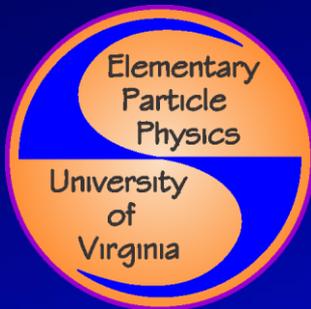
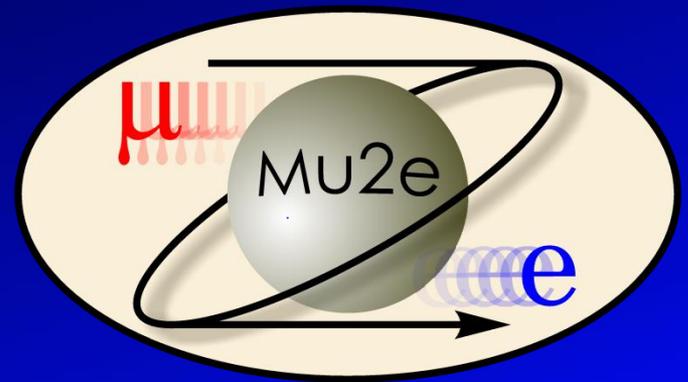


Mu2e: Probing the Frontiers of Physics Using Rare Particle Decays



E. Craig Dukes
University of Virginia
for the Mu2e Collaboration

Lake Louise Winter Institute
February 19, 2015



Planning the Future of Particle Physics

Two summers ago about a 1,000 physicists gathered together in Minneapolis for two weeks to plan the future of particle physics in the United States for the next decade.



*“Prediction is very hard,
especially about the future”*

Yogi Berra



SNOWMASS CSS 2013 ON THE MISSISSIPPI JULY 29 – AUGUST 6, 2013

ORGANIZED BY THE DIVISION OF PARTICLES AND FIELDS OF THE APS
HOSTED BY THE UNIVERSITY OF MINNESOTA

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Energy Frontier
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Michael Peskin (SLAC)
Intensity Frontier
JoAnne Hewett (SLAC),
Harry Womersley (Argonne)
Cosmic Frontier
Jonathan Feng (University of California, Irvine),
Steve Ritz (University of California, Santa Cruz)
Frontier Capabilities
William Barletta (MIT),
Humbold Gikhovene (LBNL)
Instrumentation Frontier
Marcel Demarteau (Argonne),
Howard Hitchon (Mt. Holyoke),
Ron Lipton (Fermilab)
Computing Frontier
Lothar Bauendick (Fermilab),
Steven Gottlieb (Indiana)
Education and Outreach
Marge Burdick (Fermilab),
Dan Cronin-Hennessy (Minnesota)
Theory Panel
Michael Dine (University of California, Santa Cruz)

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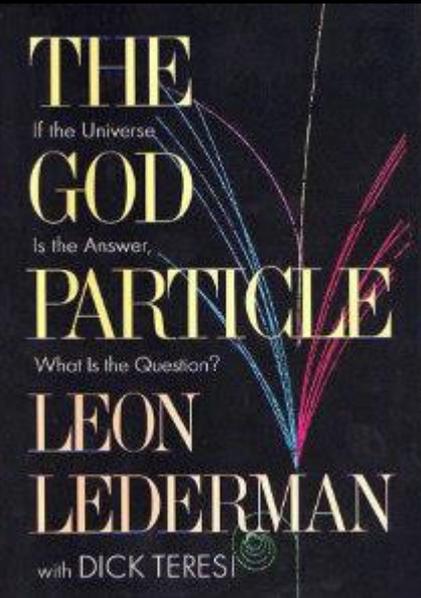


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POSTER DESIGN BY KATE SCHALLOW

