

Shedding light on modified gravity theories using voids

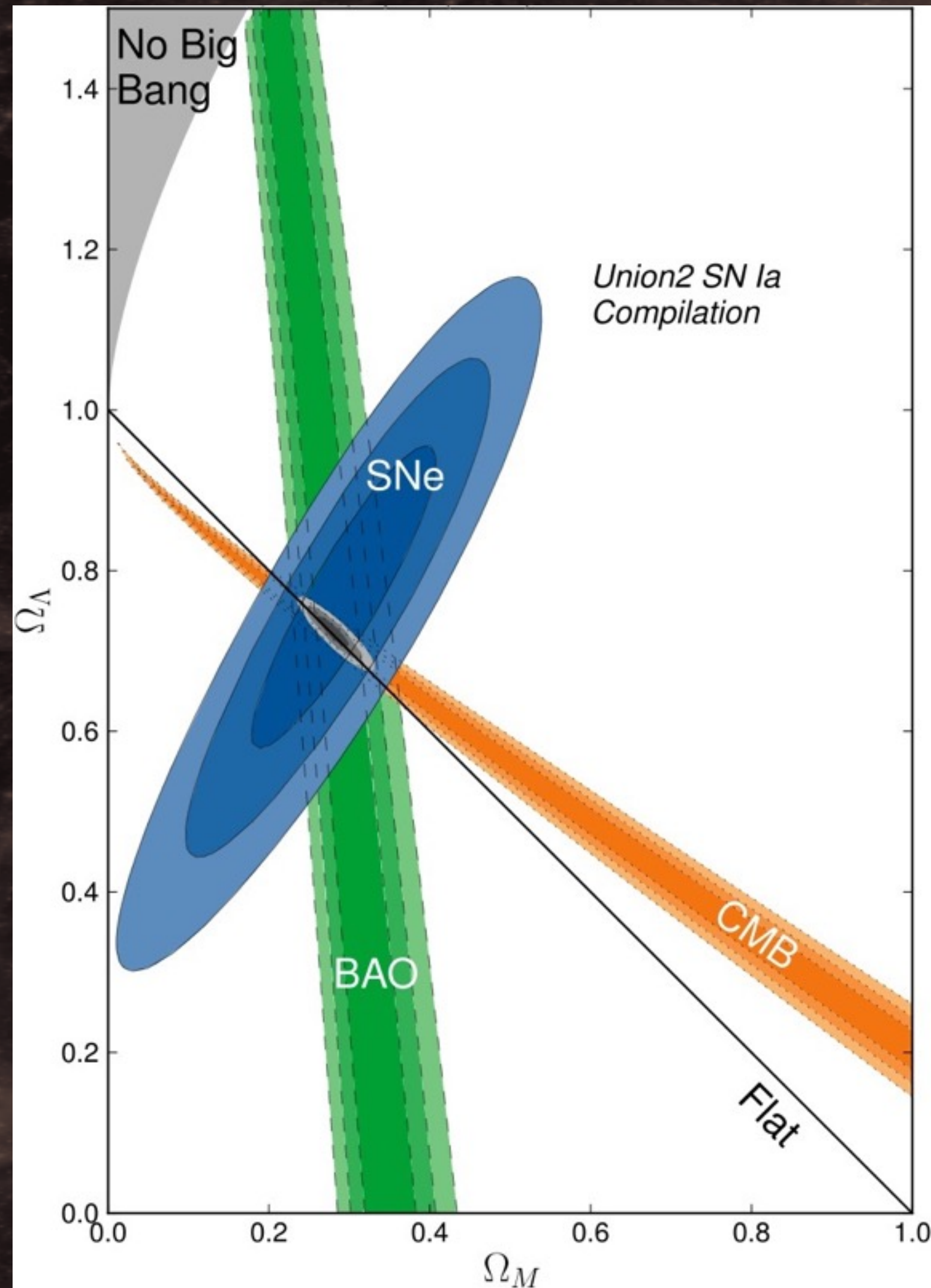
Marius Cautun

Christopher Davies, Baojiu Li (Durham),
Enrique Paillas, Joaquin Armijo, Nelson Padilla (Santiago),
Sownak Bose (Harvard), Yan-Chuan Cai (Edinburgh),

Cosmology 2018
Dubrovnik, Croatia
22 October 2018

MC+, MNRAS 476, 2018
Davies, MC+, MNRAS 480, 2018
Paillas, MC+, arXiv:1810.02864

Overview



Amanullah+
(2010)

$$R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} = -8\pi GT_{\mu\nu}$$

General Relativity + dark energy:

$$R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} = -8\pi GT_{\mu\nu} + \Lambda g_{\mu\nu}$$

Modify General Relativity:

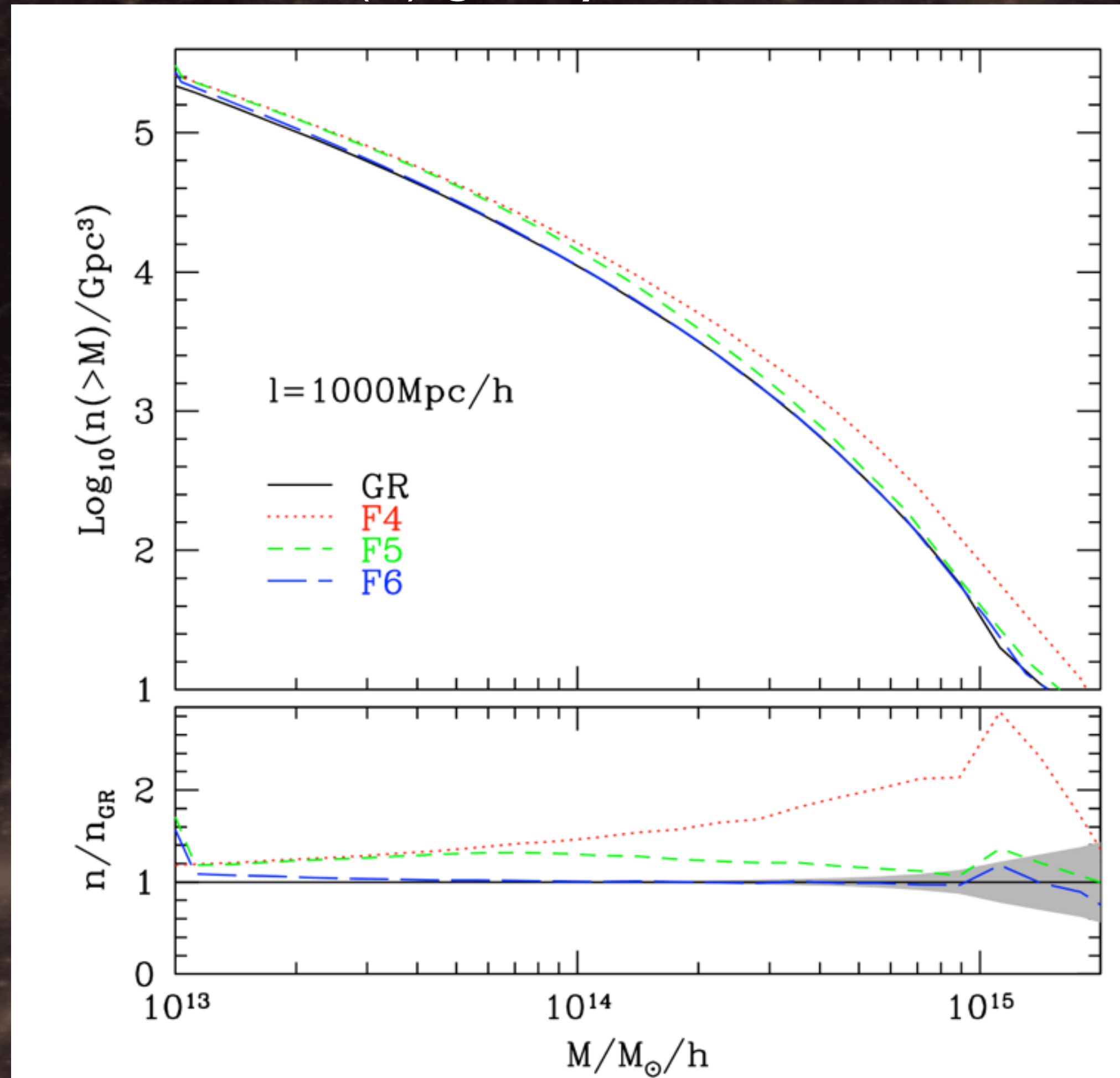
$$R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} - \Lambda g_{\mu\nu} = -8\pi GT_{\mu\nu}$$

Voids as probes of modified gravity models

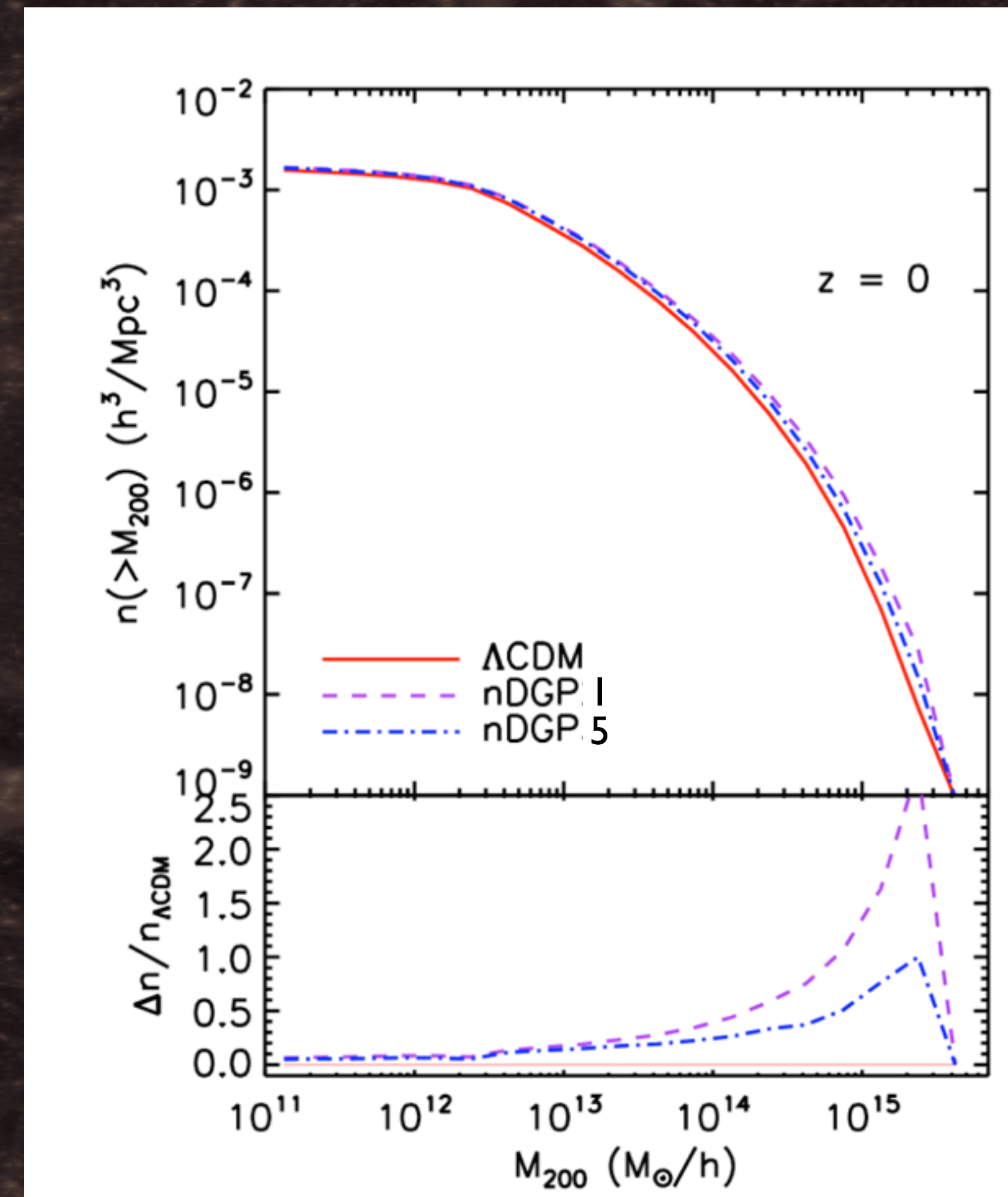
- The accelerated expansion of the Universe can be explained by modifying General Relativity, i.e. Modified Gravity (MG) models
- Such theories have additional degrees of freedom (e.g. scalar fields) that give rise to fifth forces
- Due to stringent solar system constraints, such forces must vanish in high density regions (*screening mechanisms*)
- These fifth forces can attain maximum values in low density regions, i.e. voids
- Many models have similar phenomenology, so I will focus on two such models: $f(R)$ (Hu & Sawicki 2007) and nDGP (Dvali+ 2000)

