to 0 and variances  $\sigma_a = 0.02$  and  $\sigma_b = 0.067$ )



This prescription is motivated by:

- Opacity calculations more accurate at the solar core (~2%) than at the base of the convective envelope (~7%);

- It avoids underestimating the opacity error contribution to sound speed and convective radius (sensitive to tilt and not to scale of opacity)

... but **it still remains** a very simplified description of the real situation

Neutrinos

# Hydrogen Burning: PP chain and CNO cycle

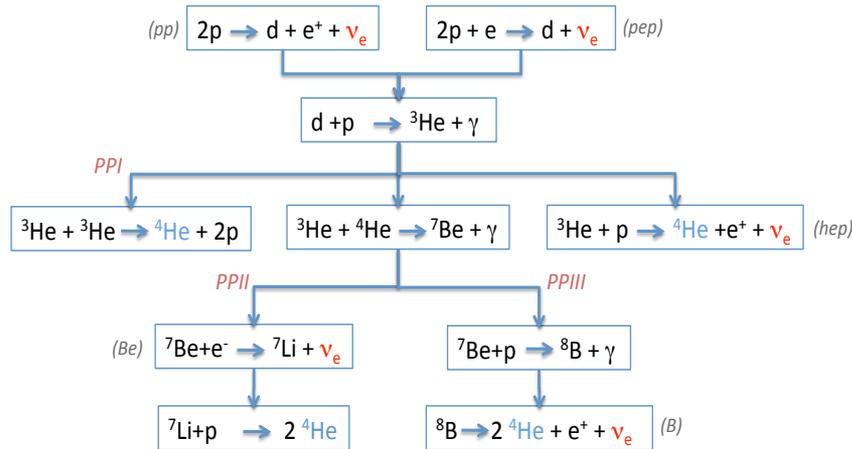
The Sun is powered by nuclear reactions that transform H into  $^4\text{He}$ :



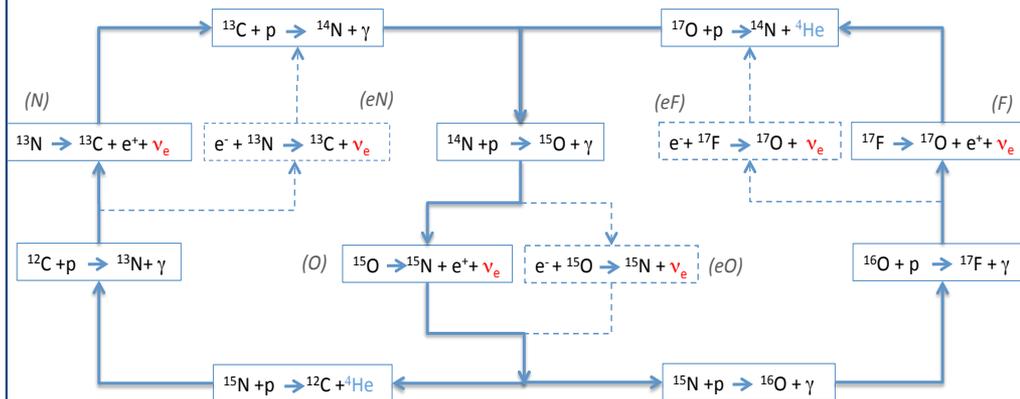
$Q = 26,7 \text{ MeV}$  (globally)

Free stream – 8 minutes to reach the earth  
Direct information on the energy producing region.

## The PP-chain



## The CN-NO (bi-)cycle

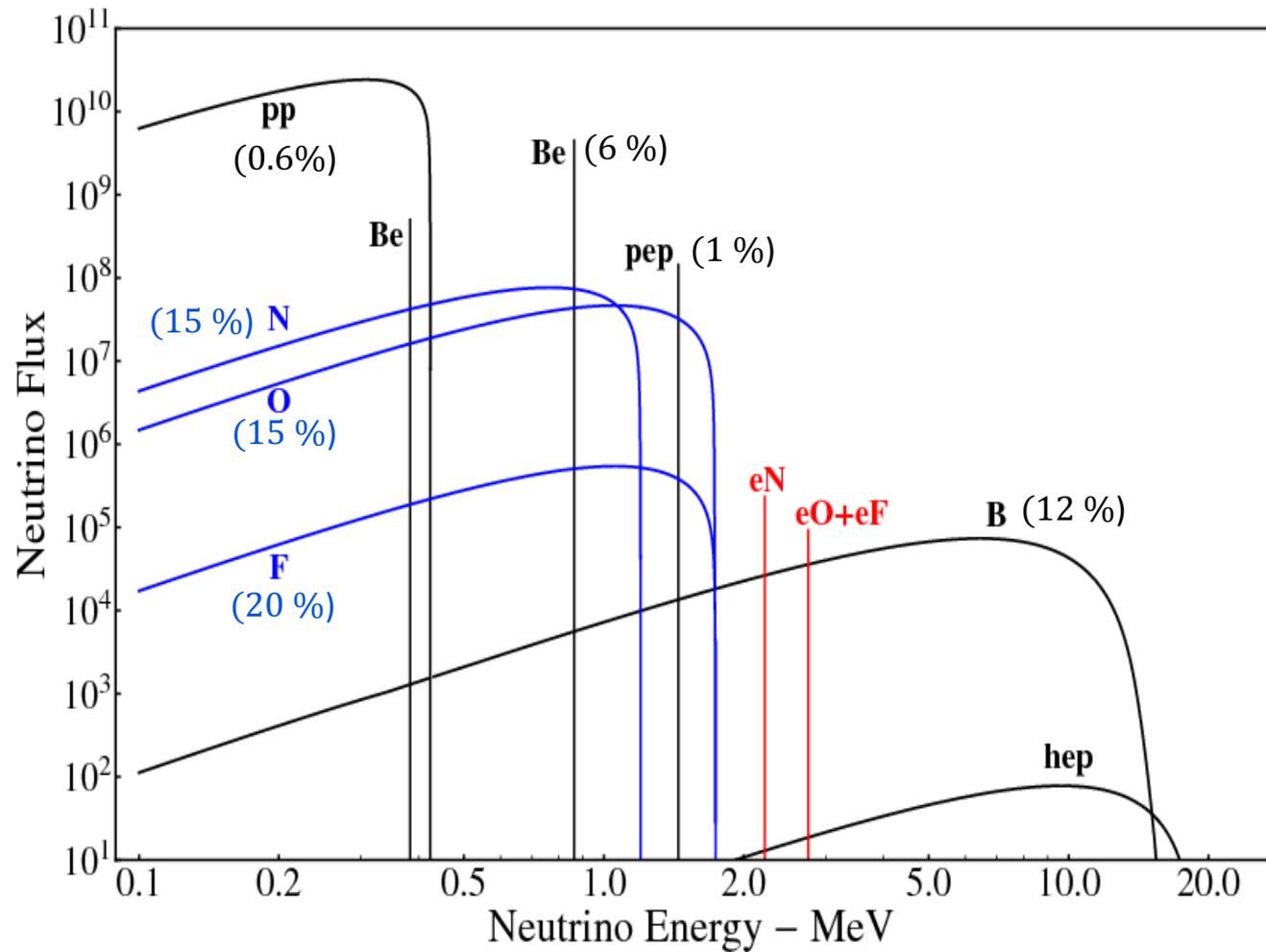


The **pp chain** is responsible for about 99% of the total energy (and neutrino) production.

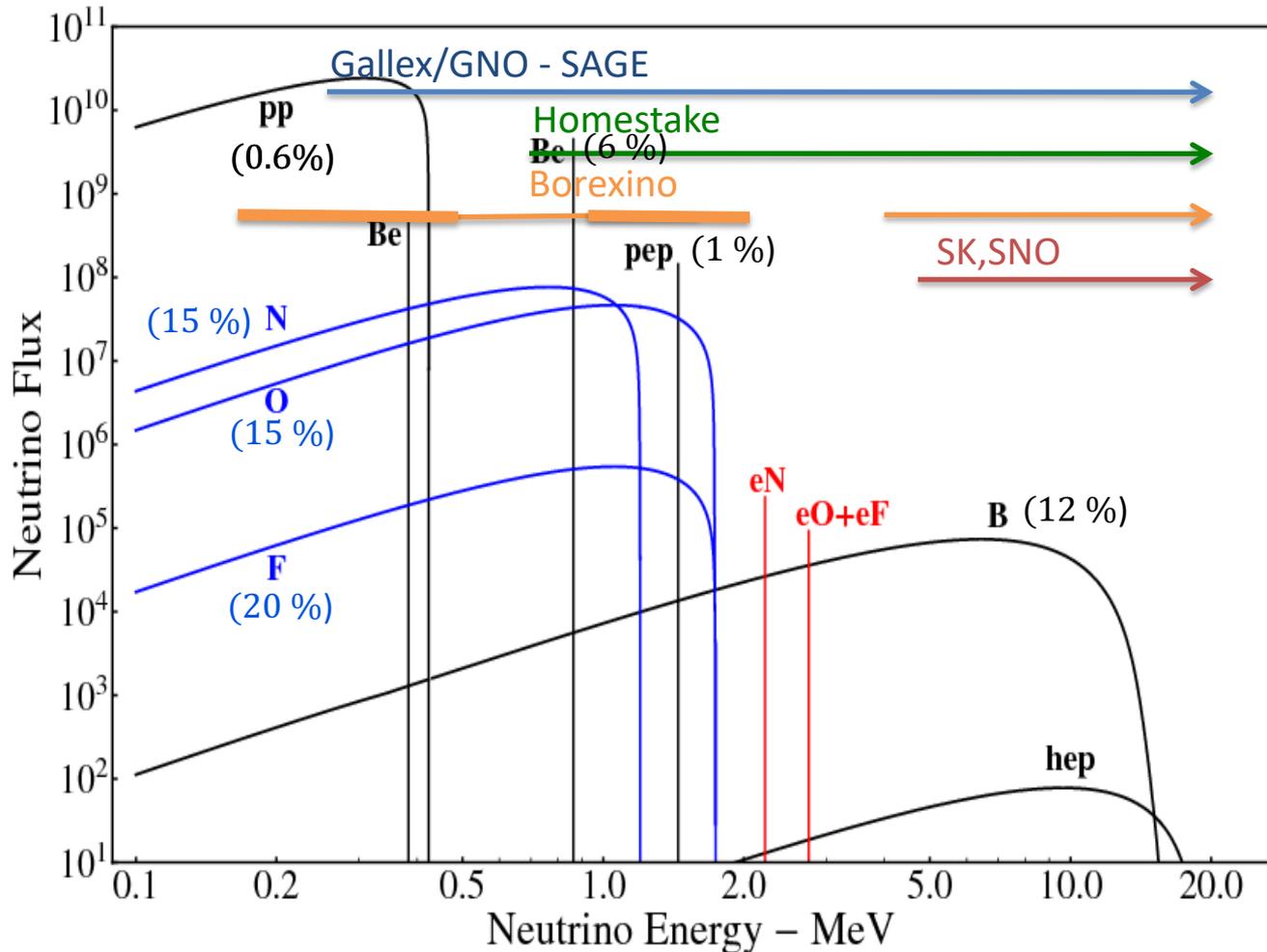
**C, N and O nuclei** are used as catalysts for hydrogen fusion.

