Detecting Dark Matter in Celestial Bodies

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TeVPA 2022, Kingston
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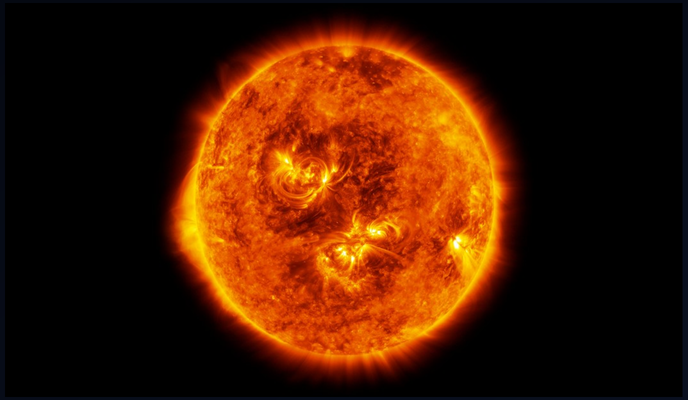
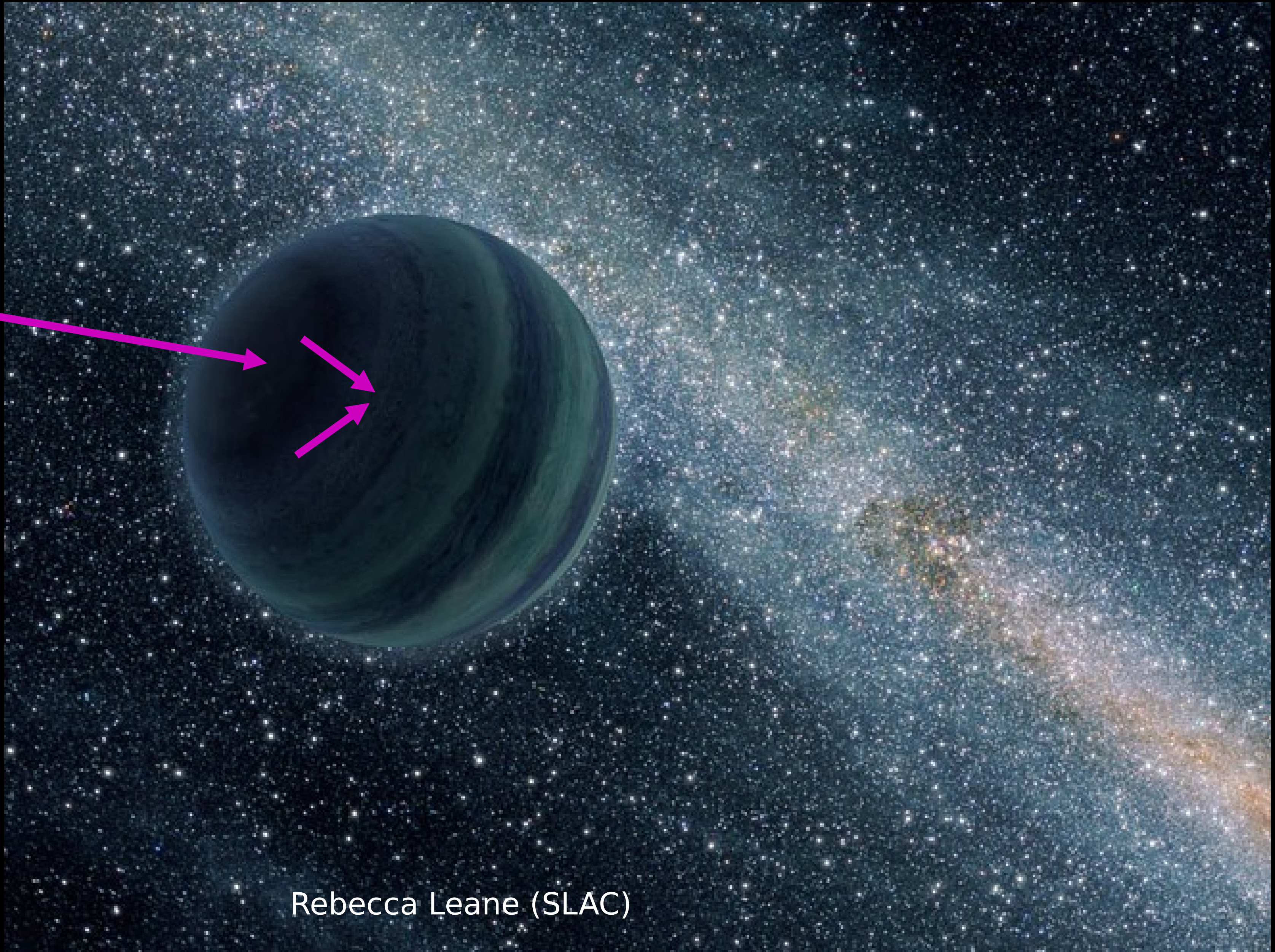
Outline

- DM capture in celestial objects
- Ideal properties of celestial objects
- Search locations
- Heating Searches
  - Telescopes, new technologies
  - Earth, White Dwarfs, Neutron Stars, Exoplanets
- Neutrino and Gamma-Ray Searches
  - Telescopes, new technologies
  - Sun, Jupiter, populations of celestial bodies
- Interesting things I don't have time to mention



DM capture in celestial bodies



Steigman, Sarazin, Quintana, Faulkner 1978
Press, Spergel 1985
Gould 1987
Griest, Seckel 1987

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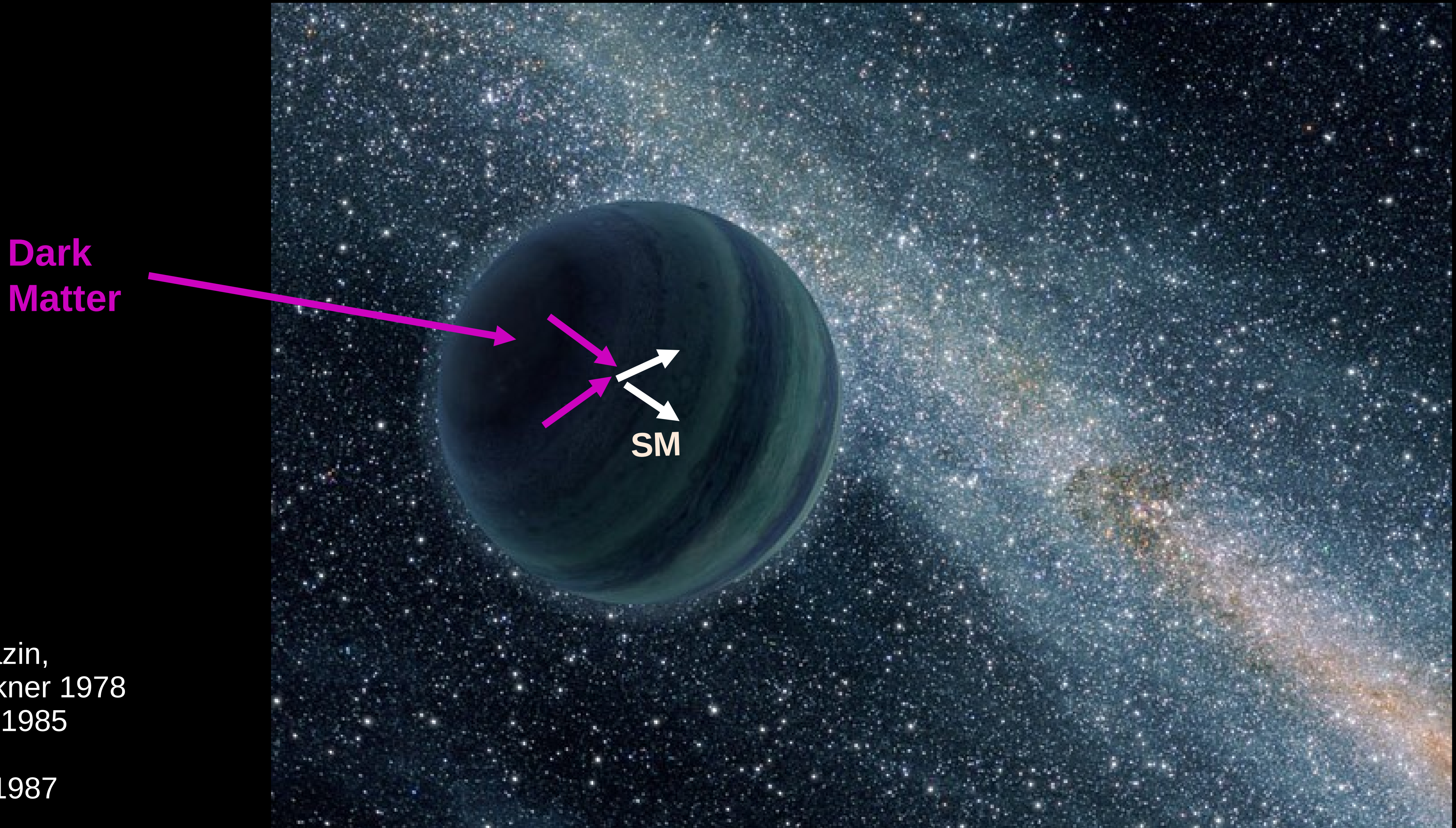
DM capture in celestial bodies



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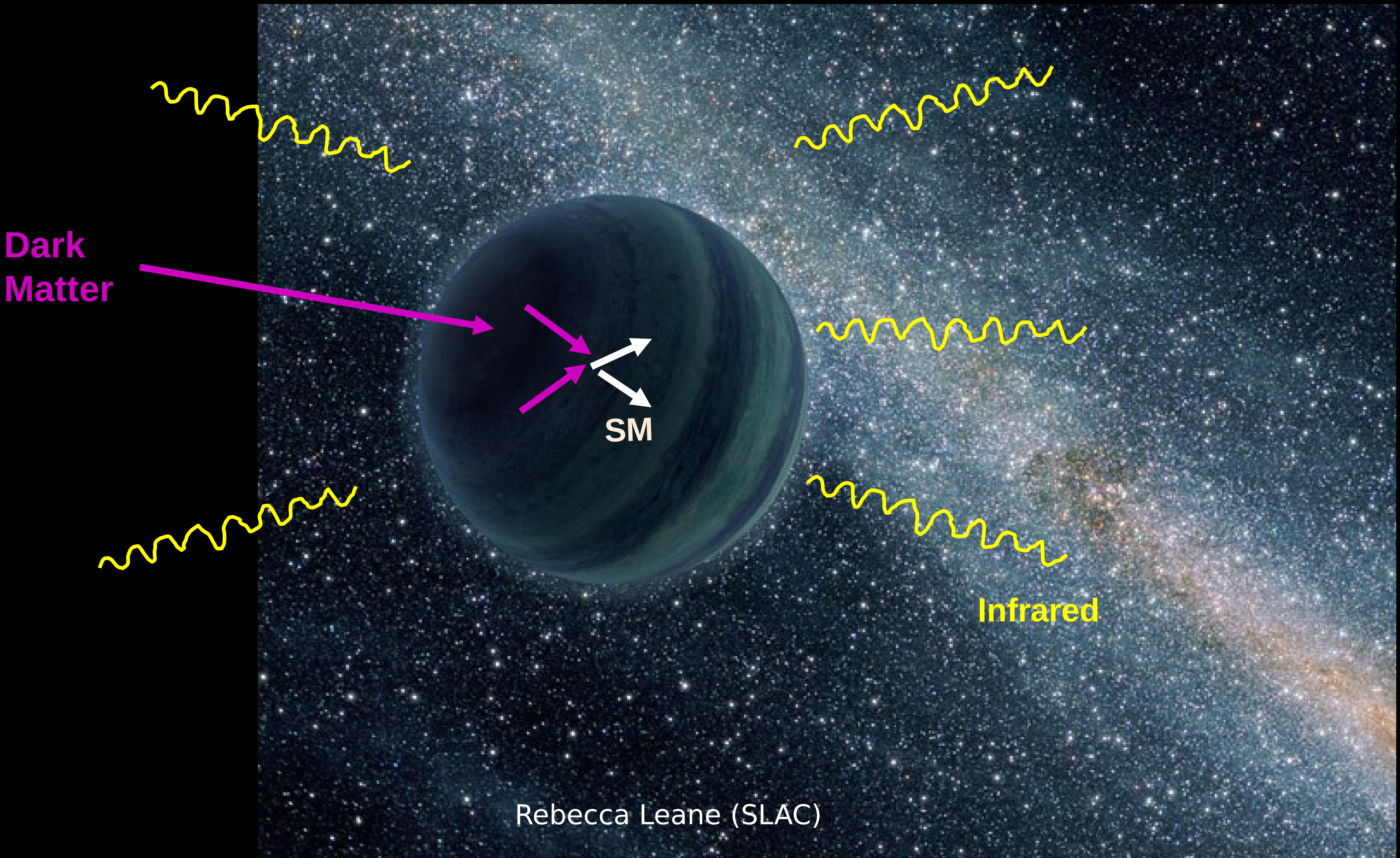
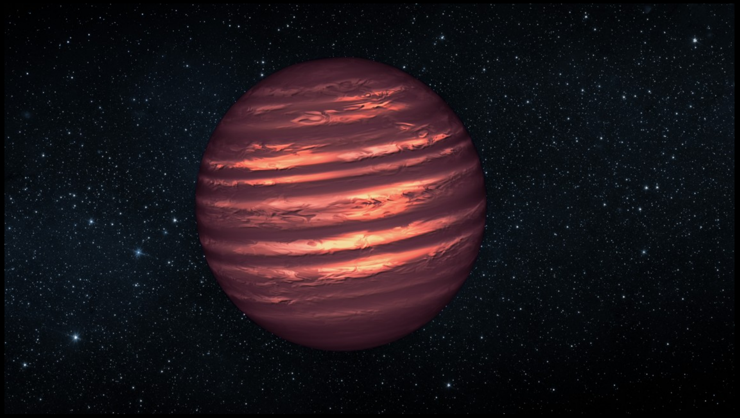
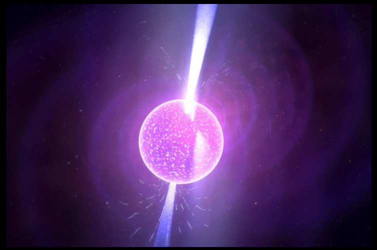
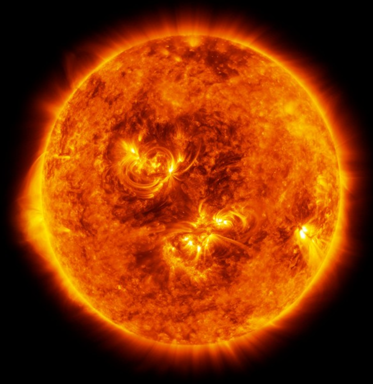
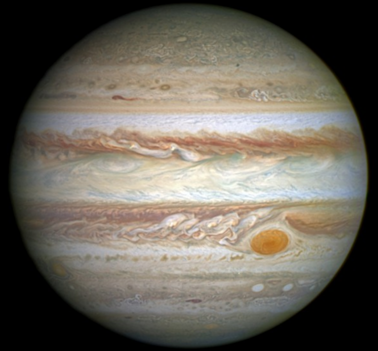
DM capture in celestial bodies



Optimal Celestial Target?

- Radius: Larger amount of DM captured, larger annihilation signal
- Density: Optical depth → lower cross section sensitivities
- Core temperature: Gives kinetic energy to DM, if high, more evaporation

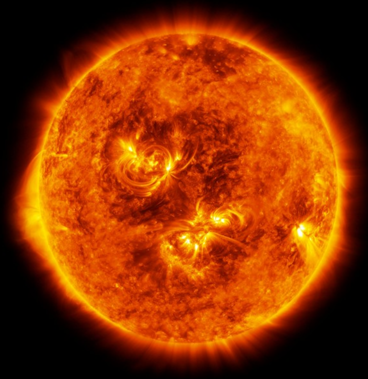





Optimal Celestial Target + Location

Signal detectability matters!

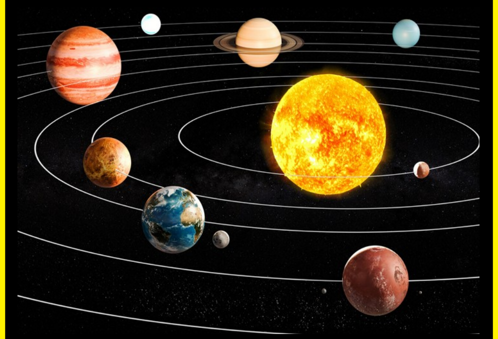
- Telescope sensitivity to a given flux size?
  - If larger amount of DM captured, larger annihilation signal
  - Further away → $1/R^2$ suppression
  - Larger objects easier to detect further away

- Background expectation?



Search Locations

Local Position



Age: ~5 Gyr
Distance: ~100 pc

DM density/velocity:
~0.4 GeV/cm^3
~230 km/s

Globular Clusters

Age: ~12 Gyr
Distance: 2 kpc

DM density/velocity*:
~100 GeV/cm^3, 2 pc
~10 km/s

Galactic Center

Age: ~8 Gyr (varies)
Distance: 8 kpc

DM density/velocity*:
~100 GeV/cm^3, 0.1kpc
~30-100 km/s

Best features:

- ✓ High DM density
- ✓ Low DM velocity
- ✓ Close proximity
- ✓ Old environment
- ✓ Low dust

Recap so far

- Lots of celestial objects: unique temperatures, radii, and densities
  - Different objects optimal for different cross sections or DM masses
- Variety of search locations
  - Beneficial environment features: DM density, velocity, proximity, age
- Variety of DM signatures in celestial objects
  - Now will consider DM heating, for many objects and locations!

Dark Matter Annihilation:
H e a t i n g

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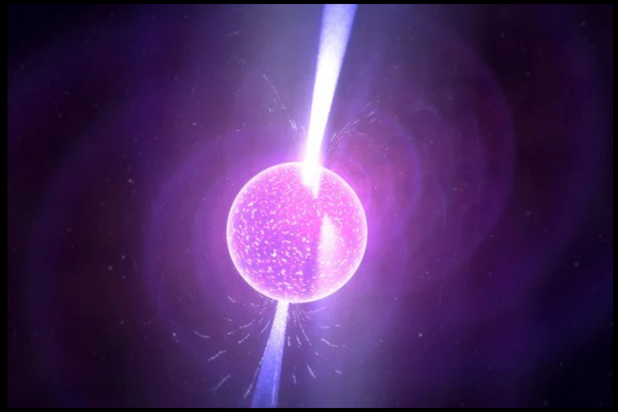
Good heating candidates

Earth

Mars

Jupiter

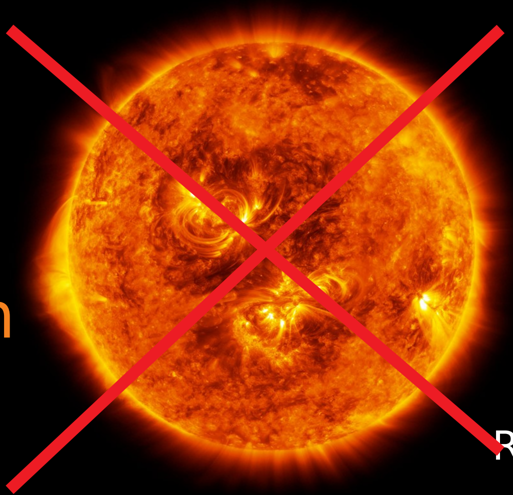
Neutron Stars



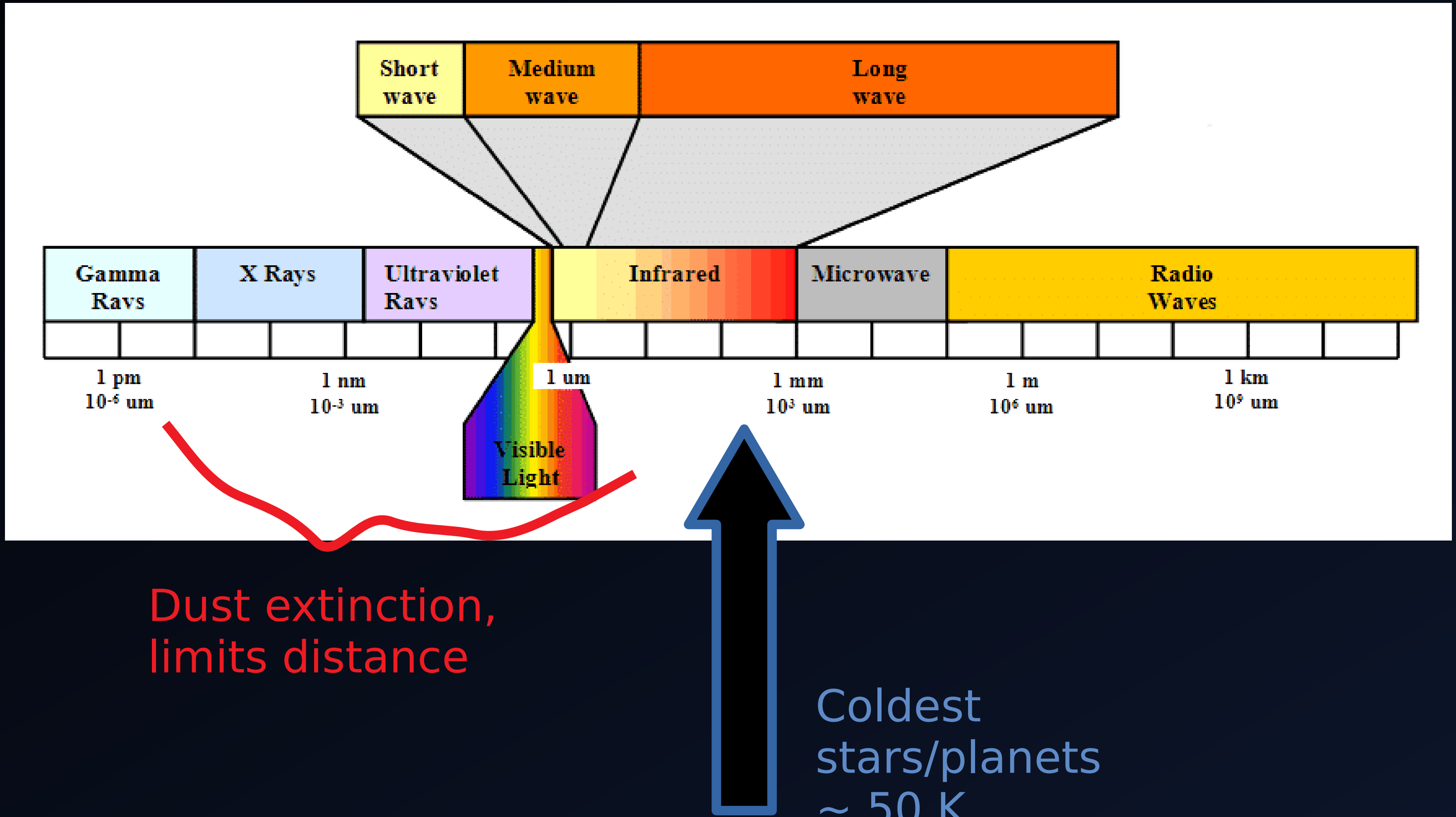
Exoplanets

White Dwarfs

NOT GOOD:
Sun, other main sequence stars



Detecting Dark Matter Heating



Dust extinction,
limits distance

Coldest
stars/planets
~ 50 K

Rebecca Leane (SLAC)

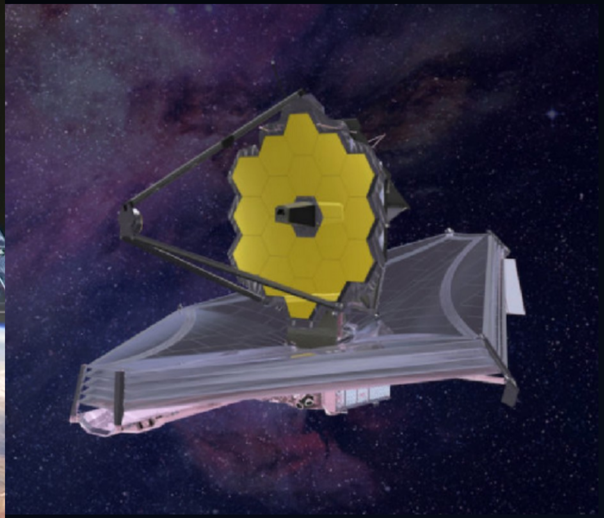
Detecting Dark Matter Heating

Hubble

Near-infrared
Optical
Ultraviolet

~0.12-2 microns

Data obtained
~31 years elapsed

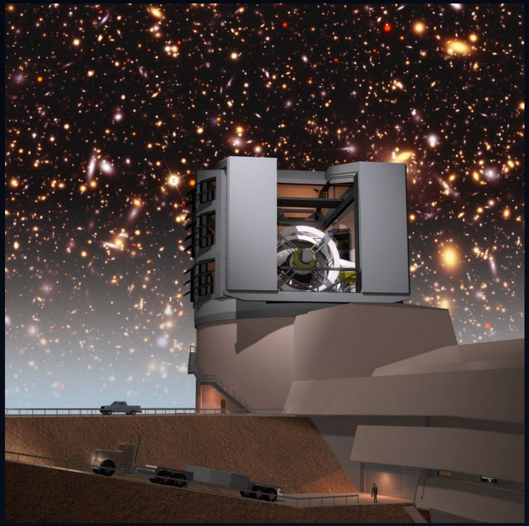


Webb

Full Infrared
Optical

~0.5 – 28 microns

Awaiting Data
Launch 2021

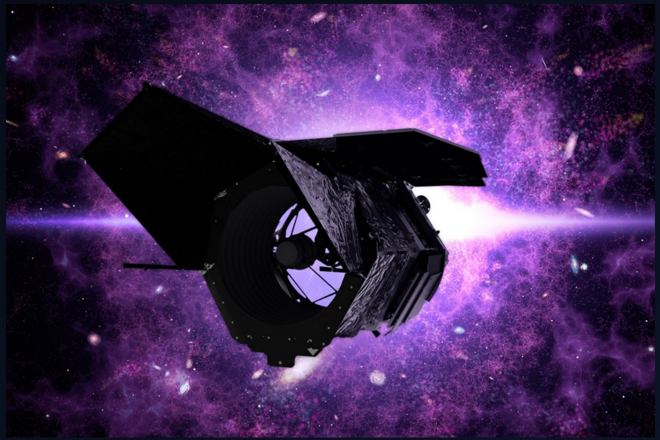


Rubin

Near-infrared
Optical

~0.32–1.06 microns

Awaiting Data
First light 2022/23



Roman

Near-infrared
Optical

~0.5 – 2 microns

Awaiting Data
Launch 2025

EARTH

Freese 1985
Krauss, Srednicki, Wilczek 1986
Gaisser, Steigman, Tilav 1986
Gould 1987, 1988, 1991, 1992
Gould, Frieman, Freese 1989
Gould, Alam 2001
Starkman, Gould, Esmailzadeh, Dimopoulos 1990
Mack, Beacom, Bertone 2007
Bramante, Buchanan, Goodman, Lodhi 2019
Acevedo, Bramante, Goodman,
Kopp, Opferkuch 2020

