Probing Fuzzy Dark Matter in Neutrino Experiments

based on :

VB, J. Kopp, J. Liu, P. Prass, X. Wang arXiv:1705.09455

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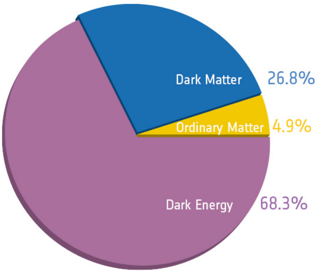
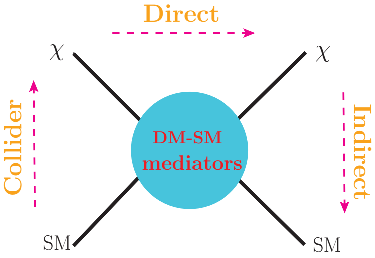
Dark Matter - the mystery of this century

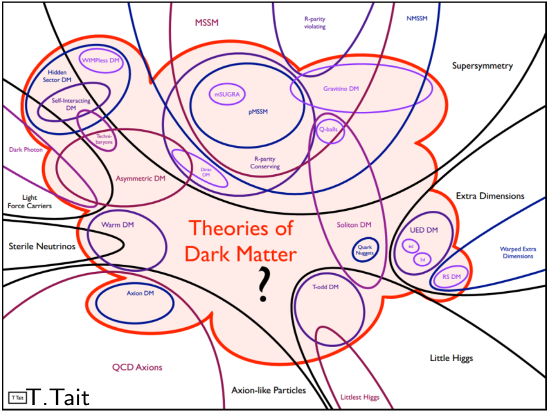


- not charged under $U(1)_{EM}$ and $SU(3)_C$
- stable or long lived
- not in SM particle list



- Direct detection
  - nuclear recoils from DM scattering
- Collider searches
  - typical signal:missing energy + mono object
- Indirect detection
  - classified by annihilation product: γ, ν, e^+...

Dark Matter - vast number of candidates



- in this talk we will focus on the low end of DM spectrum – fuzzy DM with mass $\mathcal{O}(10^{-22})$ eV
- we consider both scalar and vector ultralight DM

Properties of "non-standard" Fuzzy DM candidate

Fuzzy DM can address:

- "core vs. cusp problem"– DM density profile discrepancy between measurements and simulations

→ DM delocalization
(huge Compton wave length $\lambda = 2\pi/m_\phi \simeq 0.4\,\text{pc} \times (10^{-22}\text{eV}/m_\phi))$

- "missing satellites problem"–lower than expected abundance of dwarf galaxies

→ higher probability for tidal disruption of DM subhalos and suppression of the matter power spectrum at small scales (Hui et al. 1610.08297)

- "too big to fail problem"– apparent failure of many of the most massive Milky Way subhalos to host visible dwarf galaxies

→ Fuzzy DM predicts fewer such subhalos (Marsh et al. 1307.1705)

- admittedly, better treatment of baryonic physics in simulations (1602.05957,1202.0554) may solve these puzzles but the possibility that DM physics plays a crucial role is not excluded

DM production

Arias et al. 1201.5902
Nelson & Scholtz 1105.2812
Golovnev et al. 0802.2068

Misalignment mechanism

- EOM for real scalar field ϕ

$$\ddot{\phi} + 3\,H\,\dot{\phi} + m_\phi^2 \phi = 0.$$

- while $3\,H \gg m_\phi$, ϕ is "frozen"
- at $3\,H = m_\phi$ damping term stops dominating and the field can start to oscillate
- for vector DM ϕ^μ one introduces coupling to gravity $\sim R\phi_\mu\phi^\mu$
- The mass of ϕ^μ can be generated either through the Stückelberg mechanism or from spontaneous symmetry breaking in a dark Higgs sector
- we consider both polarized and unpolarized vector DM (polarization may be altered during structure formation)

Model

Relevant part of the Lagrangian:

$$\text{Scalar} \quad \mathcal{L}_{\text{scalar}} = \bar{\nu}_L^\alpha i\partial\!\!\!/ \nu_L^\alpha - \frac{1}{2} m_\nu^{\alpha\beta} \overline{(\nu_L^c)^\alpha} \nu_L^\beta - \frac{1}{2} y^{\alpha\beta} \phi \, \overline{(\nu_L^c)^\alpha} \nu_L^\beta .$$

The interaction term can be generated in a gauge invariant way by coupling ϕ to heavy right-handed neutrinos N_R (introduced in seesaw type-I)