**CNO (bi-)cycle** is responsible for about 1% of the total neutrino (and energy) budget. Important for more advanced evolutionary stages

# The solar neutrino spectrum



# The solar neutrino spectrum



The different comp. of the solar neutrinos flux have been **directly** determined with accuracy level:

pp: ~ 10%  
 pep: ~ 10%  
 $^7\text{Be}$ : ~ 3 %  
 $^8\text{B}$ : ~ 2 %  
 CNO: ~ 20%

## Recent Milestones from **Borexino**:

- $^7\text{Be}$  (and  $^8\text{B}$ ) neutrino direct detection [[PRL 2008](#)]
- pp (and pep) neutrinos direct detection [[Nature 2014, 2018](#)]
- CNO neutrinos signal identification [[Nature 2020, PRL 2022, arXiv: 2307.14636](#)]

# The importance of measuring pp-neutrinos

Assuming that the Sun is **stable**:

$$L_{\odot} + L_{\nu} (+ L_x) = \varepsilon_n + \varepsilon_g$$

Gravothermal energy prod.  $O(10^{-4} L_{\odot})$

where:

$$L_{\odot} \quad \text{Radiative Luminosity}$$
$$L_{\nu} = 4\pi D^2 \sum_i \langle E_{\nu} \rangle_i \Phi_i \quad \text{Neutrino Luminosity}$$

Additional (exotic) energy losses

$$\varepsilon_n = 4\pi D^2 \sum_i \frac{Q}{2} \Phi_i$$

Energy released by nuclear reactions ( $Q=27.3$  MeV)  
up to  $O(10^{-3} L_{\odot})$  corrections due to incomplete pp-chain and CNO-cycle  
[Vescovi et al., 2021]

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Neutrino Luminosity

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$$\underline{L_{\odot}} (+L_x) = 4\pi D^2 \sum_i \left( \frac{Q}{2} - \langle E_{\nu} \rangle_i \right) \underline{\Phi_i}$$

Radiative luminosity  
(Heat diff. time  $\approx 10^5$  year)

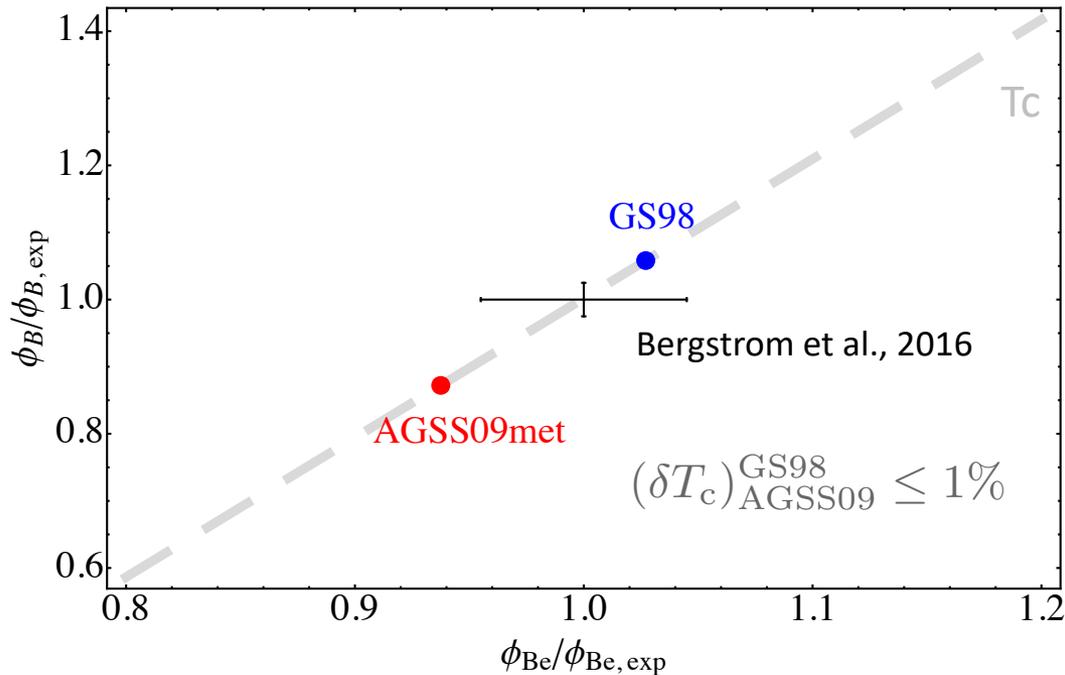
Neutrino fluxes  
 $t_{\nu} = 8$  min

**pp-neutrinos direct detection** allows us to test:

- Solar stability
- Global energy balance of the Sun
- Additional energy losses/sources

# The ${}^7\text{Be}$ and ${}^8\text{B}$ neutrino fluxes

N.Vinyoles et al. ApJ 2017 [arXiv:1611.09867v1]



${}^7\text{Be}$  and  ${}^8\text{B}$  neutrinos depend on the [core temperature  \$T\_c\$](#)  and on the [cross sections](#) that control the branching of different pp-chain terminations

$$\delta\Phi({}^7\text{Be}) = \delta S_{34} + \frac{1}{2} (\delta S_{11} - \delta S_{33}) + \beta_{\text{Be}} \delta T_c$$

$$\delta\Phi({}^8\text{B}) = (\delta S_{17} - \delta S_{e7}) + \delta S_{34} + \frac{1}{2} (\delta S_{11} - \delta S_{33}) + \beta_{\text{B}} \delta T_c$$

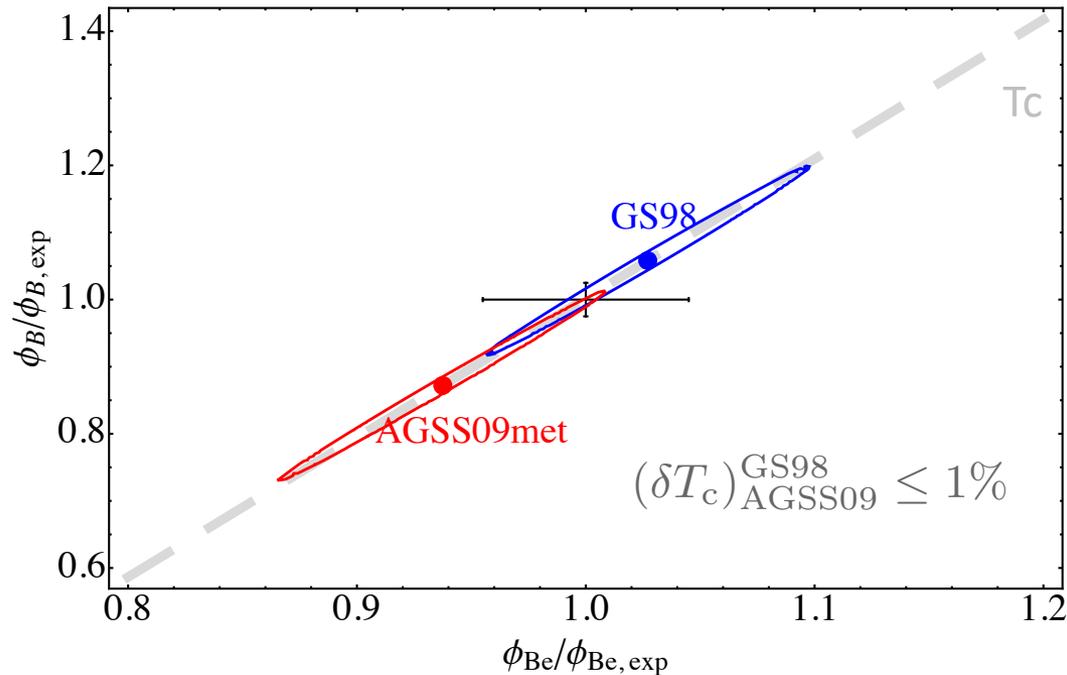
$$\left\{ \begin{array}{l} \beta_{\text{Be}} = \gamma_{34} + (\gamma_{11} - \gamma_{33})/2 \sim 11 \\ \beta_{\text{B}} = \beta_{\text{Be}} + \gamma_{17} + 1/2 \simeq 24 \end{array} \right.$$

N.B. The core temperature is a function of surface composition and enviromental parameters

$$\delta T_c = f(\delta X_i, \delta(\text{opa}), \delta(\text{diffu}), \dots)$$

# The ${}^7\text{Be}$ and ${}^8\text{B}$ neutrino fluxes

N.Vinyoles et al. ApJ 2017 [arXiv:1611.09867v1]

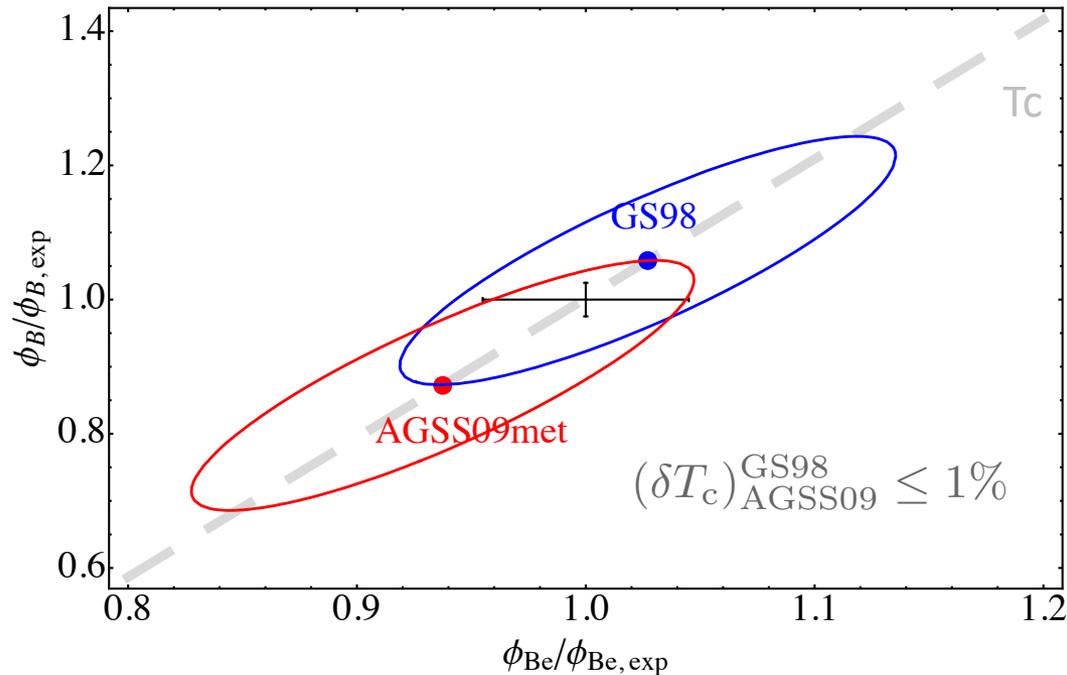


**Theoretical uncertainties dominate the error budget.** These are due to:

- Surface composition
- Environmental parameters: opacity (few %), diffusion coeff. (15%), etc
- Nuclear cross section:  $S_{17}$ (4.7%),  $S_{33}$ (5.2%),  $S_{34}$ (5.4%) dominant error sources