+ more

Category: Rocky planet
Core temp: ~10^3 K
Escape Velocity: ~11 km/s

EARTH

Available data: **20,000 bore holes drilled throughout crust**

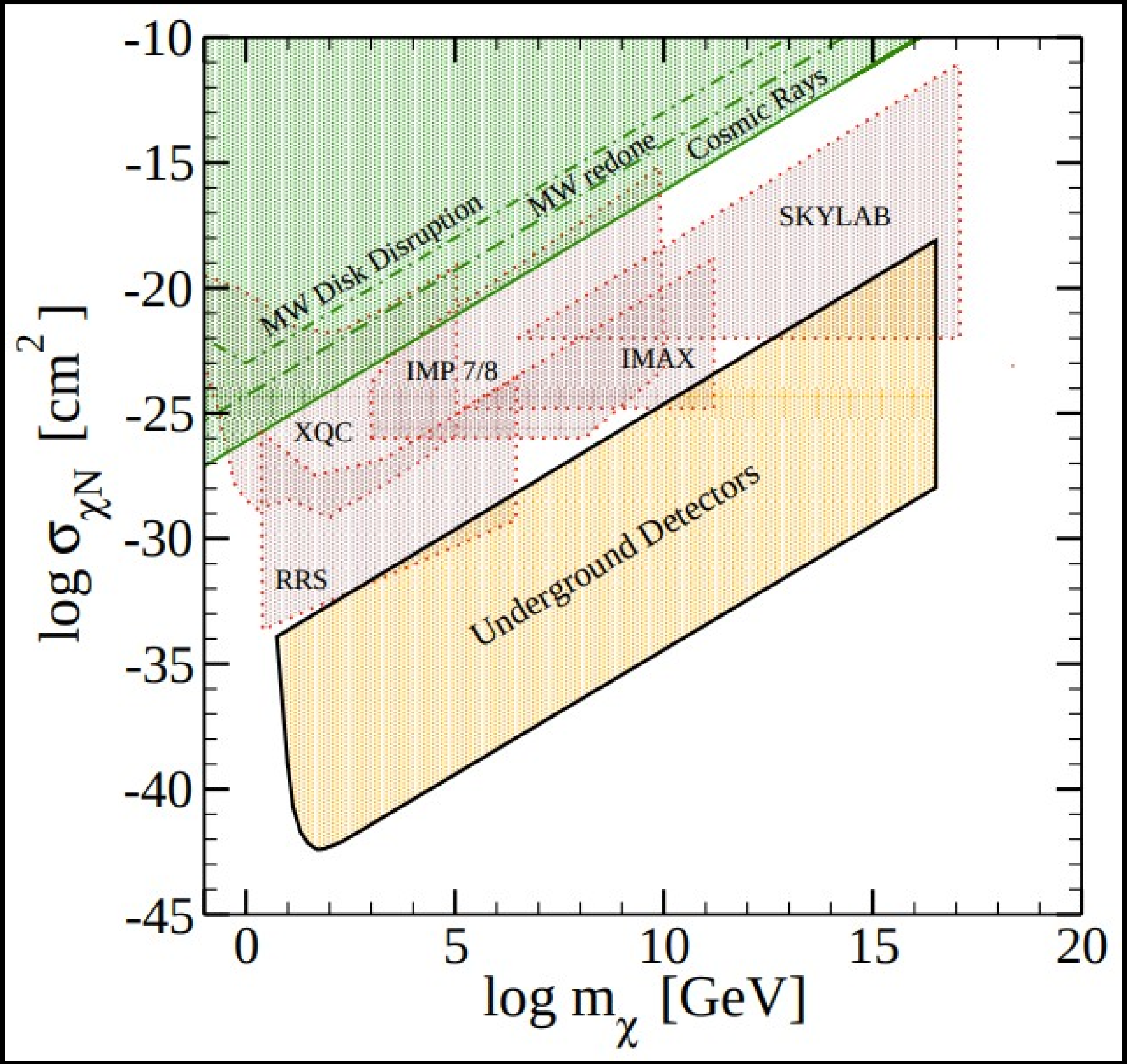
+ Geologists extensively studied Earth’s internal heat

+ Temperature gradient in borehole is recorded, multiplied by the thermal conductivity of the relevant material yields a heat flux

Benefits:
+ Systematics low
+ Data now
+ Best proximity

Limitation: Higher DM evaporation mass, cross sec reach

EARTH



Mack, Beacom, Bertone 0705.4298

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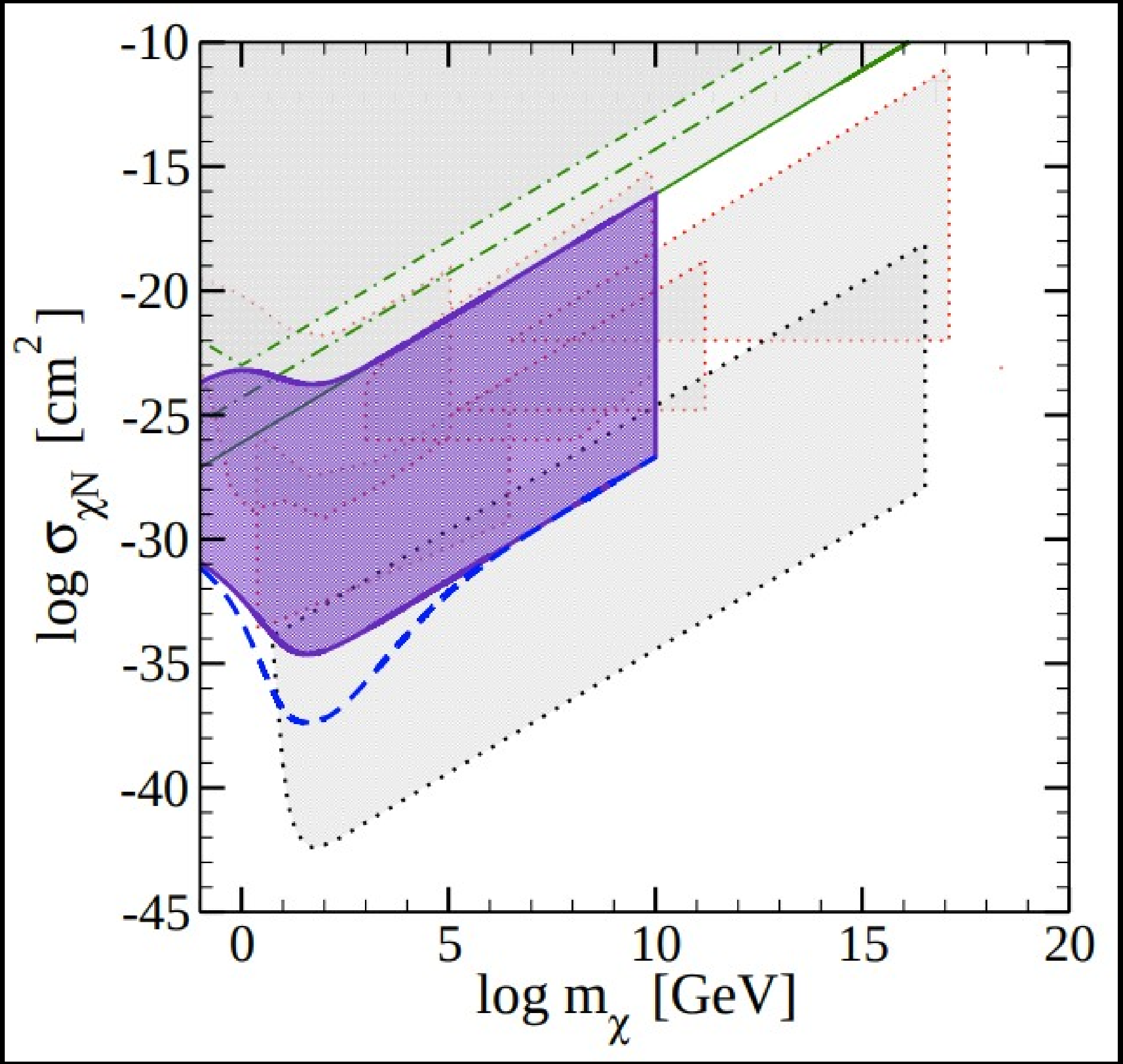
EARTH



Mack, Beacom, Bertone 0705.4298

See also Bramante, Buchanan, Goodman, Lodhi 1909.11683 (incl Mars)

WHITE DWARFS

Moskalenko, Wai, 2006, 2007
Bertone, Fairbairn 2007
McCullough, Fairbairn 2010
Hooper, Spolyar, Vallinotto, Gnedin 2010
Amaro-Seoane, Casanellas, Schoedel, Davidson, Cuadra 2015
Bramante 2015
Graham, Rajendran, Varela 2015
Graham, Janish, Narayan, Rajendran, Riggins, 2018
Dasgupta, Gupta, Ray 2019
Acevedo, Bramante 2019
Horowitz 2020
Panotopoulos, Lopes 2020
Curtin, Setford, 2020
Bell, Busoni, Ramirez-Quezada, Robles, Virgato 2021

Composition: Mostly Carbon + Oxygen
Mass: ~1 Solar mass
Radius: ~1 Earth radius
Escape velocity: ~10^3 km/s

Origin: Collapse of main sequence stars w/ mass less ~8-10 solar mass, supported against grav collapse by electron degeneracy pressure

WHITE DWARFS

Available data: **Hubble measurements of Messier 4 globular cluster**

Limitations:

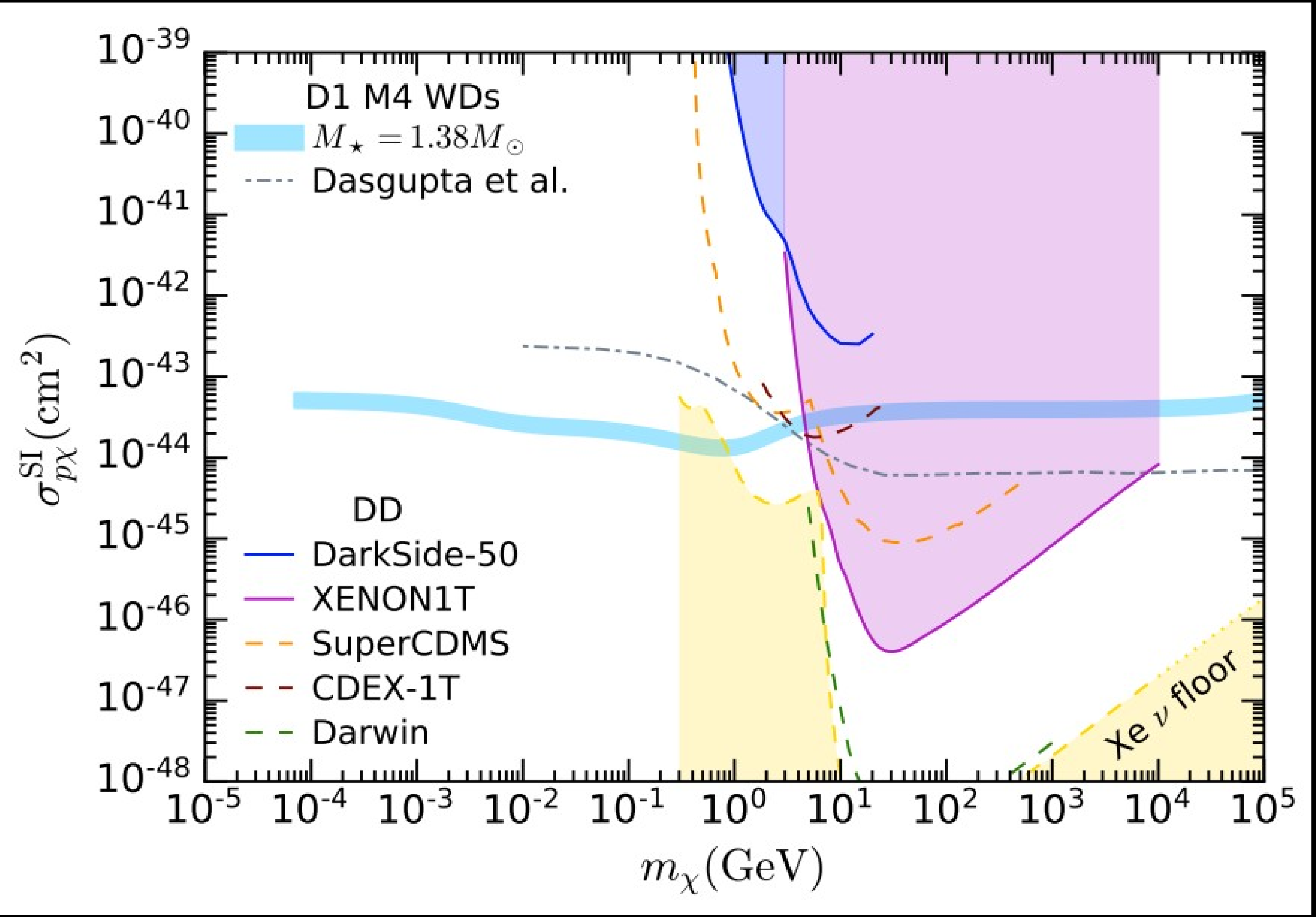
+ High surface temperature, want high DM density locations
+ DM density NOT known for M4
+ Candidates needed for Galactic Center

Benefits:

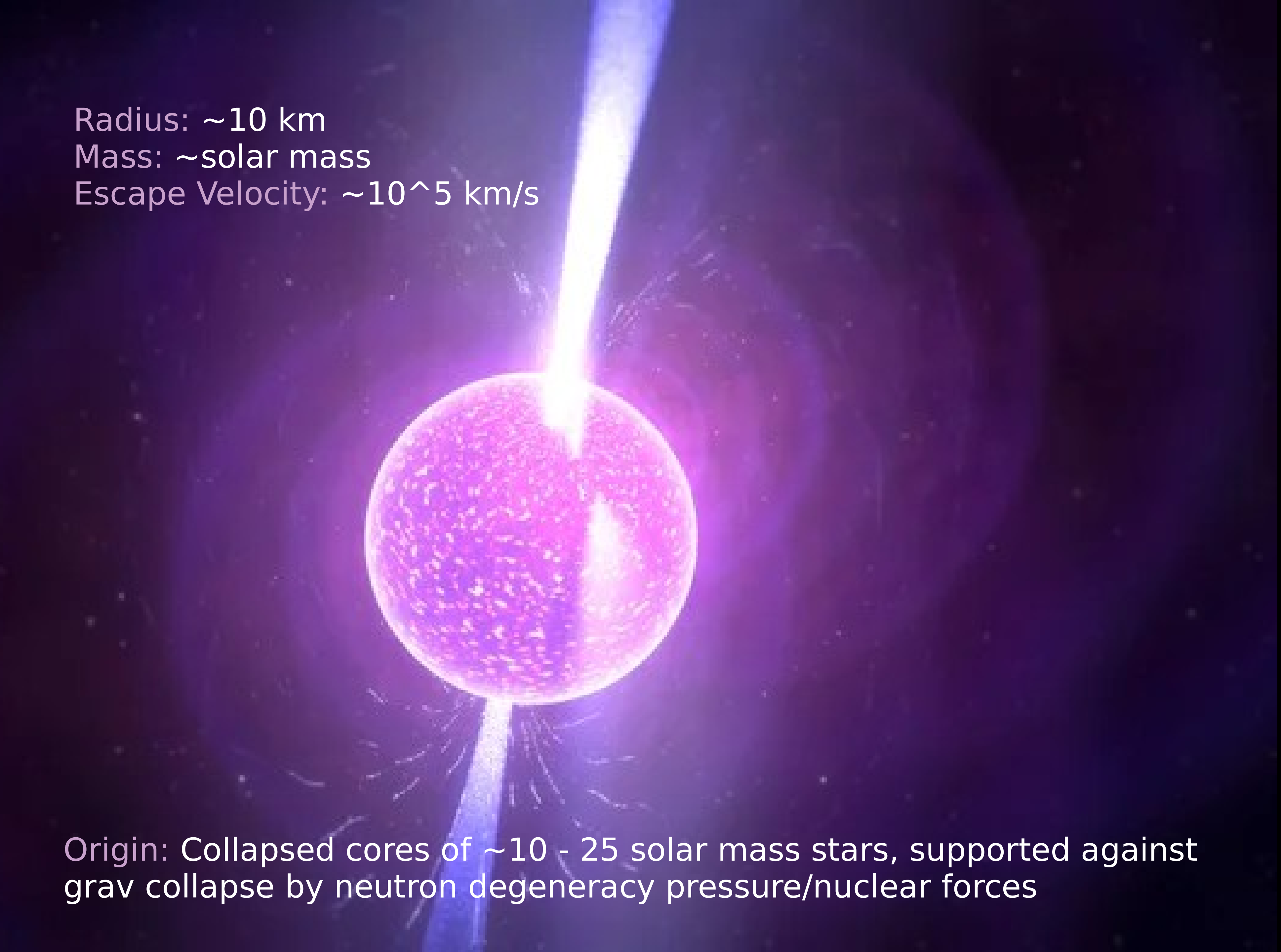
+ Do exist in globular cluster cores
+ M4 data now!
+ Low evaporation masses
+ Better cross section sensitivity than Earth

WHITE DWARFS

Bell, Busoni, Ramirez-Quezada, Robles, Virgato 2021



NEUTRON STARS



Radius: ~10 km
Mass: ~solar mass
Escape Velocity: ~10^5 km/s

Origin: Collapsed cores of ~10 - 25 solar mass stars, supported against grav collapse by neutron degeneracy pressure/nuclear forces

Gould, Draine, Romani, Nussinov 1989
Goldman, Nussinov 1989
Starkman, Gould, Esmailzadeh, Dimopoulos 1990
Bertone, Fairbairn 2007
Kouvaris 2007
Gonzalez, Reisenegger 2010
Kouvaris, Tinyakov 2011
McDermott, Yu, Zurek 2011
Bramante, Fukushima, Kumar 2013
Bell, Melatos, Petraki 2013
Bramante, Linden 2014
Bertoni, Nelson, Reddy 2014
Bramante, Elahi 2015
Baryakhtar, Bramante, Li, Linden, Raj 2017
Bramante, Delgado, Martin 2017
Raj, Tanedo, Yu 2017
Chen, Lin 2018
Jin, Gao 2018
Garani, Genolini, Hambye 2018
Acevedo, Bramante, Leane, Raj 2019
Hamaguchi, Nagata, Yanagi 2019
Camargo, Queiroz, Sturani 2019
Joglekar, Raj, Tanedo, Yu 2019
Garani, Heeck 2019
Bell, Busoni, Robles 2019
Keung, Marfatia, Tseng 2020
Bell, Busoni, Robles 2020
Bai, Berger, Korwar, Orlofsky 2020
Bell, Busoni, Motta, Robles, Thomas, Virgato 2020
Leane, Linden, Mukhopadhyay, Toro 2021

+ even more

NEUTRON STARS

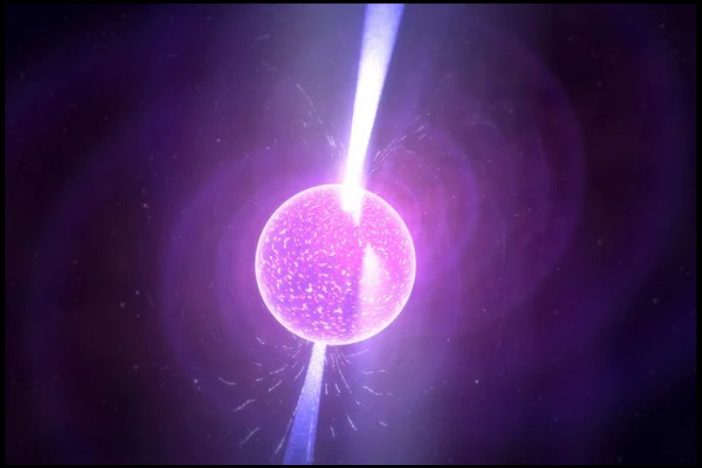
Available data:
None yet, potentially use upcoming infrared telescopes

Limitations:

+ NS are small, so need to use target close by
+ No yet known candidates
+ Exposure times required can be large

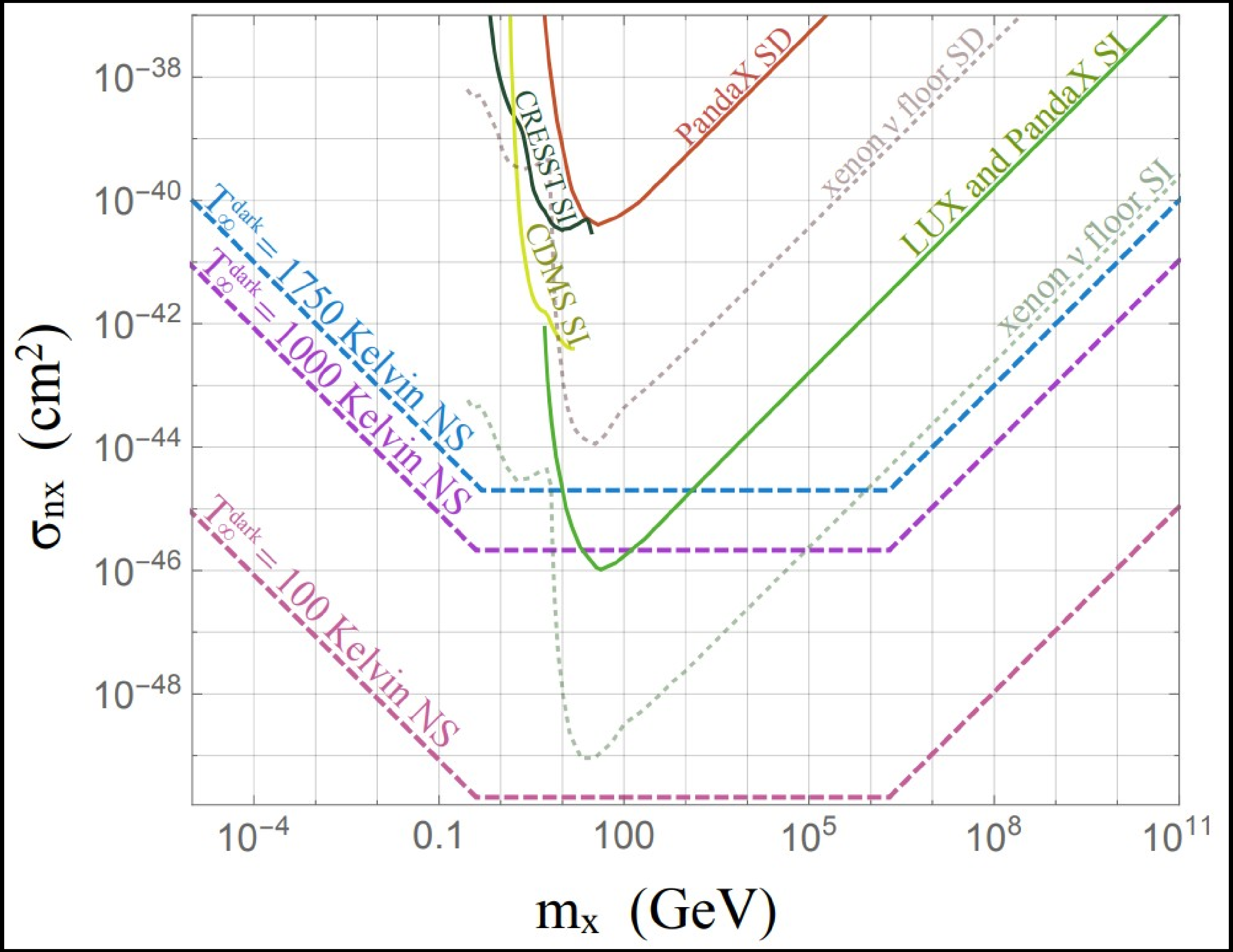
Benefits:

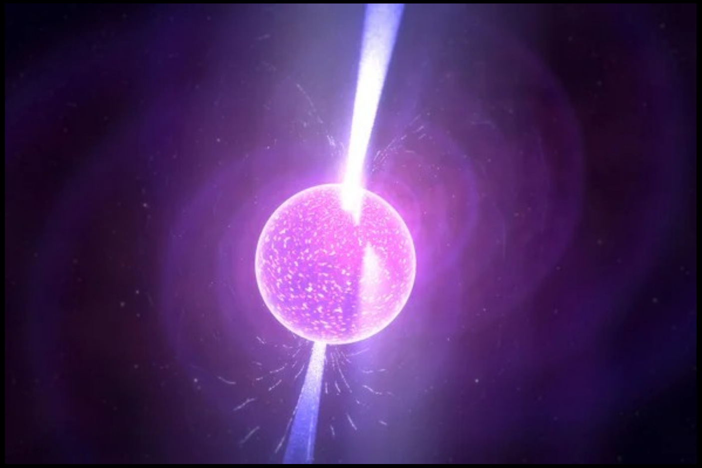
+ Superior cross section sensitivity
+ Kinetic heating boost in rate
+ Broad class of particle models