Context of Snowmass was Two Profound Discoveries



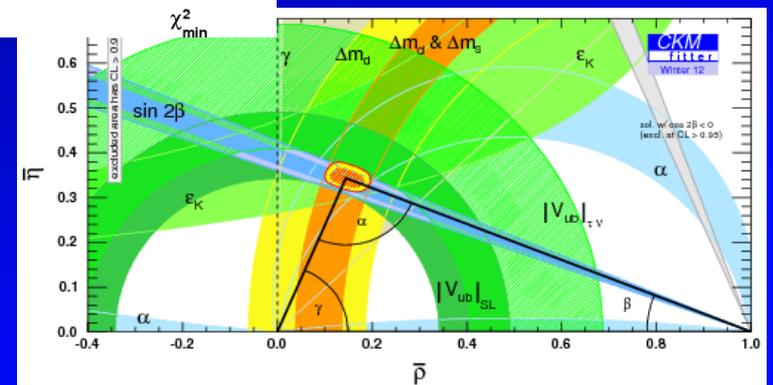
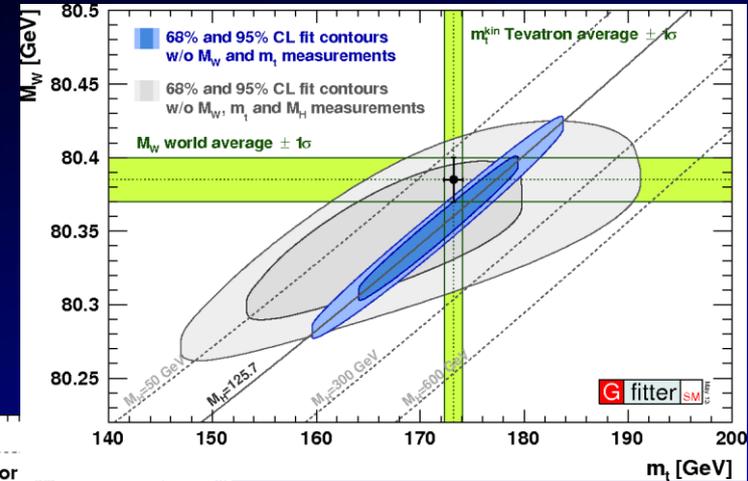
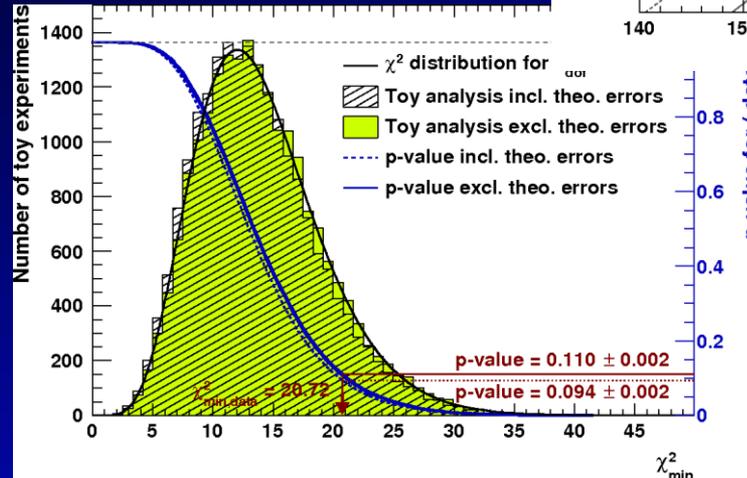
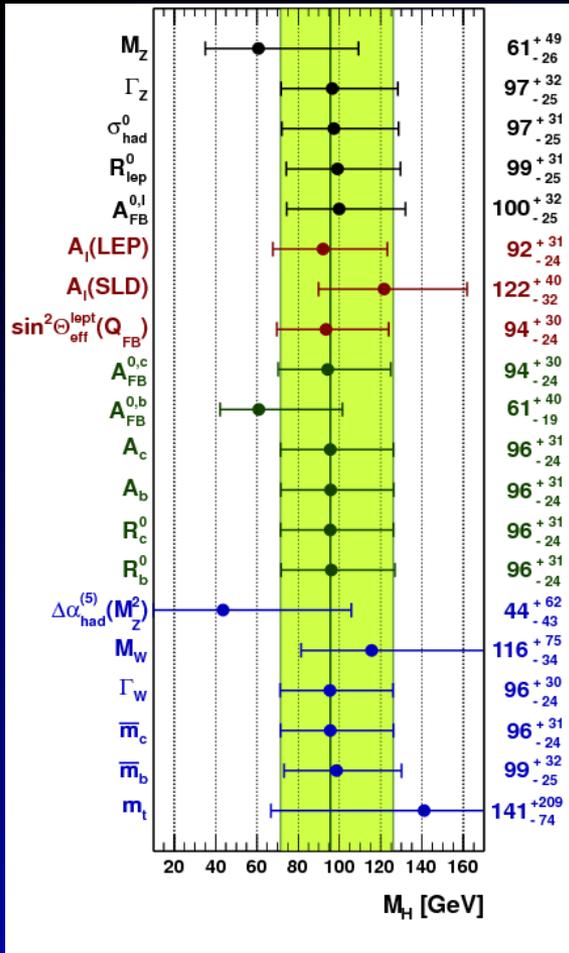
The discovery of the Higgs boson by the LHC
The particle that gives mass to all of the fundamental particles.