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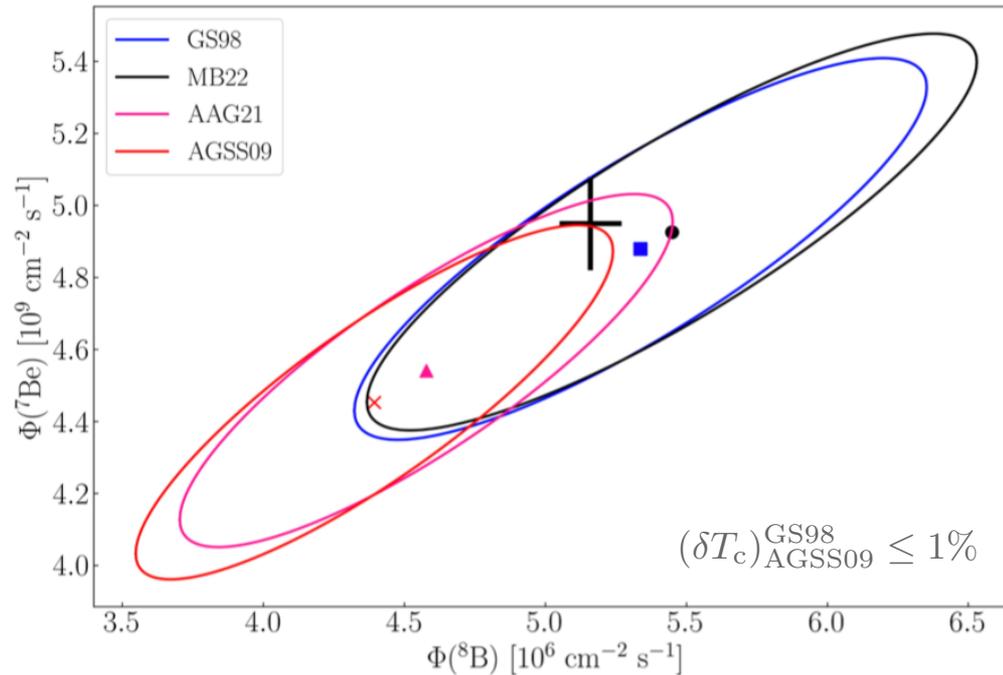
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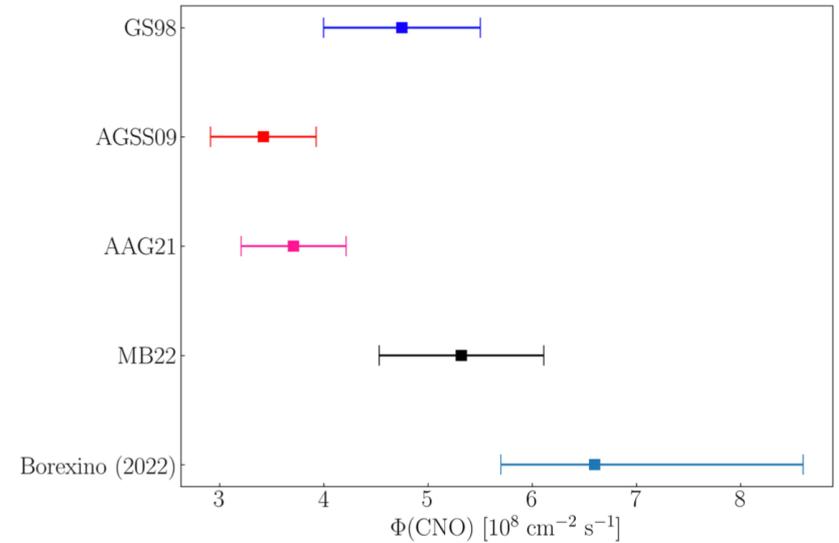
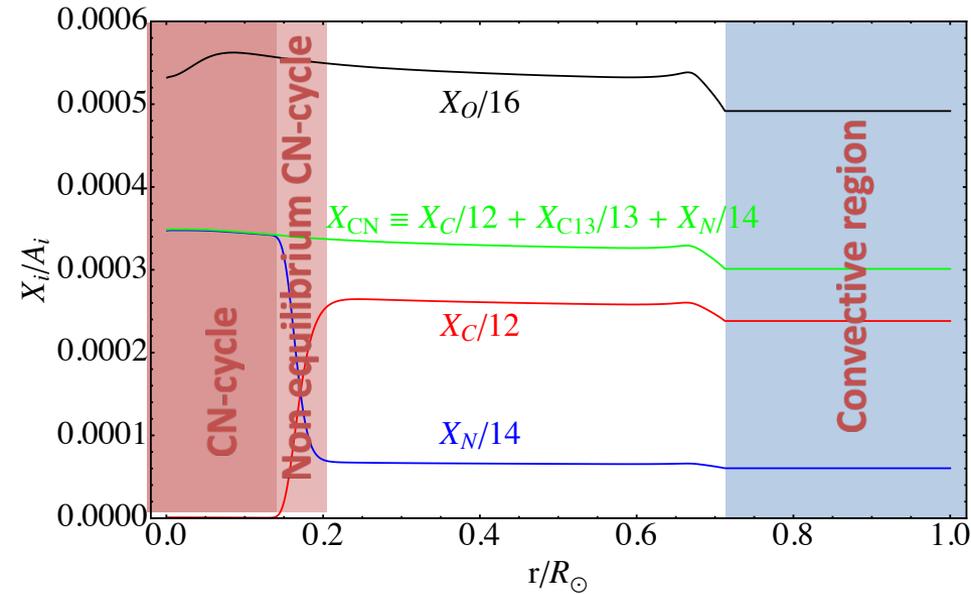


At the moment,  **${}^7\text{Be}$  and  ${}^8\text{B}$  neutrinos:**

- constrain the core temperature at  $< 1\%$  level
- do not determine the core composition with suff. accuracy

**N.B.**  ${}^7\text{Be}$  and  ${}^8\text{B}$  neutrinos alone do not break composition-opacity degeneracy

# CNO neutrino fluxes



- CNO neutrino fluxes also directly depend on the carbon+nitrogen in the core of the Sun ( $X_{CN}$ )

Assuming equal C and N fractional variations  
(i.e.  $\delta X_N^{\text{core}} = \delta X_C^{\text{core}} \equiv \delta X_{CN}^{\text{core}}$ ):

$$\delta\Phi(^{15}\text{O}) = \delta X_{CN}^{\text{core}} + \beta_O \delta T_c + \delta S_{114}$$

$$\delta\Phi(^{13}\text{N}) = \delta X_{CN}^{\text{core}} + \beta_N \delta T_c + f \delta S_{114}$$

$$\beta_O = 20$$

$$\beta_N = f \beta_O = 15$$

$$f \simeq 0.7$$

# Removing composition-opacity degeneracy

The combined measurement of pp-chain and CNO-cycle neutrinos can be used to directly **infer the solar core composition**. *Indeed:*

- The (strong) dependence on  $T_c$  (and opacity) can be eliminated by using  **$^8\text{B}$ -neutrinos as solar thermometer**;
- The additional dependence of CNO-neutrinos on  $X_{\text{CN}}$  can be used to infer core composition

*In practical terms, one can form a weighted ratio of e.g.  $^8\text{B}$  and  $^{15}\text{O}$  neutrino fluxes that is:*

- Essentially independent on environmental parameters (including opacity);
- **Directly proportional to Carbon+Nitrogen abundance in the solar core**

Serenelli et al., PRD 2013

See also (application to BX obs. rate):

Agostini et al, EPJ 2021

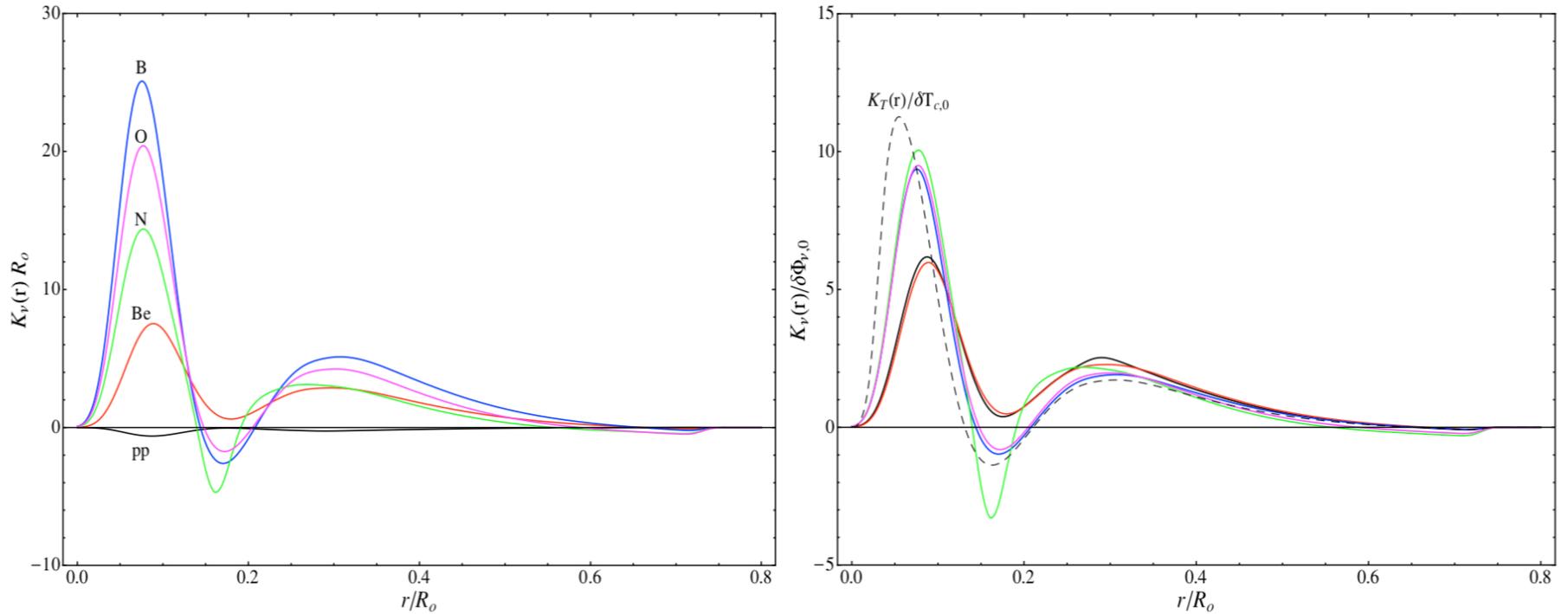
Villante & Serenelli, Frontiers 2021

$$\delta\Phi(^{15}\text{O}) - x \delta\Phi(^8\text{B}) \simeq \delta X_{\text{CN}}^{\text{core}} + \delta S_{114} - x \left( \delta S_{17} - \delta S_{e7} + \delta S_{34} + \frac{\delta S_{11}}{2} - \frac{\delta S_{33}}{2} \right)$$

$$x = \frac{\beta_{\text{O}}}{\beta_{\text{B}}} \sim 0.8$$

# Solar neutrino fluxes - opacity kernels

F.L. Villante – *Astrophys.J.*724:98-110,2010



**Figure 6.** Left panel: the solar neutrino kernels  $K_\nu(r)$  defined in Equation (32). Right panel: the solid lines are the normalized solar neutrino kernels  $K_\nu(r)/\delta\Phi_{\nu,0}$ . The dashed line shows the normalized kernel  $K_T(r)/\delta T_{c,0}$  defined in Equation (36), that describes the response of the solar central temperature to localized opacity modifications.