Baryakhtar, Bramante, Li, Linden, Raj 2017

NEUTRON STARS





See also Bell, Busoni, Motta, Robles, Thomas, Virgato 2020

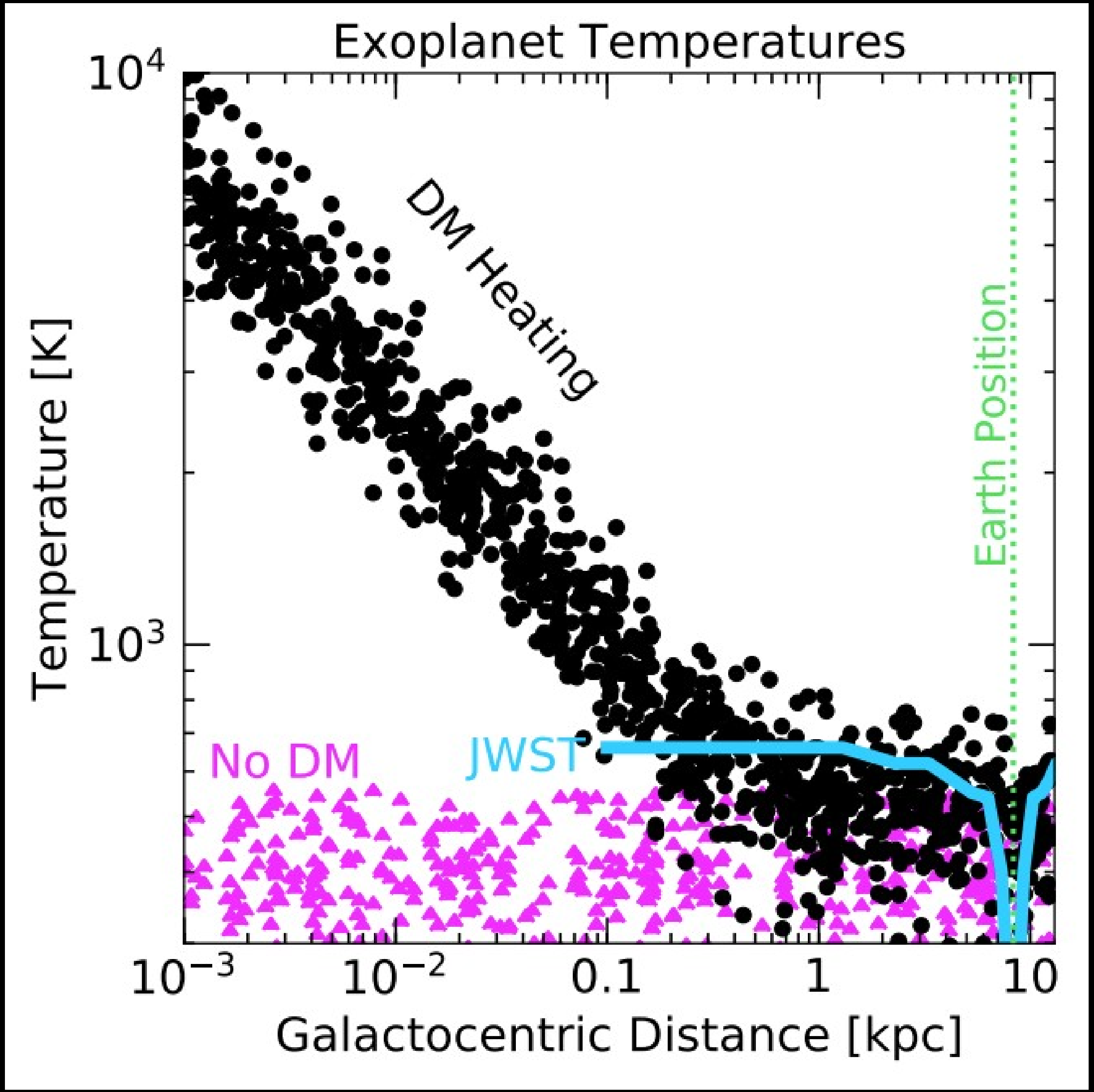
EXOPLANETS

Adler 2009

Hooper, Steffen 2011

Leane, Smirnov 2020

EXOPLANETS



Leane + Smirnov, 2020

Exoplanets can potentially be used to map the Galactic DM density

Rebecca Leane (SLAC)

EXOPLANETS

Available data:
Little yet, use upcoming infrared telescopes

Limitations:

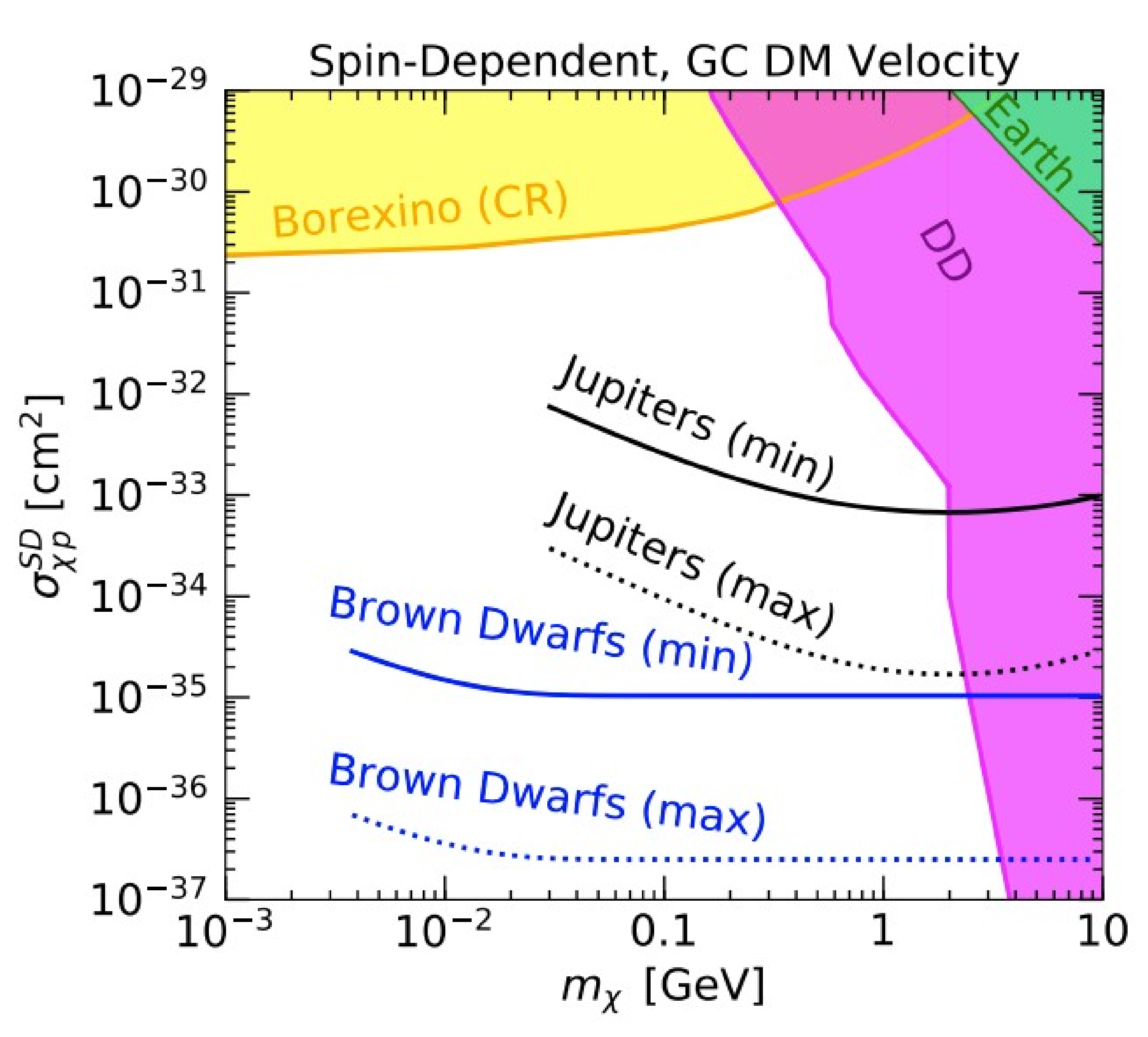
+ Having enough acceptable candidates
+ Not robustly known interiors
+ Cooling systematics

Benefits:

+ Large statistics; some candidates already exist
+ Cold (good signal over background)
+ Large radii, easier to detect than NS
+ Low evaporation masses
+ Potential probe of DM density profile

Leane + Smirnov, 2020

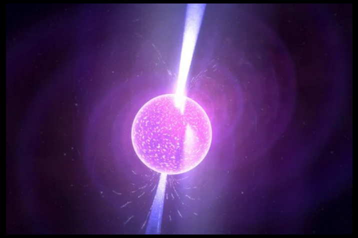
Exoplanet cross section sensitivity



Leane + Smirnov, 2020

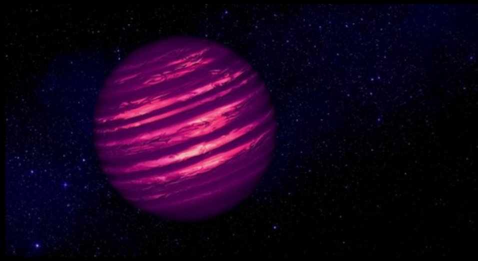
Actions for successful discovery/exclusion

- Neutron stars:
  - Find a candidate close by and old enough! (FAST radio search)
  - Enough observing time granted



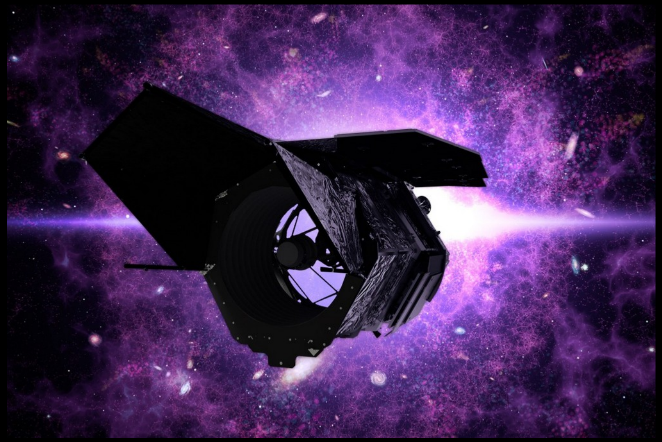
- White dwarfs:
  - Understand astrophysical uncertainties in clusters
  - More candidates

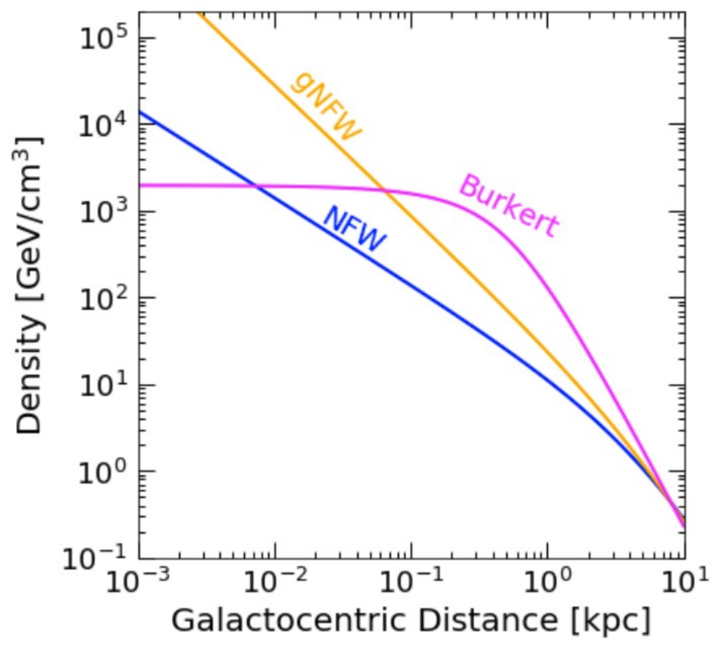
- Exoplanets:
  - Large statistical sample obtained to overcome systematics
  - Detailed studies of atmosphere effects including DM



Sensitivity to DM halo parameters

- Direct probe of unknown DM density profile
- How many exoplanets do we need to detect?
- What level of precision do we need to measure exoplanet:
  - Radii?
  - Temperatures?
  - Masses?

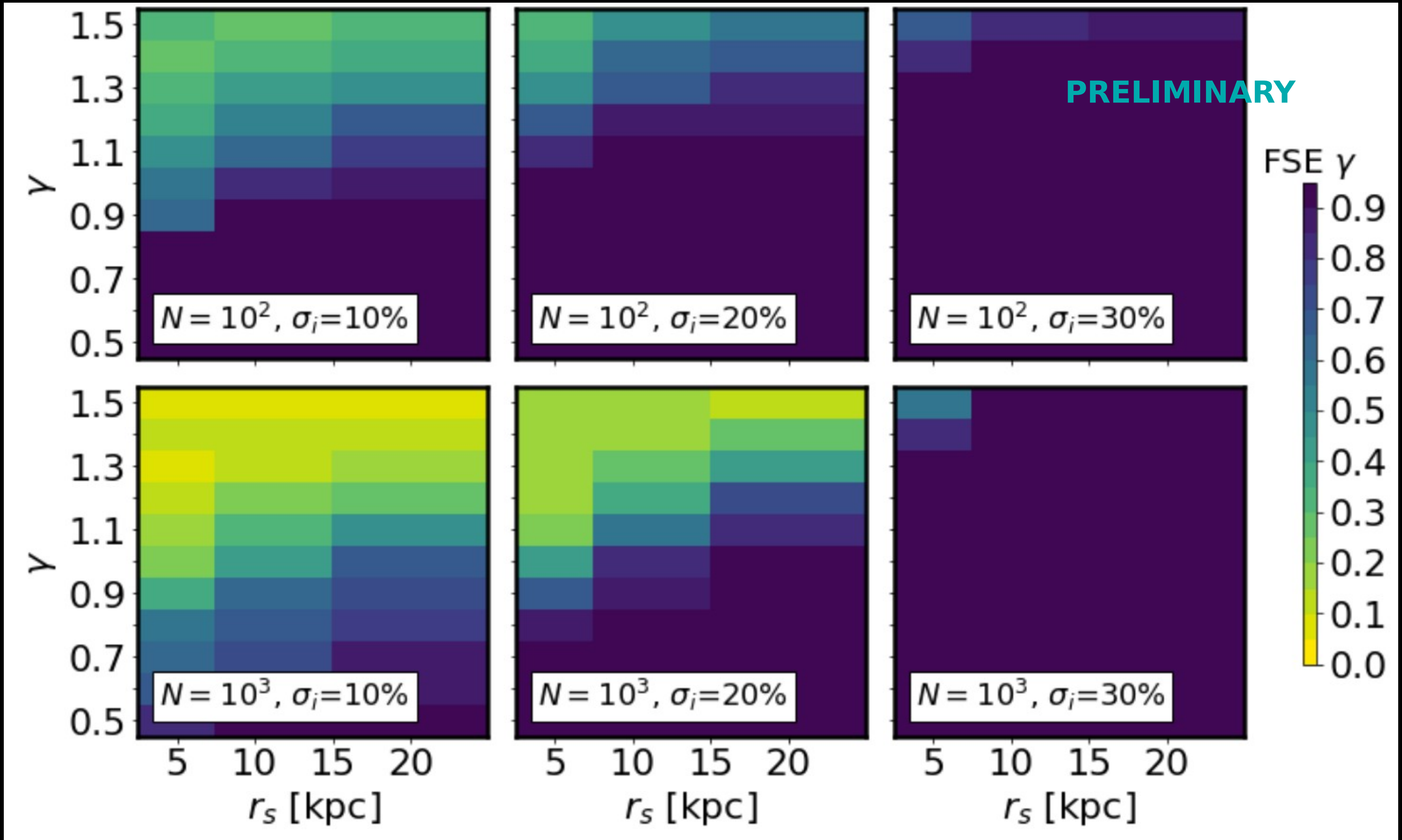




$$\rho_\chi(r) = \frac{\rho_0}{(r/r_s)^\gamma (1 + (r/r_s))^{3-\gamma}}$$

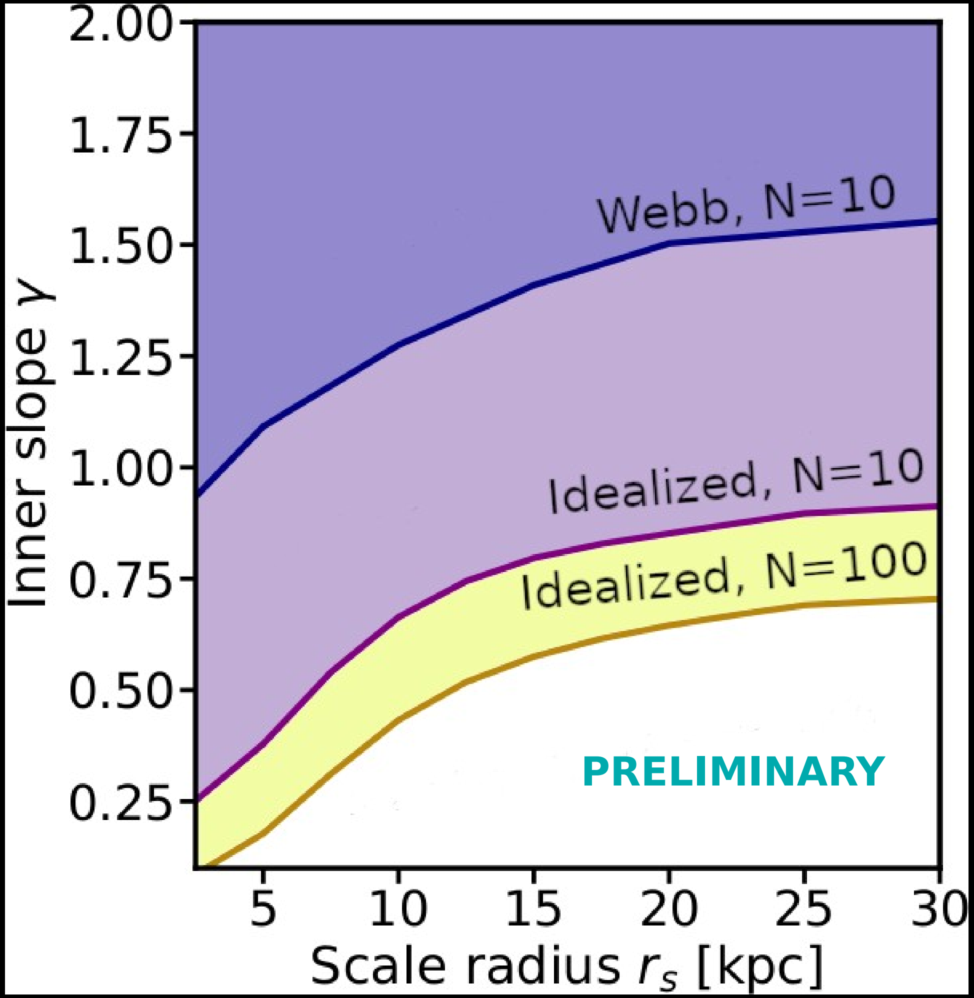
Benito, Leane, Poder, Smirnov (to appear)

Sensitivity to DM halo parameters



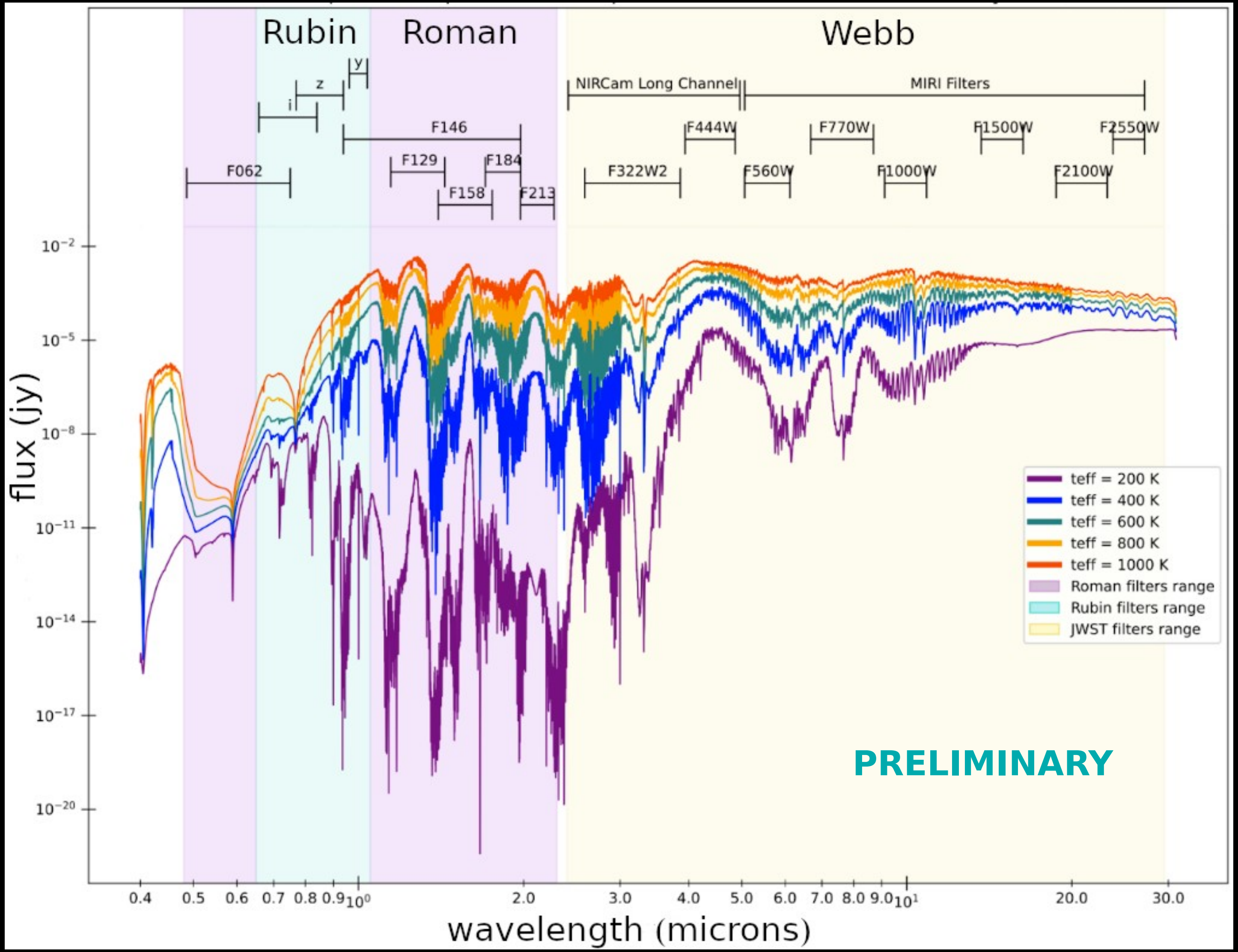
Benito, Leane, Poder, Smirnov
(to appear)

Sensitivity to DM halo parameters



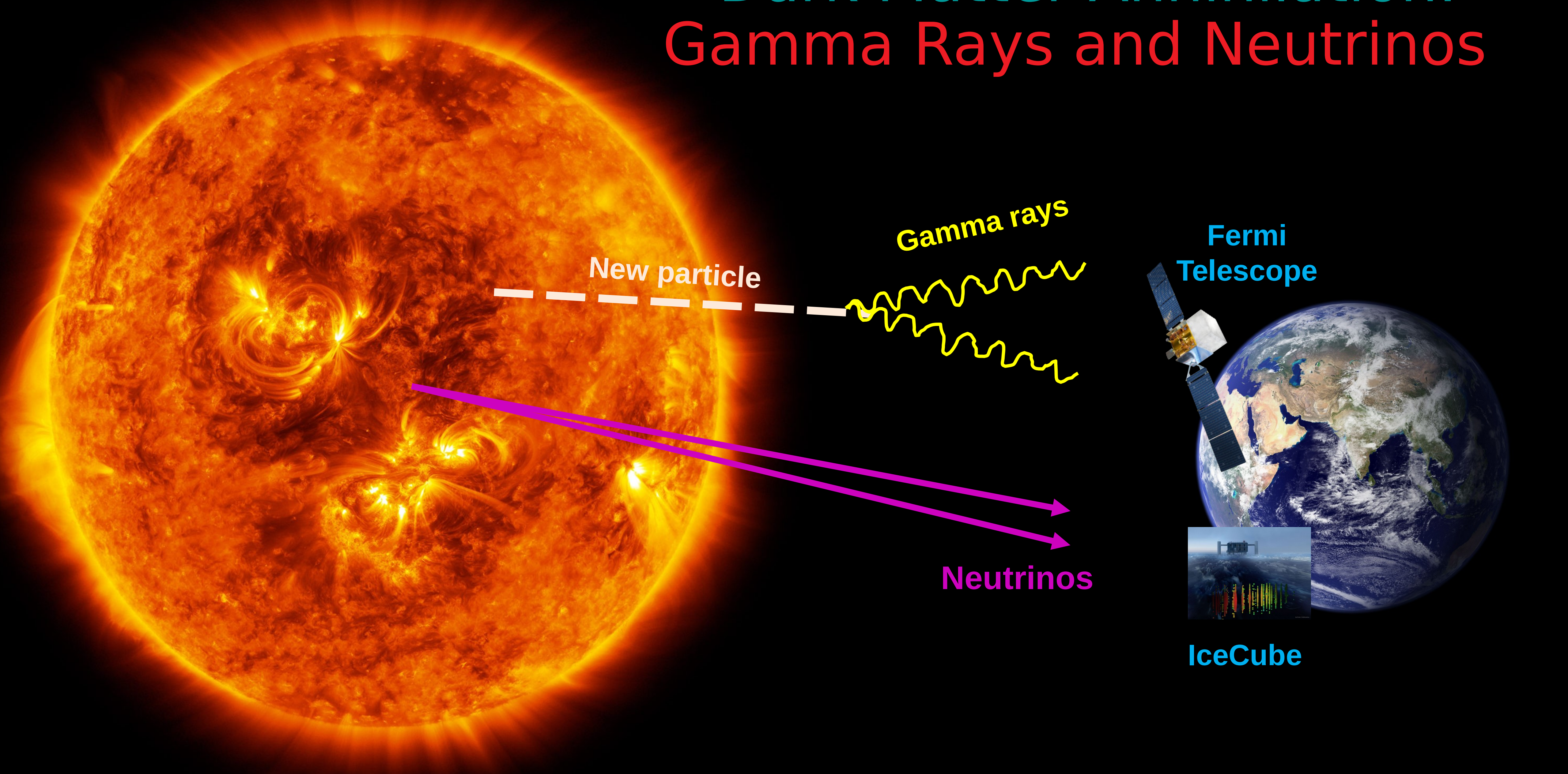
Benito, Leane, Poder, Smirnov
(to appear)