- we assume $y = y_0(m_\nu/0.1\text{eV})$

$$\text{Vector} \quad \mathcal{L}_{\text{vector}} = \bar{\nu}_L^\alpha i\partial\!\!\!/ \nu_L^\alpha - \frac{1}{2} m_\nu^{\alpha\beta} \overline{(\nu_L^c)^\alpha} \nu_L^\beta + g Q^{\alpha\beta} \phi^\mu \bar{\nu}_L^\alpha \gamma_\mu \nu_L^\beta .$$

- ϕ^μ as the $L_\mu - L_\tau$ symmetry gauge boson with couplings $Q^{\alpha\beta} = \text{diag}(0, 1, -1)$
- if $L_\mu - L_\tau$ breaking occurs at TeV scale, with $m_\phi \sim 10^{-22}$ we require coupling $g \sim 10^{-30}$ which can be probed

Model II

- alternatively, ϕ^μ could couple to the SM via mixing with a much heavier $L_\mu - L_\tau$ gauge boson K^μ (term $\epsilon \phi^{\mu\nu} K_{\mu\nu}$)
- $g \sim m_\phi / m_K$
- neutrino masses generated by introducing 3 RH neutrinos $N_1 \sim (1,1,0)(0), \quad N_2 \sim (1,1,0)(+1), \quad N_3 \sim (1,1,0)(-1)$

$$\mathcal{L}_{yuk} = \frac{1}{2} a \bar{N}_1^c N_1 + \frac{1}{2} b \left(\bar{N}_2^c N_3 + \bar{N}_3^c N_2 \right) + \lambda_e \bar{L}_e \tilde{H} N_1 + \lambda_\mu \bar{L}_\mu \tilde{H} N_2 + \lambda_\tau \bar{L}_\tau \tilde{H} N_3 + h.c. + \lambda_S^{12} \bar{N}_1^c N_2 S + \lambda_S^{13} \bar{N}_1^c N_3 S^\star + h.c.$$

$$m_D = \begin{pmatrix} m_{\nu_e} & 0 & 0 \\ 0 & m_{\nu_\mu} & 0 \\ 0 & 0 & m_{\nu_\tau} \end{pmatrix}, \quad m_R = \begin{pmatrix} a & s & t \\ s & 0 & b \\ t & b & 0 \end{pmatrix},$$

$m_{\nu_j} \equiv \lambda_i v / \sqrt{2}$, $s \equiv \lambda_S^{12} v_X$ and $t \equiv \lambda_S^{13} v_X$.

$$m_\nu \simeq -m_D \cdot m_R^{-1} \cdot m_D \,.$$

MSW Potential

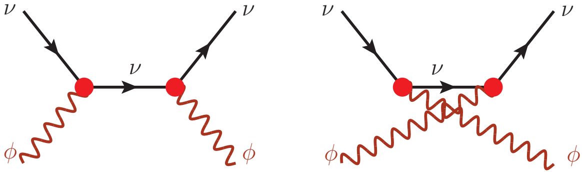
- Coherent Forward Scattering of Neutrinos on Fuzzy DM
- scalar DM

$$V_{\text{eff}} = \frac{1}{2E_\nu}\left(\phi\,(y\,m_\nu + m_\nu\,y) + \phi^2 y^2\right), \qquad \phi = \frac{\sqrt{2\rho_\phi}}{m_\phi}\cos(m_\phi t)\ ,$$

- vector DM

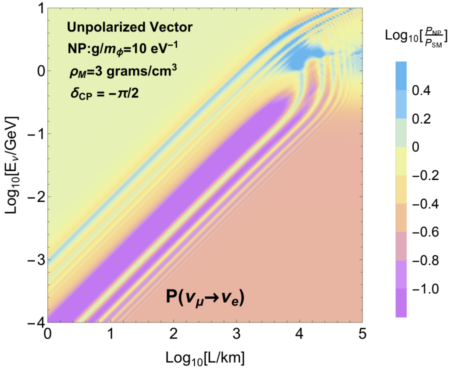
$$V_{\text{eff}} = -\frac{1}{2E_\nu}\left(2(p_\nu\cdot\phi)gQ + g^2Q^2\phi^2\right). \qquad \phi^\mu = \frac{\sqrt{2\rho_\phi}}{m_\phi}\xi^\mu\cos(m_\phi t)\,.$$