# Probing solar composition with neutrinos

By considering

$$\frac{R_{\text{CNO}}^{\text{Bx}}}{R_{\text{CNO}}^{\text{SSM}}} = \frac{R_{15\text{O}}^{\text{Bx}}}{R_{15\text{O}}^{\text{SSM}}} = \frac{\Phi_{15\text{O}}^{\text{Bx}}}{\Phi_{15\text{O}}^{\text{SSM}}} = 1.35^{+0.24}_{-0.16}$$

Borexino CNO neutrino signal  
(scaled to GS98 prediction)

[Borexino: PRL 2022, arXiv: 2307.14636]

$$\frac{\Phi_{8\text{B}}^{\text{global}}}{\Phi_{8\text{B}}^{\text{SSM}}} = 0.96 \pm 0.027$$

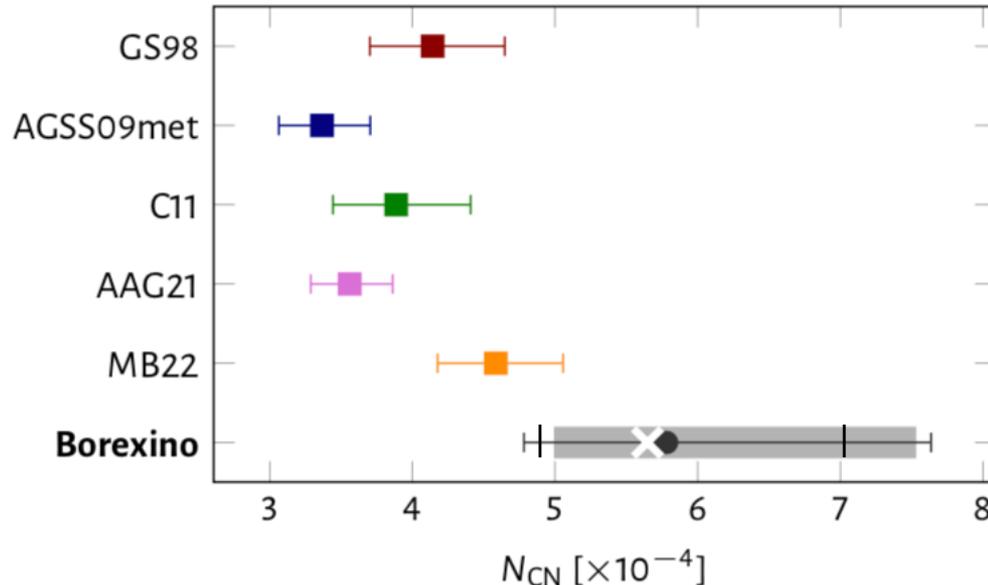
$^8\text{B}$  flux determined from global analysis  
(scaled to GS98 prediction)

One obtains:

$$\frac{(N_{\text{C}} + N_{\text{N}})/N_{\text{H}}}{[(N_{\text{C}} + N_{\text{N}})/N_{\text{H}}]^{\text{SSM}}} = 1.35 \times (0.96)^{-0.769} \times$$

$$\times [1 \pm ({}^{+0.18}_{-0.12}(\text{CNO}) \pm 0.097(\text{nucl}) \pm 0.023(^8\text{B}) \pm 0.005(\text{env}) \pm 0.027(\text{diff}) \pm 0.022(\text{O/N}))]$$

↳ Note: reduced error wrt Borexino, PRL 2022



N.B.

This determination is robust wrt to environmental parameters variations (including opacity).

Only limited by nuclear reaction uncertainties:

$$S_{114} \rightarrow 7.6 \%$$

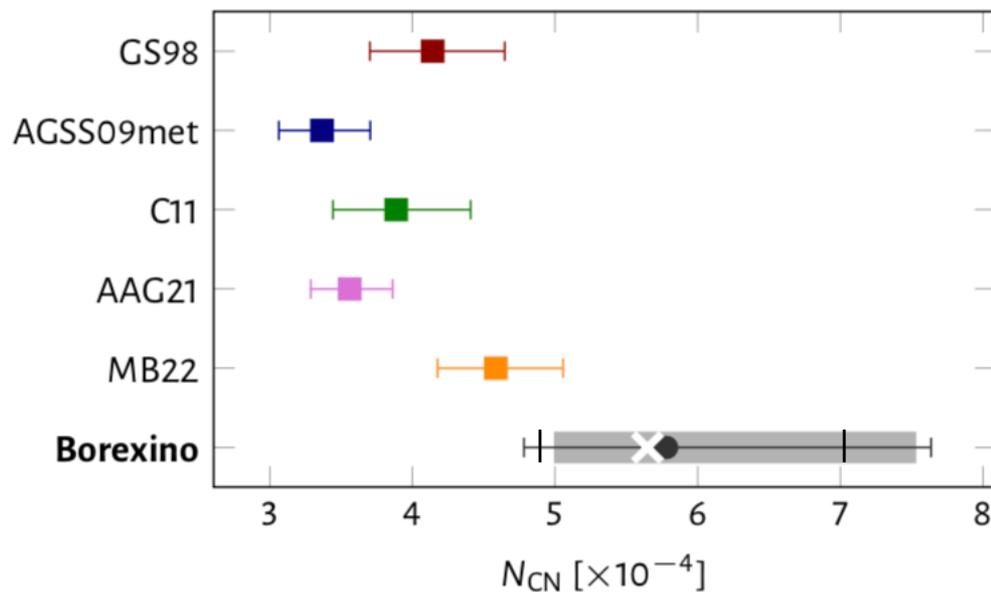
$$S_{17} \rightarrow 3.5 \%$$

$$S_{34} \rightarrow 3.4 \%$$

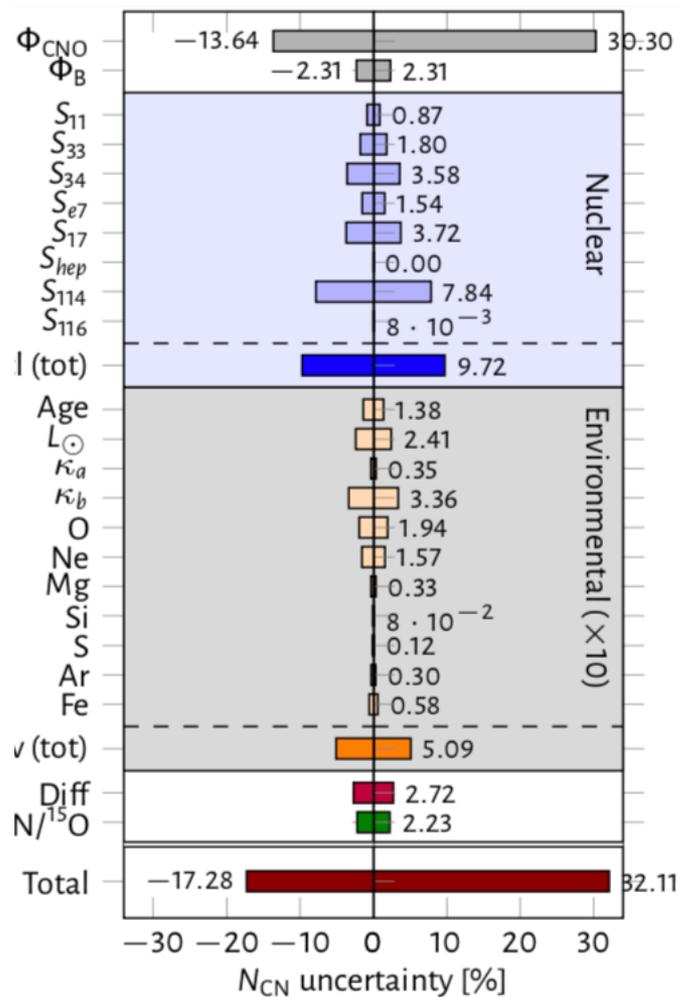
# Probing solar composition with neutrinos

$$\frac{(N_C + N_N)/N_H}{[(N_C + N_N)/N_H]^{SSM}} = 1.35 \times (0.96)^{-0.769} \times$$

$$\times [1 \pm ({}^{+0.18}_{-0.12}(\text{CNO}) \pm 0.097(\text{nucl}) \pm 0.023({}^8\text{B}) \pm 0.005(\text{env}) \pm 0.027(\text{diff}) \pm 0.022(\text{O/N}))]$$



## Error contributions

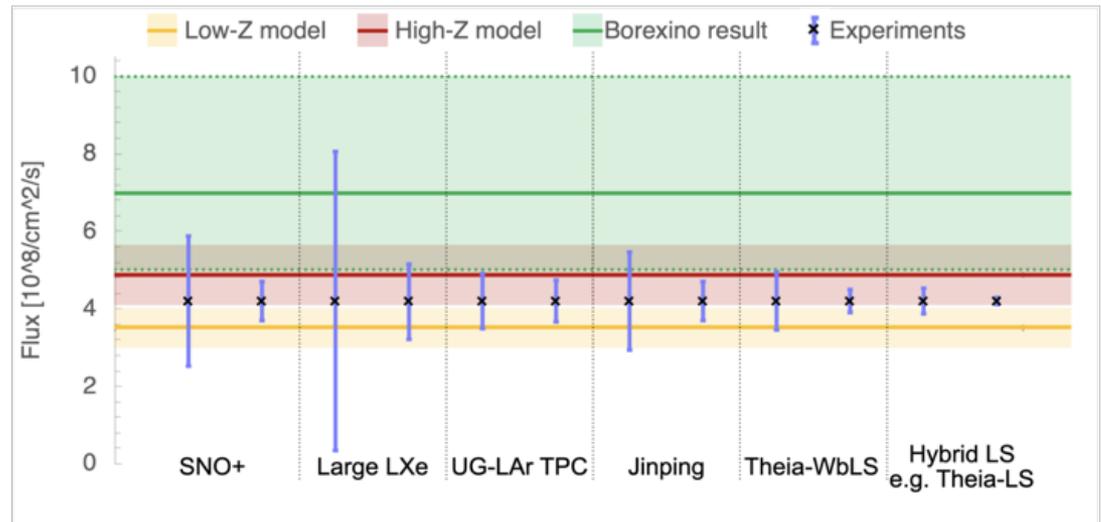


# Future perspectives

Borexino has opened the way to CNO neutrino detection

Improvements on the experimental side will be provided in the future by planned detectors, e.g.:

- SNO+
- JUNO
- Jinping
- Hyper-Kamiokande
- THEIA
- DUNE
- Dark Matter experiments
- .....