Halo mass function in MG models

f(R) gravity models



nDGP gravity models



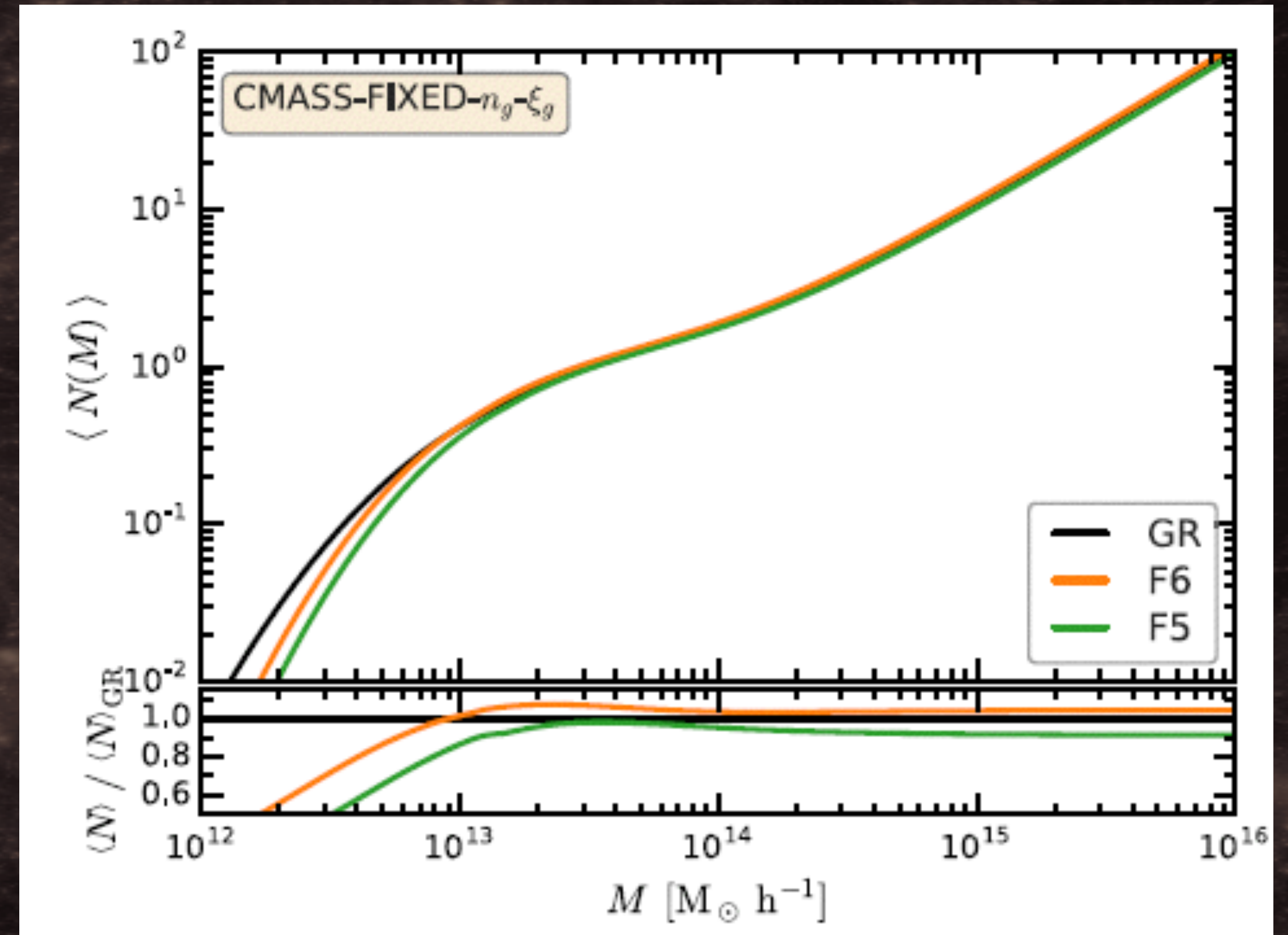
Cai+
(2015);

Falck+
(2018)

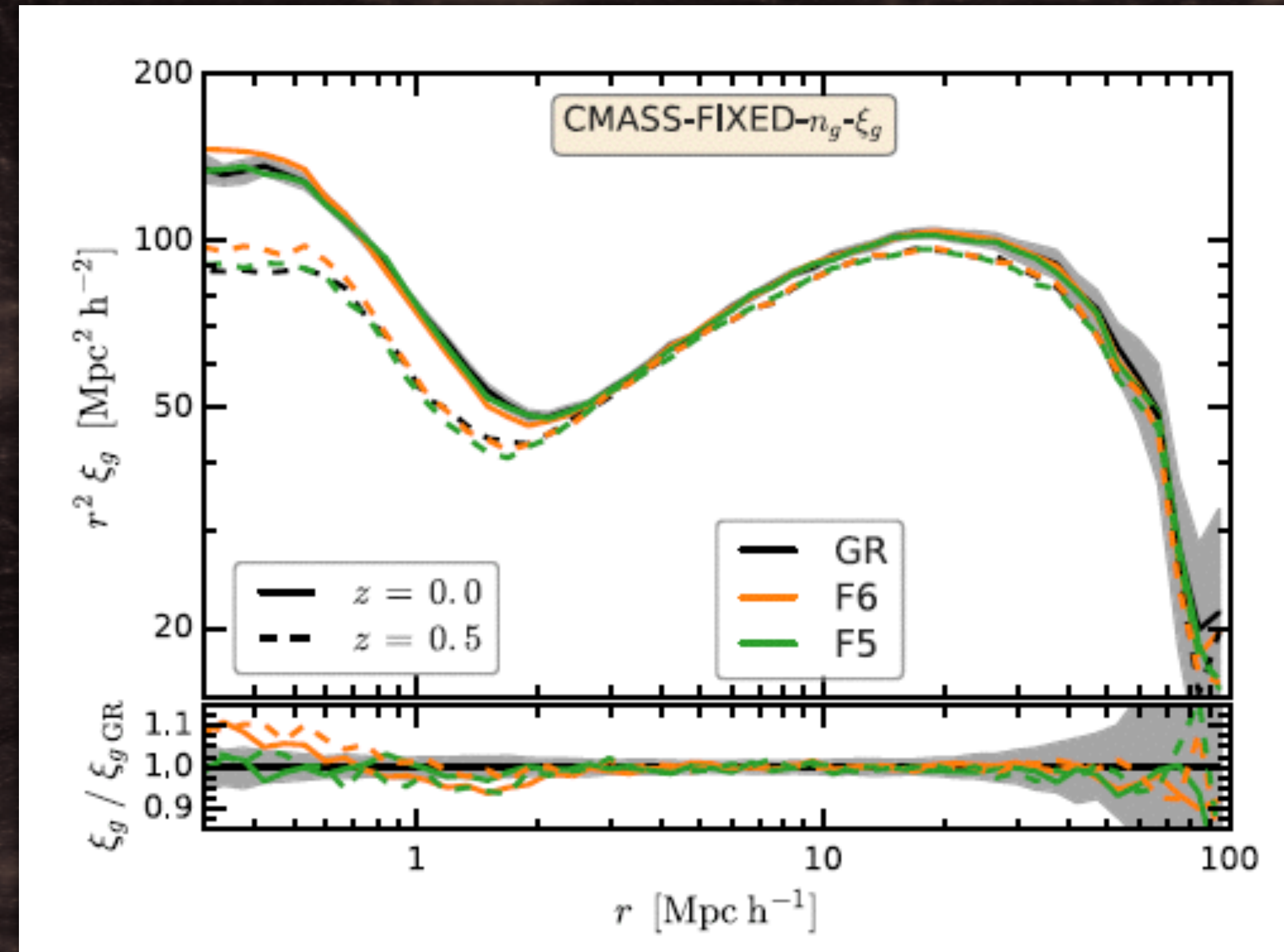
Tracer field

Distribution of galaxies as realistic as possible:

- HOD galaxies with number density and 2-pt correlation function as in SDSS CMASS sample.
- Match the 2-pt correlation function of galaxies in GR and MG models.



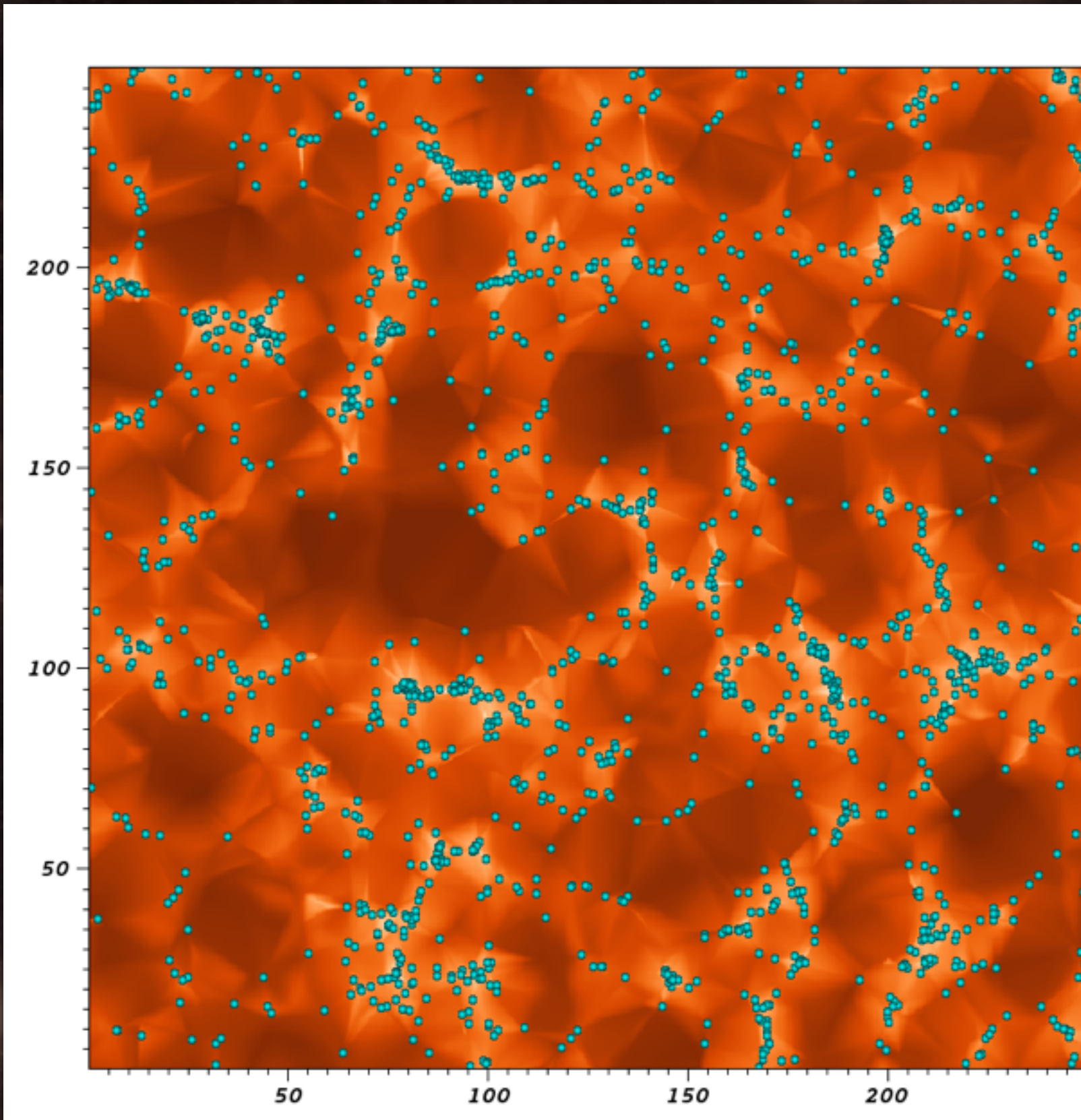
Same galaxy clustering



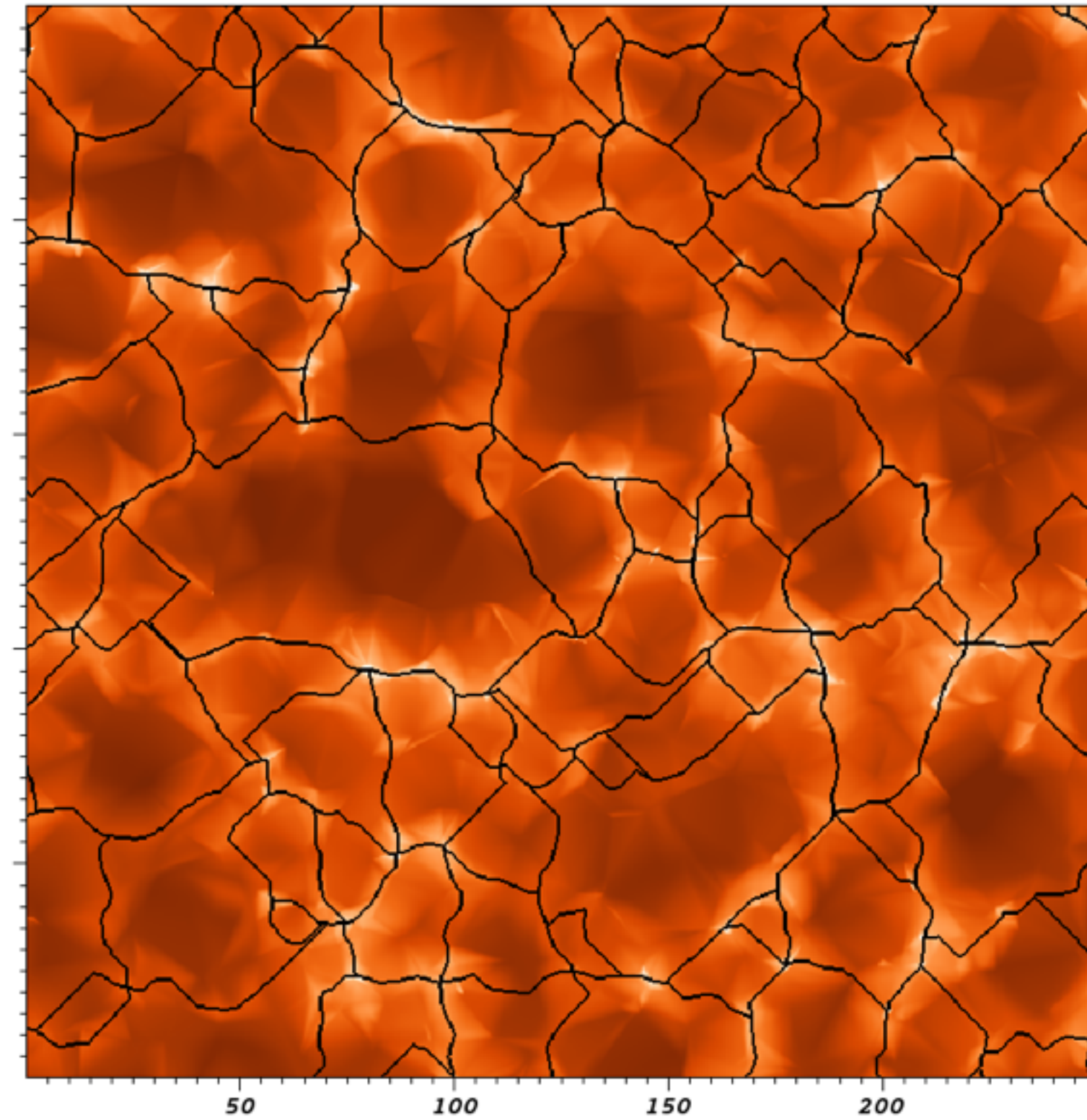
MC+(2018)