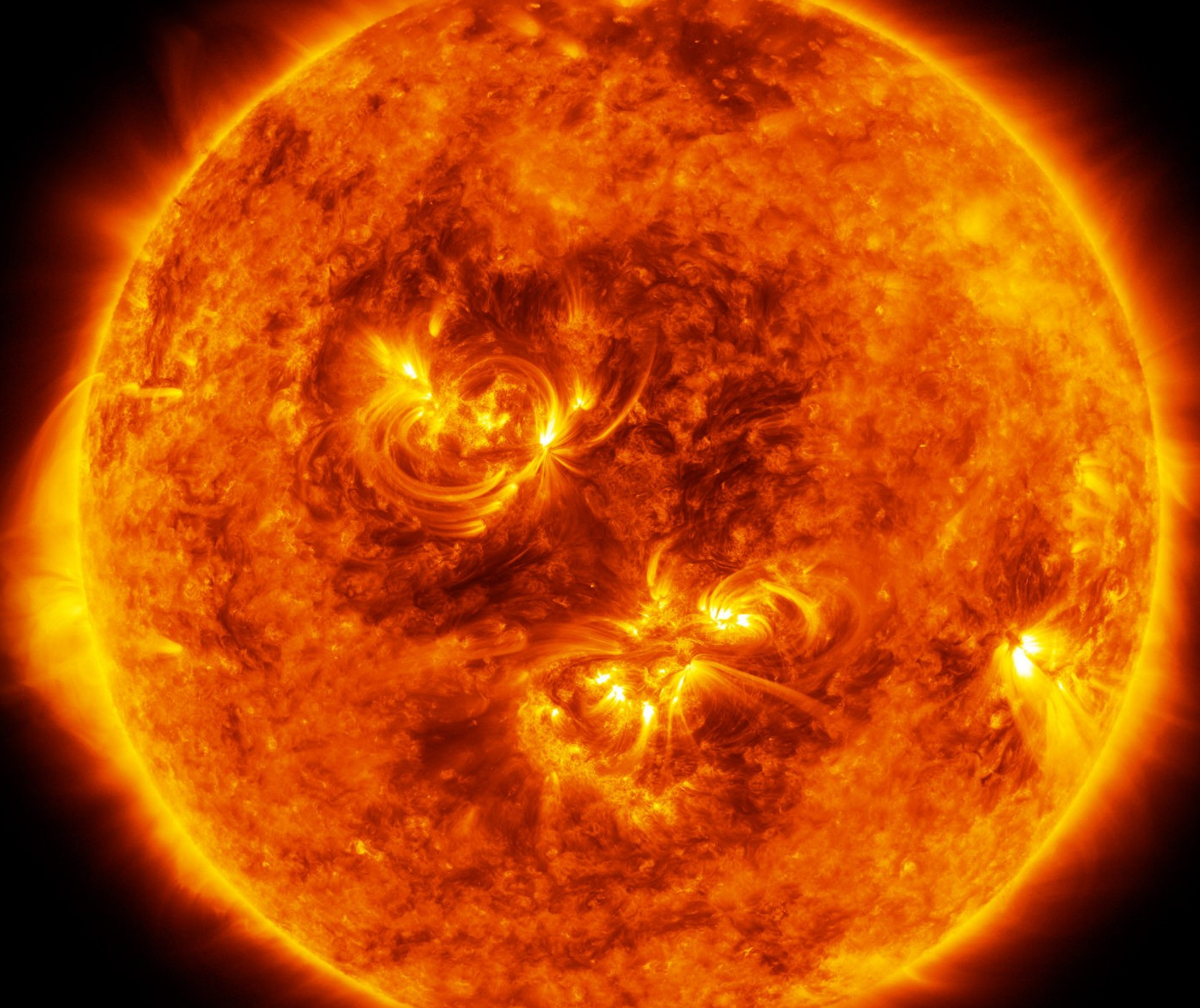
Atmospheric Modeling + DM Heating



Ongoing w/ Paylor, Macintosh+

Dark Matter Annihilation:
Gamma Rays and Neutrinos





THE SUN

Press, Spergel 1985

Krauss, Freese, Press, Spergel 1985

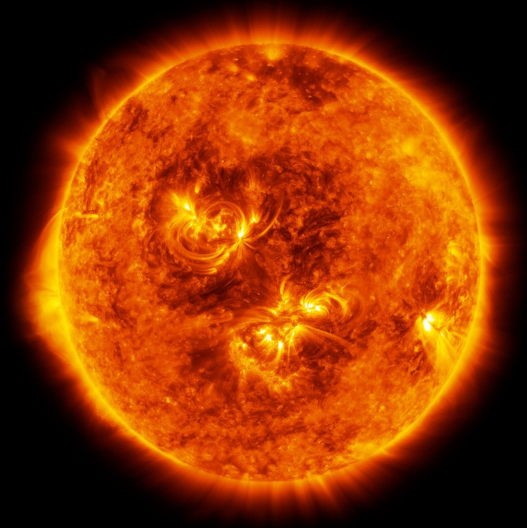
Silk, Olive, Srednicki, 1985

Stats: Hot, big, close

Escape velocity: 615 km /s

THE SUN

Available data:
Gamma-ray data (e.g. Fermi, HAWC)
Neutrino data (e.g. SuperK, IceCube)



Limitations:

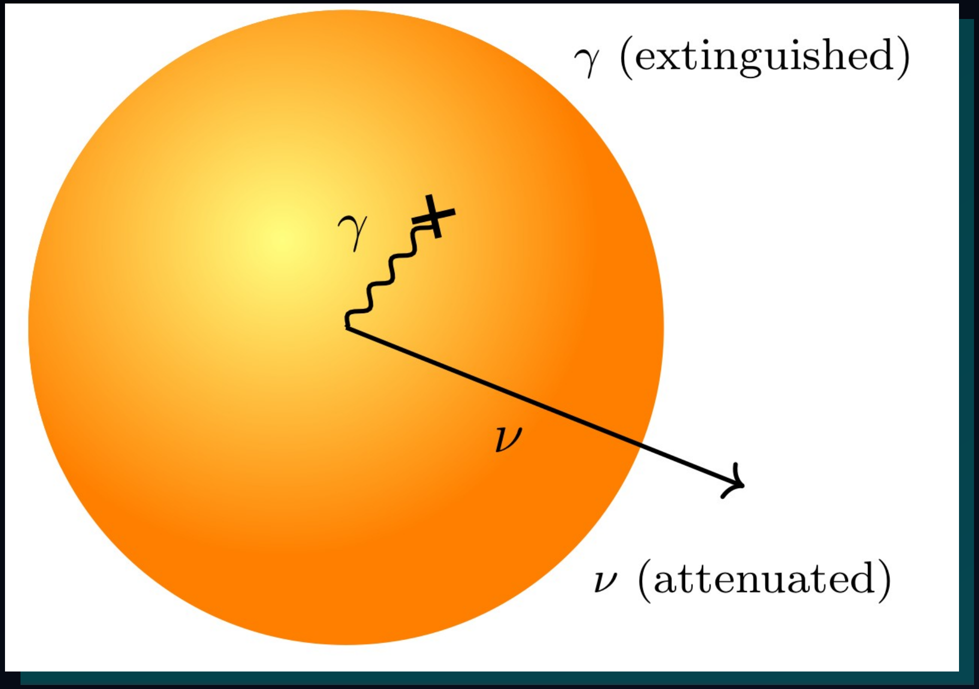
+ Hot
+ Higher DM evaporation (~GeV mass)

Benefits:

+ Huge
+ Proximity
+ Excellent data

THE SUN

- DM can be captured by scattering with solar matter, then annihilate to neutrinos



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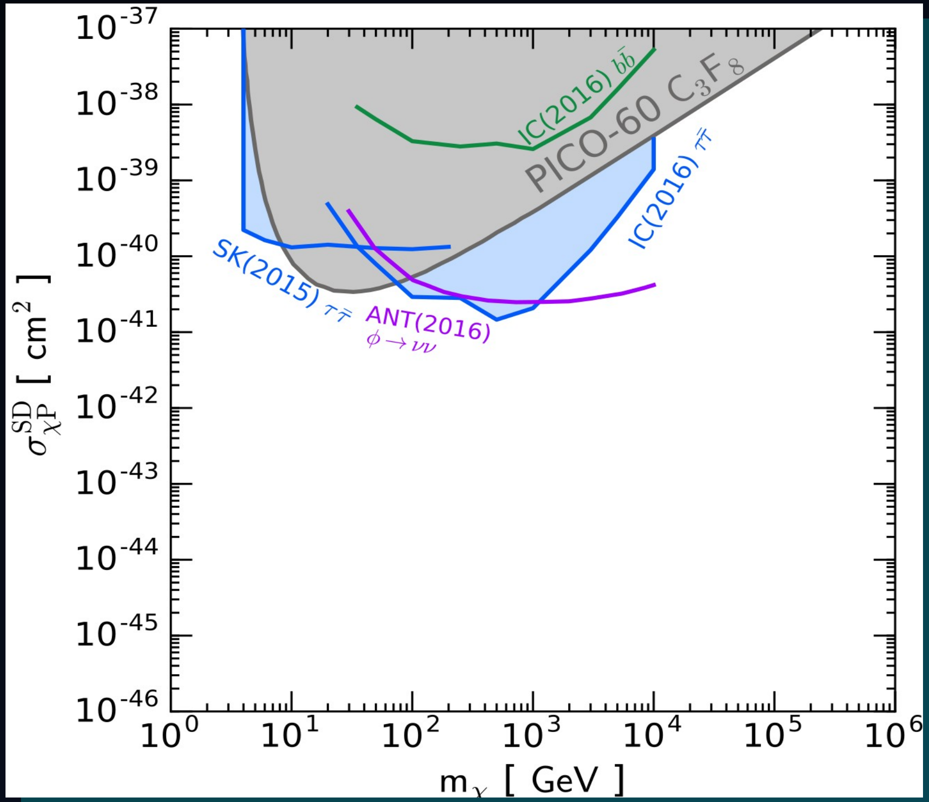
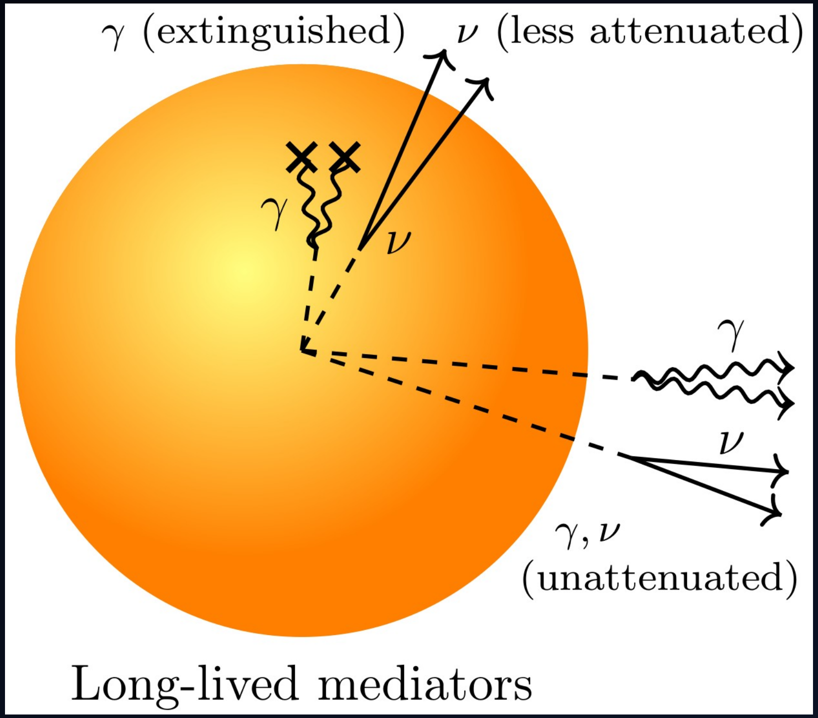


THE SUN

- DM can be captured by scattering with solar matter
- If DM annihilates to long-lived particles, neutrino signal is boosted



Schuster+ '10
Batell+ '10
Meade+ '10

THE SUN

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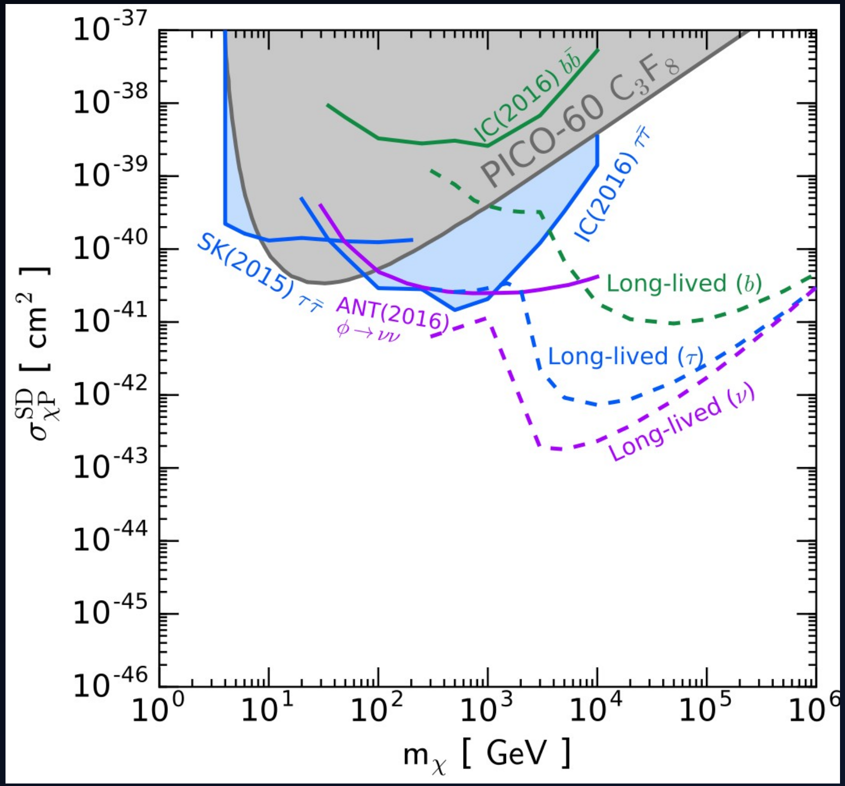

Leane, Ng, Beacom, 2017

THE SUN

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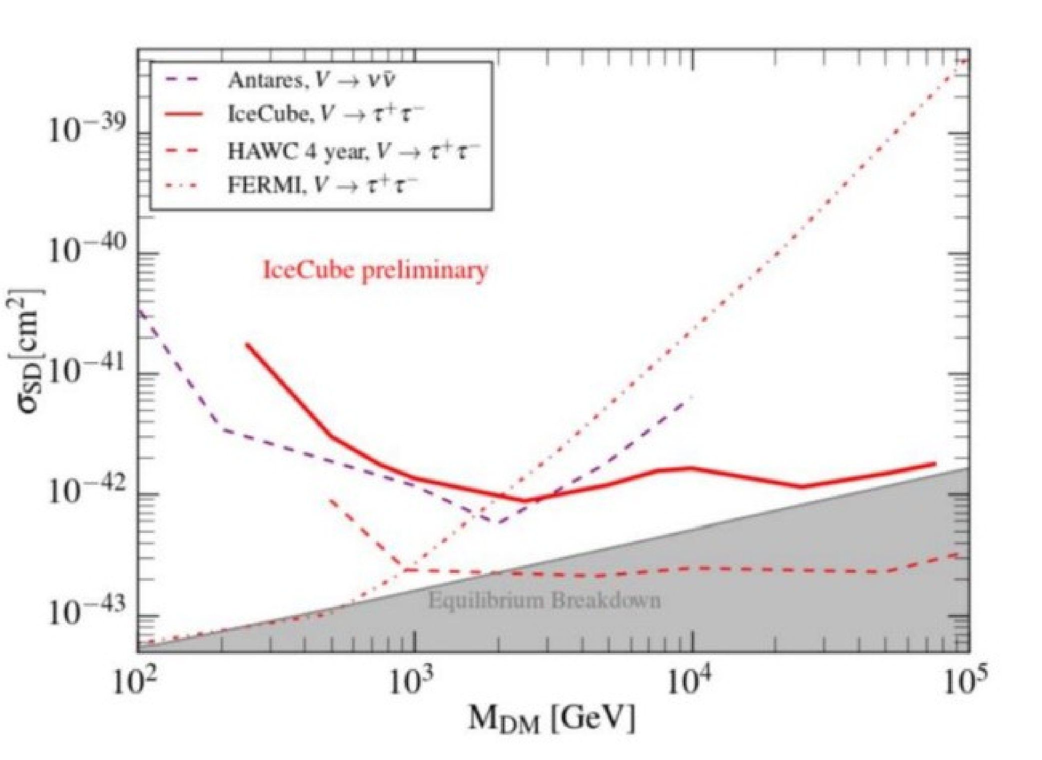
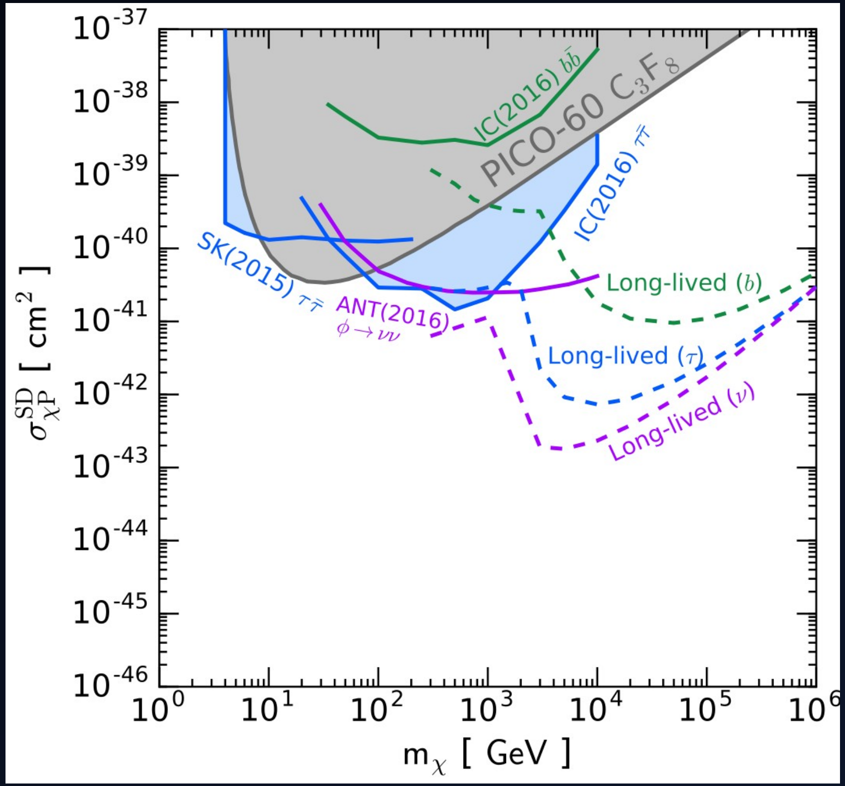


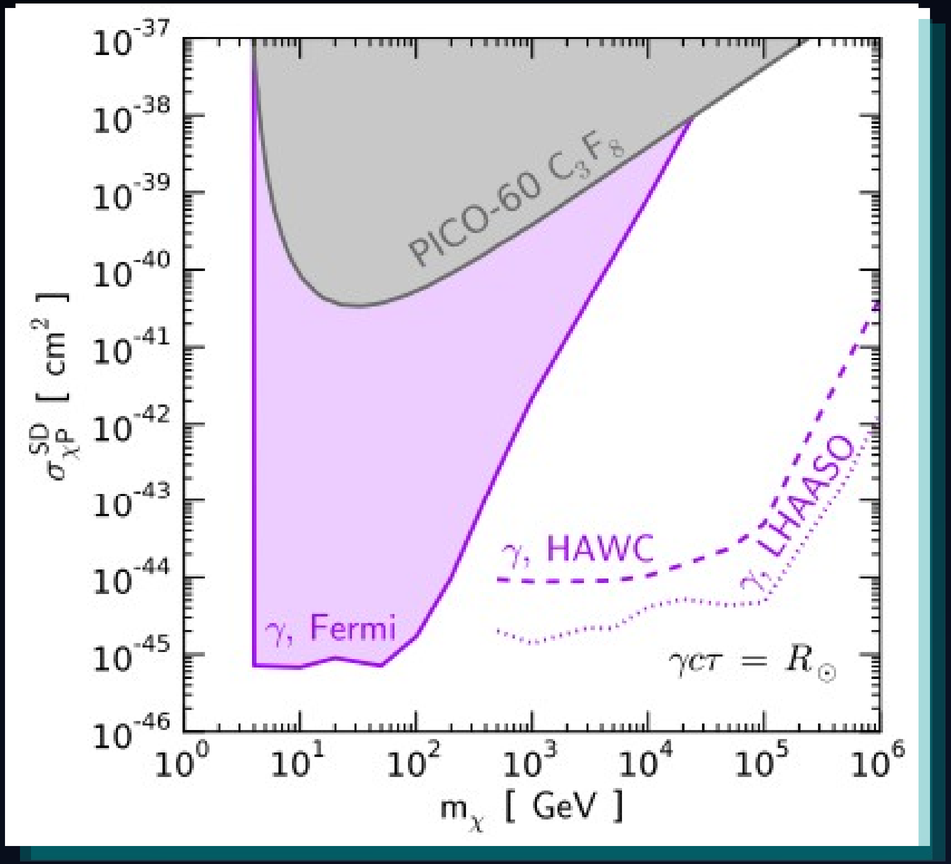
IceCube 2022



Leane, Ng, Beacom, 2017

THE SUN

- Long-lived particle scenario, excellent gamma-ray sensitivity



Leane, Ng, Beacom (PRD ‘17)

THE SUN

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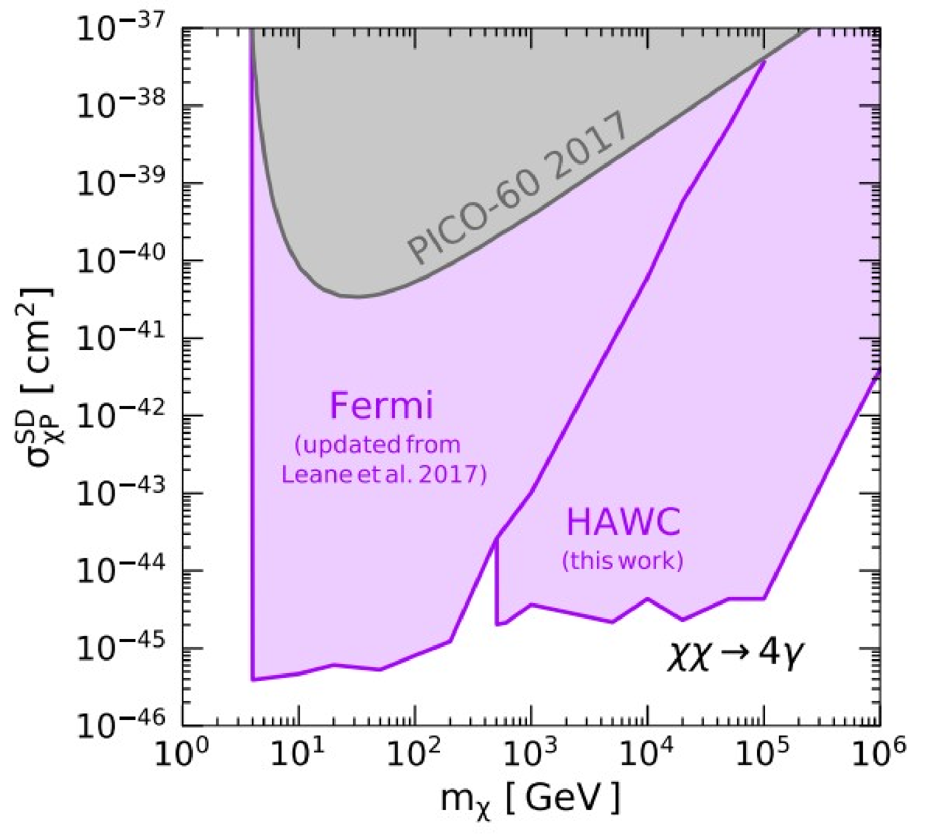
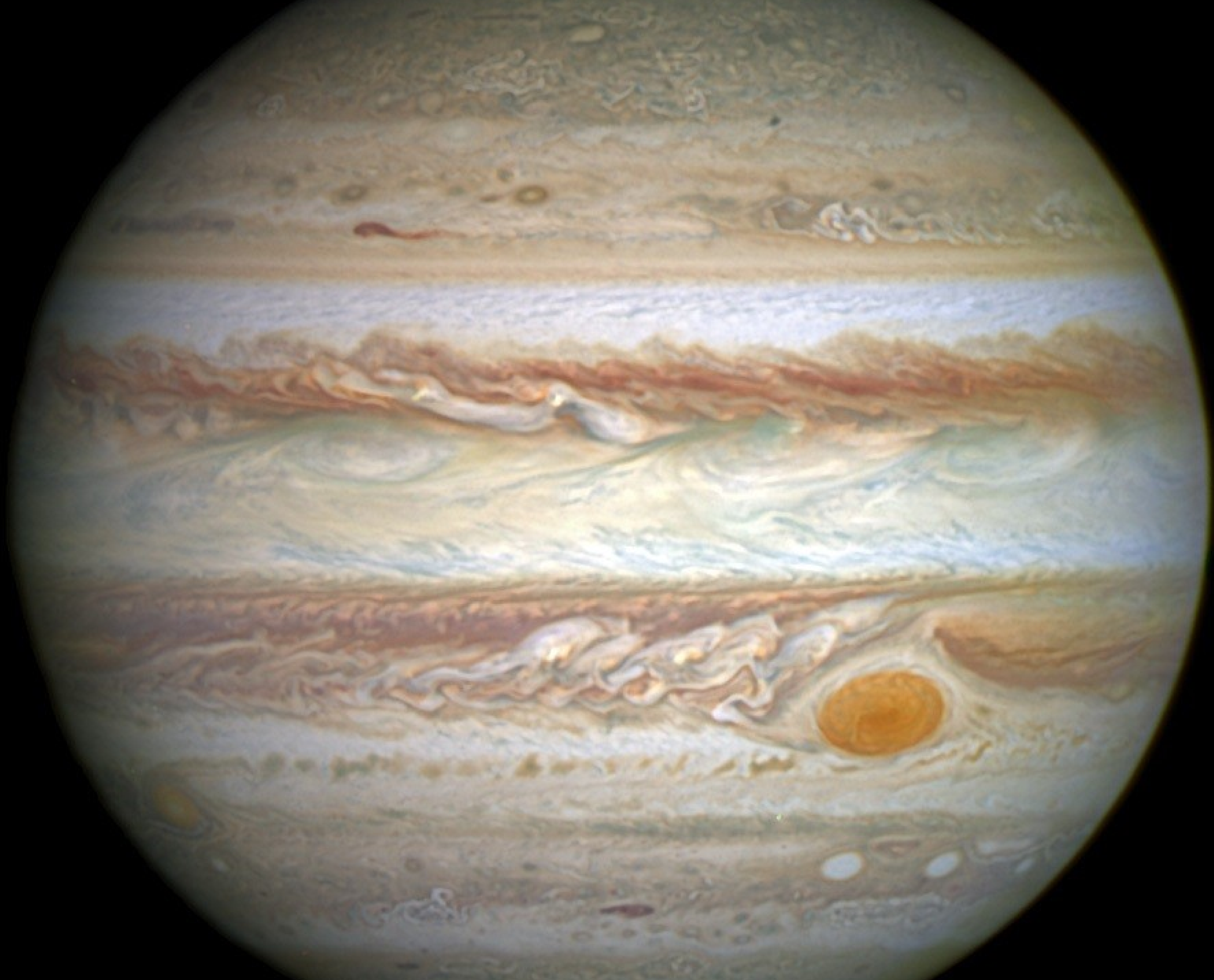


Leane, Ng, Beacom (PRD '17)
Beacom, Leane, Linden, Ng, Peter, Zhou
Un Nisa + HAWC Collaboration (PRD '18)



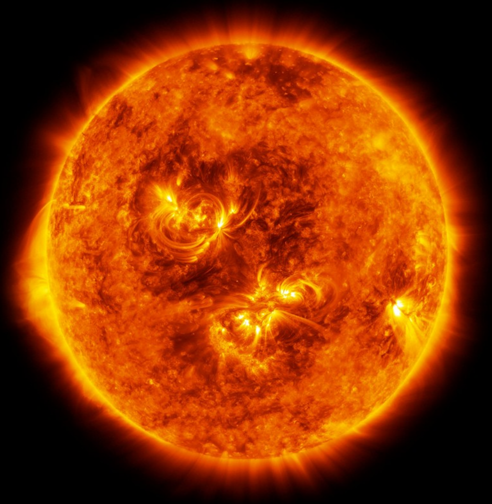
JUPITER

Kawasaki, Murayama, Yanagida 1992

Adler 2009

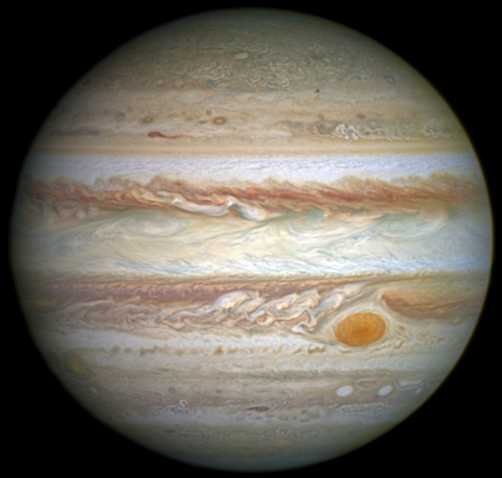
Leane, Linden 2021

Why Jupiter?