The Standard Model of particle physics is now complete.



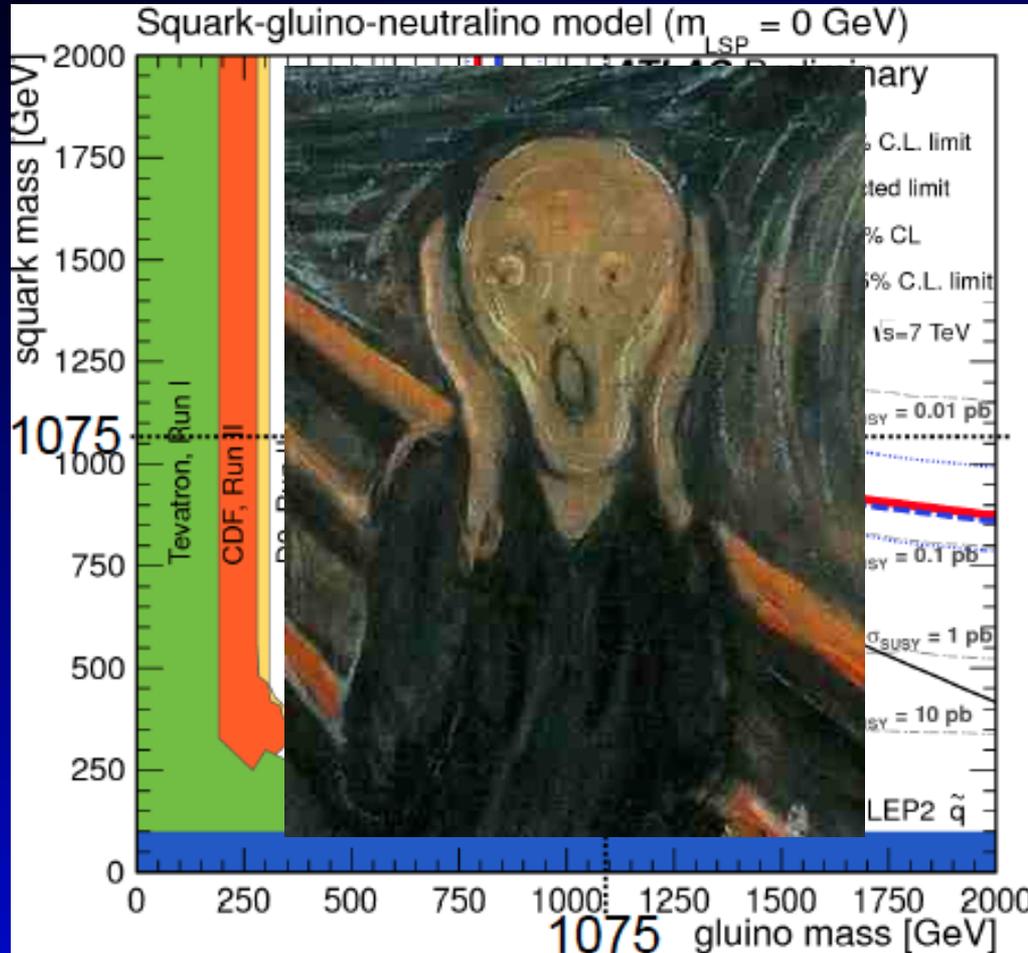
CERN, July 4, 2012

Standard Model and Data Agree Exquisitely Well!



Second Profound Discovery

The non-discovery of Supersymmetry by the LHC



Current State of Theoretical Physics



We have a Standard Model that explains (almost) everything we see in our accelerator-based particle experiments.