Void identification

Distribution of galaxies



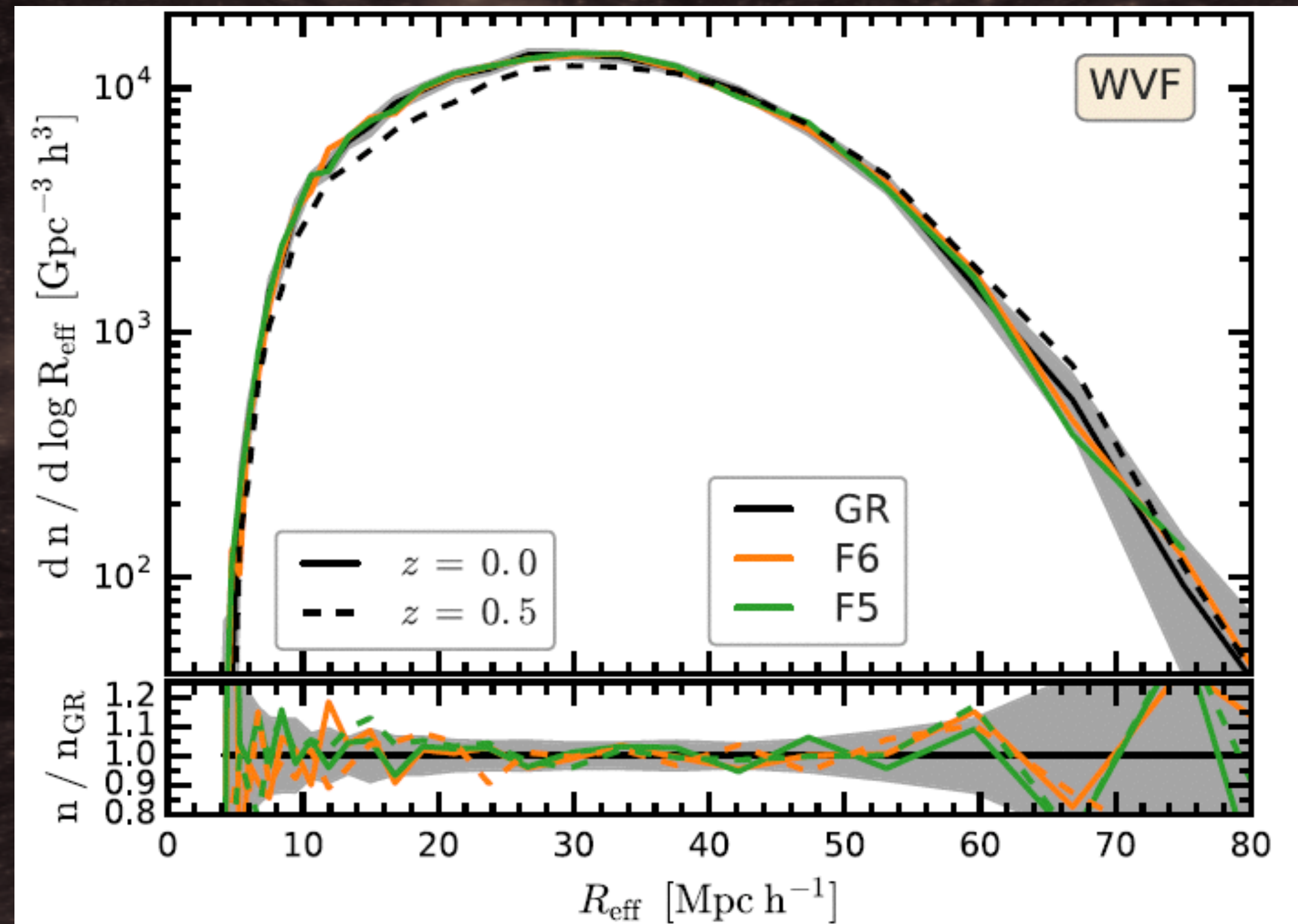
Distribution of voids



Watershed void
finder
(Platen + 2007)

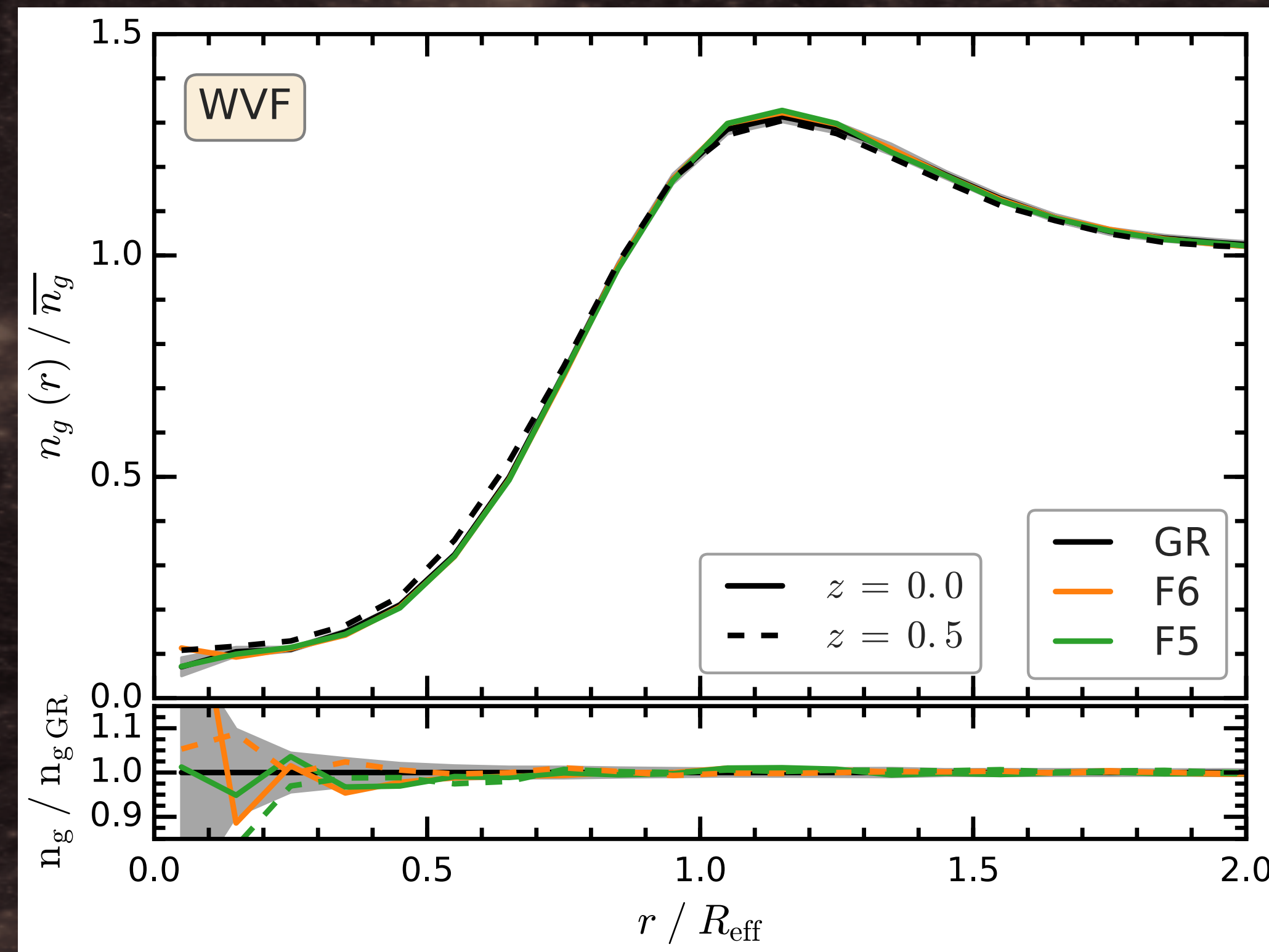
MC+(2018)

Distribution of void sizes



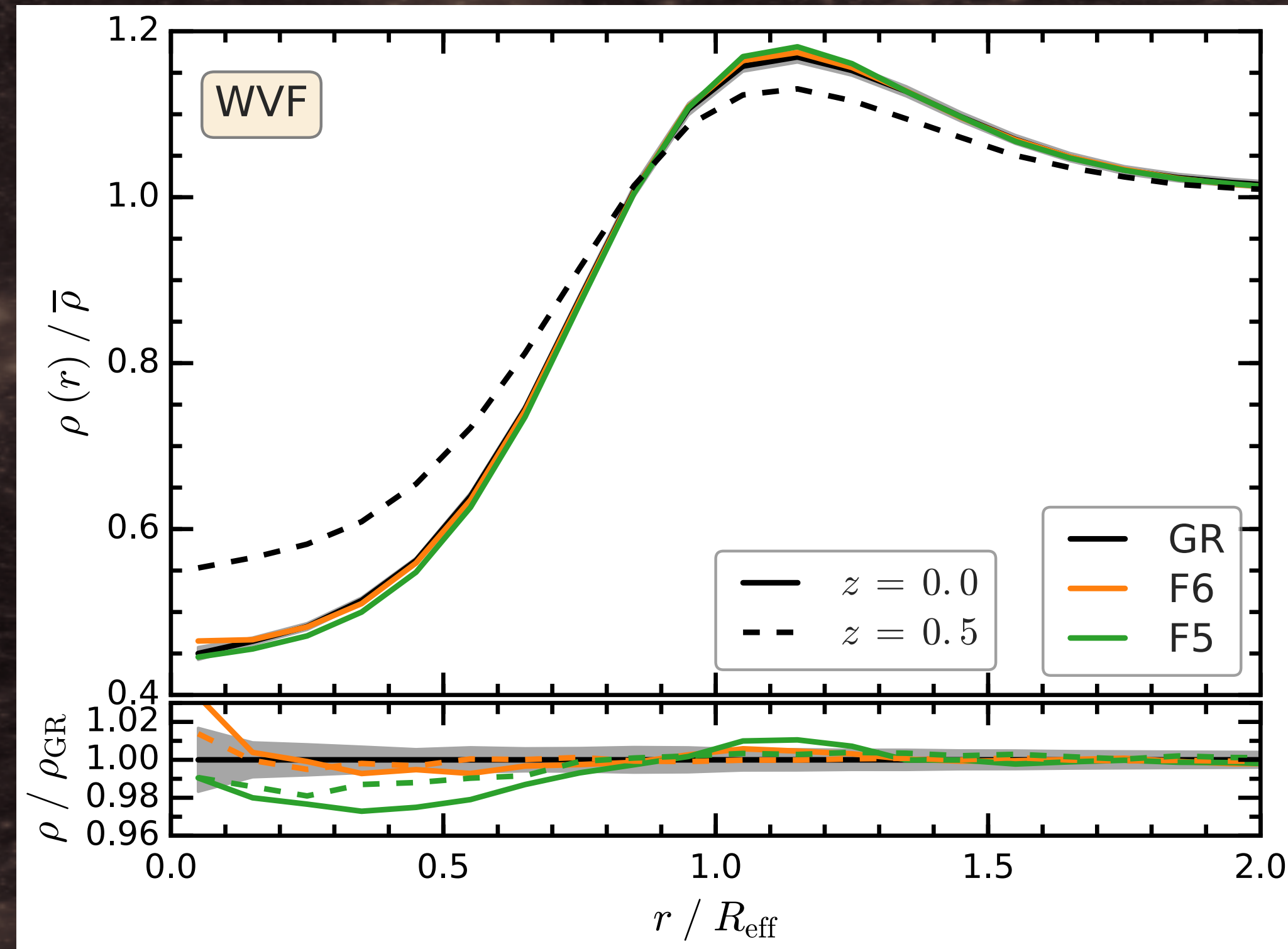
MC+(2018)

