

Dark Gravity confronted with SN, BAO and the CMB

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Dark Gravity

Regularly updated review not on arXiv (too many versions !) but here :

www.darksideofgravity.com/DG.pdf
DG : Gravity with its Dark side

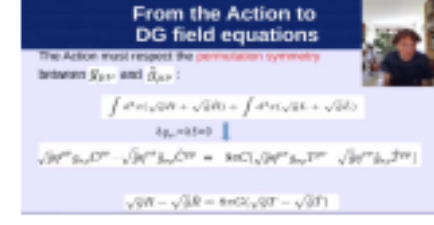
- DG : Theoretical review (part 1)

<https://www.youtube.com/watch?v=FkehW60m9fc>



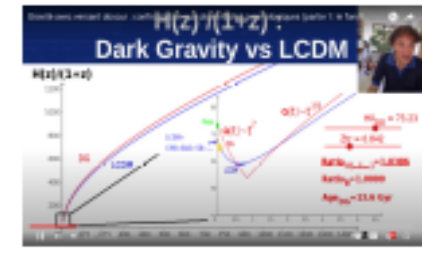
- DG : Theoretical review (part 2)

<https://www.youtube.com/watch?v=VYqkX1UyEo>



- DG : Confrontation with cosmological observables (part 1: background)

<https://www.youtube.com/watch?v=ikbWTctPDkU>



- DG : Confrontation with cosmological observables (part 2: fluctuations)

https://www.youtube.com/watch?v=_cL5m4A7SU



From background dependence to Dark Gravity (DG)

How far can we go ?

GR : $g_{\mu\nu}$

DG : $g_{\mu\nu}$ and $\eta_{\mu\nu}$

$$\text{Riem}(\eta_{\mu\nu})=0$$

$\Rightarrow g_{\mu\nu}$ has a twin, « the inverse metric » $\tilde{g}_{\mu\nu}$

$$\tilde{g}_{\mu\nu} = \eta_{\mu\rho}\eta_{\nu\sigma} [g^{-1}]^{\rho\sigma}$$

$\Rightarrow (g_{\mu\nu}, \tilde{g}_{\mu\nu})$ is a Janus field



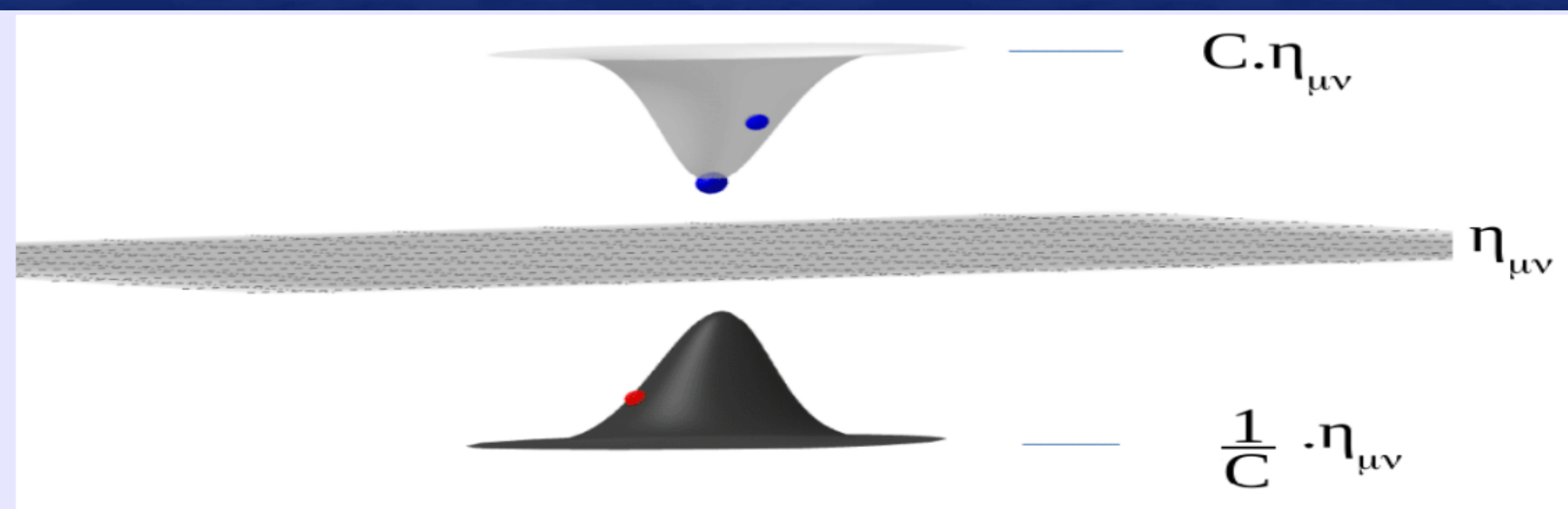
The Action must respect the permutation symmetry between $g_{\mu\nu}$ and $\tilde{g}_{\mu\nu}$:

$$\int d^4x (\sqrt{g}R + \sqrt{\tilde{g}}\tilde{R}) + \int d^4x (\sqrt{g}L + \sqrt{\tilde{g}}\tilde{L})$$

$$\delta g_{\mu\nu} \Rightarrow \delta S = 0$$

$$\sqrt{g}\eta^{\mu\sigma}g_{\sigma\rho}G^{\rho\nu} - \sqrt{\tilde{g}}\eta^{\nu\sigma}\tilde{g}_{\sigma\rho}\tilde{G}^{\rho\mu} = -8\pi G(\sqrt{g}\eta^{\mu\sigma}g_{\sigma\rho}T^{\rho\nu} - \sqrt{\tilde{g}}\eta^{\nu\sigma}\tilde{g}_{\sigma\rho}\tilde{T}^{\rho\mu})$$

The static isotropic solution



- Antigravity without run away !
- Asymptotic C matters : GR corresponds to C infinite

C=1

DG:

$$g_{ii}(r) = A = e^{2MG/r} \approx 1 + 2\frac{MG}{r} + 2\frac{M^2G^2}{r^2}$$

$$-g_{00}(r) = \frac{1}{A} = e^{-2MG/r} \approx 1 - 2\frac{MG}{r} + 2\frac{M^2G^2}{r^2} - \frac{4M^3G^3}{3r^3}$$

C=∞

RG (Schwarzschild) :

$$g_{ii}(r) = \left(1 + \frac{MG}{2r}\right)^2 \approx 1 + 2\frac{MG}{r} + \frac{3M^2G^2}{2r^2}$$