ARNP – Orebi Gann et al. in press

Note that: some minor components (hep and ecCNO) of the solar neutrino flux are still undetected

- **ecCNO neutrinos:** A challenge for gigantic ultra-pure LS detectors (Villante, PLB 2015)  
Expt. requirements: *as clean (and deep) as Borexino; as large as JUNO*

# Conclusions

- Solar neutrino physics entered the precision era.
- Borexino has opened the way to CNO neutrino detection
- Some unsolved puzzles could be addressed → (Present and future) CNO neutrino measurements, combined with precise determinations of  $^8\text{B}$  and  $^7\text{Be}$  fluxes, can shed light on the [solar abundance problem](#)
- To exploit the full potential of future measurements → improvements in the SSM constitutive physics are needed [[nuclear cross sections and radiative opacities](#)]

[Solar Fusion Cross Sections III \(INT-22-82W\)](#)  
July 2022, UC Berkeley, Berkeley, CA, USA

Thank you

# Standard Solar Models

**Stellar structure equations** are solved, starting from a ZAMS model to present solar age (we neglect rotation, magnetic fields, etc.):

$$\begin{aligned}\frac{\partial m}{\partial r} &= 4\pi r^2 \rho \\ \frac{\partial P}{\partial r} &= -\frac{G_N m}{r^2} \rho \\ P &= P(\rho, T, X_i) \\ \frac{\partial l}{\partial r} &= 4\pi r^2 \rho \epsilon(\rho, T, X_i) \\ \frac{\partial T}{\partial r} &= -\frac{G_N m T \rho}{r^2 P} \nabla\end{aligned}\quad \nabla = \text{Min}(\nabla_{\text{rad}}, \nabla_{\text{ad}}) \rightarrow \begin{cases} \nabla_{\text{rad}} &= \frac{3}{16\pi ac G_N} \frac{\kappa(\rho, T, X_i) l P}{m T^4} \\ \nabla_{\text{ad}} &= (d \ln T / d \ln P)_s \simeq 0.4 \end{cases}$$

**Chemical evolution** driven by nuclear reaction, diffusion and gravitational settling, convection

**Standard input physics** for equation of states, nuclear reaction rates, opacity, etc.

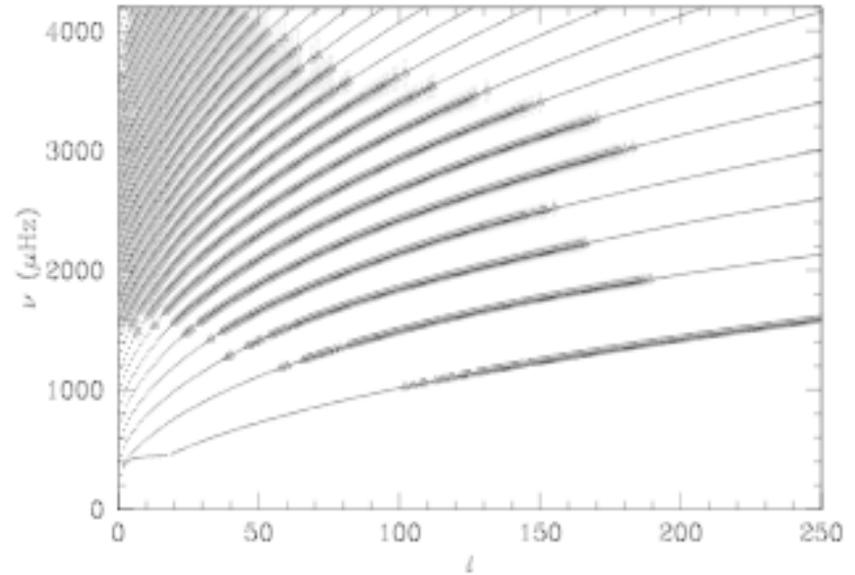
Free-parameters (**mixing length**,  $Y_{\text{ini}}$ ,  $Z_{\text{ini}}$ ) adjusted to match the observed properties of the Sun (**radius**, **luminosity**,  $Z/X$ ).

Note that equations are non-linear  $\rightarrow$  Iterative method to determine mixing length,  $Y_{\text{ini}}$ ,  $Z_{\text{ini}}$

# The solar abundance problem

The **downward revision** of heavy elements photospheric abundances leads to SSMs which **do not correctly reproduce helioseismic observables**

Oscillation frequencies of the sun



*360 days of observation of the MDI instrument (errors multiplied by 5000)*

The Sun is a non radial oscillator. The observed oscillation frequencies can be used to determine the properties of the Sun. Linearizing around a known solar model:

$$\frac{\delta\nu_{nl}}{\nu_{nl}} = \int_0^R dr K_{u,Y}^{nl}(r) \frac{\delta u}{u}(r) + \int_0^R dr K_{Y,u}^{nl}(r) \delta Y + \frac{F(\nu_{nl})}{\nu_{nl}}$$

*squared isothermal sound speed*

*Related to temperature stratification in the sun*

*surface helium abundance*

See Basu & Antia 07  
for a review

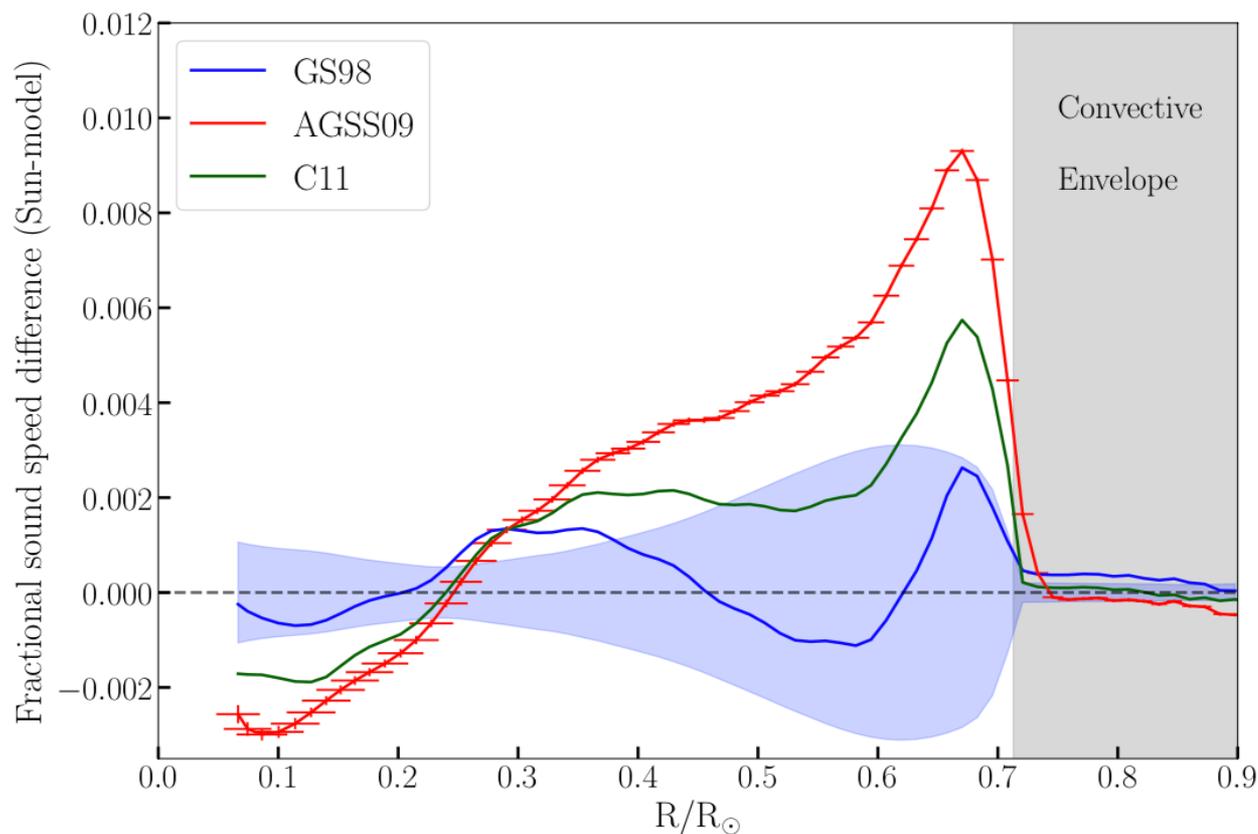
# The solar abundance problem

| Model      | $R_{CZ}/R_{\odot}$ | $Y_{\text{S}}$ |
|------------|--------------------|----------------|
| MB22-phot  | 0.7123             | 0.2439         |
| MB22-met   | 0.7120             | 0.2442         |
| AAG21      | 0.7197             | 0.2343         |
| AGSS09-met | 0.7231             | 0.2316         |
| GS98       | 0.7122             | 0.2425         |
| C11        | 0.7162             | 0.2366         |