Galaxy density profiles



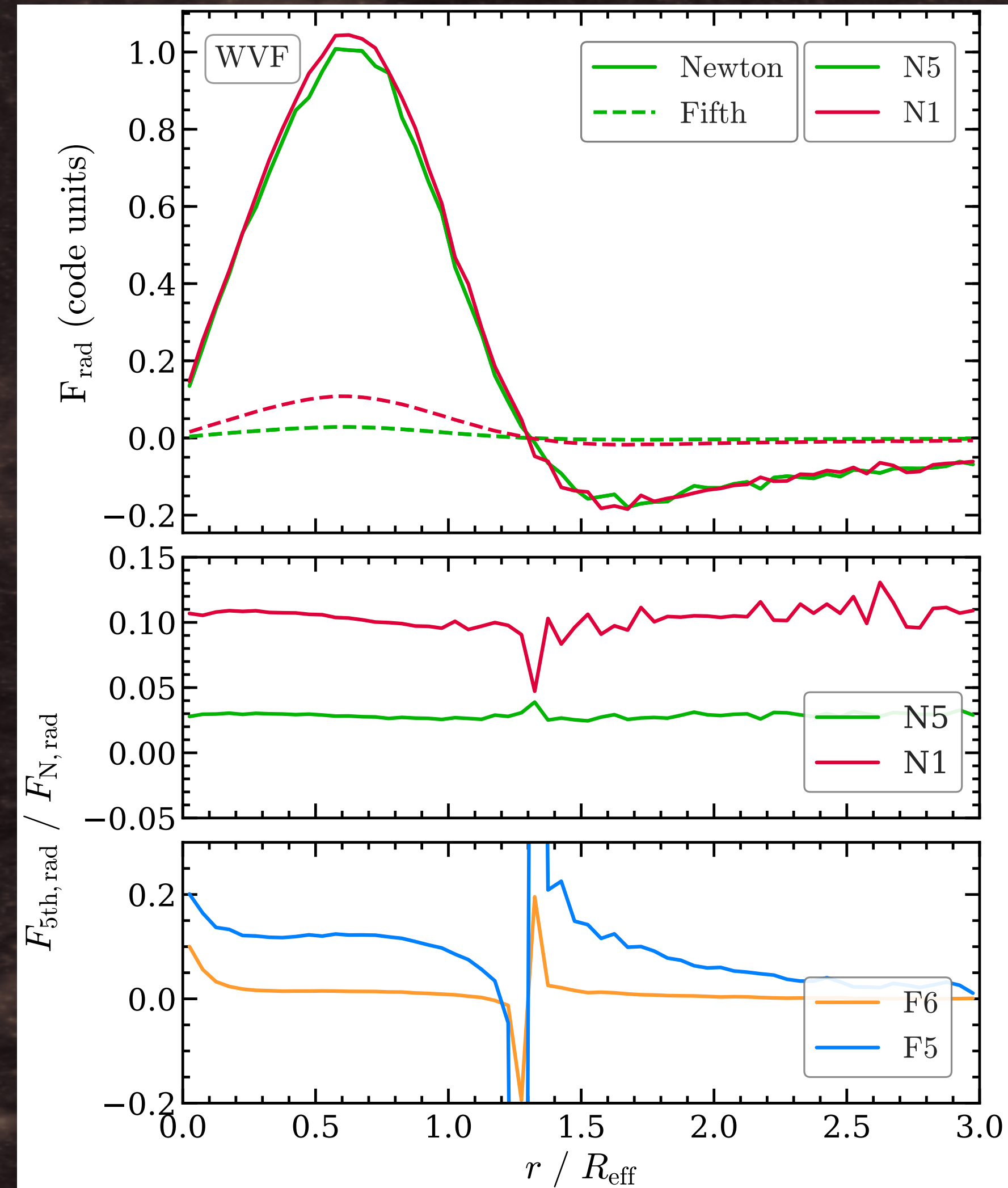
MC+(2018)

Mass density profiles



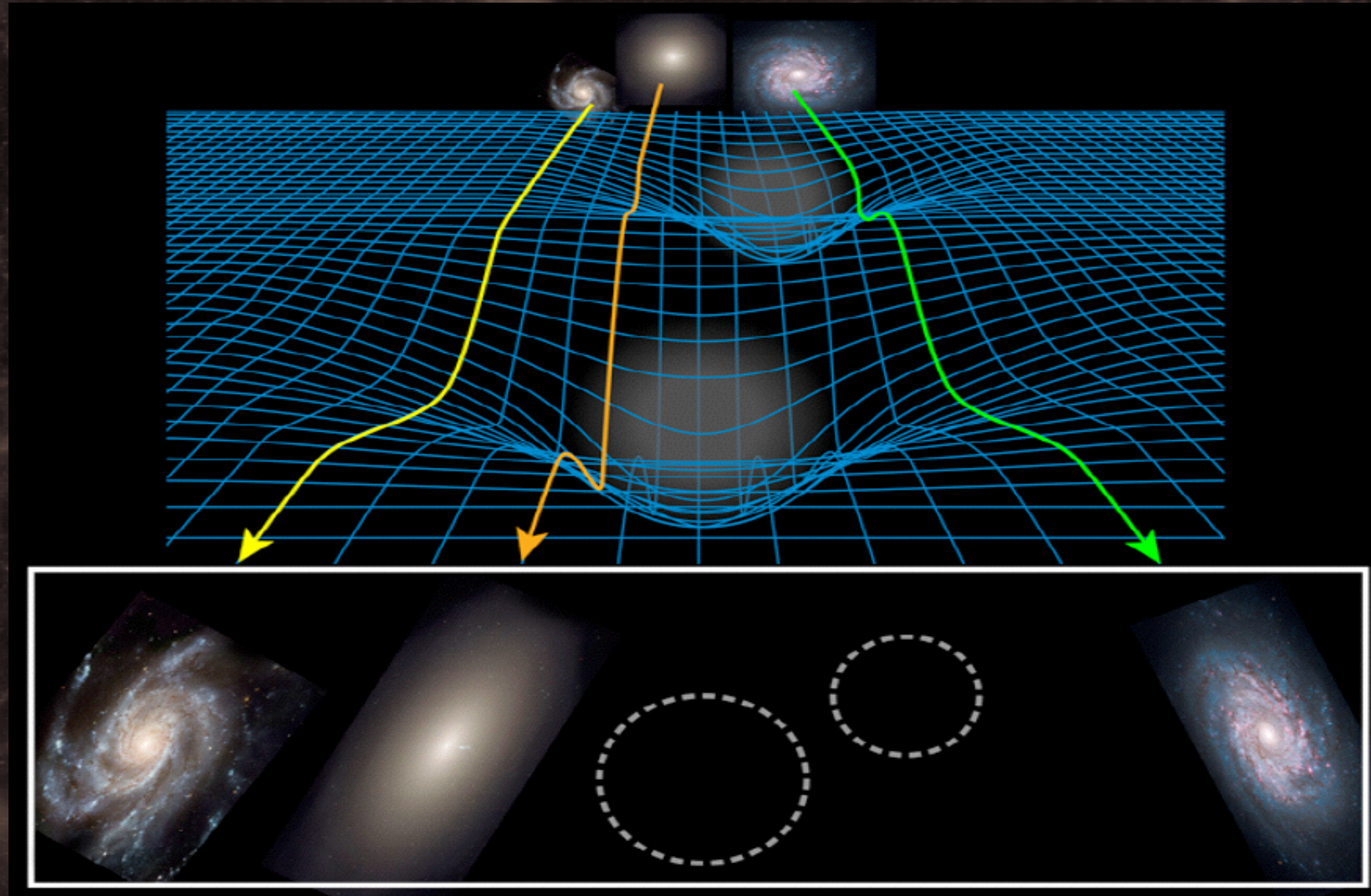
MC+(2018)

Why more underdone voids?



Paillas + (2018)







Gravitational lensing



Gravitational lensing

Convergence

Shear

	< 0	> 0
κ		
$\text{Re}[\gamma]$		
$\text{Im}[\gamma]$		

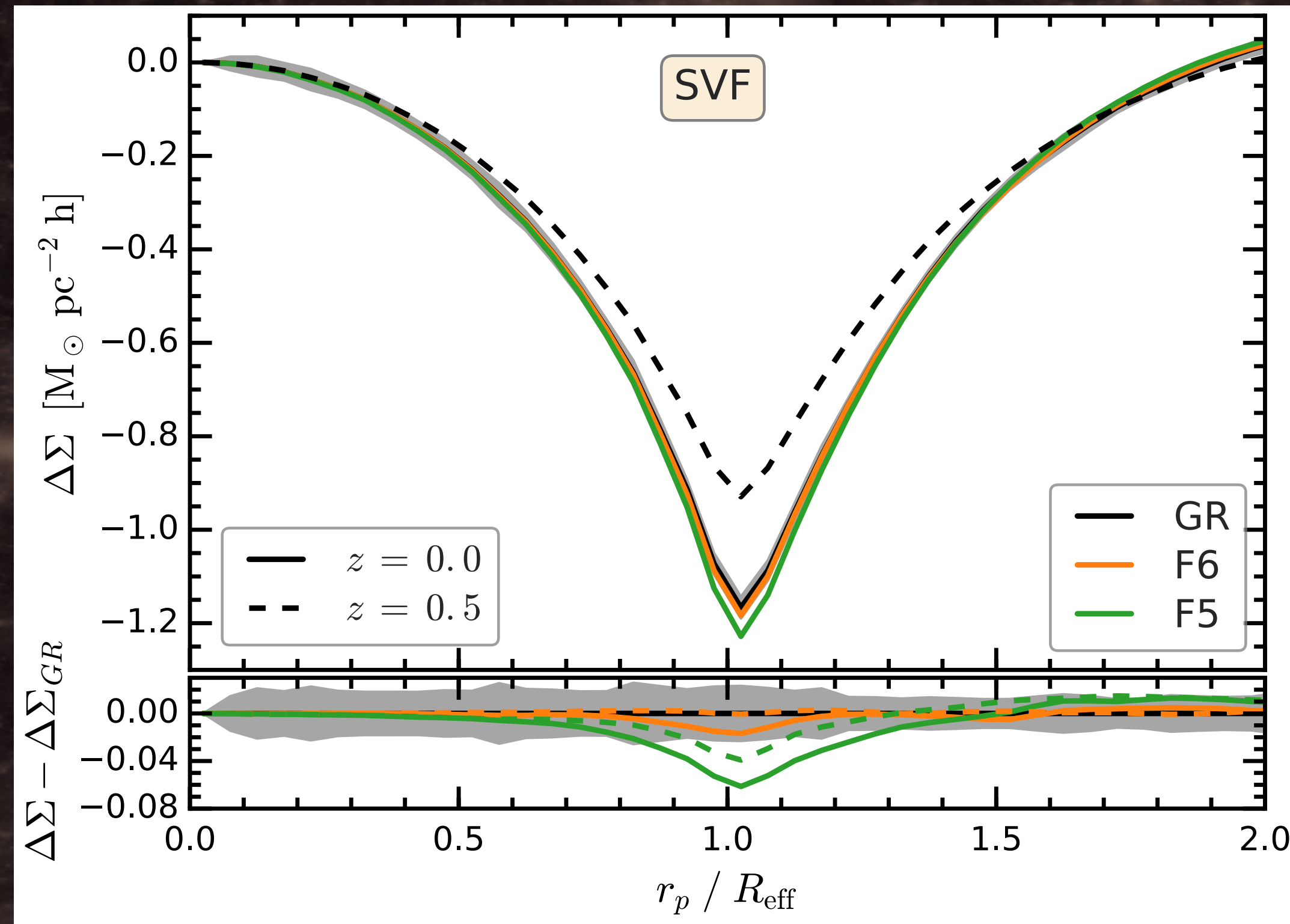
- magnification of background galaxies
- proportional to the line-of-sight projected density

- change in the shape of background galaxies
- proportional to the line-of-sight projected density

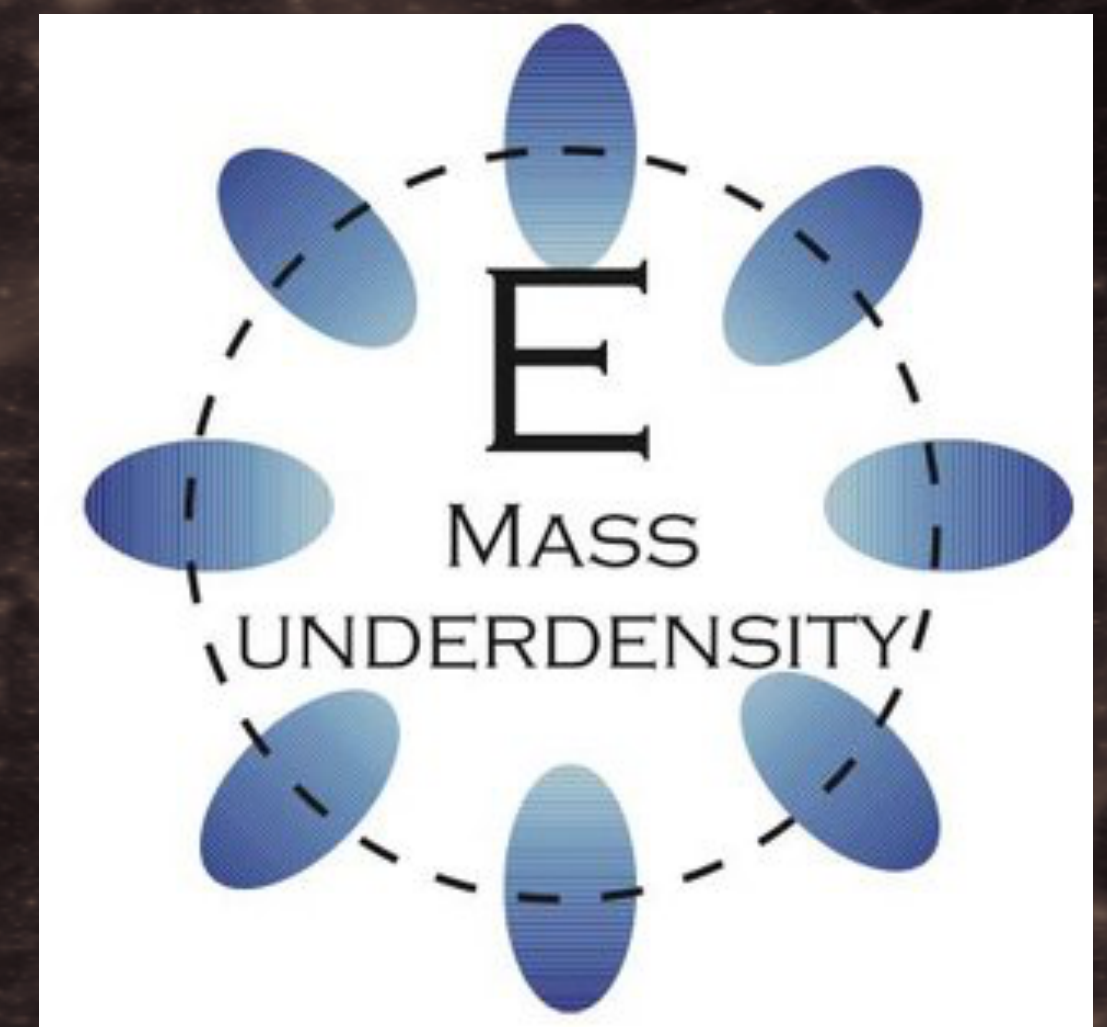
$$\gamma_t = \bar{\kappa}(< r) - \kappa(r)$$