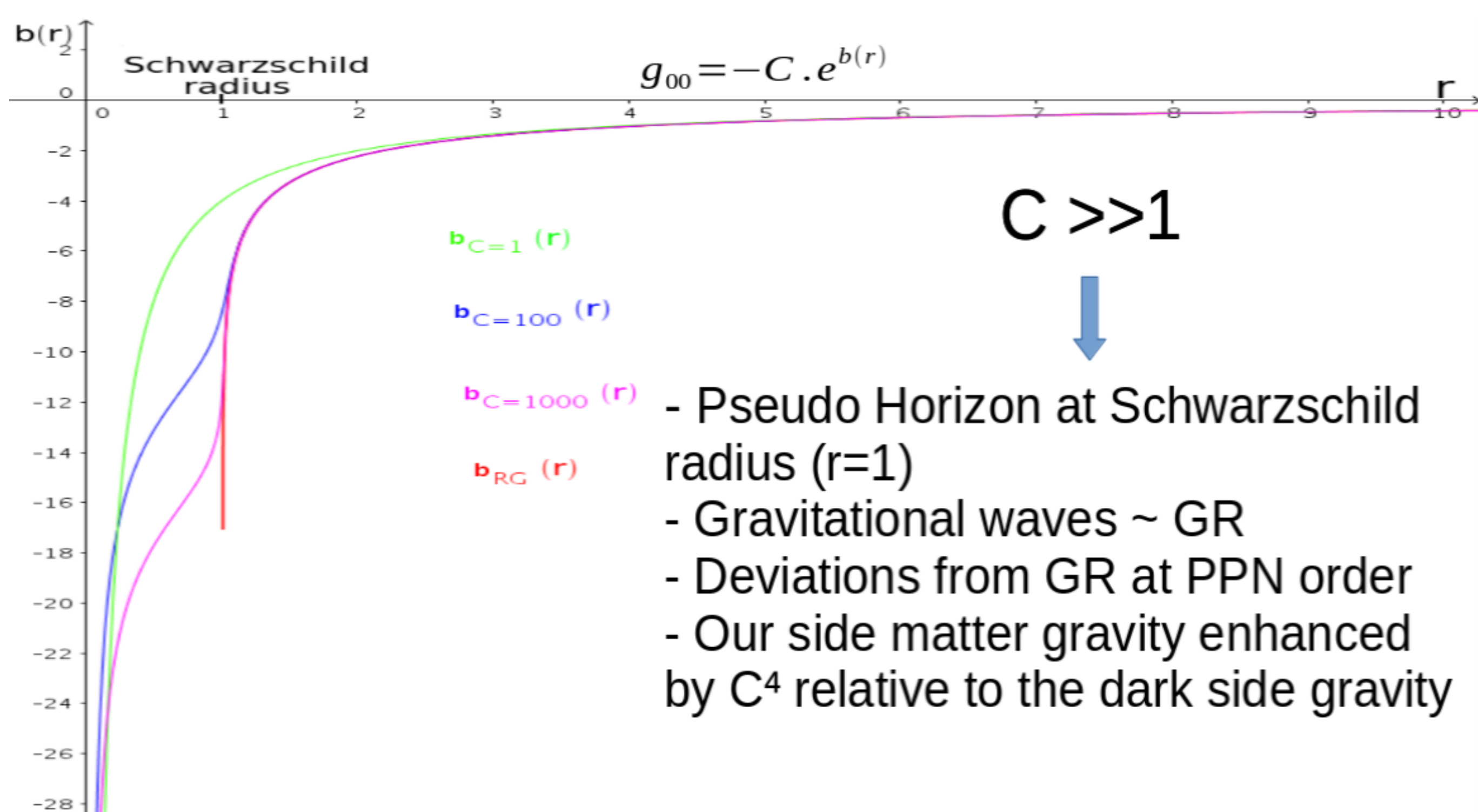
$$-g_{00}(r) = \left(\frac{1 - \frac{MG}{2r}}{1 + \frac{MG}{2r}}\right)^2 \approx 1 - 2\frac{MG}{r} + 2\frac{M^2G^2}{r^2} - \frac{3M^3G^3}{2r^3}$$

- No Horizon
- Zero Gravitational Waves

$$\tilde{h}_{\mu\nu} = -h_{\mu\nu} + O(h^2)$$

$$2(R_{\mu\nu}^{(1)} - \frac{1}{2}\eta_{\mu\nu}R_{\lambda}^{(1)\lambda}) = -8\pi G(T_{\mu\nu} - \tilde{T}_{\mu\nu} + \tilde{t}_{\mu\nu} - \tilde{t}_{\mu\nu}) \rightarrow 0$$