Sun

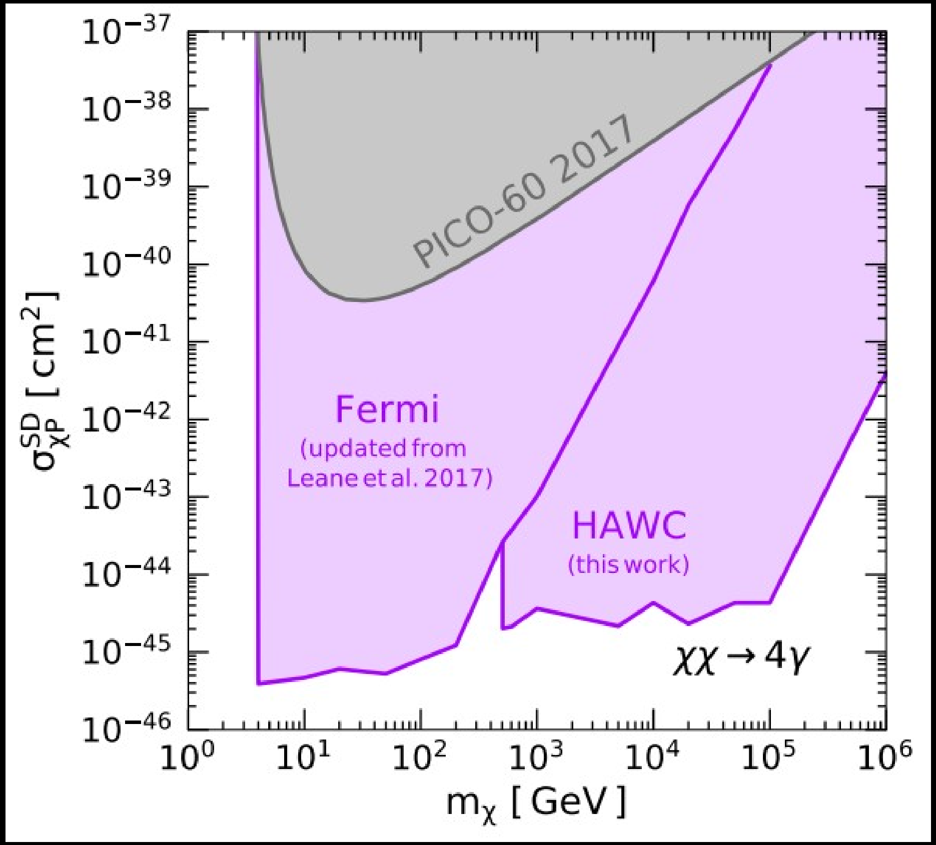
BIG
Hot



Jupiter

BIG
Cold

Solar Comparison



Sun

Long-Lived Mediator Limits

Leane, Ng, Beacom (PRD '17)
Leane + HAWC Collaboration (PRD '18)

Jupiter

Cooler than the Sun:
MeV-DM mass sensitivity!

Jupiter in Gamma Rays

What does Jupiter look like in gamma rays?
No one had ever really checked!

If we find gammas, they could be from:

+ acceleration of cosmic rays in Jovian magnetic fields

+ interaction of cosmic rays with Jupiter's atmosphere

...or something exotic (dark matter)!

Fermi Analysis of Jupiter

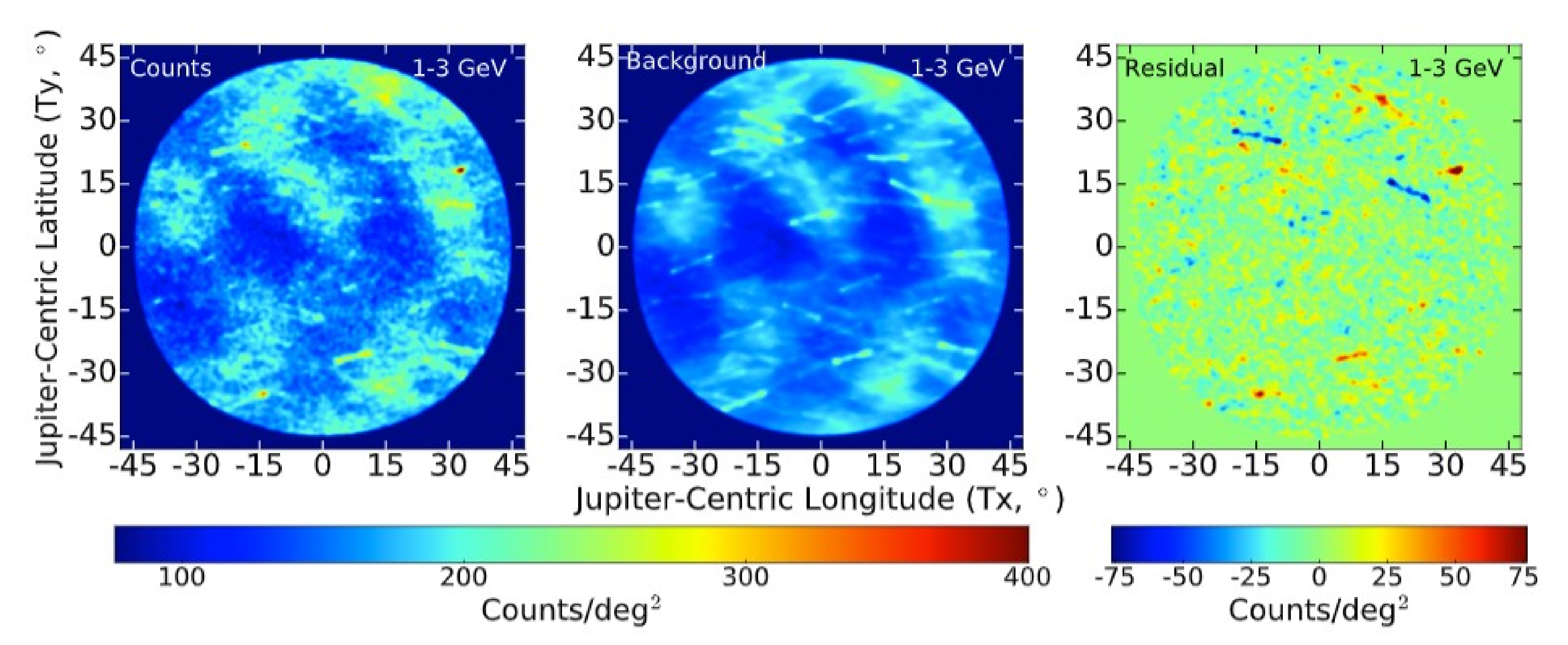
+ Analyze 12 years of Fermi data, 10 MeV – 10 GeV

+ Select photons within 45 degrees of Jupiter's orbit

+ Data-driven background model from Jupiter orbit when it is not there

+ Subtract "on" and "off" map events

Jupiter in Gamma Rays



Leane + Linden '21

New dark matter limits

Some assumptions:

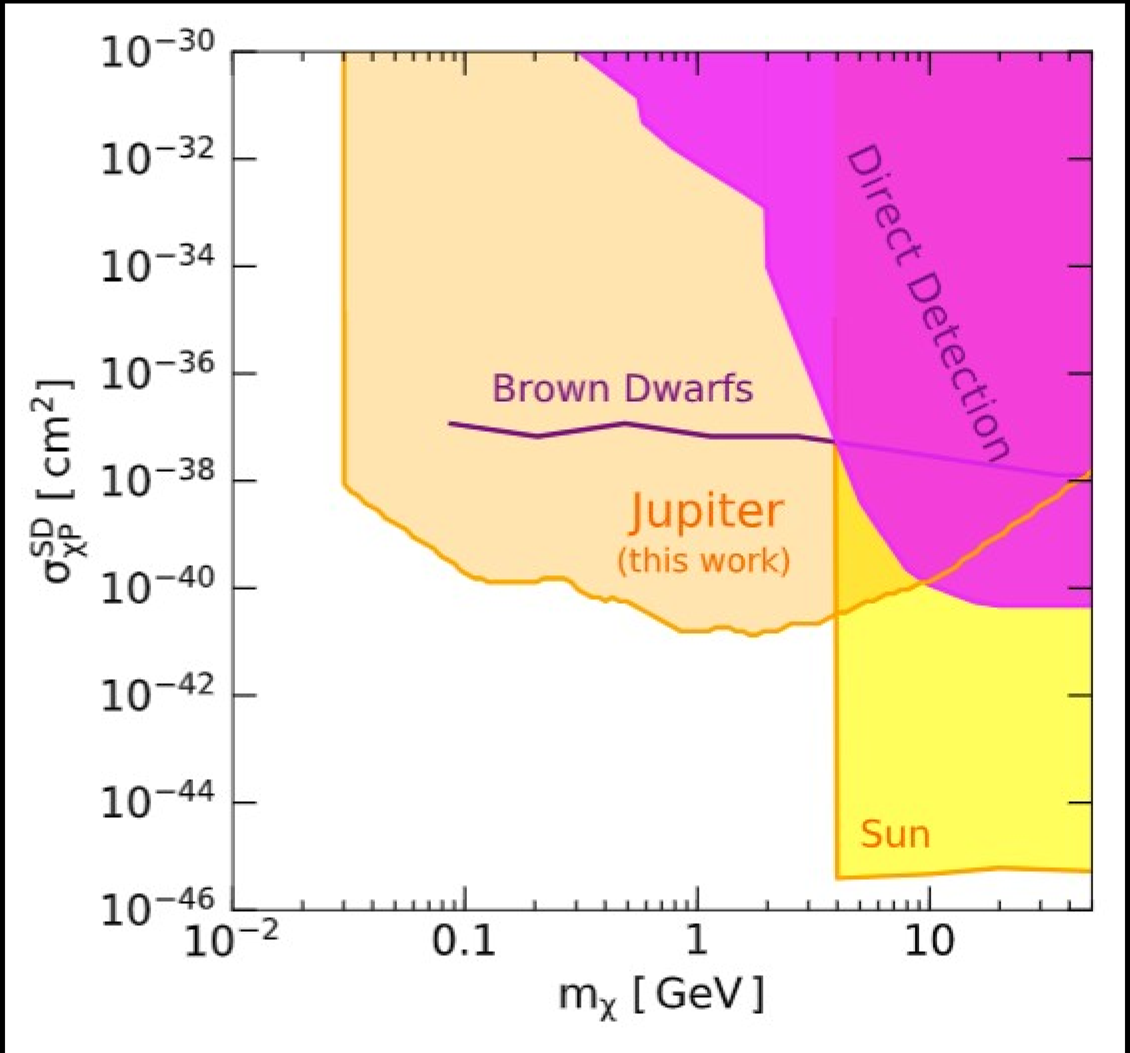
+ direct decay to gammas
(but other final states possible)

+ mediator decay length
> Jupiter radius

+ equilibrium

+ low mass end model dependent

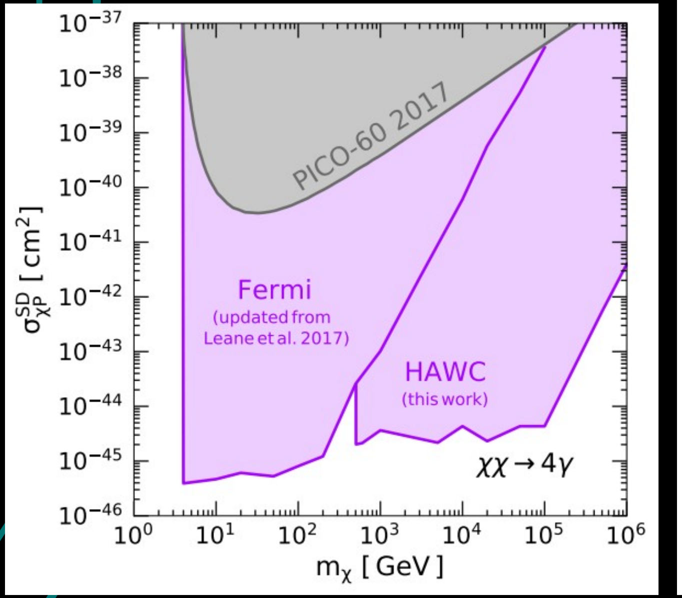
Not guaranteed for all models!



Rebecca Leane (SLAC)

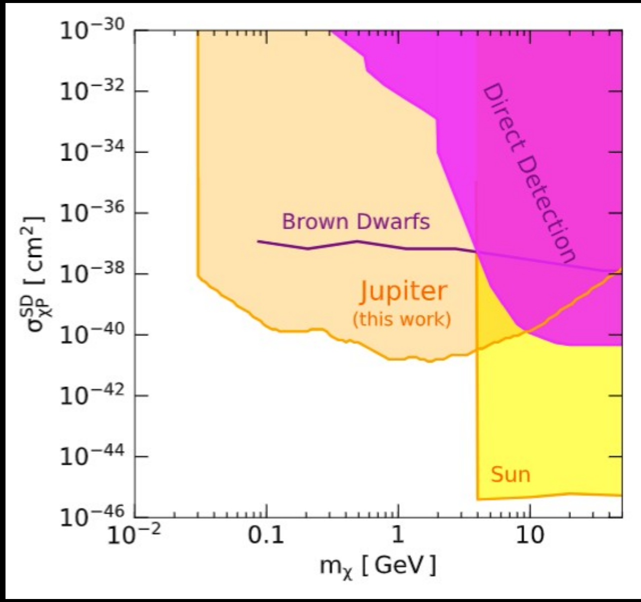
Leane + Linden ‘21

Optimal Celestial Target for Gammas?



Sun

Leane, Ng, Beacom 2017
Leane + HAWC Collaboration 2018



Jupiter

Leane, Linden 2021

?

Neutron Star

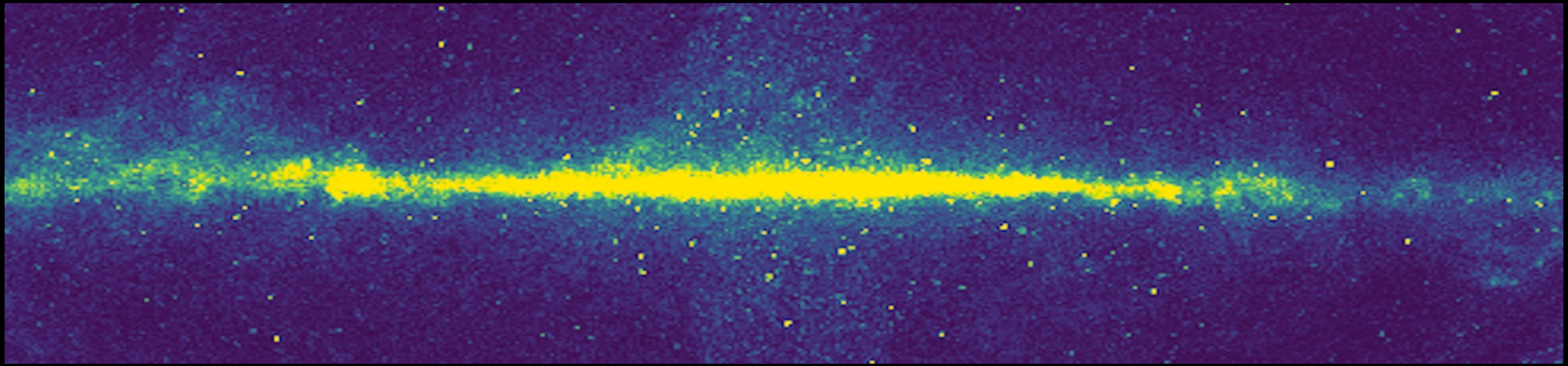
?

Brown Dwarf

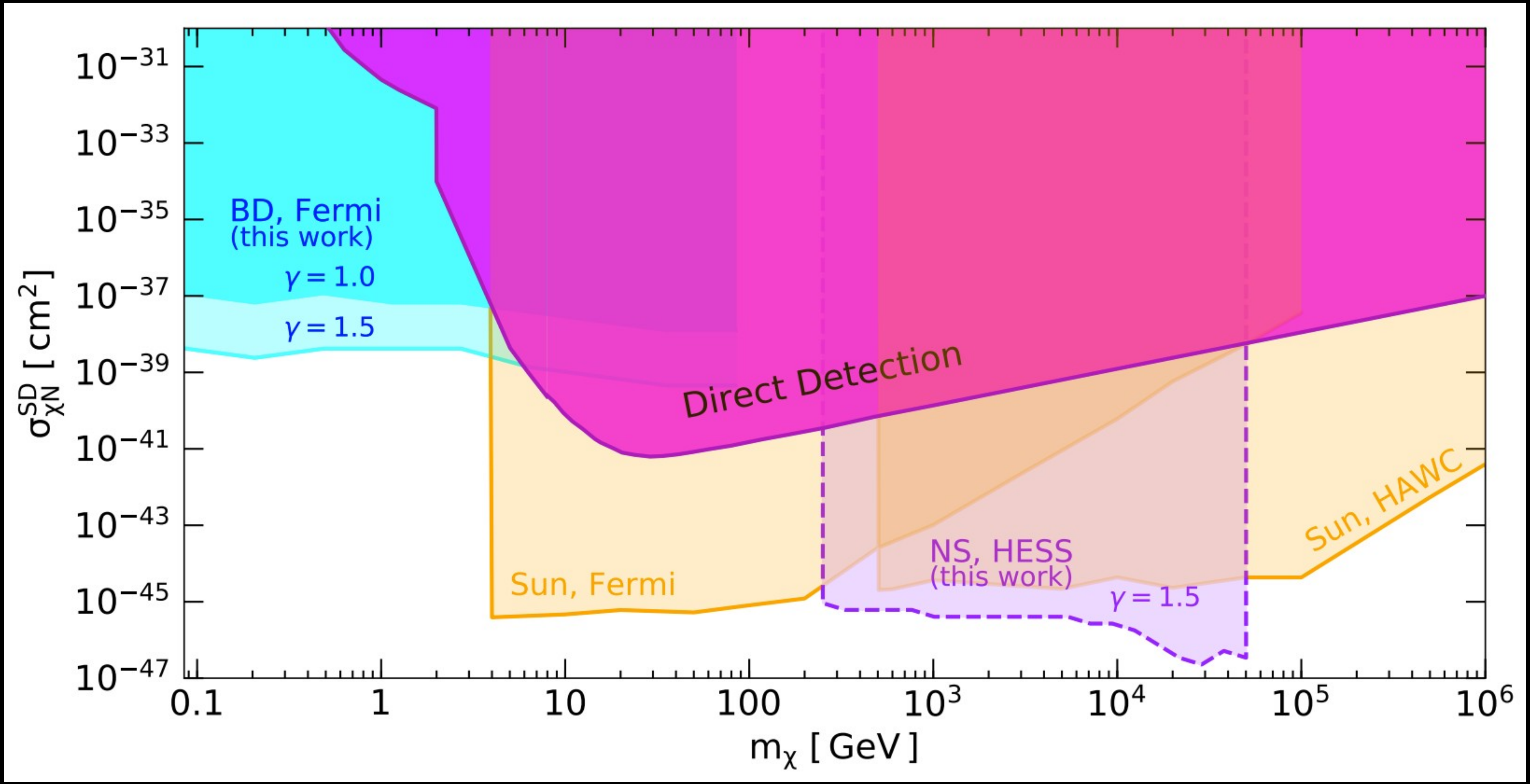
Galactic Center Population Signal

- Use all the neutron stars, all the brown dwarfs
  - Compare with Fermi and H.E.S.S. data for Galactic Center
  - No model assumptions on mediator, other than must escape

- Our new signal follows matter density: DM density * stellar density
  - DM Halo annihilation scales with DM density squared



New Limits w/ Brown Dwarfs and Neutron Stars



Leane, Linden, Mukhopadyay, Toro 2021

Interesting things I didn’t mention…

- EoS effects on NSs, gravitational waves

Panotopoulos, Lopes 2017
Ellis et al 2018
Nelson, Reddy, Zhou, 2018
Collier, Croon, Leane, 2022

- DM in Pop III stars

Freese, Spolyar, Aguirre 2008
Freese, Gondolo, Sellwood, Spolyar 2008

- Stellar evolution effects

Taoso et al 2010
Frandsen, Sarkar 2010
Zentner, Hearin 2011

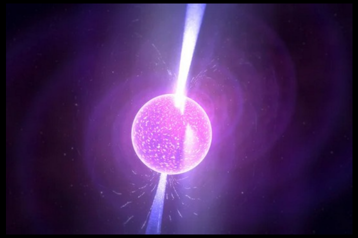
- Creation of black holes, destruction of stars

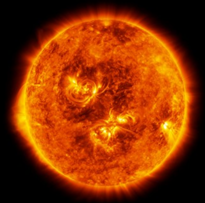
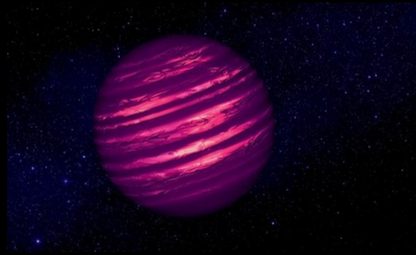
Gould, Draine, Romani, Nussinov 1989

- Evaporation of black holes, neutrinos

Acevedo, Bramante, Goodman, Kopp, Opferkuch 2020

Summary



- Celestial bodies are playgrounds for discovering DM!
- Heating and neutrino/gamma-ray detection possible
- Earth, Sun, and Jupiter now already have strong constraints
- Exoplanets, Planets, White Dwarfs and Neutron Stars may provide new DM sensitivities
- New technologies and searches coming soon, also, hopefully DM!


EXTRA SLIDES

Exoplanet Search Targets

Earths + Super Earths:
Mass: 0.001– 0.01 M_{jup}
Radius: ~0.1 - 1 R_{jup}

ideal

Jupiters + Super Jupiters:
Mass: 1 – 13 M_{jup}
Radius: ~1 R_{jup}

Brown dwarfs:
Mass: 13 – 75 M_{jup}
Radius: ~1 R_{jup}
Very dense!

Rogue Planets:
Cold and all alone!