Tangential shear profiles

$$\gamma_t(r) = \frac{\Delta\Sigma(r)}{\Sigma_c}$$



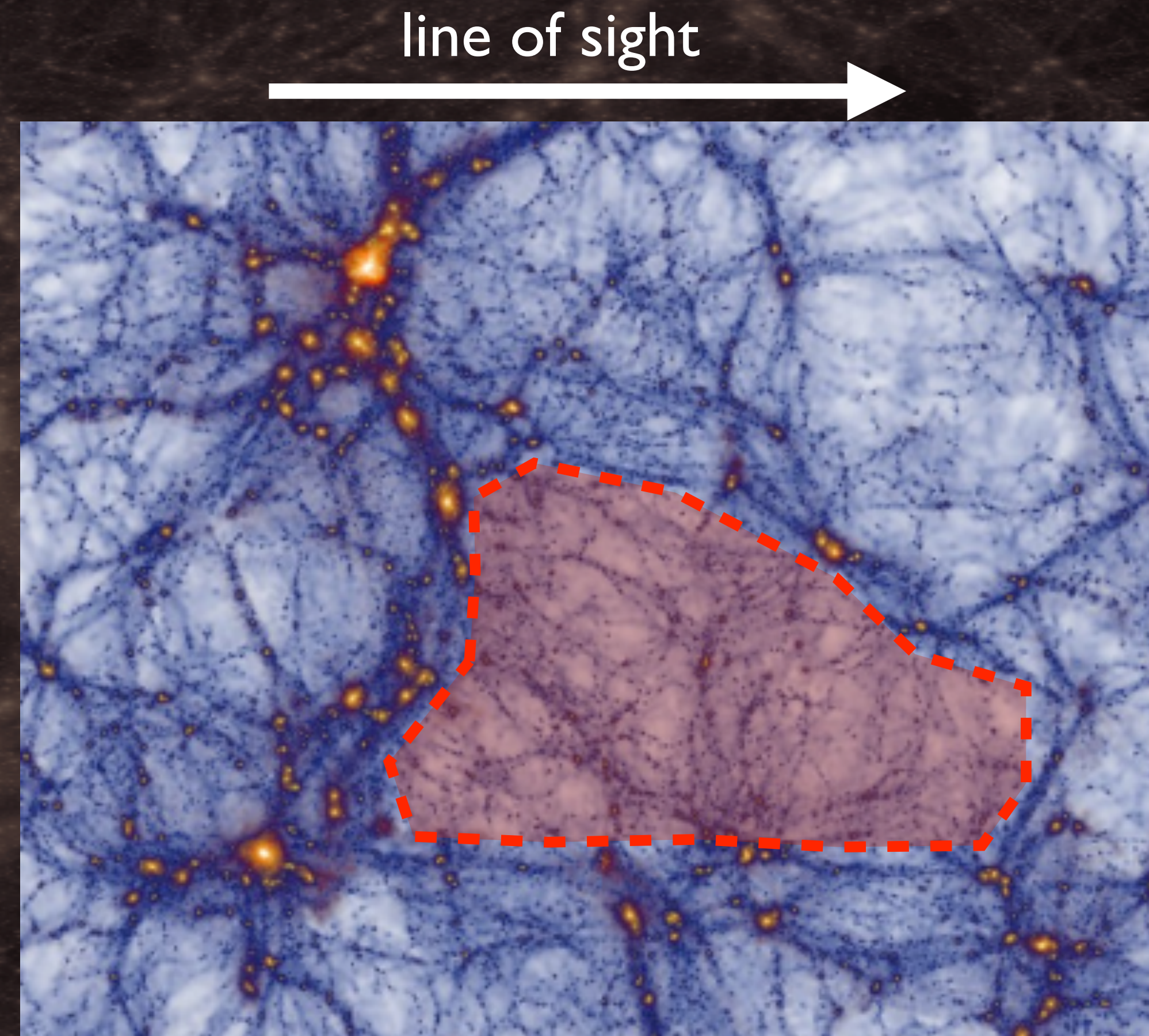
MC+(2018)



Why is weak lensing by voids so weak?

$$\gamma_t(r) = \frac{\Delta\Sigma(r)}{\Sigma_c} = \frac{\bar{\Sigma}(<r) - \Sigma(r)}{\Sigma_c}$$

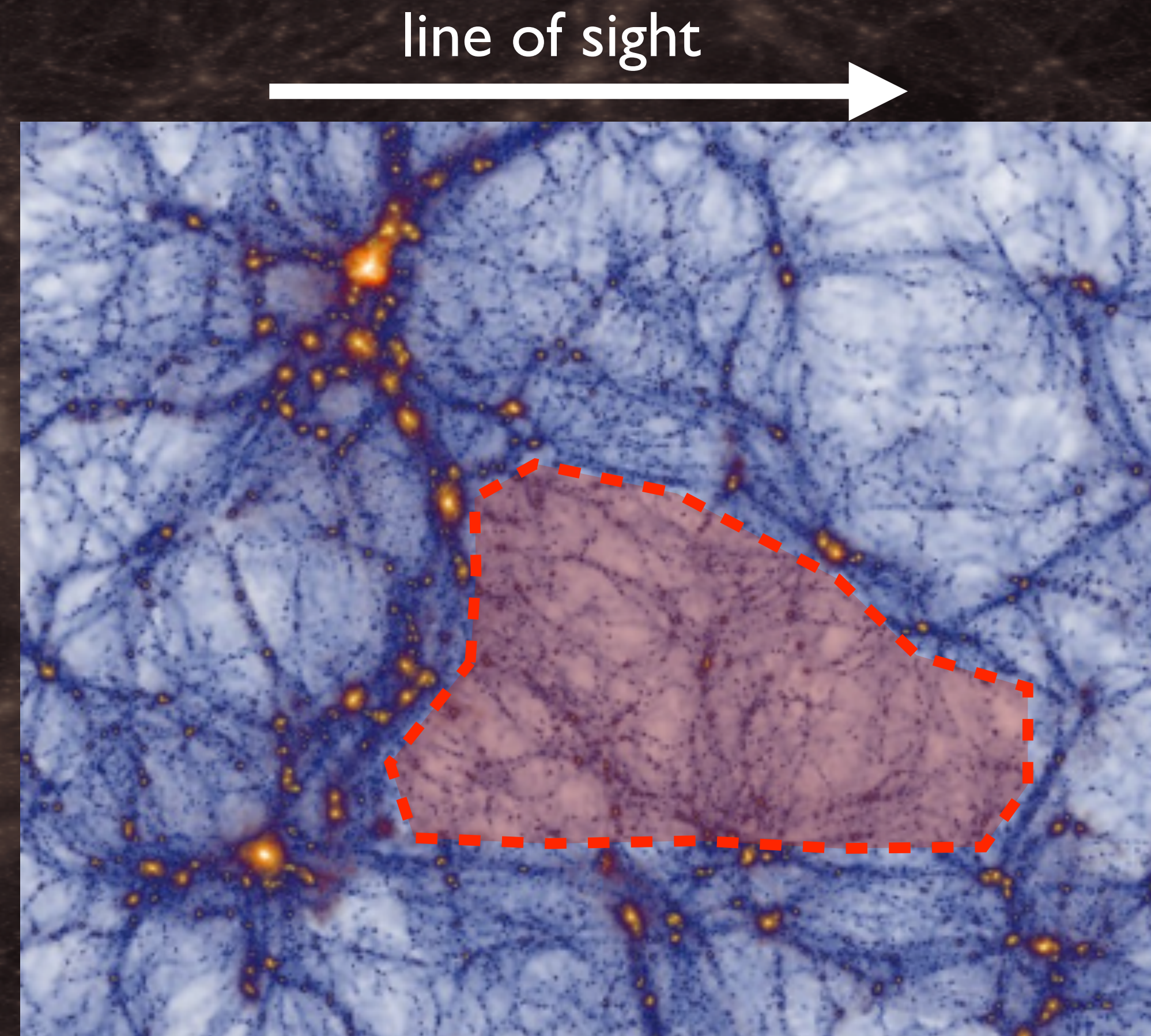
- The lensing signal depends on the projected line-of-sight density



Why is weak lensing by voids so weak?

$$\gamma_t(r) = \frac{\Delta\Sigma(r)}{\Sigma_c} = \frac{\bar{\Sigma}(<r) - \Sigma(r)}{\Sigma_c}$$

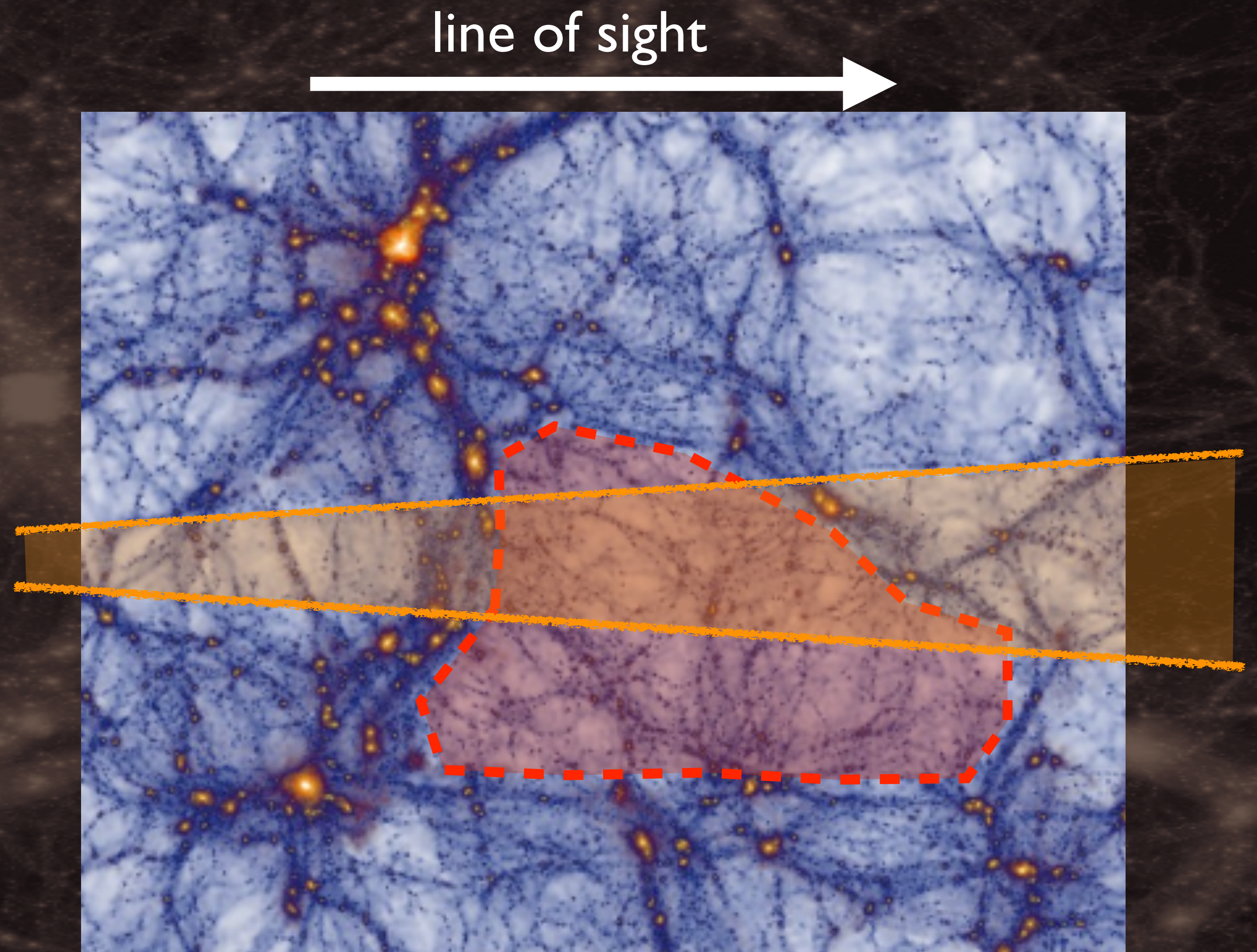
- The lensing signal depends on the projected line-of-sight density
- High density regions on the void edge and also along the line of sight partially compensate the emptiness of the void



Why is weak lensing by voids so weak?