- $V_{\mu\mu}^{(T,U)} = V_{\tau\tau}^{(T,U)} = \frac{g^2\rho_\phi}{E_\nu m_\phi^2}\cos^2(m_\phi t)$
- for polarized DM we evaluate $p_\nu\cdot\phi$ assuming the polarization axis to be parallel to the ecliptic plane



Methods

- We have implemented the potential in GLoBES Huber et al. 0701187,0407333
- the time dependence of matter potential induces time dependent oscillation probabilities
- we evaluate the oscillation probabilities at several fixed times and interpolate using a second order polynomial in $cos(m_\phi t)$

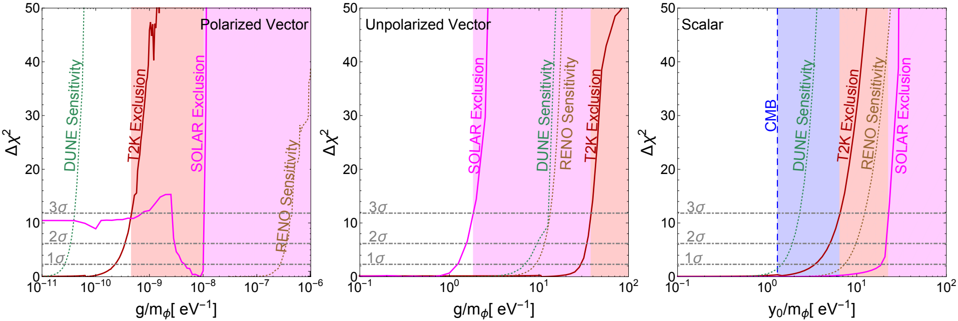


$$P(\overset{(-)}{\nu}_\alpha \to \overset{(-)}{\nu}_\beta) = P_0^{\alpha\beta}(E_\nu) + P_1^{\alpha\beta}(E_\nu) \cdot V(t) + P_2^{\alpha\beta}(E_\nu) \cdot V(t)^2 + \ldots$$

- the probability is then averaged in a given time interval T

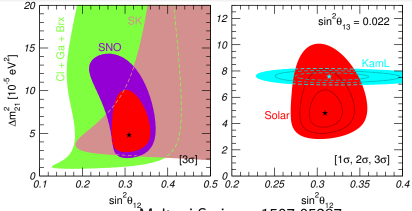
$$\bar{P}(E) = \frac{1}{T}\int_0^T dt P(E_\nu, t)$$

Constraints

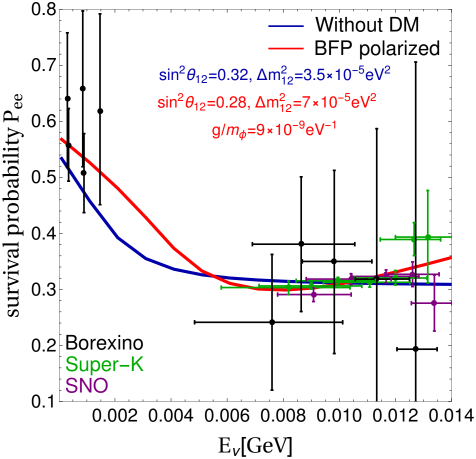


- for vector DM, the sensitivity is more than ten orders of magnitude better in the polarized case
- for scalar and polarized vector DM acceleration-based experiments give stronger limits and sensitivities
- for unpolarized vector DM, experiments at lower energies are better (energy dependence of the potential)

Impact on Solar and Astrophysical neutrinos



Maltoni,Smirnov 1507.05287
Gonzalez-Garcia, Maltoni 1307.3092



Solar

- adiabatic evolution in the sun Sun
- survival probability of electron flavor $P_{ee}(E_\nu) = \sum_i |U_{ei}^{\odot}|^2 \, |U_{ei}^{\oplus}|^2$
- fitted data from Borexino, Super-K and SNO

Astrophysical

- obtaining constraints from optical depth $\tau_\nu(E_\nu) = \sigma_{\nu\phi}(E_\nu) X_\phi m_\phi^{-1}$ with $X_\phi \equiv \int_{\mathrm{l.o.s}} dl \, \rho_\phi$
- much weaker limits in comparison to oscillation exp.

Summary

- fuzzy DM is an interesting alternative to WIMP
- fuzzy neutrinophilic DM has recently received attention (Berlin 1608.01307, Krnjaić et al. 1705.06740)
- we have demonstrated that unique opportunities exist at current and future neutrino oscillation experiments to probe interactions between neutrinos and ultra-light DM particles
- possible connections with LHCb anomalies