Most commonly Jupiter-sized up to brown dwarf sized

Calculating Exoplanet Temperatures

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$$\Gamma_{\rm heat}^{\rm tot} = \Gamma_{\rm heat}^{\rm ext} + \Gamma_{\rm heat}^{\rm int} + \Gamma_{\rm heat}^{\rm DM} = 4\pi R^2 \, \sigma_{\rm SB} \, T^4 \, \epsilon$$

- External heat: assume zero, means we need exoplanets either very far from their host, or not bound at all (rogue planets)

- Internal heat: determined by cooling rate over time, choose old exoplanets (e.g. 1-10 gigayears old) to minimize internal heat

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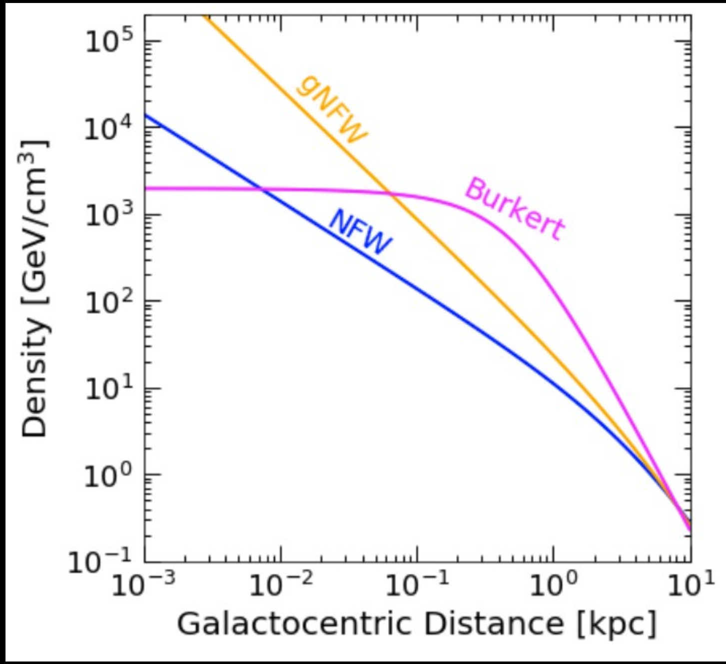
<u>Heat power from DM:</u>

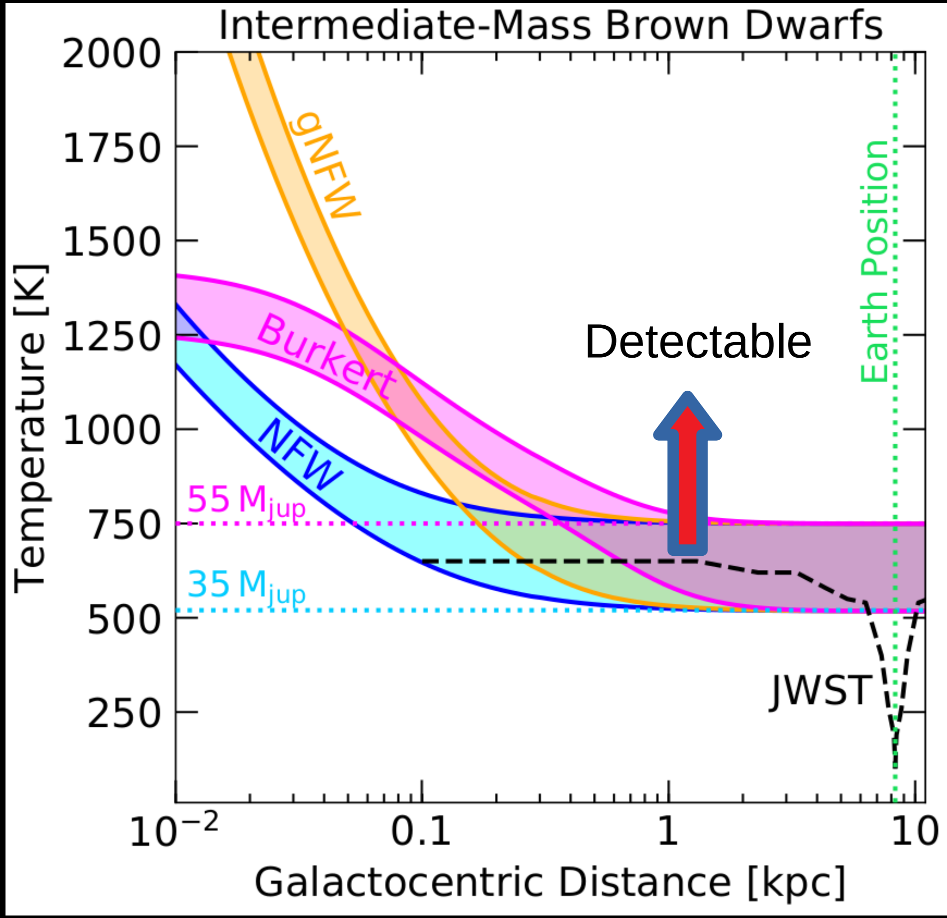
- DM density throughout Galaxy
- DM halo velocity
- Exoplanet escape velocity



Exoplanet temperatures vs sensitivity

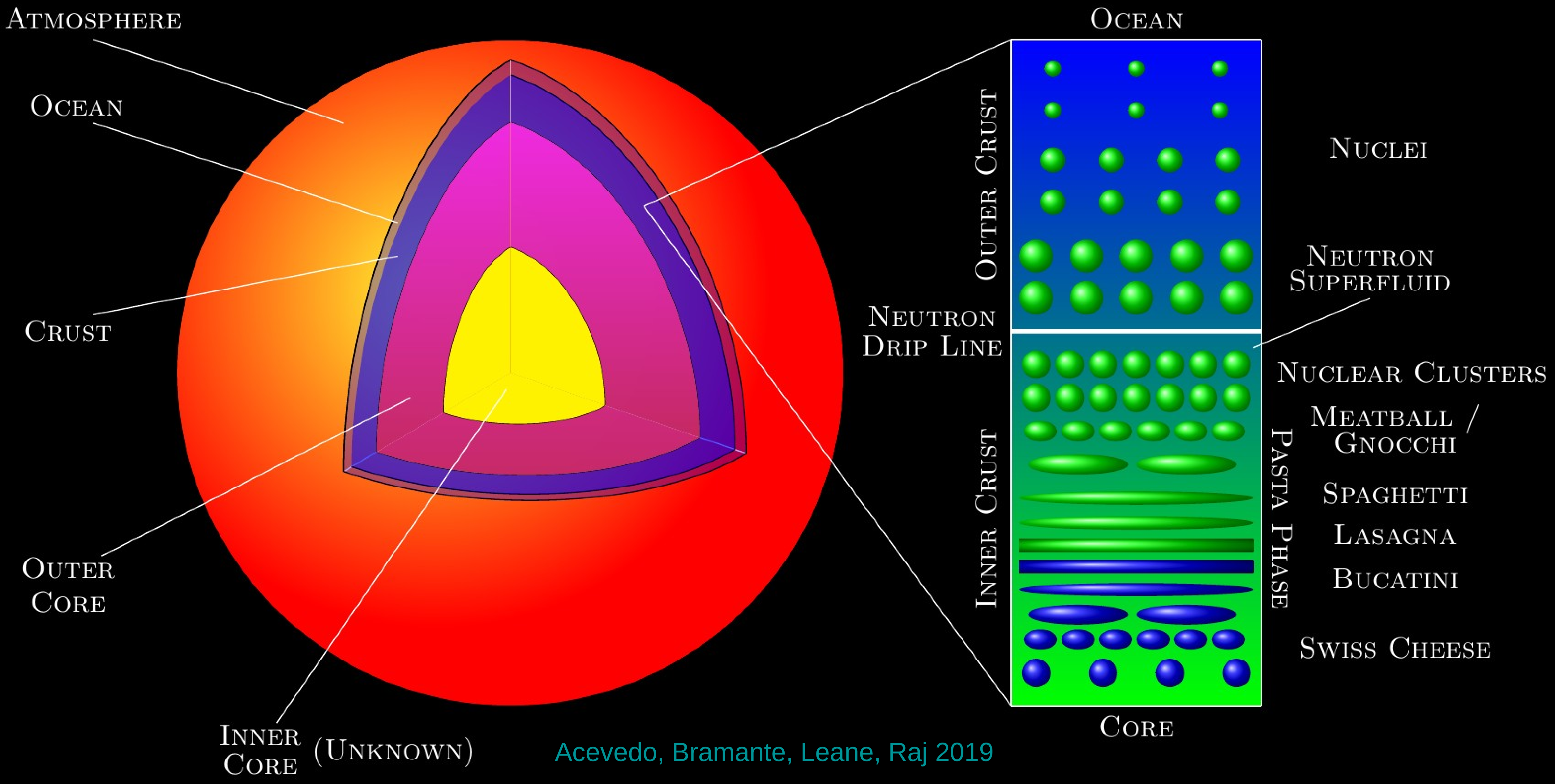
- NFW, gNFW, Burkert are DM profiles, shaded area is exoplanet mass range
- Sensitivity truncates at ~0.1kpc, due to stars per pixel, and dust scattering

35 Mjup – 55 Mjup



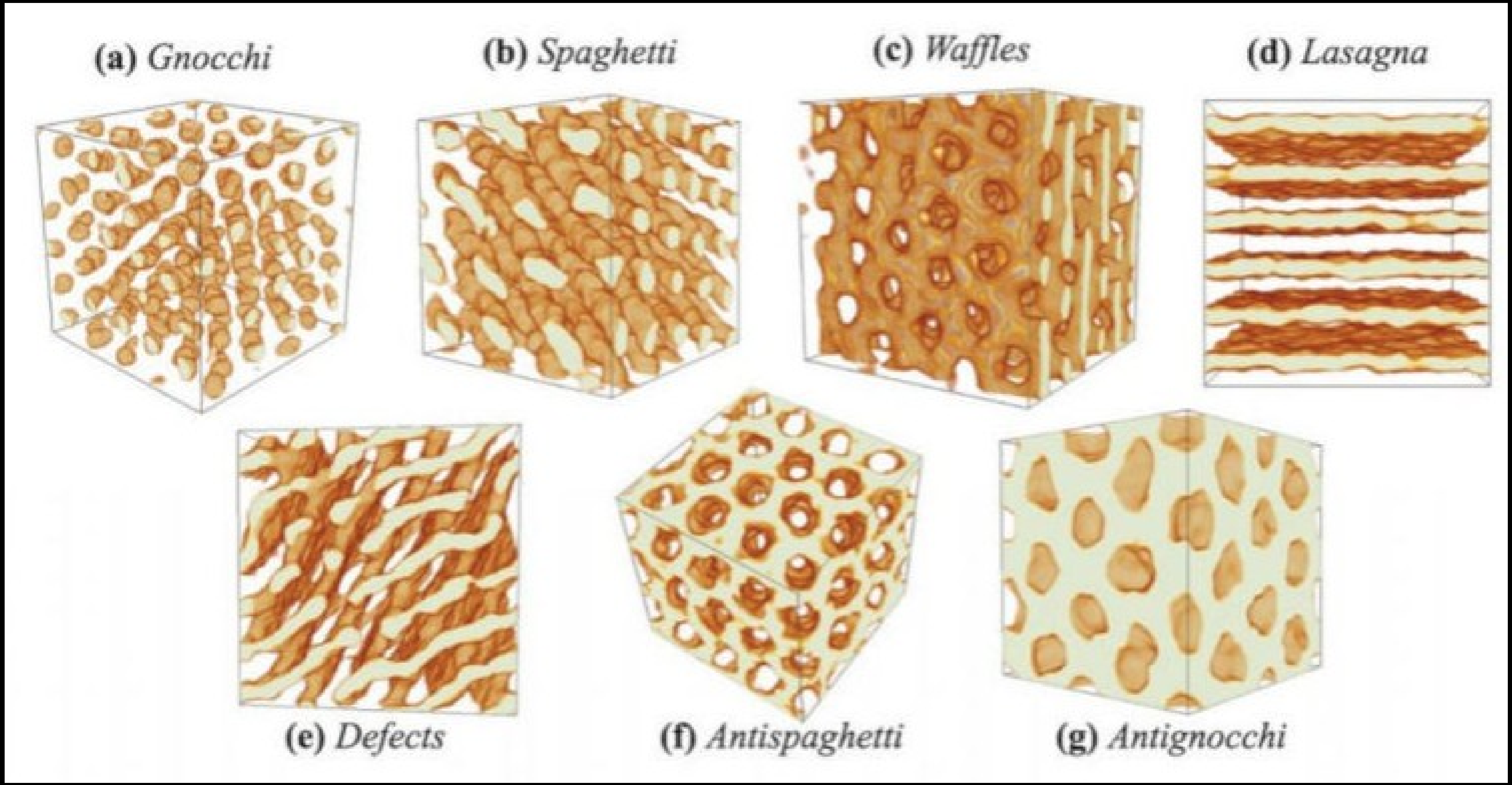
Leane + Smirnov, 2020

INSIDE NEUTRON STARS



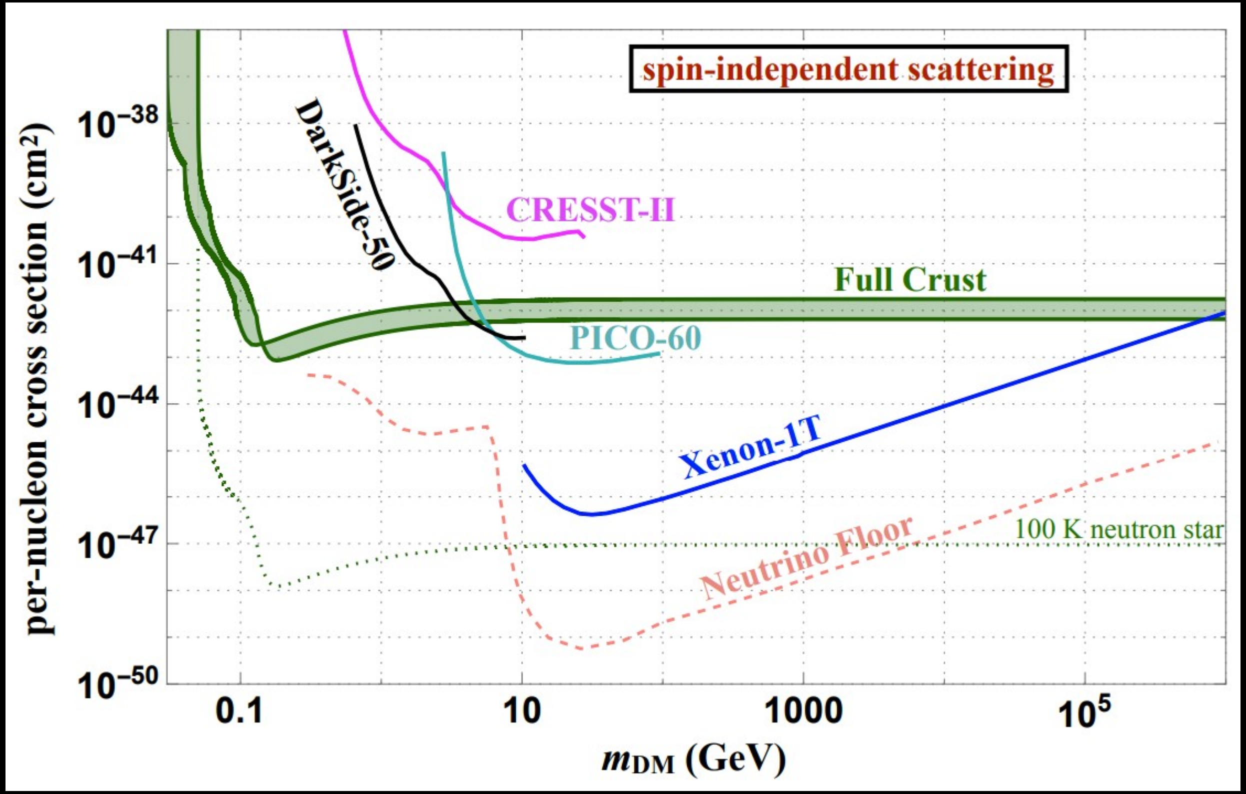
Acevedo, Bramante, Leane, Raj 2019

NUCLEAR PASTA



Caplan, Schneider, Horowitz '18

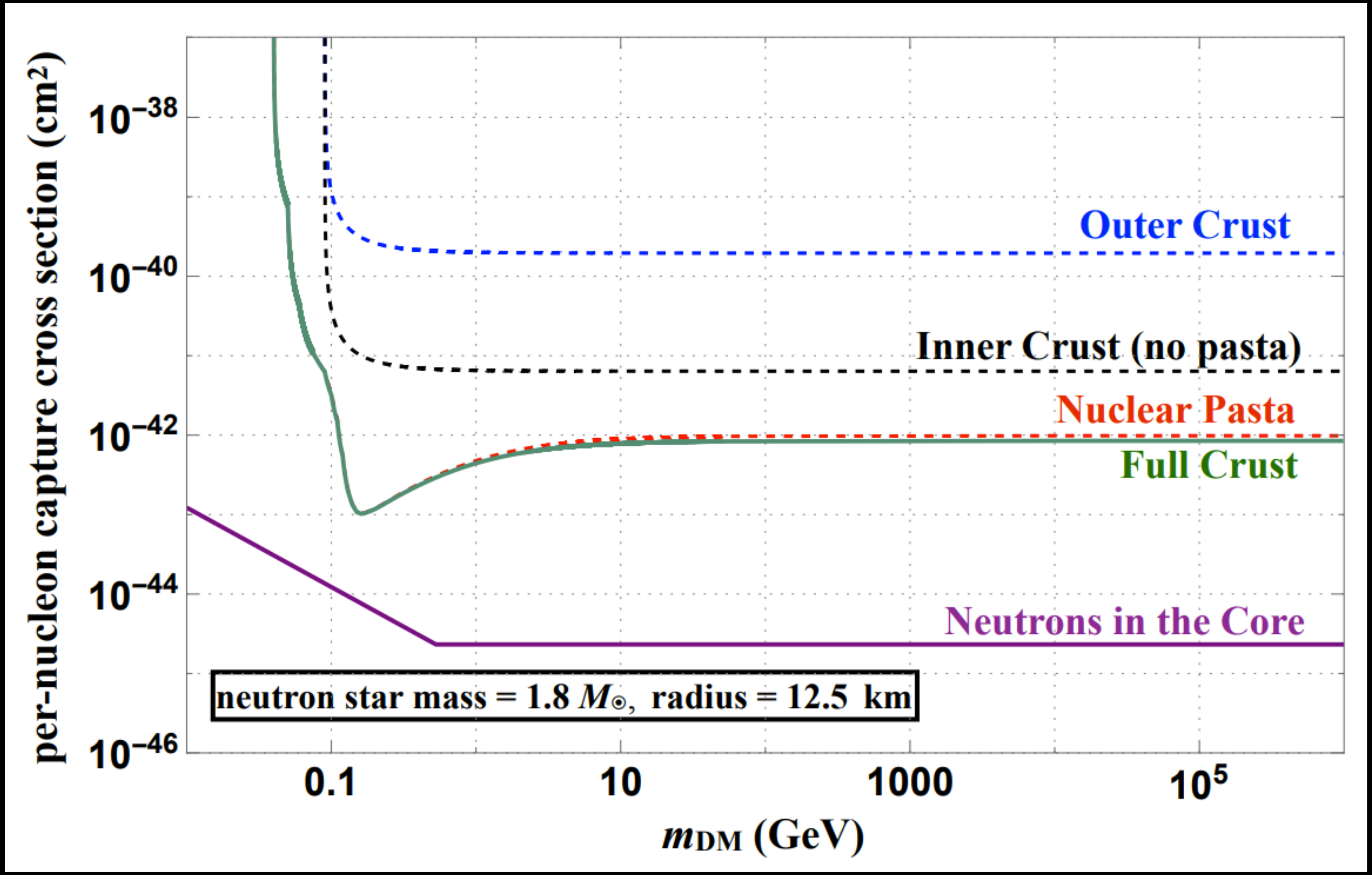
PASTA CAN BEAT DIRECT DETECTION



Acevedo, Bramante, Leane, Raj, 2019

Low + high masses, velocity suppressed, spin-dependent, inelastic DM

DARK MATTER – NEUTRON STAR INTERACTIONS



$T_\infty^{\rm crust} = 1620$ K

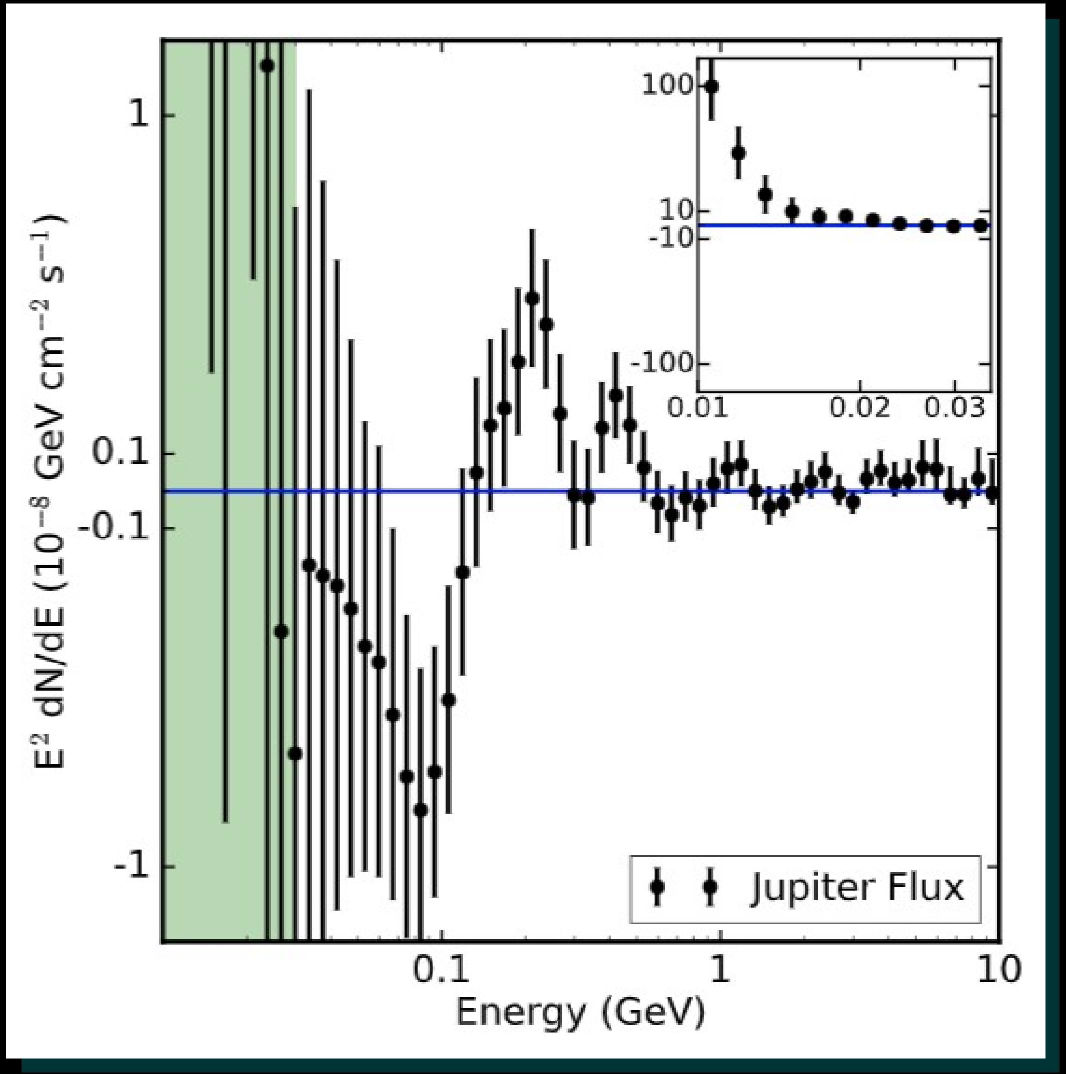
Acevedo, Bramante, Leane, Raj, 2019

Jupiter Flux Limits

+ For range of power-law spectra, statistical sig of Jupiter emission never exceeds ~1.5σ

+ In low energy bins, "5σ" excess, but important systematics not there

+ Motivates follow-up with MeV telescopes: AMEGO, e-ASTROGAM



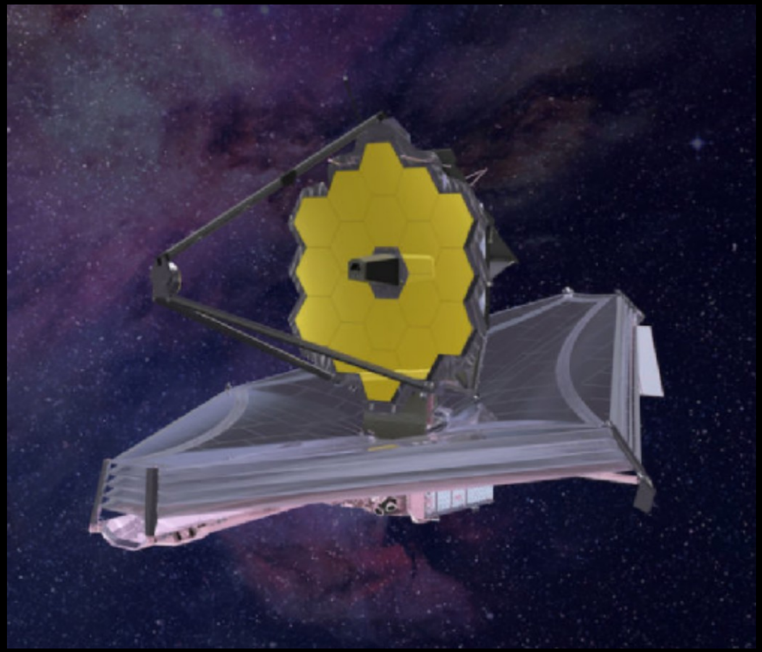
Leane, Linden, 2021

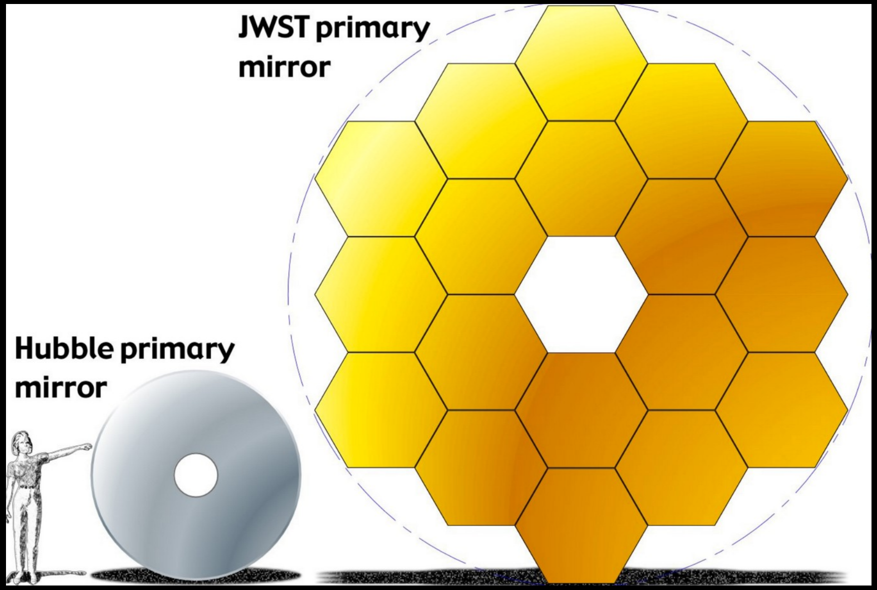
Telescope Sensitivity

- Use James Webb Space Telescope (planned launch Oct 2021)

- Infrared sensitivity (~0.5 – 28 microns)

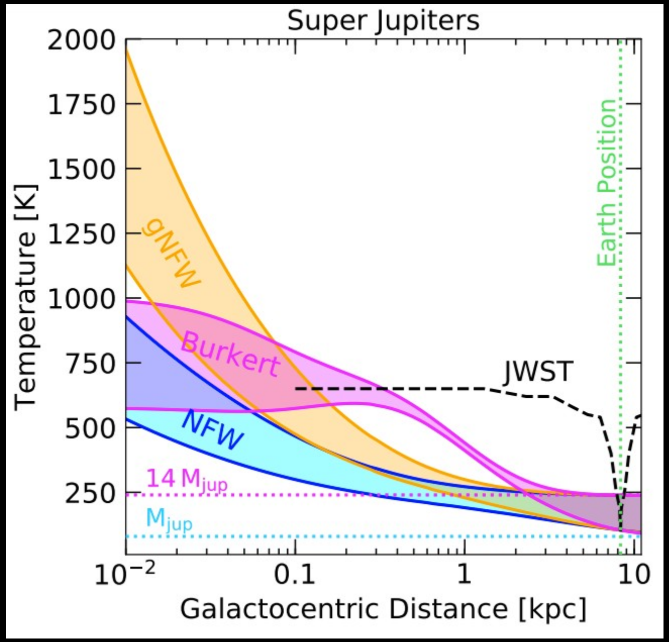
- Has many instruments and filters, relevant choice for maximum sensitivity depends on peak wavelength