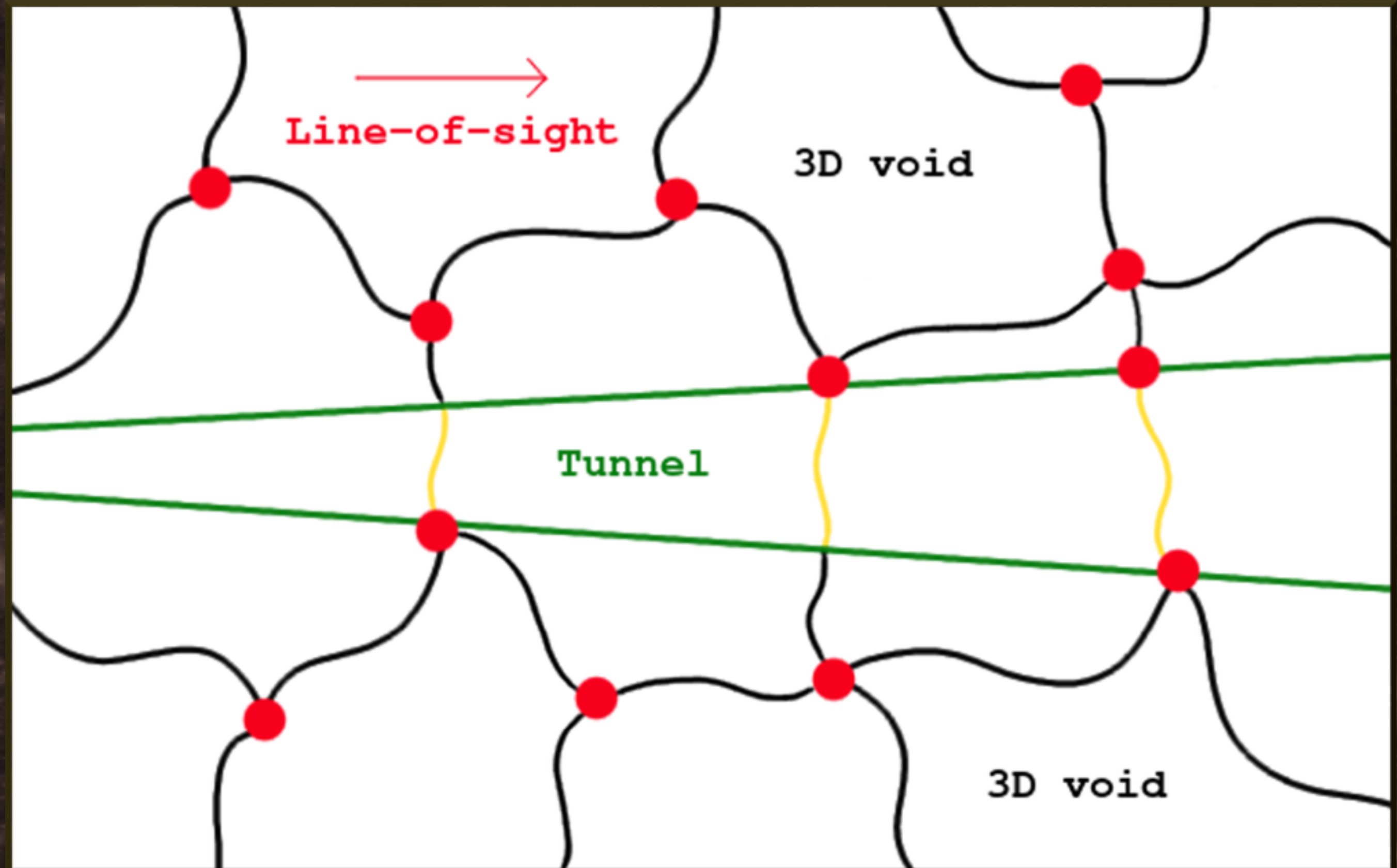
$$\gamma_t(r) = \frac{\Delta\Sigma(r)}{\Sigma_c} = \frac{\bar{\Sigma}(<r) - \Sigma(r)}{\Sigma_c}$$

- The lensing signal depends on the projected line-of-sight density
- High density regions on the void edge and also along the line of sight partially compensate the emptiness of the void
- Choose line of sights through voids that do not overlap with high density regions
—> line-of-sight under densities

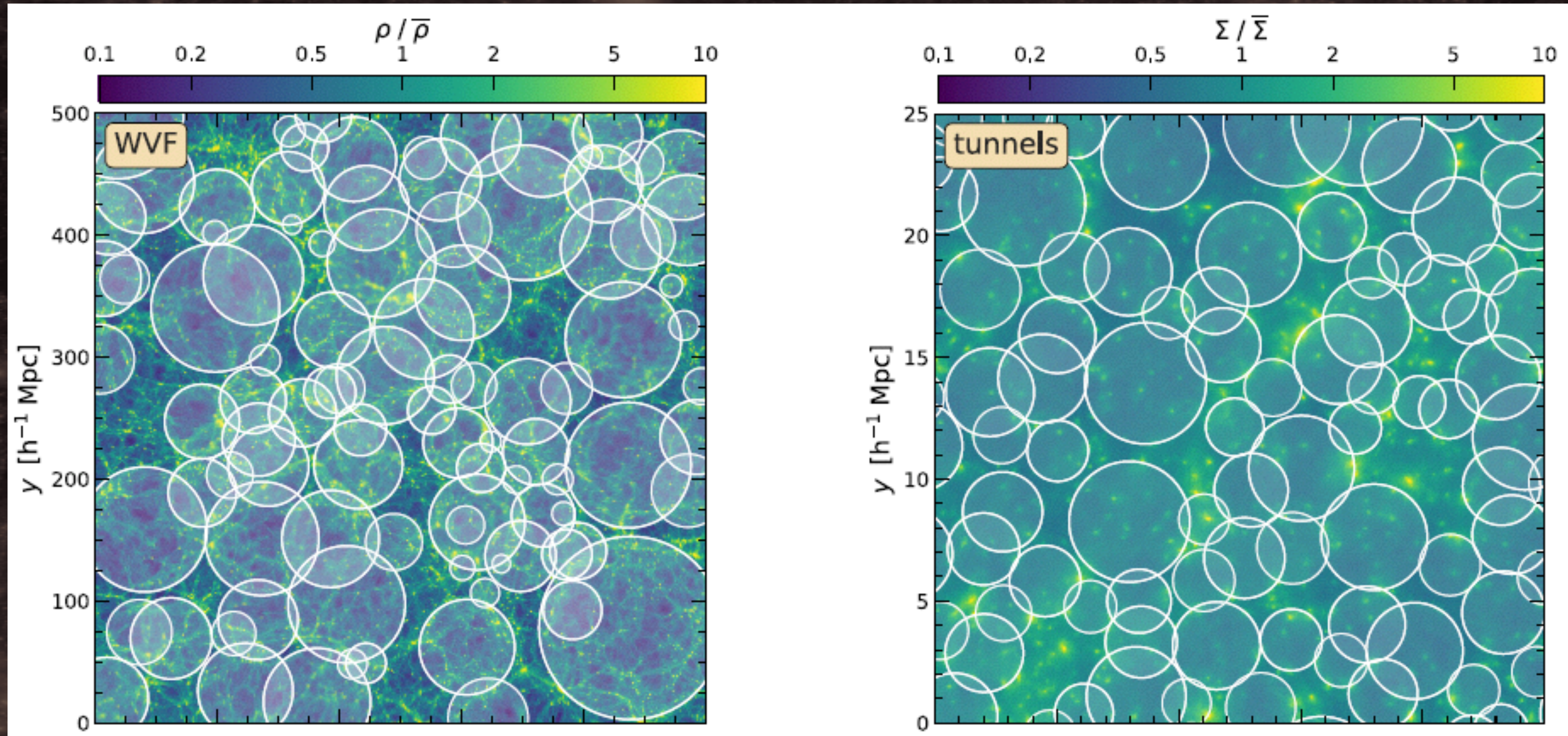


Voids in the projected galaxy distribution

Tunnels
(MC + 2018)

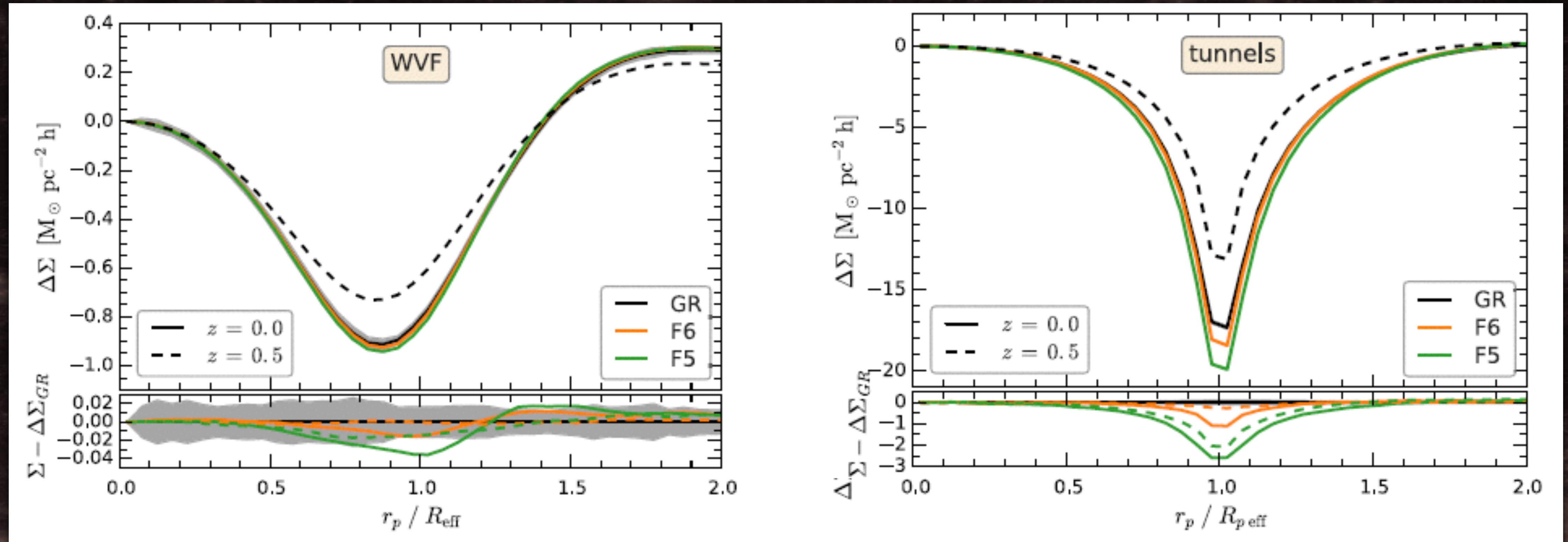


3D versus 2D void finders



MC+(2018)

Void lensing profiles



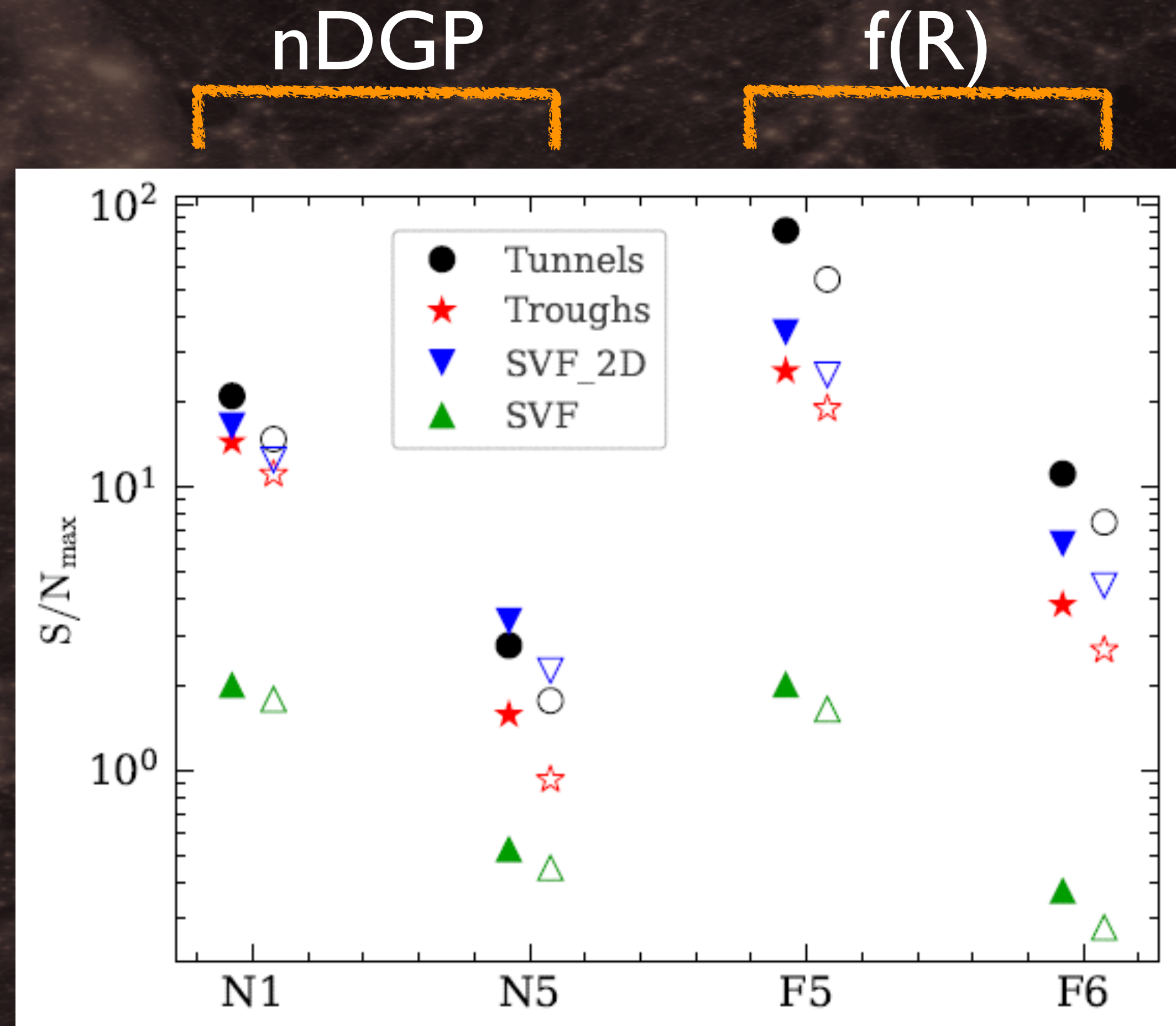
MC+(2018)