*Helioseismic determinations*

$$R_b/R_{\odot} = 0.713 \pm 0.001$$

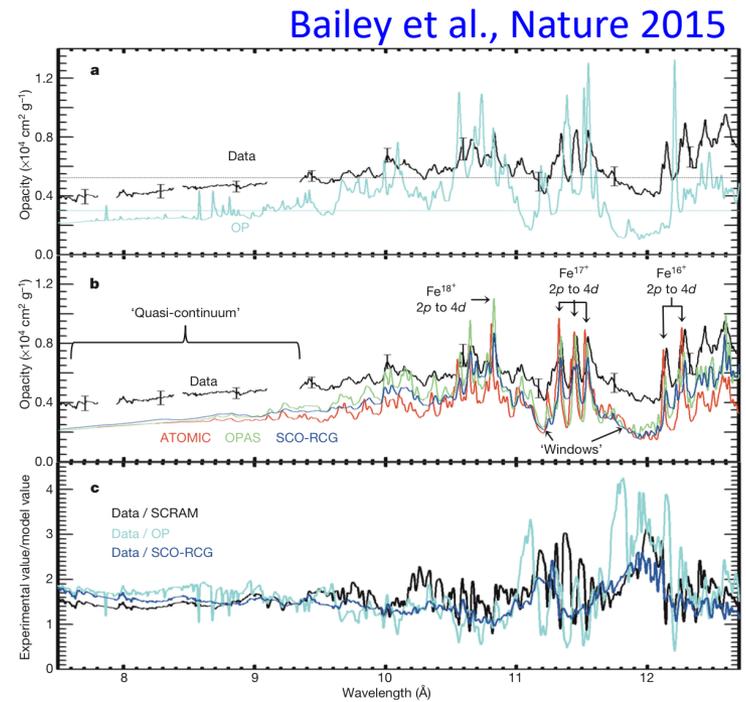
$$Y_b = 0.2485 \pm 0.0035$$



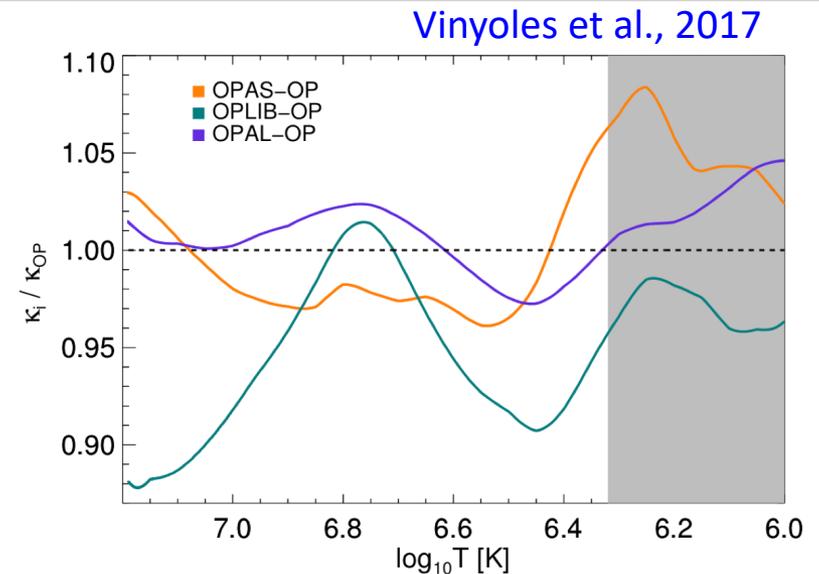
Magg et al. 2022

# Wrong opacity?

- Opacity is being measured at stellar interiors conditions (Bailey et al., Nature 2015);
- Monochromatic opacity is higher than expected for iron (up to a factor 2);
- Total opacity (integrated over the wavelength and summed over the composition) is increased by about 7%



- Different opacity tables may differ “locally” by a large amount (up to 10%) and with a complicated pattern



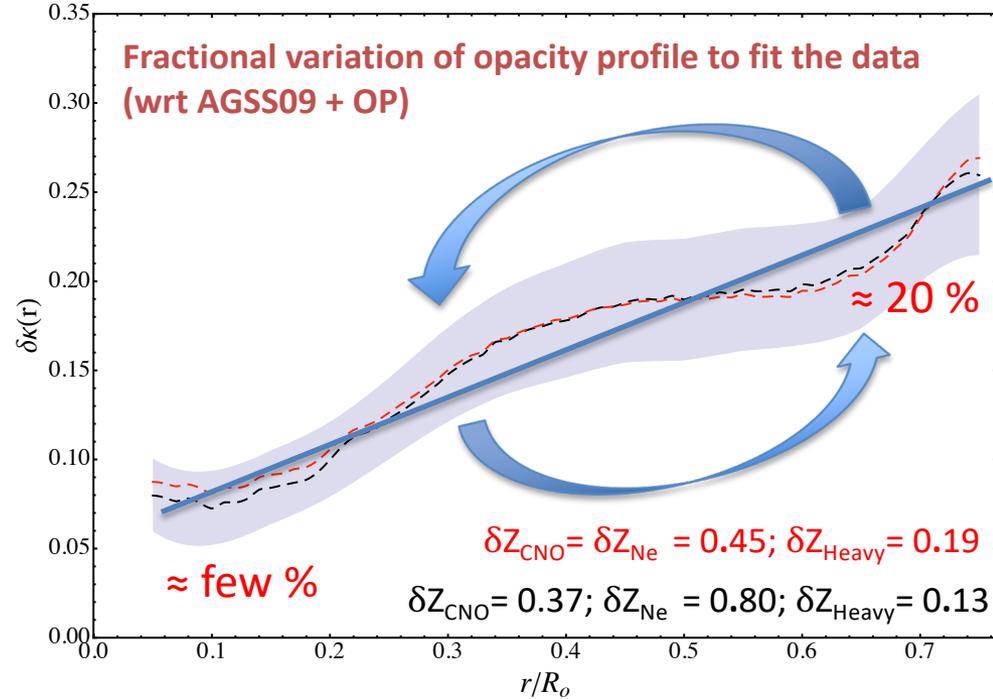
# The solar opacity profile

F.L. Villante and B. Ricci - *Astrophys.J.* 714:944-959, 2010  
 F.L. Villante – *Astrophys.J.* 724:98-110, 2010  
 F.L. Villante, A. Serenelli et al., *Astrophys.J.* 787 (2014) 13

The “optimal” opacity profile of the Sun can be determined from obs. data

Note that:

- The sound speed and the convective radius determine **the tilt** of  $\delta\kappa(r)$  (but not **the scale**)
- The surface helium and the neutrino fluxes determine **the scale** for  $\delta\kappa(r)$



The interpretation is however complicated by the **opacity-composition degeneracy**. Which fraction of the required  $\delta\kappa(r)$  has to be ascribed to **intrinsic** ( $\delta\kappa_I(r)$ ) and/or **composition** opacity changes?

$$\delta\kappa(r) = \delta\kappa_I(r) + \sum_j \frac{\partial \ln \kappa(r)}{\partial \ln Z_j} \delta z_j$$

Opacity table “errors”  
 Non standard effects (WIMPs in solar core)  
 ...

different admixtures  $\{\delta z_j\}$  can do equally well the job

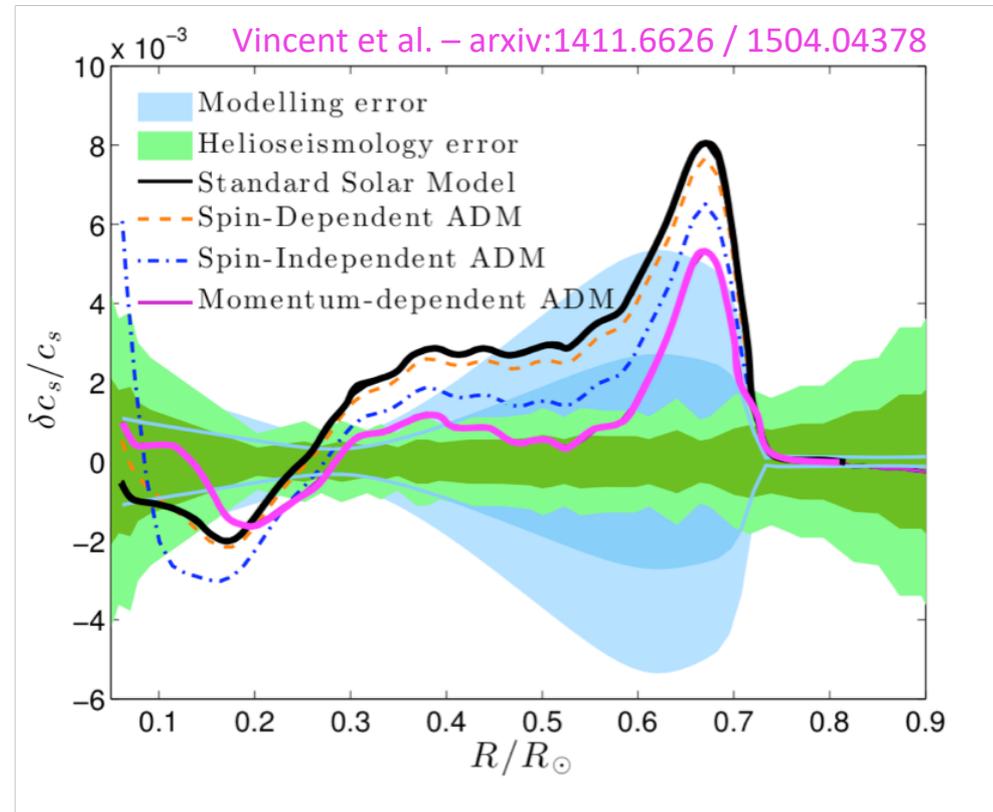
# Asymmetric DM

DM accumulation in the solar core:

- Additional energy transport;
- **Reduction** of the “effective opacity”;
- Modification of temperature profile;

Agreement with helioseismic data can be improved. However:

- DM accumulation do not provide the optimal opacity profile;
- Potential tension with neutrino fluxes and surface helium;
- **Caveat:** DM evaporation not accounted for (relevant for few GeV masses)



$$\sigma = \sigma_0 \left( \frac{q}{q_0} \right)^2 \quad \begin{cases} m_{\chi} = 3 \text{ GeV} \\ \sigma_0 = 10^{-37} \text{ cm}^2 \\ q_0 = 40 \text{ MeV} \end{cases}$$

# Wrong chemical evolution?

Helioseismic observables and neutrino fluxes are sensitive to **the metallicity of the radiative interior of the Sun.**

The observations determine **the chemical composition of the convective envelope** (2-3% of the solar mass).

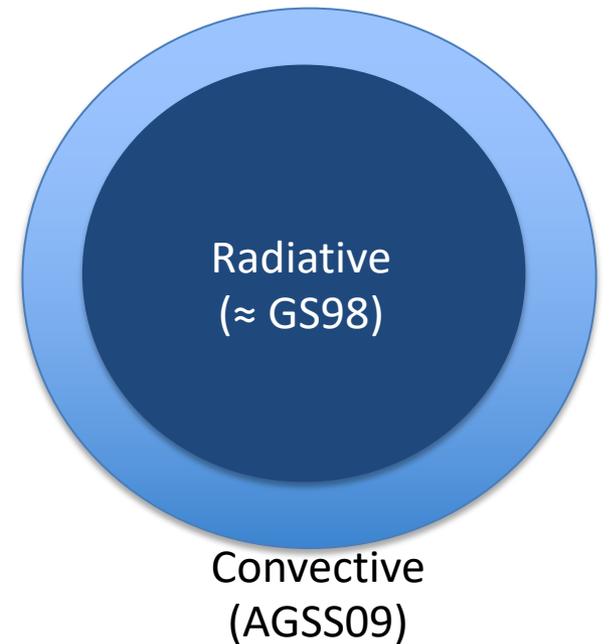
Difference between AGSS09 and GS98 correspond to  $\approx 40M_{\oplus}$  of metal, when integrated over the Sun's convective zone.