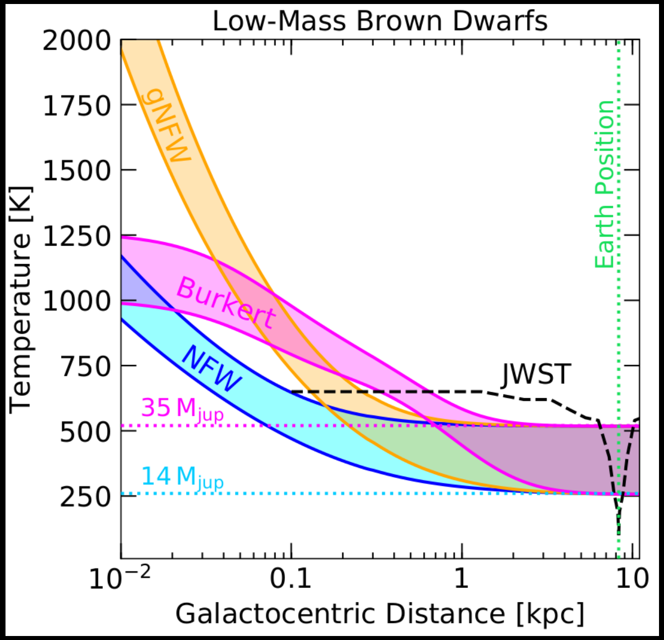


Exoplanet masses vs sensitivity

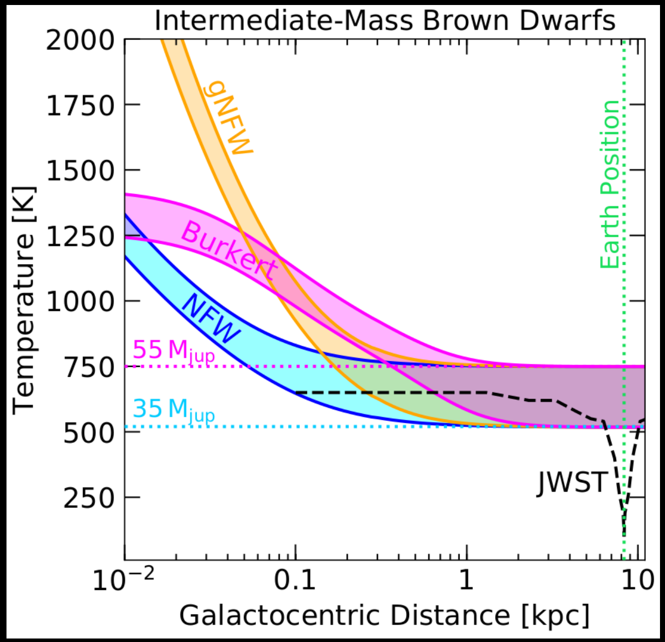
Mjup – 14 Mjup



14 Mjup – 35 Mjup



35 Mjup – 55 Mjup



Lower masses:
DM heat > internal heat at all positions

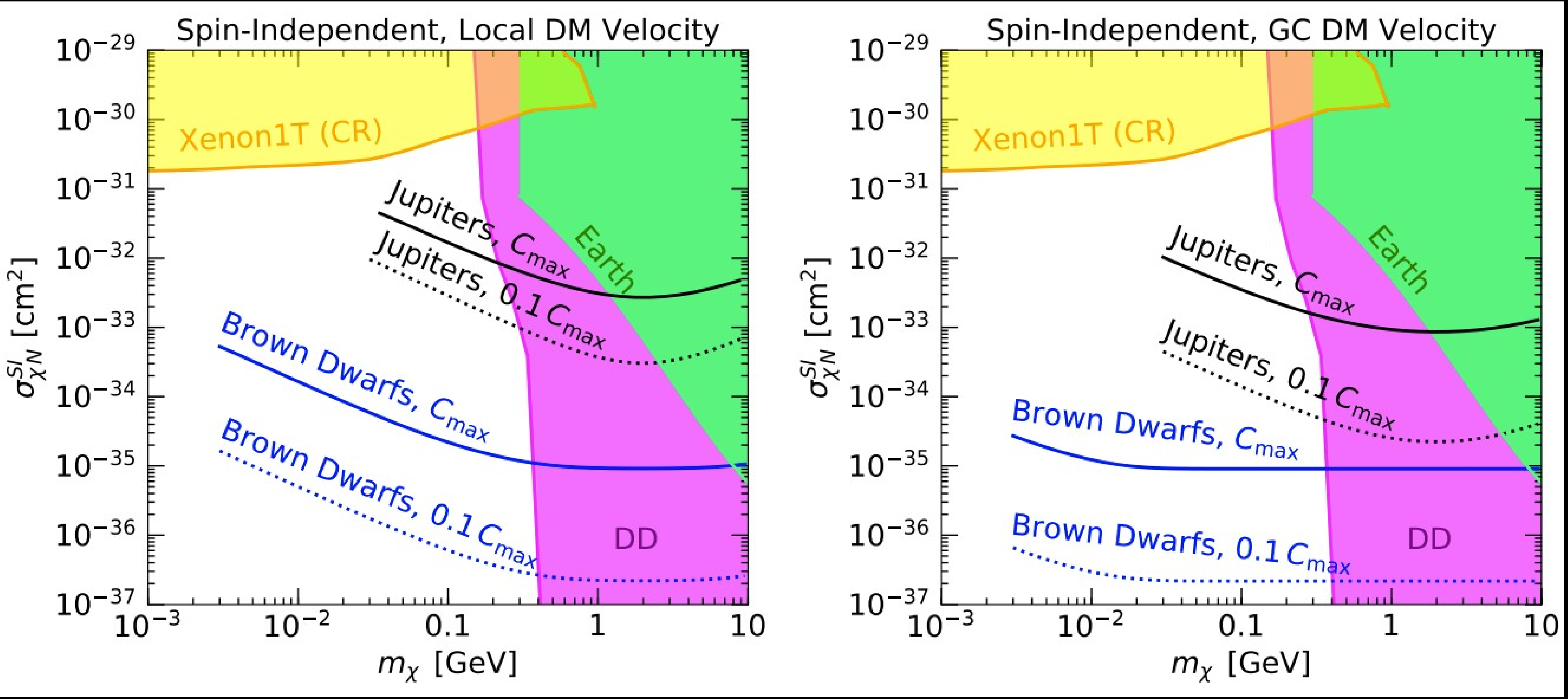
Higher masses:
Strongest signal towards Galactic Center, local DM heating signal difficult to outperform internal heat

Prospects for these searches?

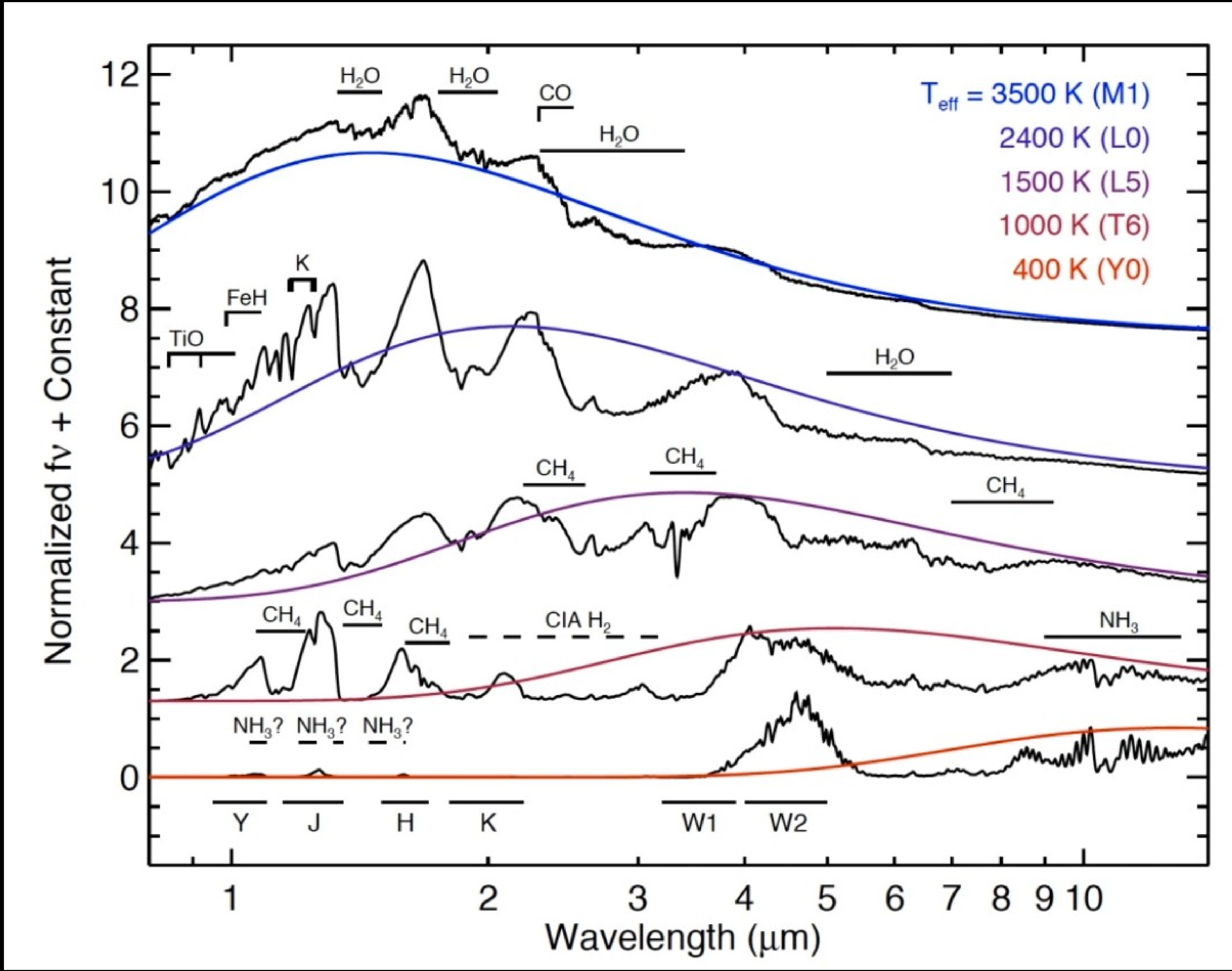
Planet	Radius ($R_{\rm jup}$)	Mass ($M_{\rm jup}$)	Distance	Orbit	Temp (No DM)	Temp (with DM)	Ref
Epsilon Eridani b	1.21	1.55	3 pc	3.4 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[84]
Epsilon Indi A b	1.17	3.25	3.7 pc	11.6 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[85]
Gliese 832 b	1.25	0.68	4.9 pc	3.6 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[86]
Gliese 849 b	1.23	1.0	8.8 pc	2.4 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[87]
Thestias	1.19	2.3	10 pc	1.6 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[88]
Lipperhey	1.16	3.9	12.5 pc	5.5 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[89]
HD 147513 b	1.22	1.21	12.8 pc	1.3 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[90]
Gamma Cephei b	1.2	1.85	13.5 pc	2.0 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[91]
Majriti	1.16	4.1	13.5 pc	2.5 au	$\sim$ 218 K	$\lesssim$ 650 K	[92]
47 Ursae Majoris d	1.2	1.64	14 pc	11.6 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[93]
Taphao Thong	1.2	2.5	14 pc	2.1 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[93]
Gliese 777 b	1.21	1.54	15.9 pc	4.0 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[94]
Gliese 317 c	1.21	1.54	15.0 pc	25.0 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[95]
q^1 Eridani b	1.23	0.94	17.5 pc	2.0 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[87]
HD 87883 b	1.21	1.54	18.4 pc	3.6 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[96]
ν^2 Canis Majoris c	1.24	0.87	19.9 pc	2.2 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[97]
$\rm Psi^1$ Draconis B b	1.21	1.53	22.0 pc	4.4 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[98]
HD 70642 b	1.19	1.99	29.4 pc	3.3 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[99]
HD 29021 b	1.2	2.4	31 pc	2.3 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[100]
HD 117207 b	1.2	1.9	32.5 pc	4.1 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[101]
Xolotlan	1.2	0.9	34.0 pc	1.7 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[102]
HAT-P-11 c	1.2	1.6	38.0 pc	4.1 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[103]
HD 187123 c	1.2	2.0	46.0 pc	4.9 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[104]
HD 50499 b	1.2	1.6	46.3 pc	3.8 au	$\lesssim$ 200 K	$\lesssim$ 650 K	[101]
Pirx	1.2	1.1	49.4 pc	0.8 au	$\sim$ 200 K	$\lesssim$ 650 K	[105]

- Many candidates already exist!
- Gaia may be able to see up to around 90,000 planets within 100 pc (local search)
- WFIRST/Roman expects to detect least several thousand exoplanets in the inner galaxy

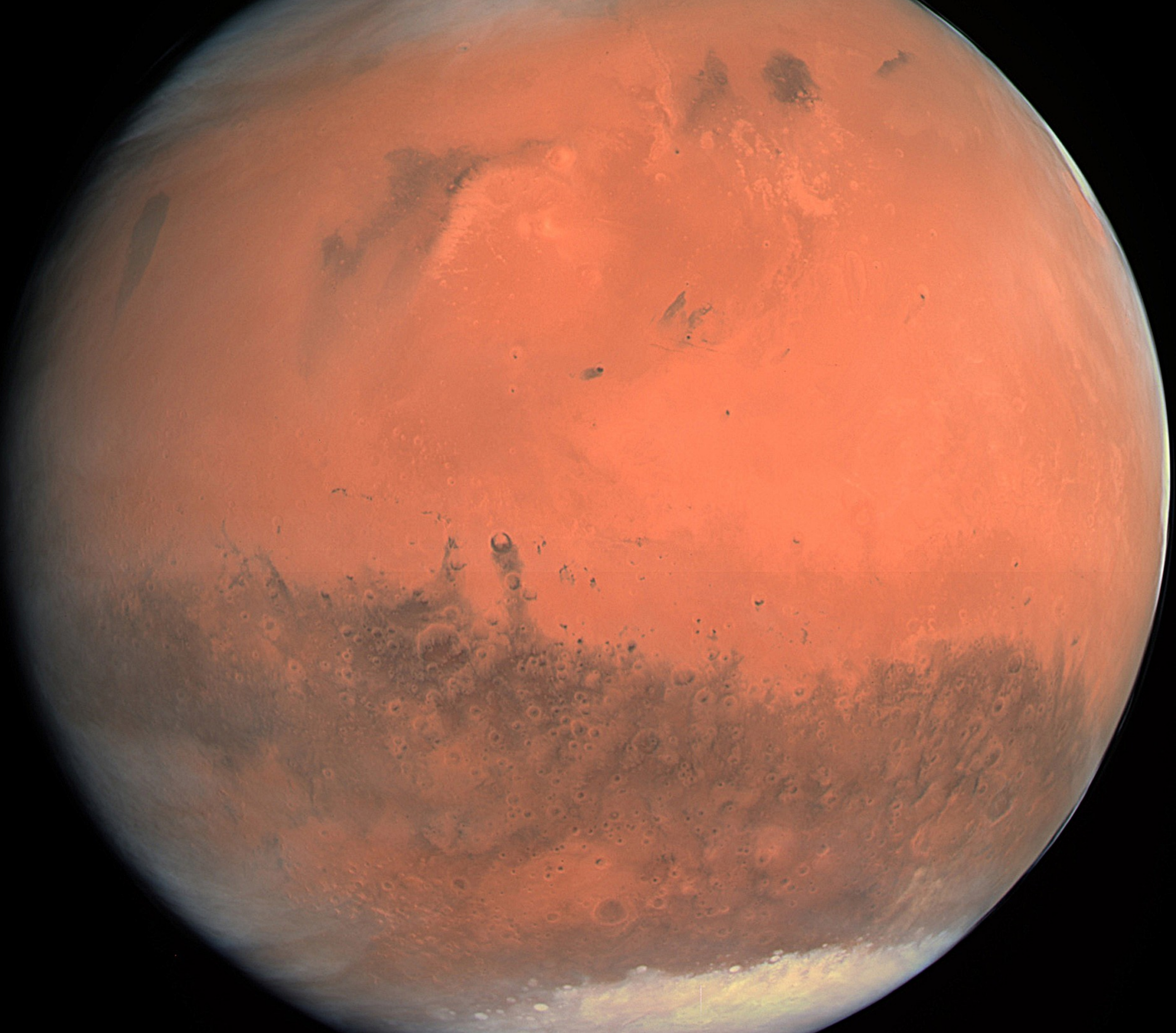
DM scattering cross section sensitivity



RKL + Smirnov, 2020



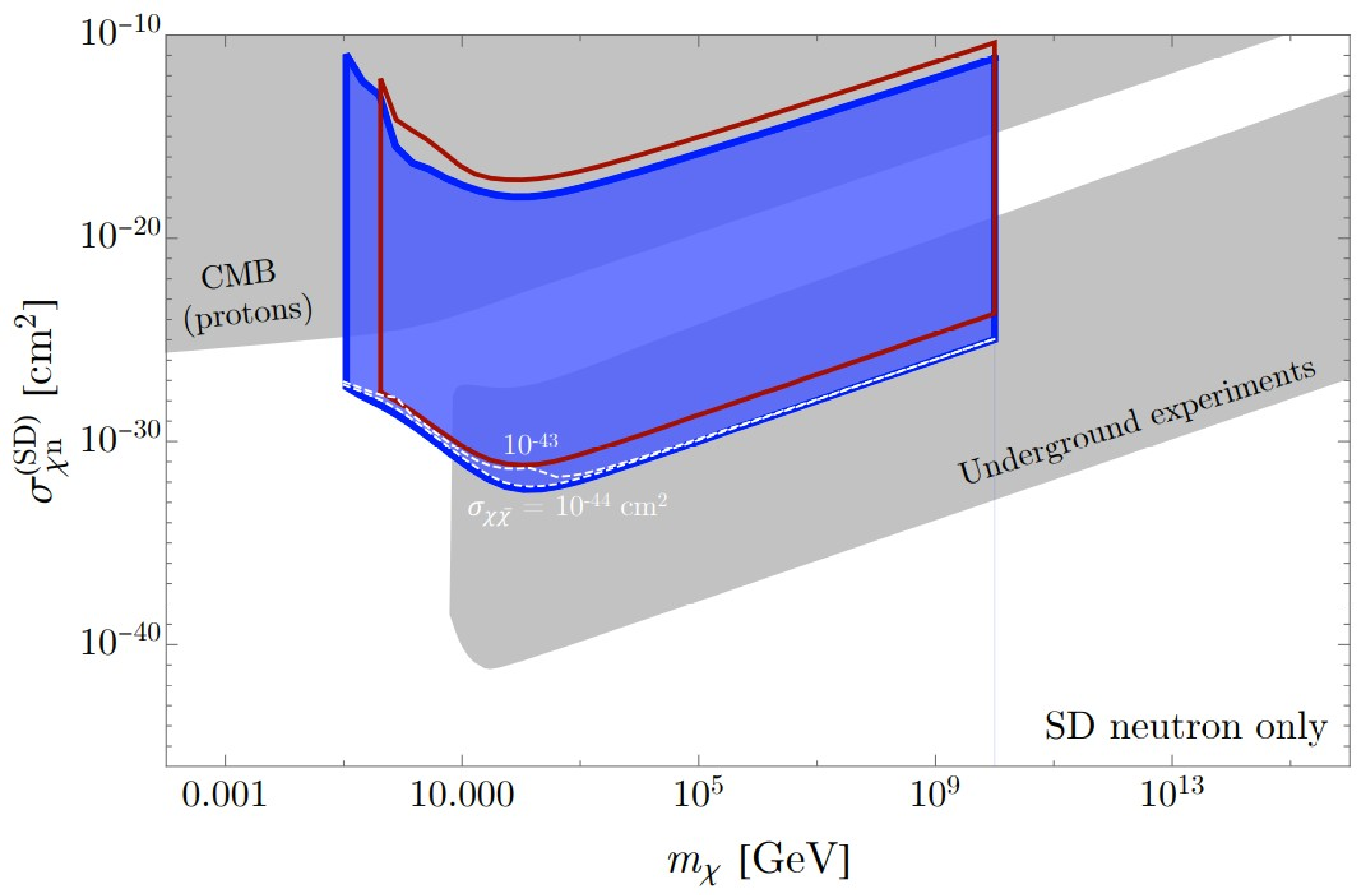
T_{eff} = 3500 K (M1)
2400 K (L0)
1500 K (L5)
1000 K (T6)
400 K (Y0)
H_2O
H_2O
CO
H_2O
TiO
FeH
K
H_2O
CH_4
CH_4
CH_4
CH_4
CH_4
CH_4
CIA H_2
NH_3
NH_3?
NH_3?
NH_3?
Y
J
H
K
W1
W2
Normalized fν + Constant
Wavelength (μm)



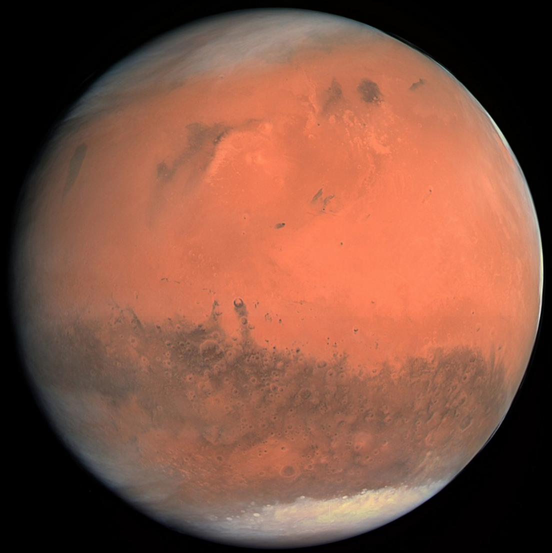
MARS

Bramante, Buchanan,
Goodman, Lodhi 2019

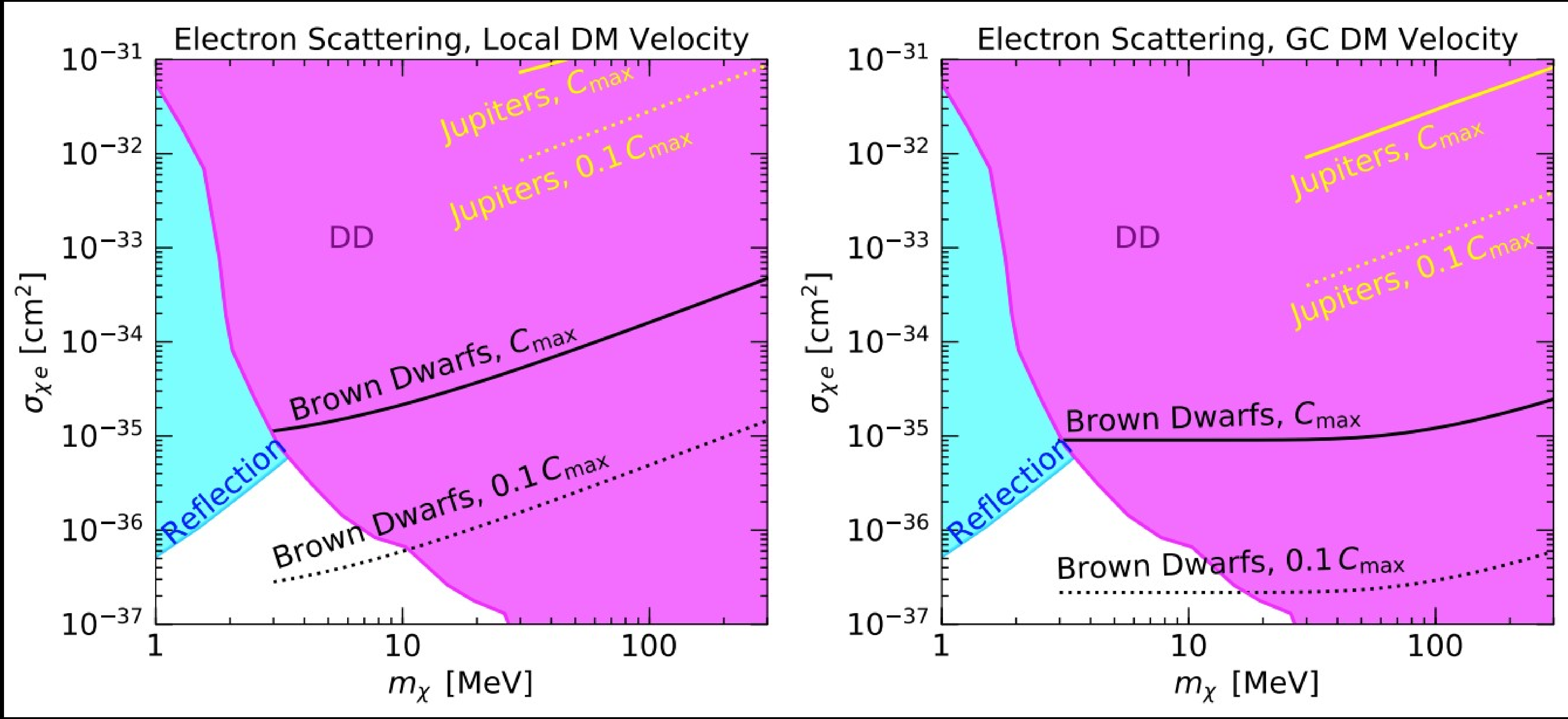
MARS



Bramante, Buchanan, Goodman, Lodhi
1909.11683



DM scattering cross section sensitivity



RKL + Smirnov, 2020

Calculating Exoplanet Temperatures

- Contributions to temperature from external heat (i.e. nearby stars), internal heat (e.g. from formation or burning processes), and dark matter:

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Heat power from DM:

- DM density throughout Galaxy:

$$\rho_\chi(r) = \frac{\rho_0}{(r/r_s)^\gamma (1 + (r/r_s))^{3-\gamma}}$$

- Relevant velocities:
  - DM halo velocity
  - Exoplanet escape velocity

$$v_{\rm esc}^2 = 2 G_N M / R$$

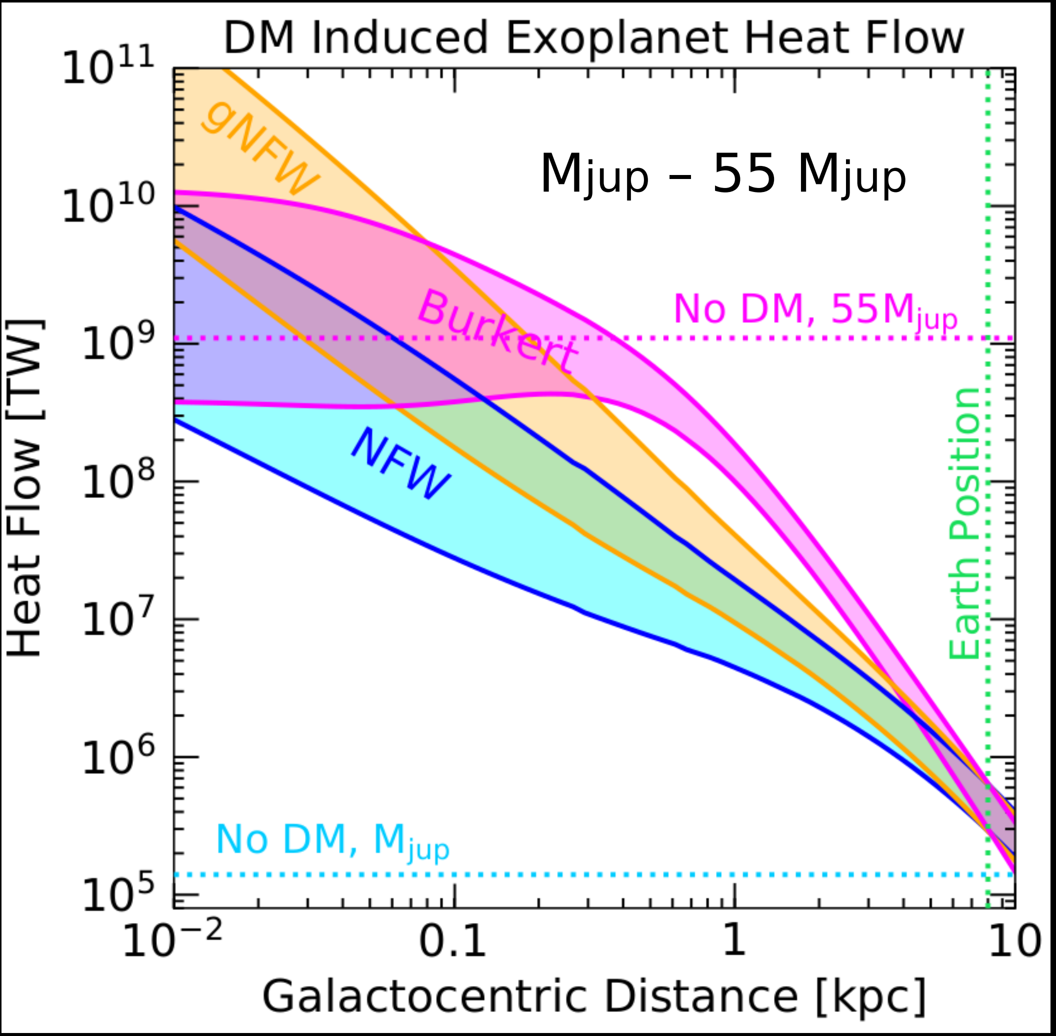
$$\Gamma_{\rm heat}^{\rm DM} = f \, \pi R^2 \rho_\chi(r) \, v_0 \left(1 + \frac{3}{2} \frac{v_{\rm esc}^2}{v_d(r)^2}\right)$$

DM Heating vs Internal Heat

RKL + Smirnov, 2020

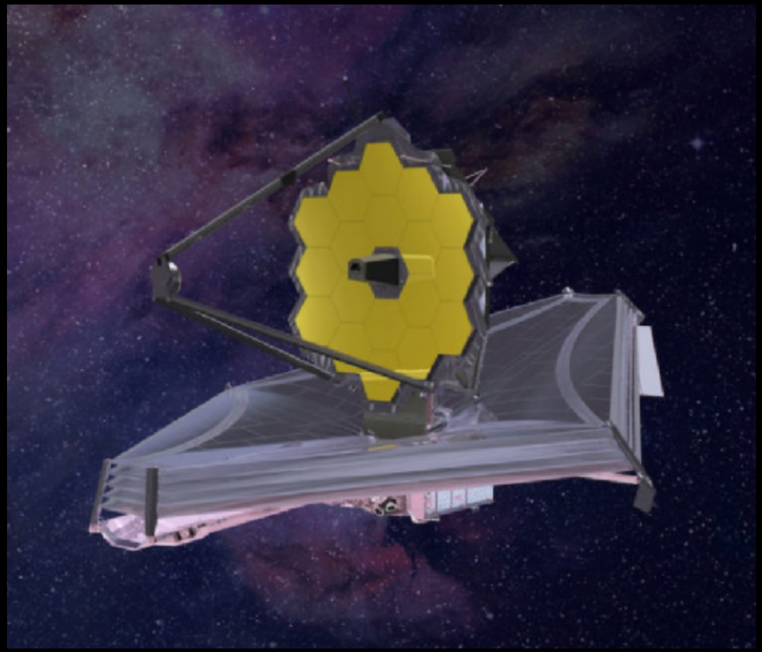
$$\Gamma^{\rm tot}_{\rm heat} = \Gamma^{\rm ext}_{\rm heat} + \Gamma^{\rm int}_{\rm heat} + \Gamma^{\rm DM}_{\rm heat} = 4\pi R^2 \, \sigma_{\rm SB} \, T^4 \, \epsilon$$

$$\Gamma^{\rm DM}_{\rm heat} = f \, \pi R^2 \rho_\chi(r) \, v_0 \left(1 + \frac{3}{2} \frac{v_{\rm esc}^2}{v_d(r)^2}\right)$$

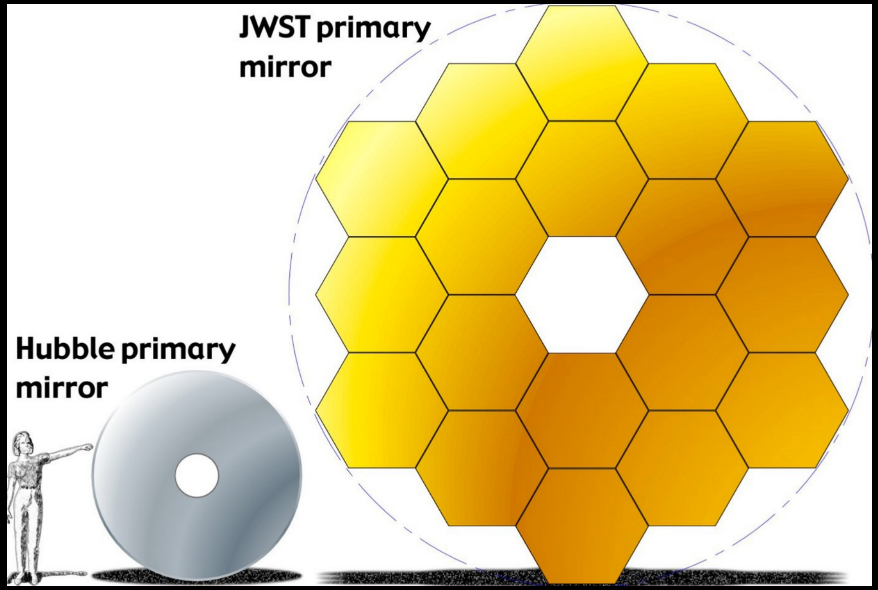


1 parsec = 3.26 light years

Telescope Sensitivity



- Use James Webb Space Telescope (planned launch Oct 2021)
- Infrared sensitivity (~0.5 – 28 microns)
- Has many instruments and filters, relevant choice for maximum sensitivity depends on peak wavelength

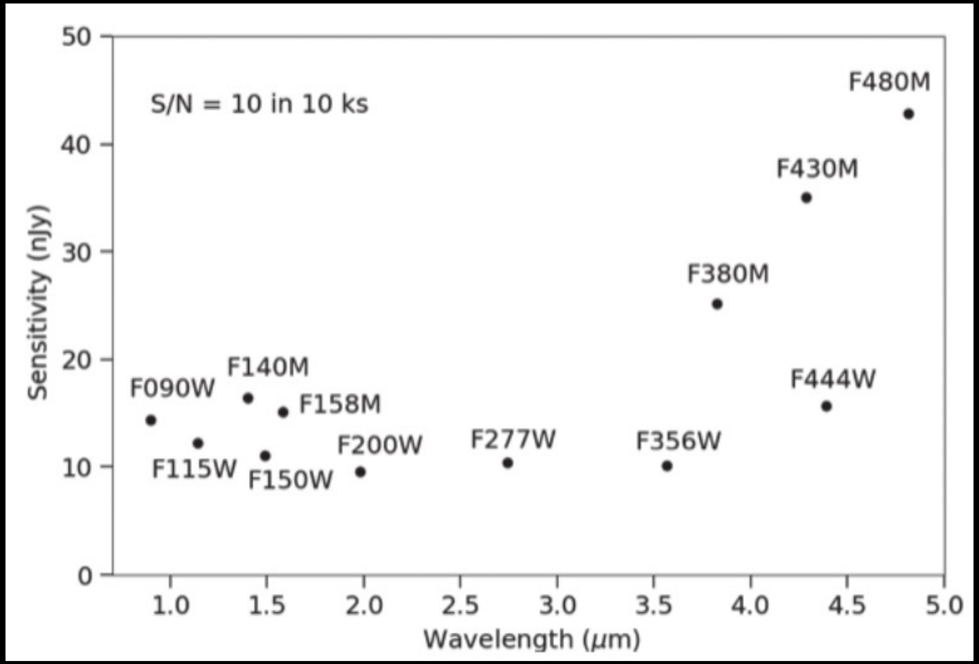


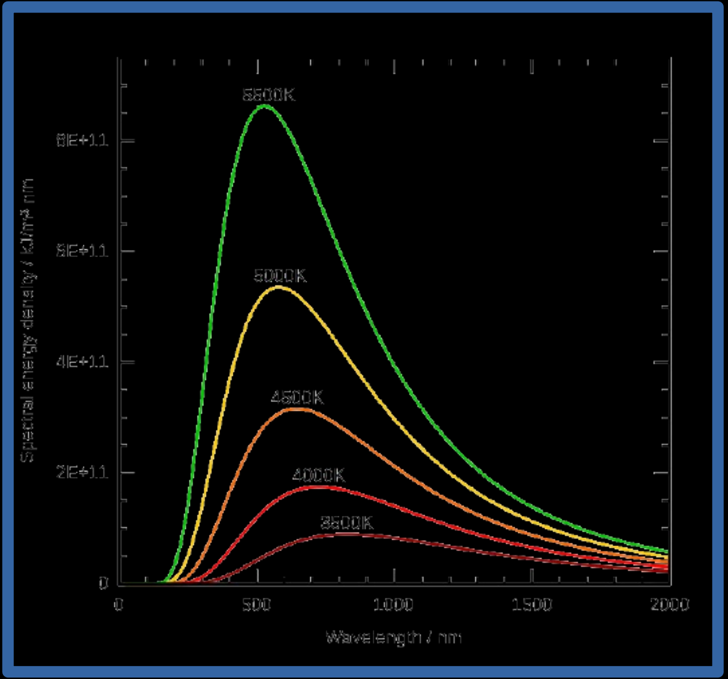
Signal with James Webb

- Can see many stars/planets at once
- Assume exoplanets radiate as a blackbody
  - Assume peak of blackbody temperature sets the sensitivity limit

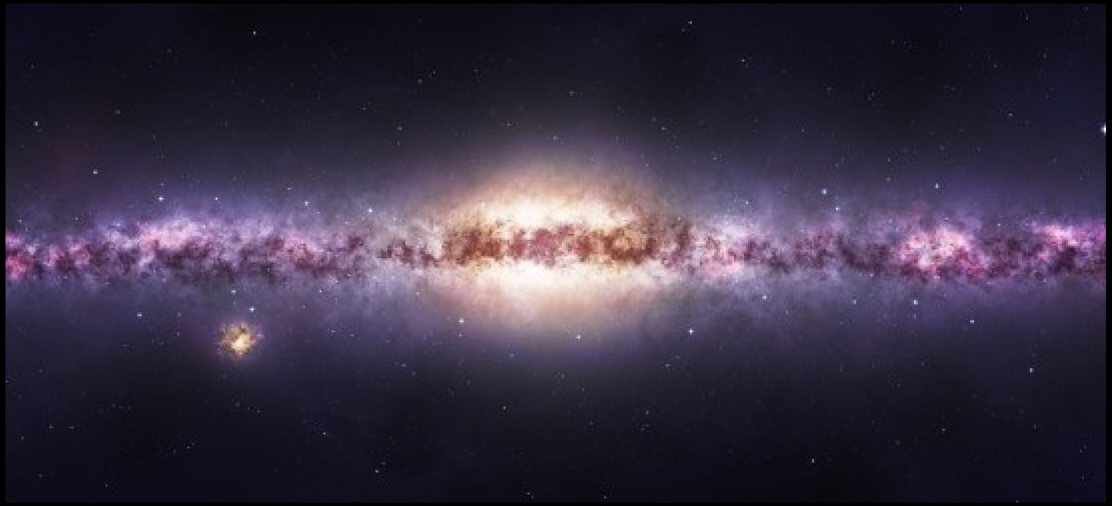
- Near-Infrared Imager and Slitless Spectrometer (NIRISS) for T > 500 K
- Mid-Infrared Instrument (MIRI) for T = 100 - 500 K

Won’t need new dedicated searches; can piggyback





Search Challenges

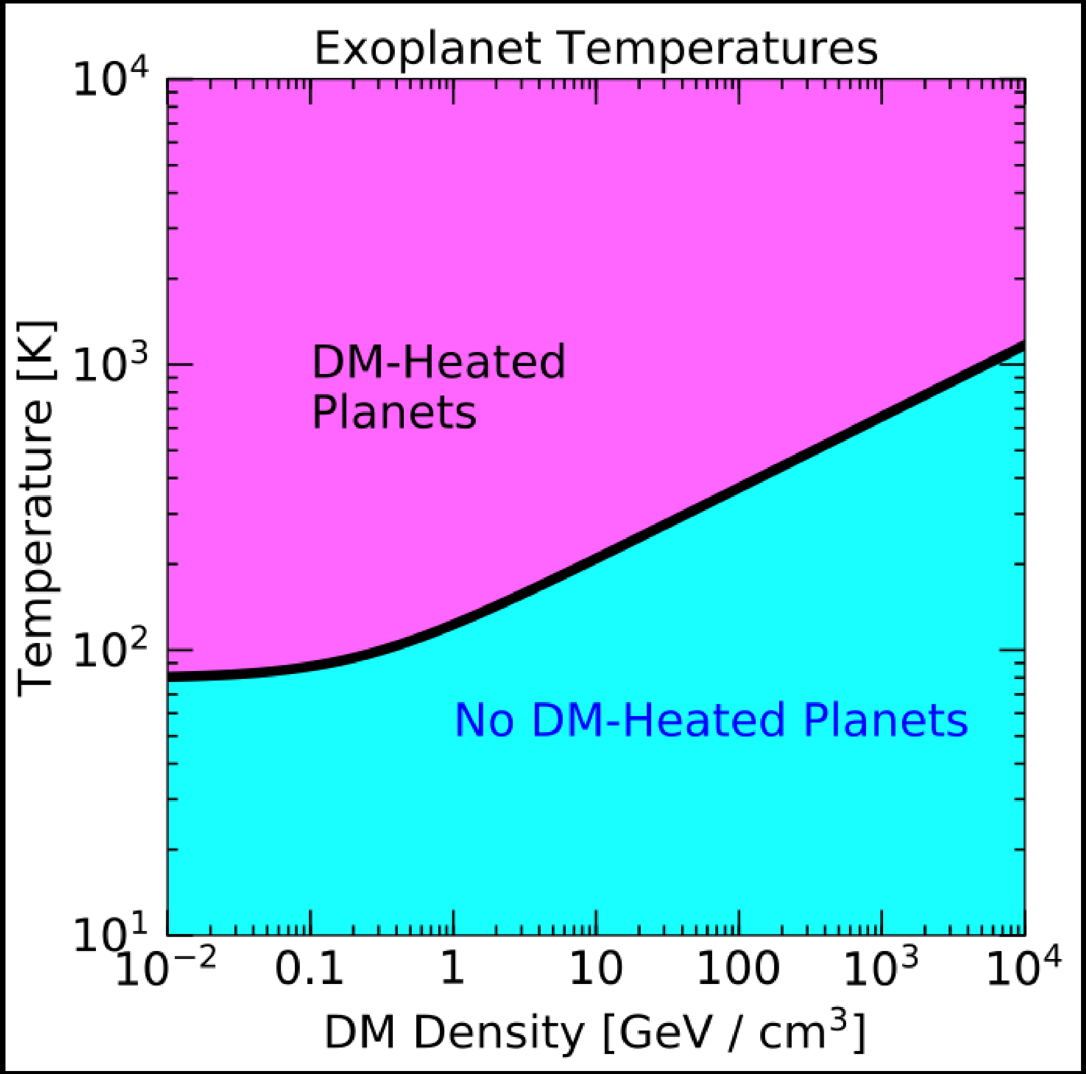


Dust backgrounds:
Rescatter some wavelengths, which can reduce intensity and shift spectrum peaks

Stellar crowding:
Stars per pixel important, can outshine exoplanet signal

Optimal sensitivity is outside 0.1 kpc (about 1 degree off the plane)

Deviations: DM-overdensities

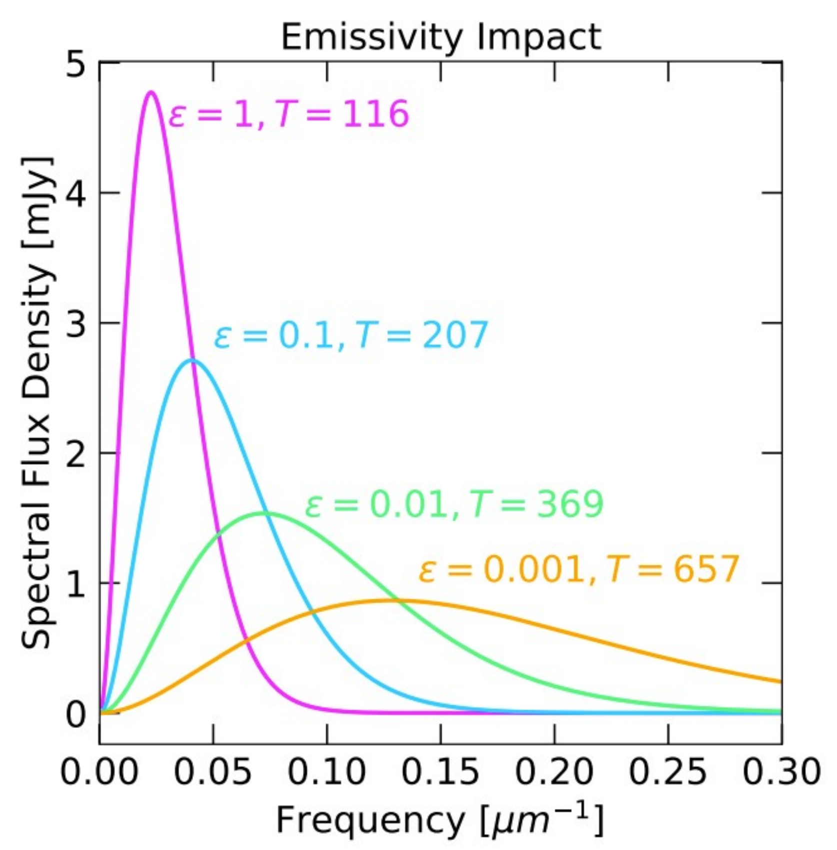


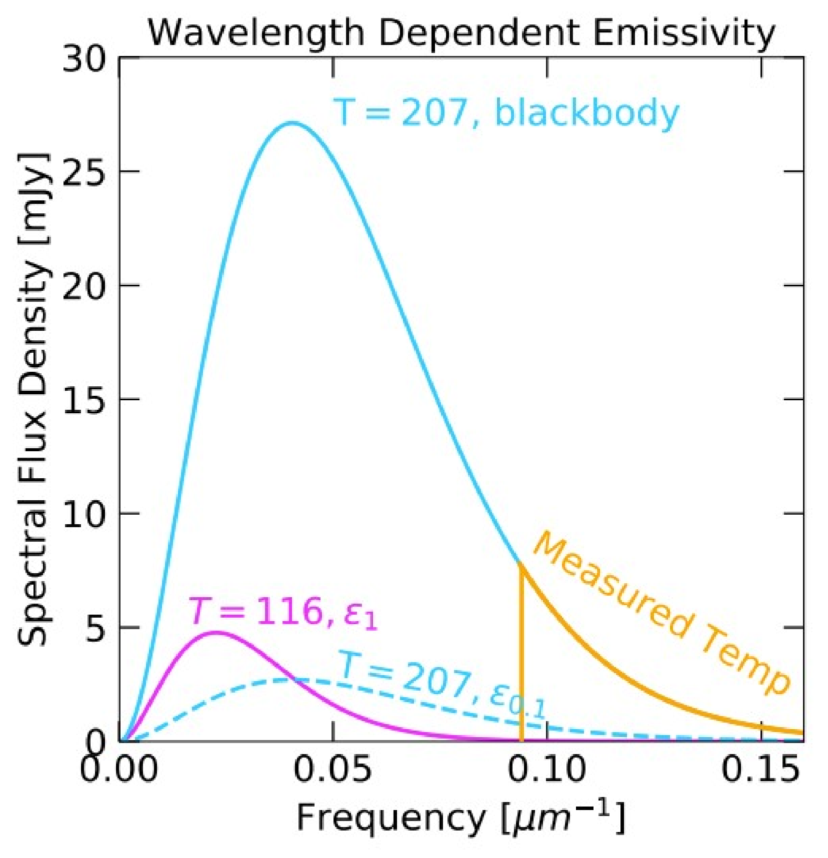
Rebecca Leane (SLAC)

Deviations: Non-Blackbody Spectra

Atmosphere effects can cause deviations from a blackbody

$$B(\nu, T) = \frac{2\nu^3 \epsilon}{\exp\left(\frac{2\pi\nu}{\mathrm{k_b T}}\right) - 1}$$

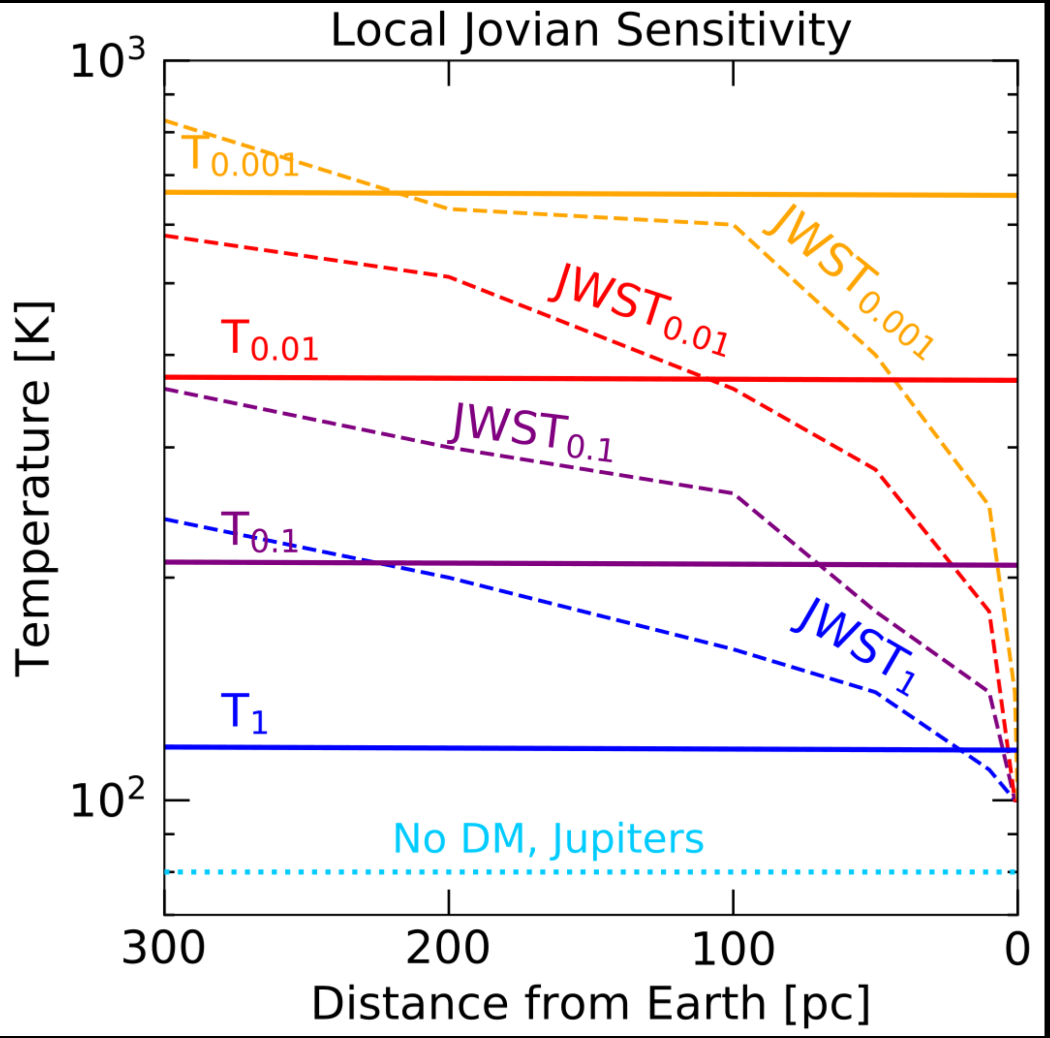




Rebecca Leane (SLAC)

Local DM-Heated Exoplanet Search

- Local fluxes easier to detect, so lower normalization from emissivity isn't a severe penalty
- Allows use of more powerful filters: best JWST filter sensitivity is with higher temps (in this case, higher wavelength peaks)
- Local exoplanets with lower emissivities can extend local sensitivity to DM heating



DM scattering cross section sensitivity

$$f = \frac{C_{\text{cap}}}{C_{\text{max}}} = \sum_{N=1}^{\infty} f_N$$

$$f_N = p(N,\tau)\left[1 - \kappa \exp\left(-\frac{3\left(v_{\text{N}}^2 - v_{\text{esc}}^2\right)}{2v_d^2}\right)\right]$$

$$\kappa = \left(1 + \frac{3}{2}\frac{v_{\text{N}}^2}{v_d^2}\right)\left(1 + \frac{3}{2}\frac{v_{\text{esc}}^2}{v_d^2}\right)^{-1}$$

Here v_d is the velocity dispersion, $v_N = v_{\text{esc}}\left(1 - \langle z\rangle\beta\right)^{-N/2}$ where the average scattering angle is $\langle z\rangle = 1/2$ [143], $\beta = 4m_\chi m_A/(m_\chi + m_A)^2$, and m_A is the mass of the target particle. The probability that the DM particle scatters N times is

$$p(N,\tau) = \frac{2}{\tau^2}\left(N_s + 1 - \frac{\Gamma(N_s + 2, \tau)}{N_s!}\right)$$

$$\tau = \frac{3}{2}\frac{\sigma}{\sigma_{\text{sat}}}$$

$$\sigma_{\text{sat}} = \pi R^2 / N_{\text{SM}}$$

$$\sigma_{\chi A}^{\text{SD}} = \sigma_{\chi N}^{\text{SD}}\left(\frac{\mu(m_A)}{\mu(m_N)}\right)^2 \frac{4(J+1)}{3J}\left[a_p\langle S_p\rangle + a_n\langle S_n\rangle\right]^2$$

$$\sigma_{\chi A}^{\text{SI}} = \sigma_{\chi N}^{\text{SI}}\left(\frac{\mu(m_A)}{\mu(m_N)}\right)^2\left[Z + \frac{a_n}{a_p}(A - Z)\right]^2$$

AGE – COOLING CURVES