S/N for distinguishing models

Predictions for future surveys including shape noise:

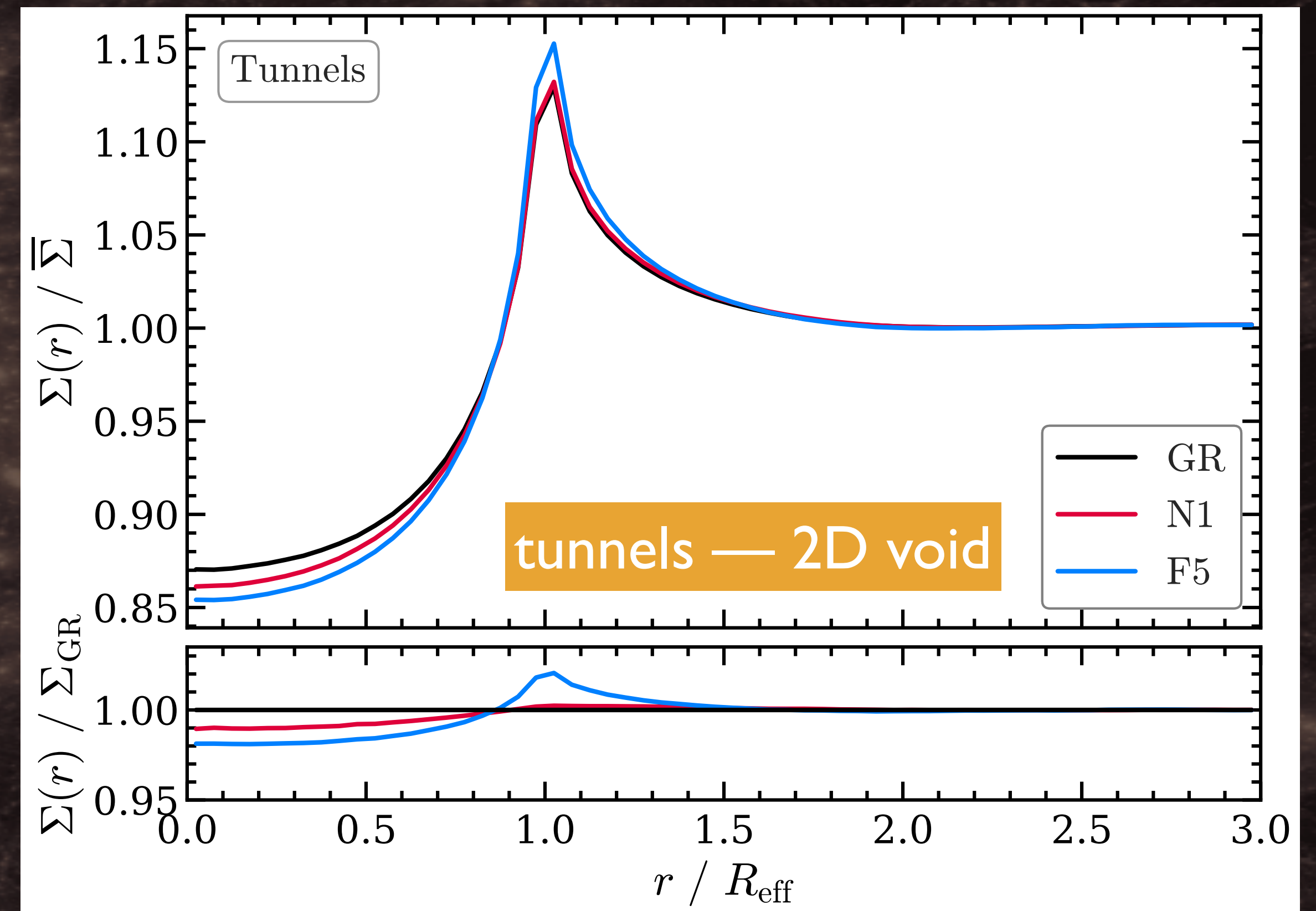
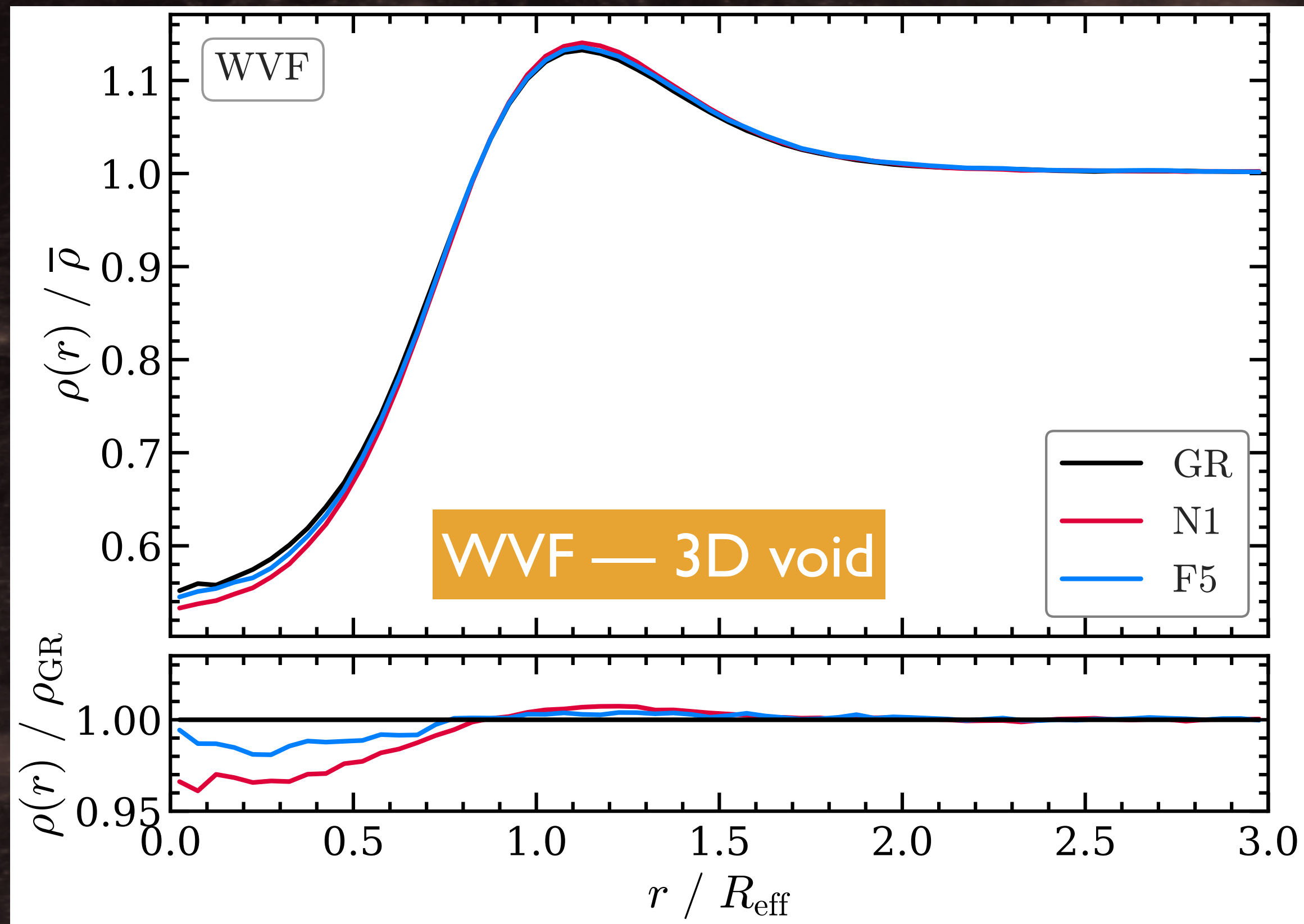
LSST (filled symbols)

Euclid (empty symbols)



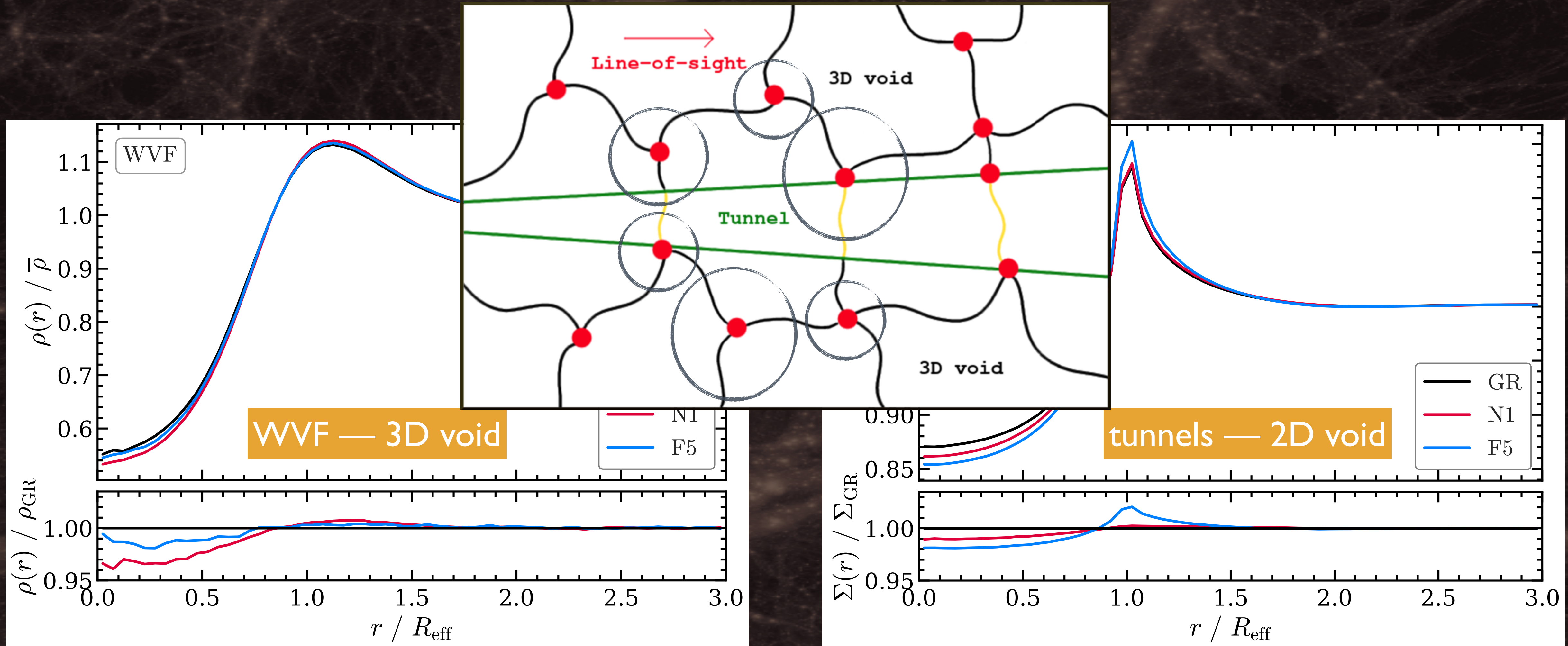
Paillas + (2018)

f(R) versus nDGP gravity



Paillas + (2018)

f(R) versus nDGP gravity

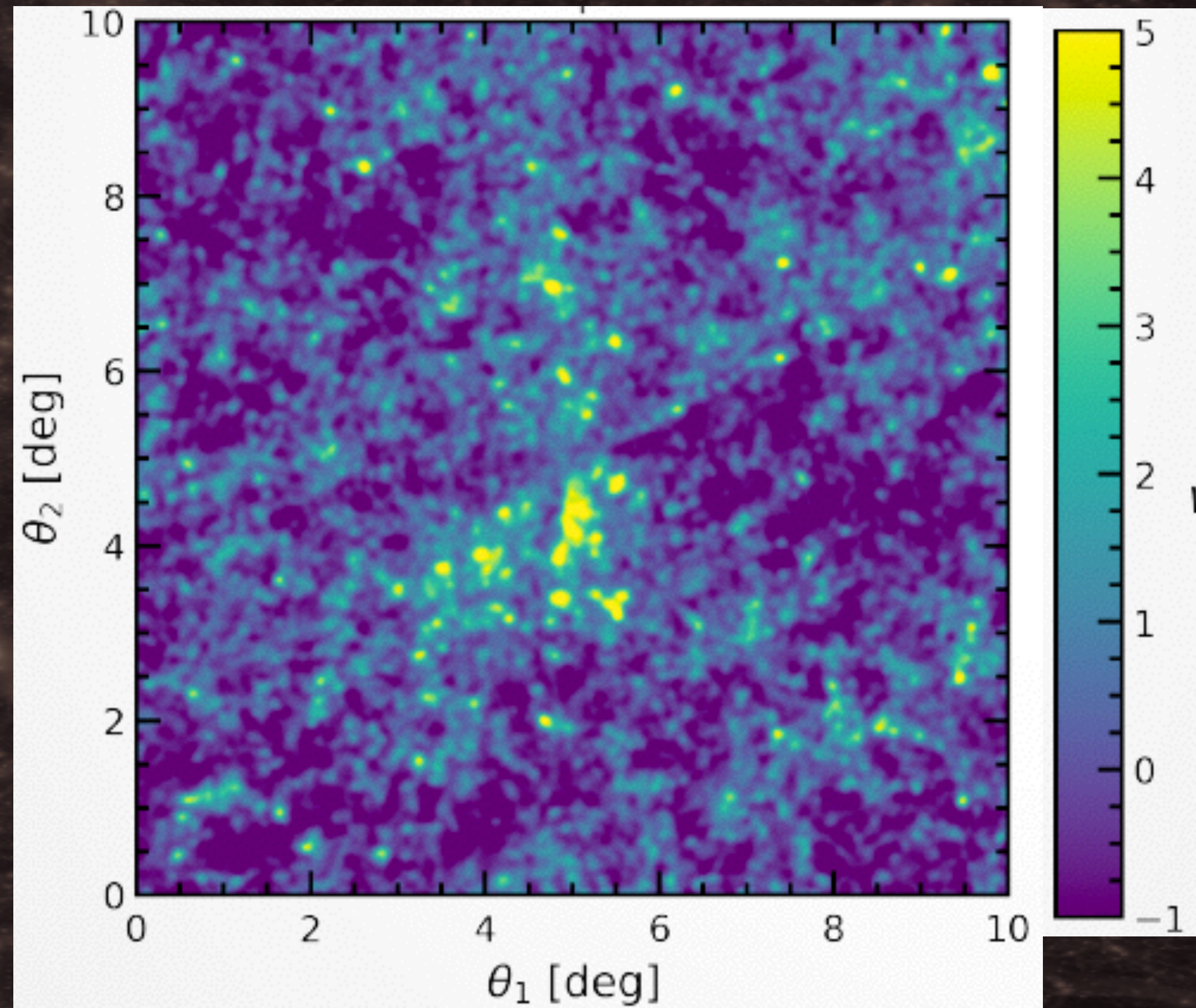


Paillas + (2018)