- Deviations from GR at PPN order only



DG Cosmology

$$g_{\mu\nu} = a^2(t)\eta_{\mu\nu}$$

$$\tilde{a}(t) = \frac{1}{a(t)}$$

$$\tilde{g}_{\mu\nu} = a^{-2}(t)\eta_{\mu\nu}$$

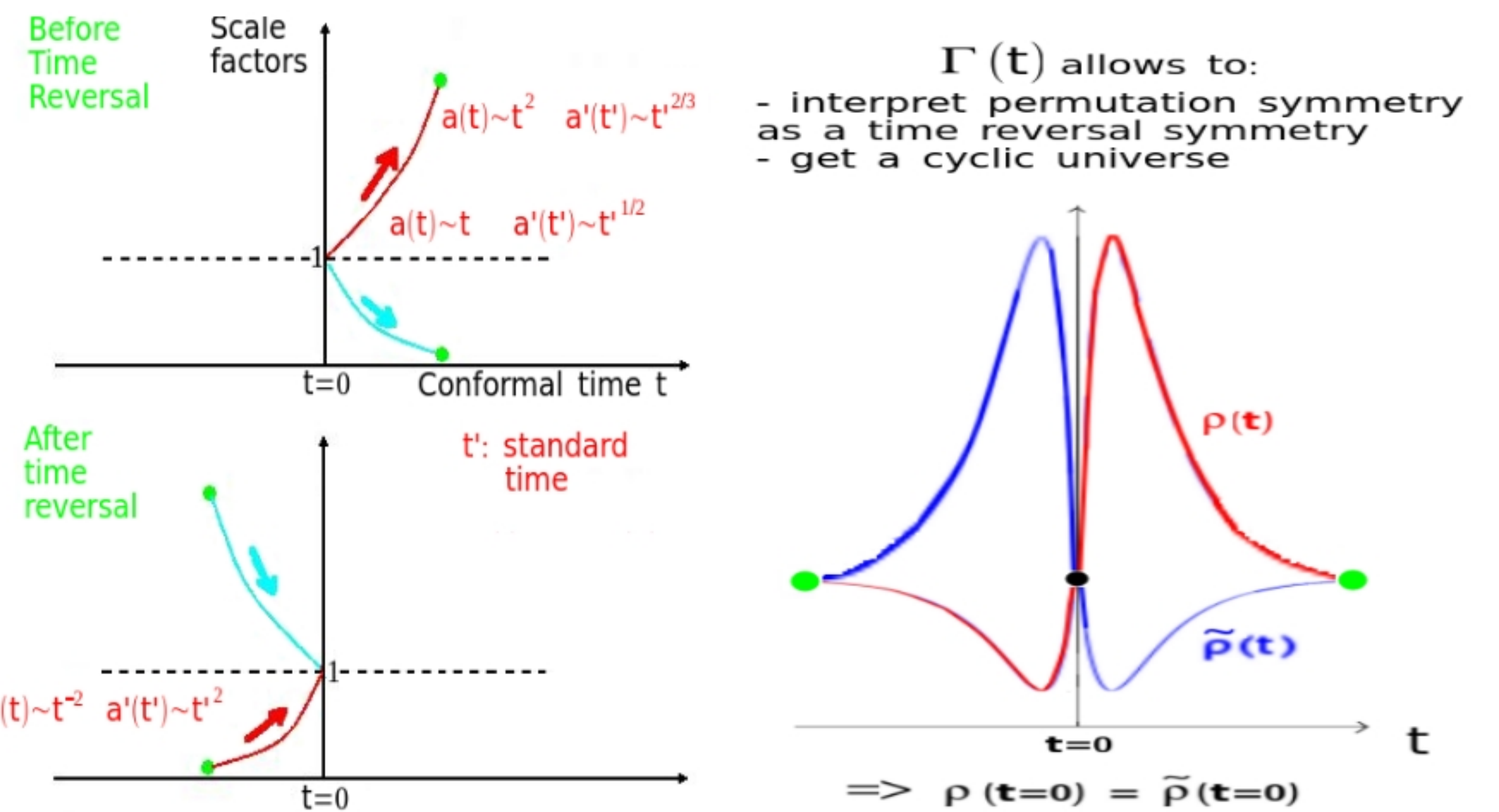
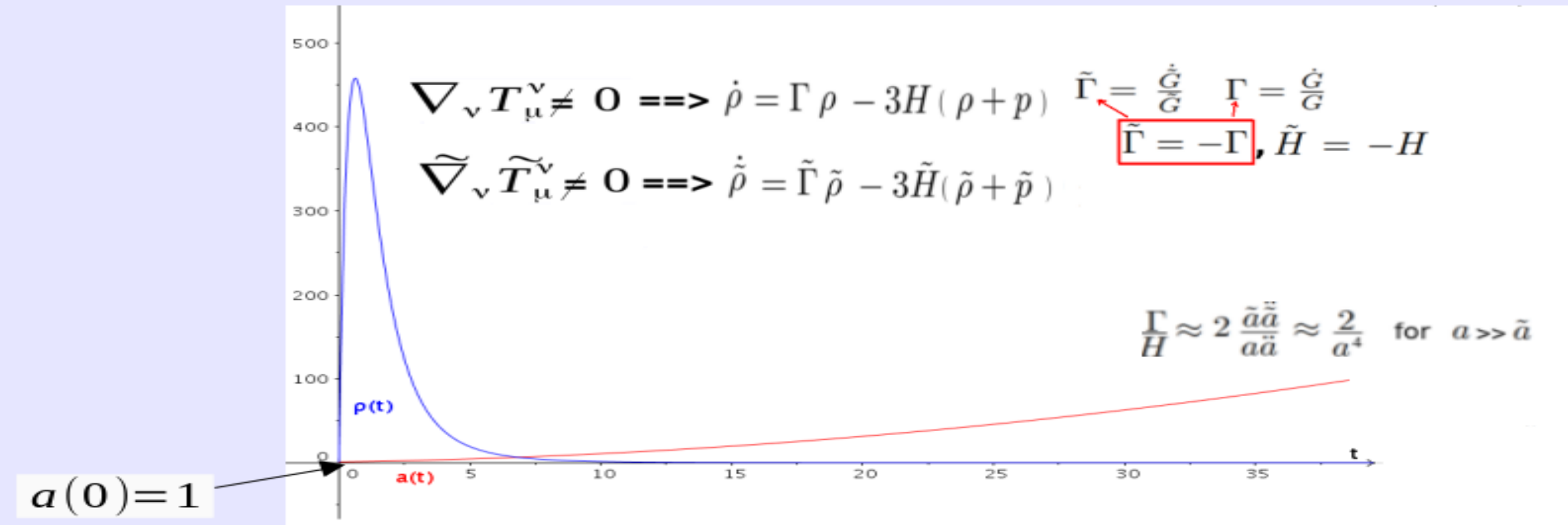
- Problem : Homogeneous & isotropic Janus solution is flat but static !