**Could this difference be accounted in non standard chemical evolution scenarios (e.g. by accretion of material with non standard composition)?**

See A. Serenelli et al. – ApJ 2011

*This is a well posed and extremely important question but ...*

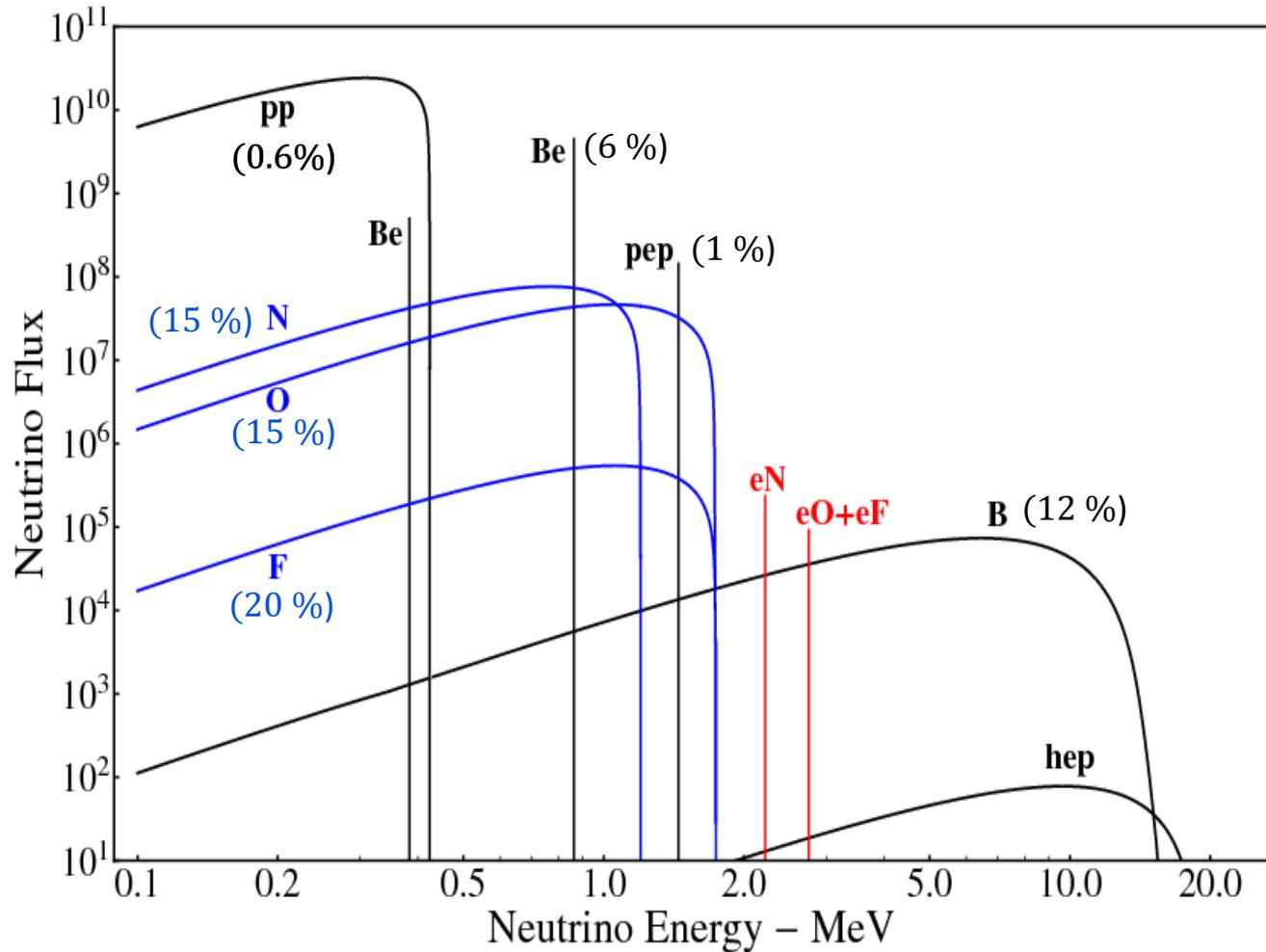
*... no satisfactory solutions have been proposed up to now, in my opinion*





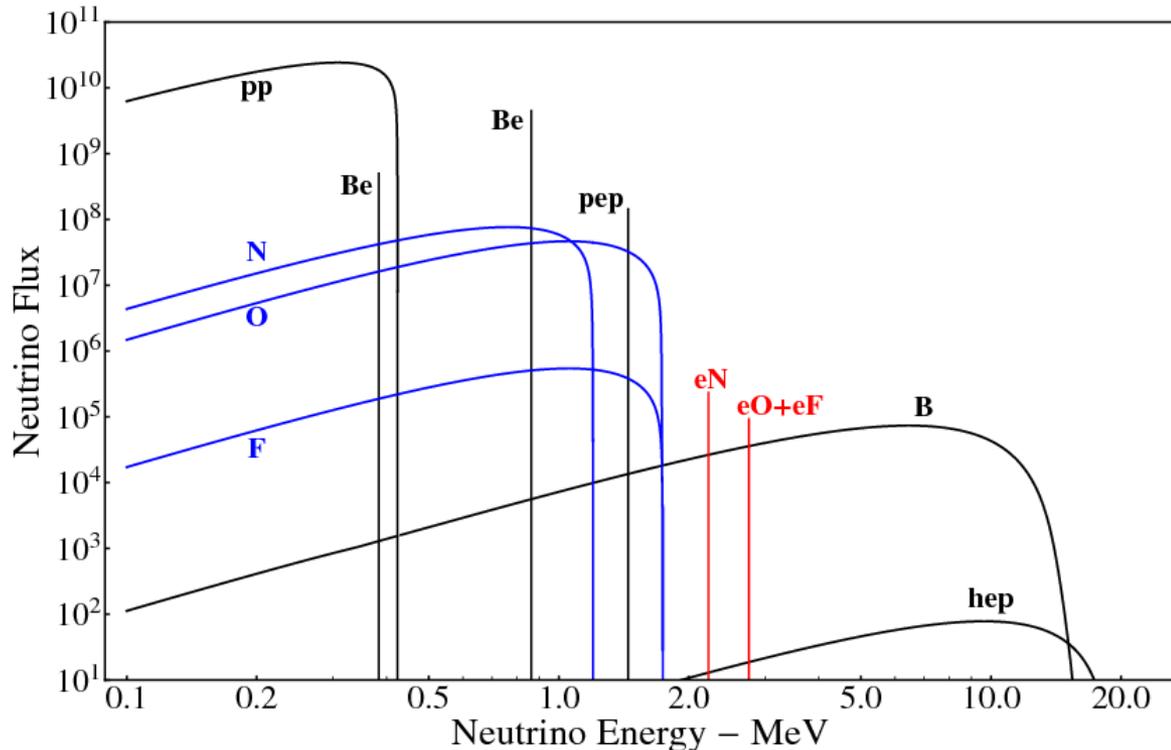
# The solar neutrino spectrum

The Sun is powered by nuclear reactions that transform H into  $^4\text{He}$ :



# ecCNO neutrinos

In the CN-NO cycle, besides the conventional CNO neutrinos (blue lines), **monochromatic ecCNO neutrinos (red lines)** are also produced by **electron capture** reactions:

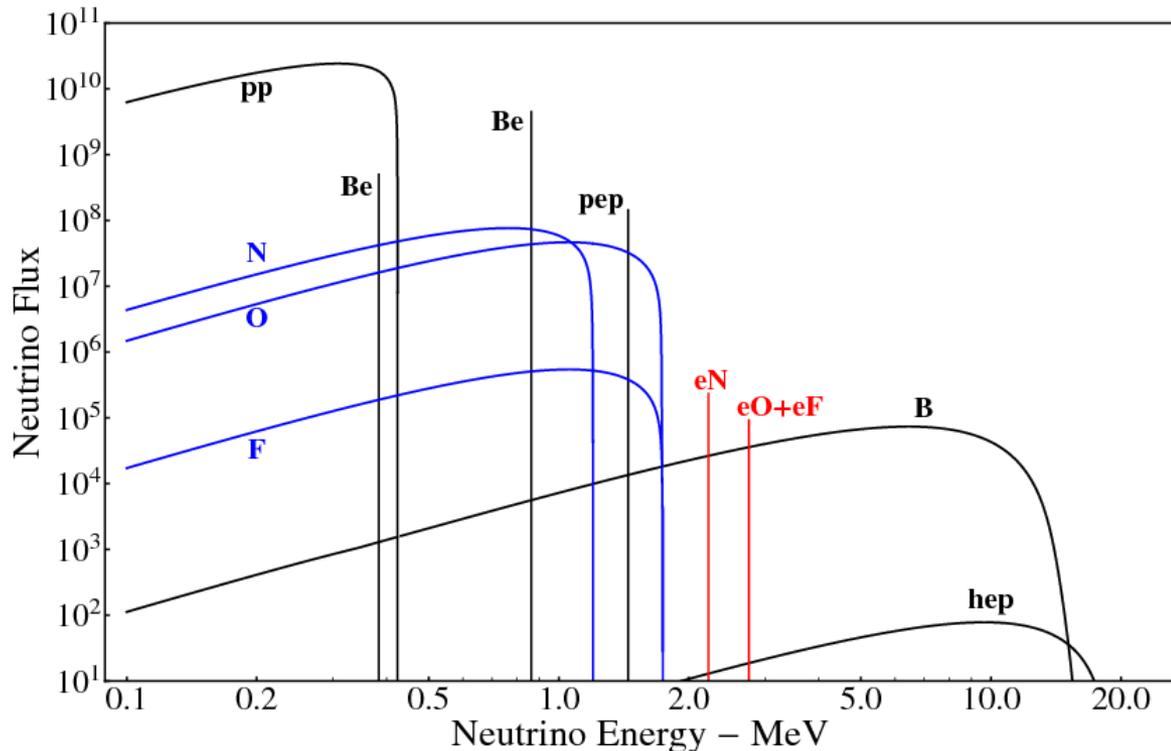


*F.L. Villante, PLB 742 (2015) 279-284  
L.C. Stonehill et al, PRC 69, 015801 (2004)  
J.N. Bahcall, PRD 41, 2964 (1990).*

# ecCNO neutrinos

The ecCNO fluxes are extremely low:  $\Phi_{\text{ecCNO}} \approx (1/20) \Phi_{\text{B}}$ . Detection is extremely difficult but could be rewarding. Indeed:

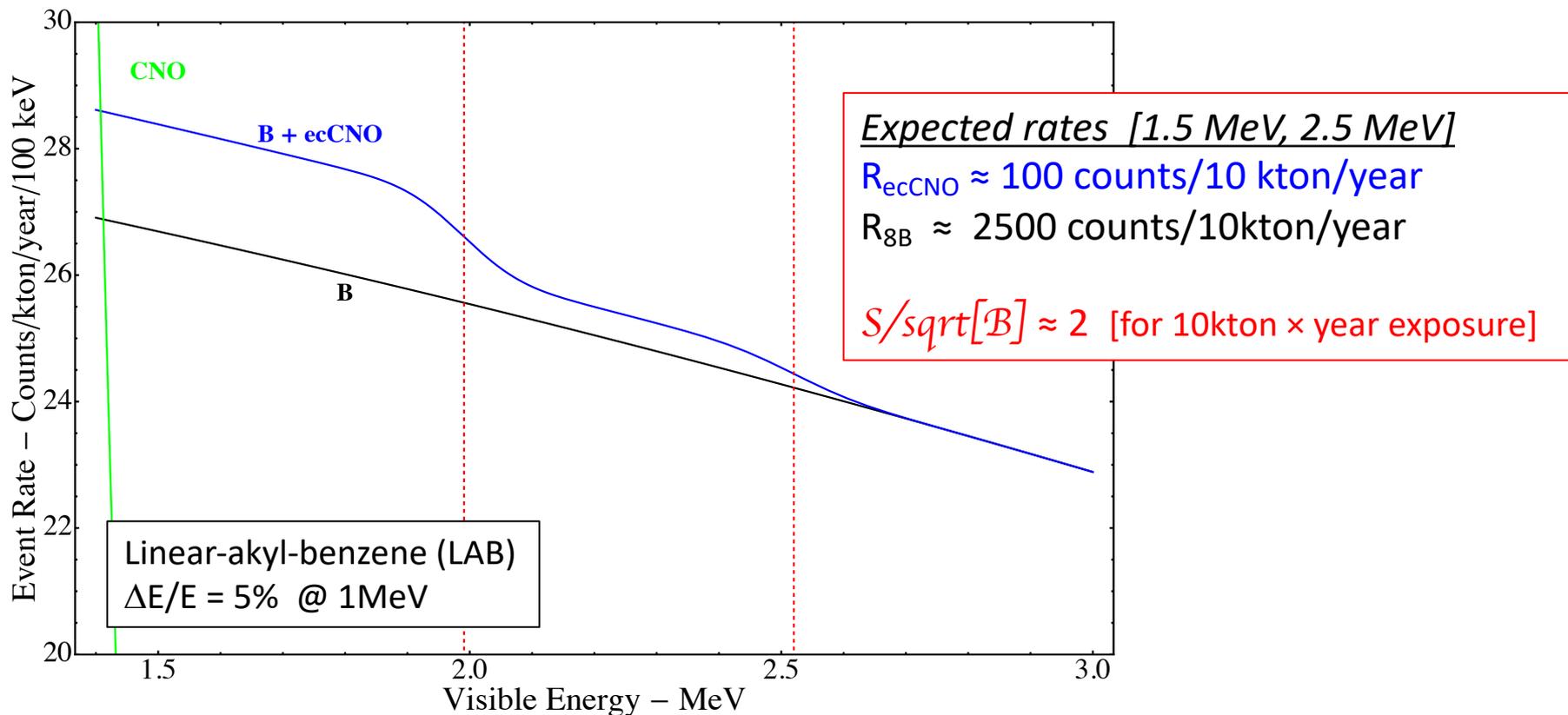
- ecCNO neutrinos are sensitive to the **metallic content of the solar core** (same infos as CNO neutrinos);
- Being monochromatic, they probe the solar neutrino **survival probability** at specific energies ( $E_{\nu} \approx 2.5 \text{ MeV}$ ) exactly **in the transition region**.



*F.L. Villante, PLB 742 (2015) 279-284  
L.C. Stonehill et al, PRC 69, 015801 (2004)  
J.N. Bahcall, PRD 41, 2964 (1990).*

# Expected rates in Liquid Scintillators

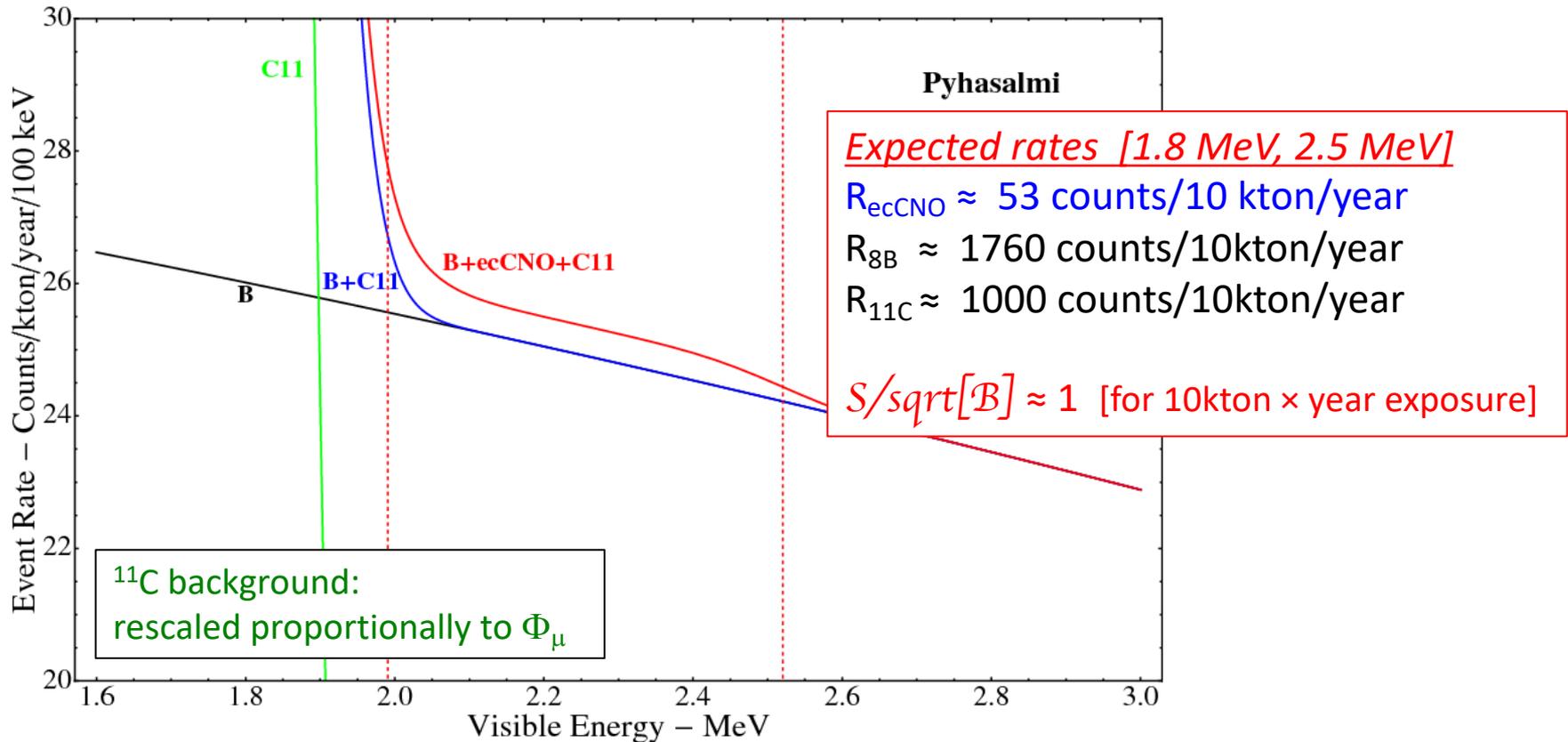
- $\nu - e$  elastic scattering of ecCNO neutrinos produces Compton shoulders (smeared by energy resolution) at 2.0 and 2.5 MeV;
- ecCNO neutrino signal has to be extracted statistically from the (irreducible)  $^8\text{B}$  neutrino background.



# Expected rates in Liquid Scintillators

Additional background sources:

- **Intrinsic:** negligible/tagged (with Borexino Phase-I radio-purity levels);
- **External:** reduced by self-shielding (Fid. mass reduced from 50 to  $\approx 20$  kton in LENA);
- **Cosmogenic:**  $^{11}\text{C}$  overlap with the observation window.



Signal comparable to stat. fluctuations for exposures 10 kton  $\times$  year or larger.

100 counts / year above 1.8 MeV in 20 kton detector  $\rightarrow 3\sigma$  detection in 5 year in LENA

# Removing composition-opacity degeneracy

The combined measurement of pp-chain and CNO-cycle neutrinos can be used to directly **infer the solar core composition**. *Indeed:*

- The (strong) dependence on  $T_c$  (and opacity) can be eliminated by using  **$^8\text{B}$ -neutrinos as solar thermometer**;
- The additional dependence of CNO-neutrinos on  $X_{\text{CN}}$  can be used to infer core composition

*In practical terms, one can form a weighted ratio of e.g.  $^8\text{B}$  and  $^{15}\text{O}$  neutrino fluxes that is:*

- Essentially independent on environmental parameters (including opacity);
- **Directly proportional to Carbon+Nitrogen abundance in the solar core**

Serenelli et al., PRD 2013

See also (application to BX obs. rate):

Agostini et al, EPJ 2021

Villante & Serenelli, Frontiers 2021

$$\begin{aligned} \varphi^{15\text{O}}/\varphi^{8\text{B}} &= X_{\text{C}}^{0.802} X_{\text{N}}^{0.204} X_{\text{D}}^{0.181} \\ &\times \begin{bmatrix} -0.866 & 0.345 & -0.689 & 0.769 & -0.791 & 0.000 & 1.046 & 0.001 \\ X_{\text{S}_{11}} & X_{\text{S}_{33}} & X_{\text{S}_{34}} & X_{\text{S}_{e7}} & X_{\text{S}_{17}} & X_{\text{S}_{hep}} & X_{\text{S}_{114}} & X_{\text{S}_{116}} \end{bmatrix} \quad (\text{nucl}) \\ &\times \begin{bmatrix} 0.313 & 0.602 & 0.018 & -0.050 \\ X_{\text{Age}} & X_{L_{\odot}} & X_{\kappa_a} & X_{\kappa_b} \end{bmatrix} \quad (\text{solar}) \\ &\times \begin{bmatrix} 0.006 & -0.003 & -0.003 & 0.001 & 0.001 & 0.001 & 0.005 \\ X_{\text{O}} & X_{\text{Ne}} & X_{\text{Mg}} & X_{\text{Si}} & X_{\text{S}} & X_{\text{Ar}} & X_{\text{Fe}} \end{bmatrix} \quad (\text{met}) \end{aligned}$$

# Probing solar composition with neutrinos

$$\frac{(N_C + N_N)/N_H}{[(N_C + N_N)/N_H]^{SSM}} = 1.35 \times (0.96)^{-0.769} \times$$

$$\times [1 \pm ({}^{+0.303}_{-0.136}(\text{CNO}) \pm 0.097(\text{nucl}) \pm 0.023({}^8\text{B}) \pm 0.005(\text{env}) \pm 0.027(\text{diff}) \pm 0.022(\text{O/N}))]$$

## Error contributions