VOLES — voids from weak lensing

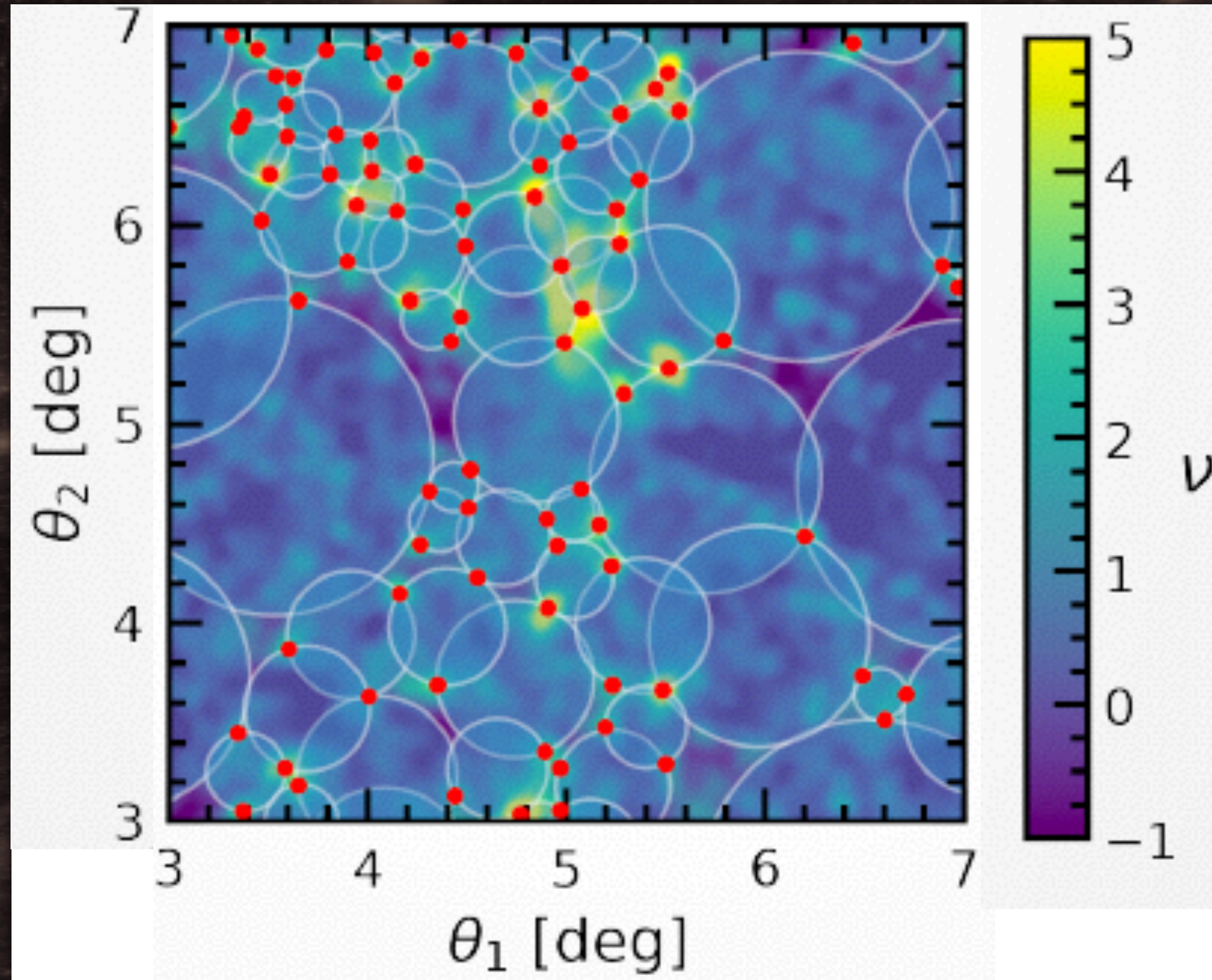
Convergence map

= (weighted) line-of-sight
projected density



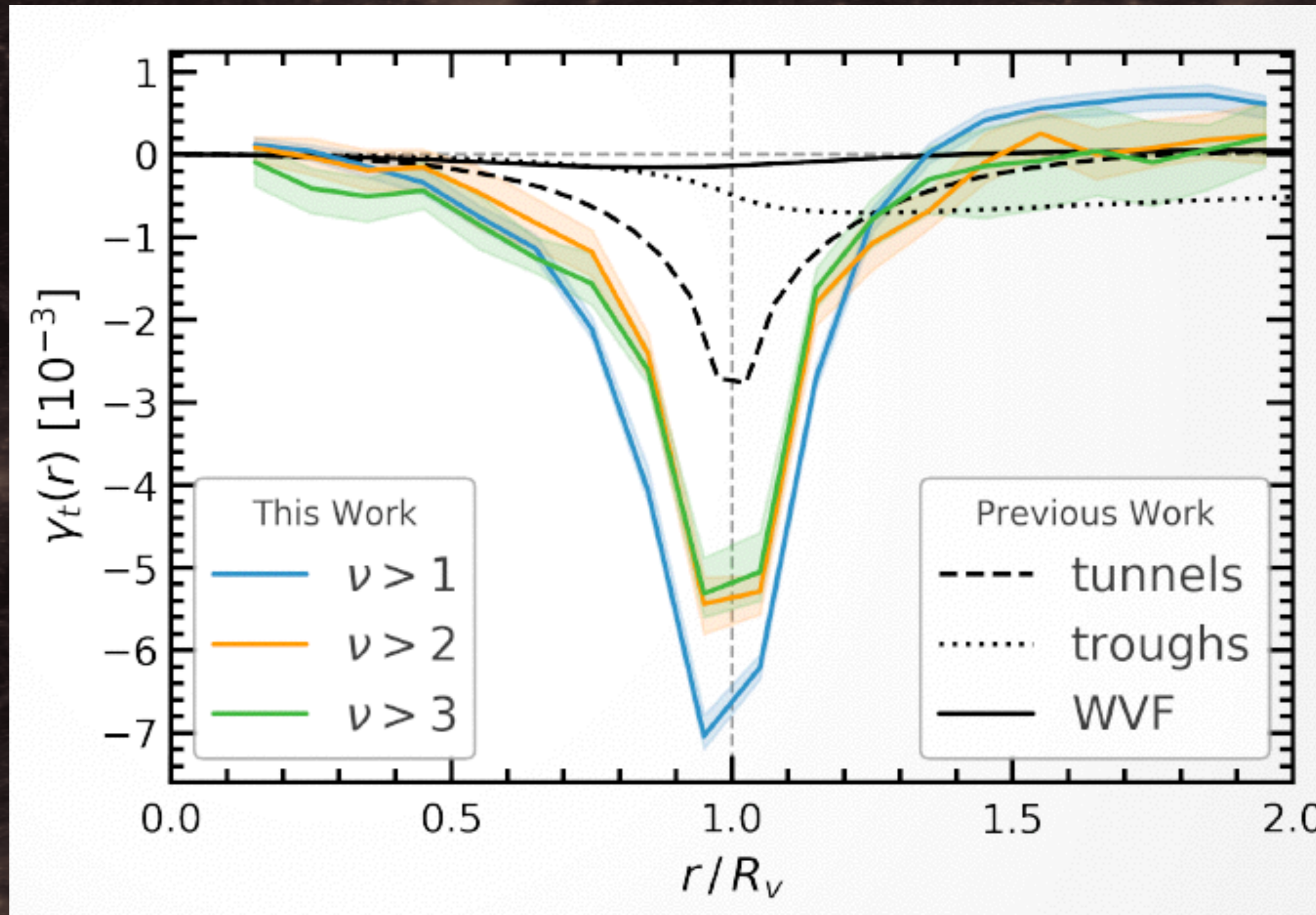
Davies + (2018)

VOLES — voids from weak lensing



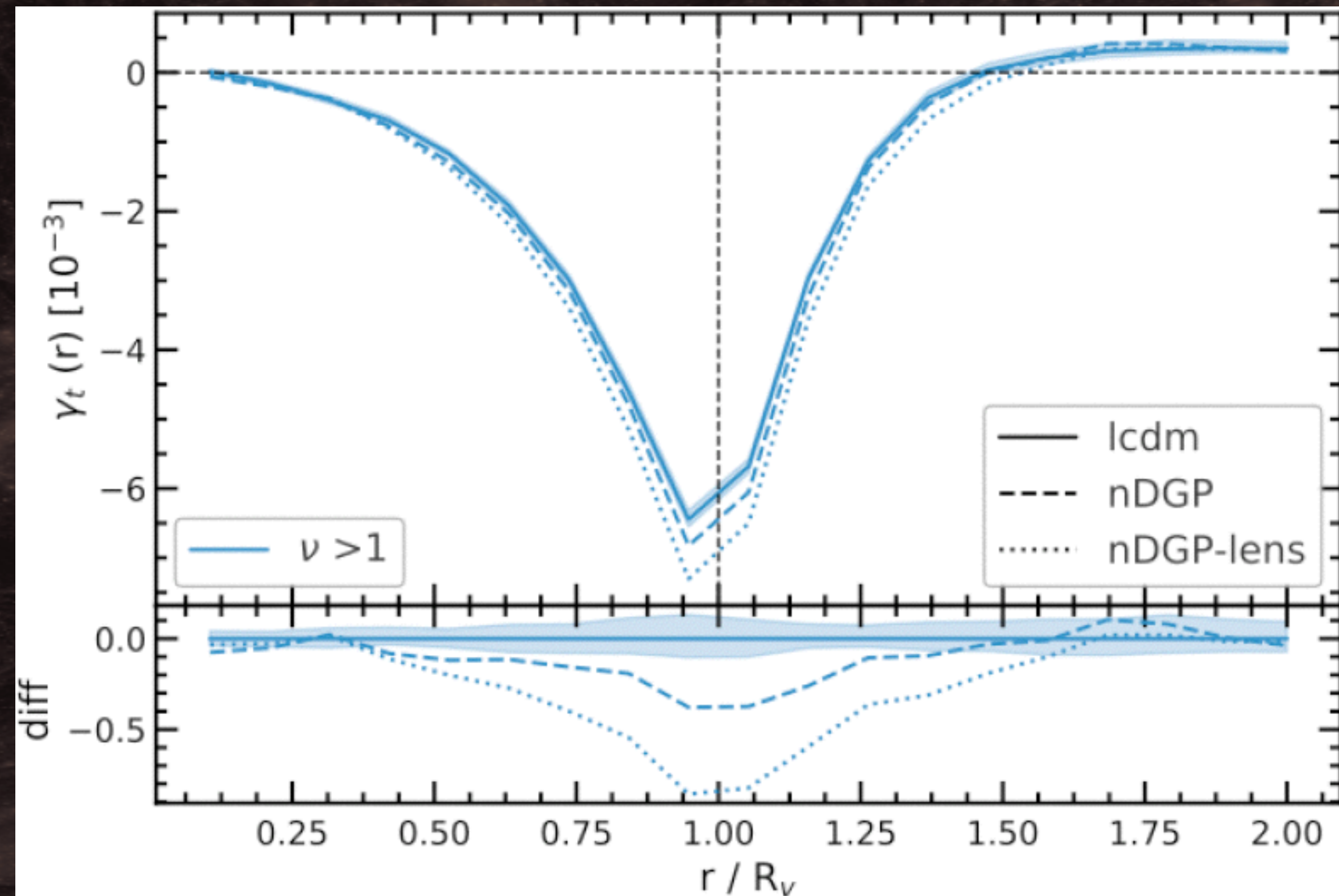
Davies + (2018)

Shear profile of VOLES



Davies + (2018)

Testing modified gravity theories using VOLES



Davies + in prep

Conclusions

- Voids are ideal probes of modified gravity theories since usually the fifth force in these models is maximal inside voids.
- The fifth force enhances the evacuation of matter from voids -> emptier voids.
- When matching the tracer distribution between GR and MG models, void statistics are the same except for the matter density profile (and RSD), and hence weak lensing by void.
- Voids identified in the weak lensing maps (VOLES) represent a new method of selecting the most underdense (line-of-sight) regions of the Universe and hence opens a new avenue for cosmological tests.

Thank you!