$$a^2 H^2 - \tilde{a}^2 \tilde{H}^2 = \frac{8\pi G}{3}(a^4 \rho - \tilde{a}^4 \tilde{\rho})$$

$$a^2(2\dot{H} + H^2) - \tilde{a}^2(2\dot{\tilde{H}} + \tilde{H}^2) = -8\pi G(a^4 p - \tilde{a}^4 \tilde{p})$$

- Solution : Introduce offshell mechanism $\Gamma(t)$: variable $\tilde{G}(t) = \frac{1}{G(t)}$

- Differential equations can be solved numerically :



- Hyp : $\rho = \tilde{\rho}$ occurred at transition redshift

triggering $\Gamma \Rightarrow a(t) \sim t^{-2}$

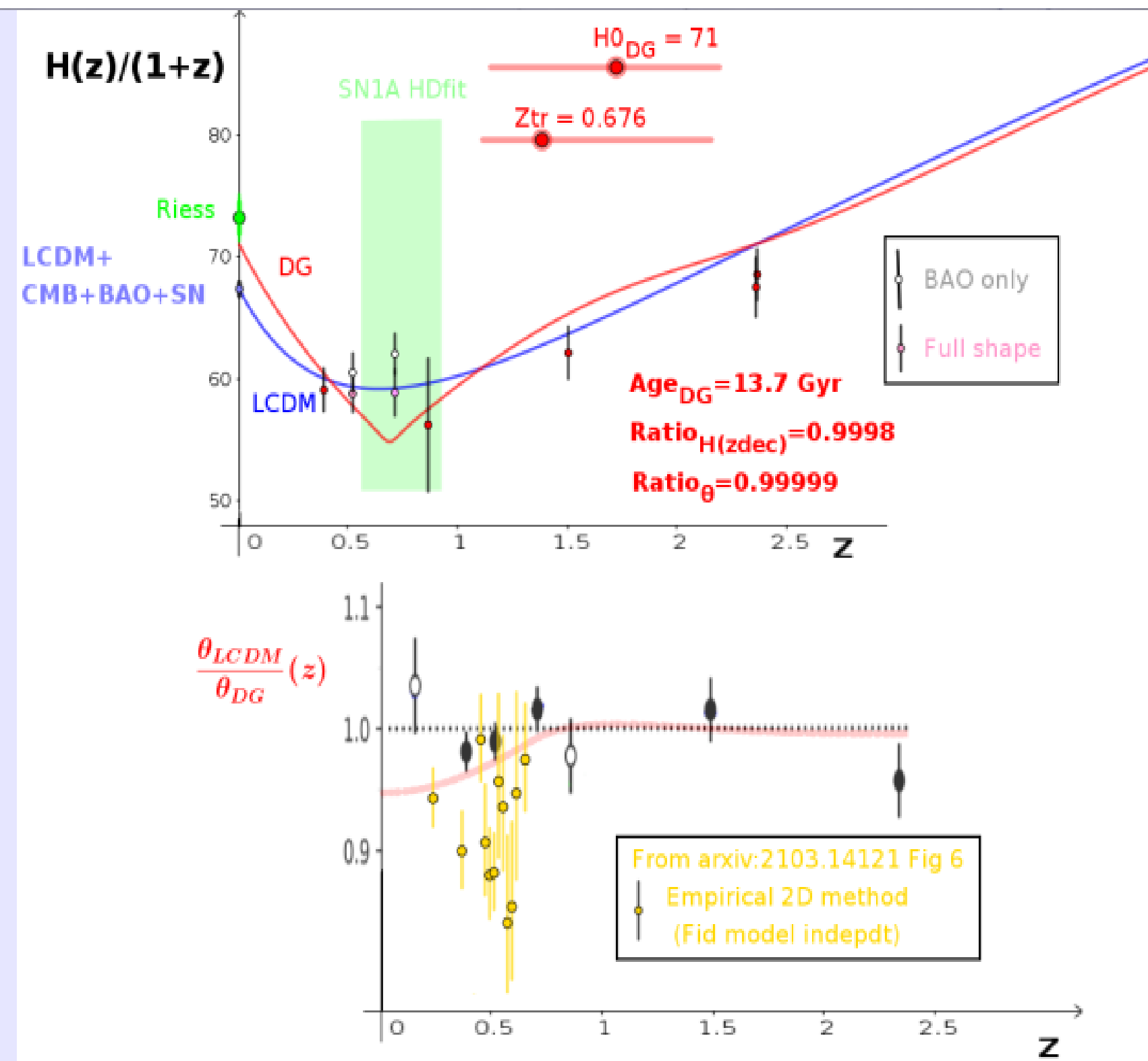
With $H'(t)$ continuous at the transition from $t^{2/3}$ to t^0 and assuming same rough universe age as within LCDM ($\approx 1/H_0$)

$$z_{tr} = \left(\frac{2/3 - \alpha}{1 - \alpha}\right)^\alpha - 1$$

$\Rightarrow z_{tr} = 0.78$ vs observed LCDM $z_{tr} = 0.67 \pm 0.1$

- Close to LCDM scale factor evolution
- Without DE
- Inflation not needed to get $k=0$
- Without Big Bang singularity
- Cosmological DM still needed
- Dark side effects only since t_r or near $t=0$

- Good agreement with BAO transverse, CMB and SN data !



- Tension in BAO LoS for BAO only method
- But BAO only is Fid Model Biased according arxiv 1811.12312 (Anselmi et al.)

Planck and matter power spectra

Planck arxiv:1807.06209
Class software arxiv:1104.2933

- Planck TT+TE+EE

$$\chi^2_{DG} = 6701.9$$

$$\chi^2_{DG} - \chi^2_{LCDM} = +4.4$$