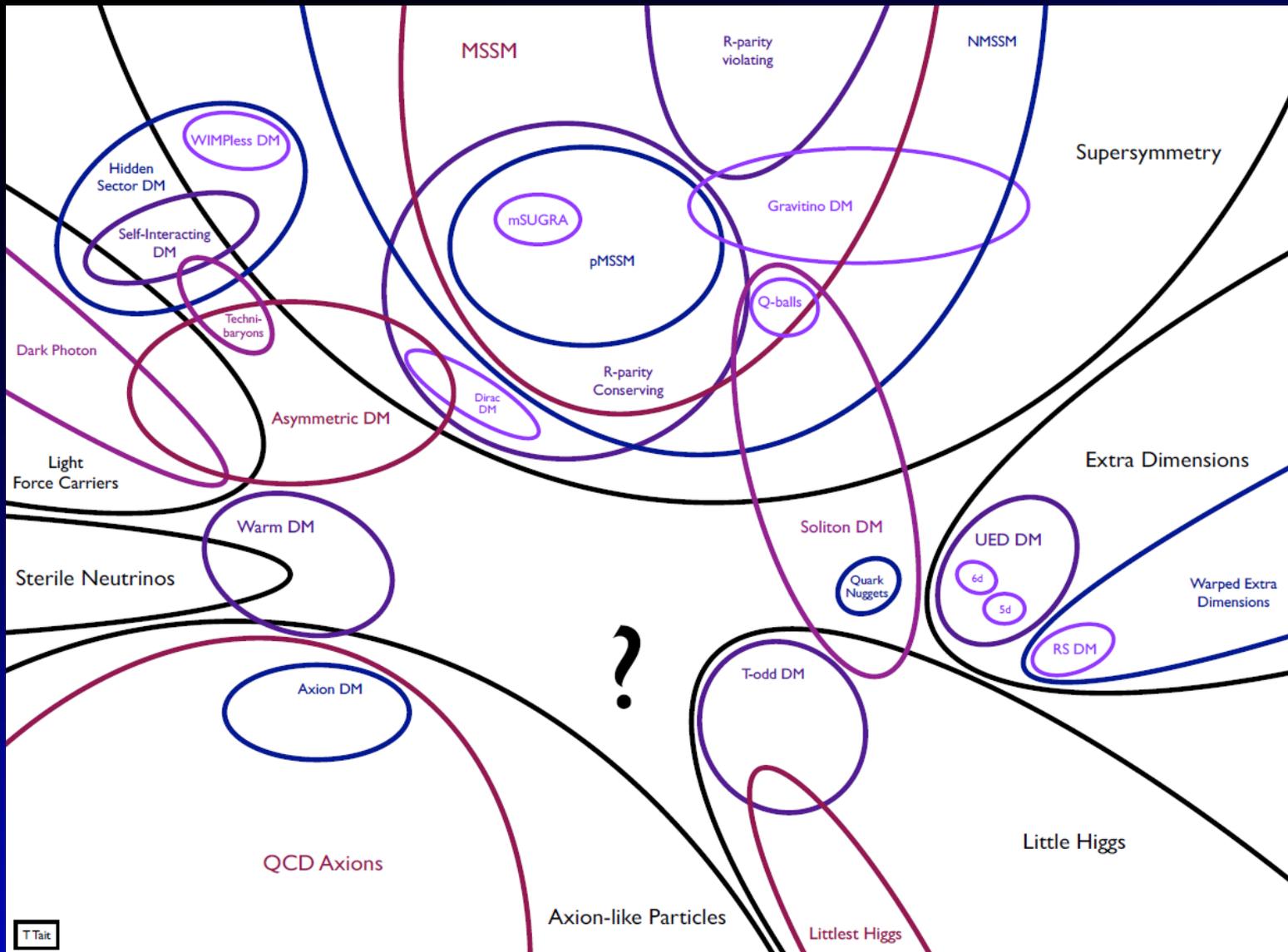
It does not explain very gross features we unambiguously observe in the universe: dark matter, dark energy, antimatter asymmetry.

“If you ask most theorists working on particle physics, they’re in a state of confusion.”

Neil Turok, Director Perimeter Institute

How do we proceed from here?

No Lack of Theoretical Ideas, but Little Guidance

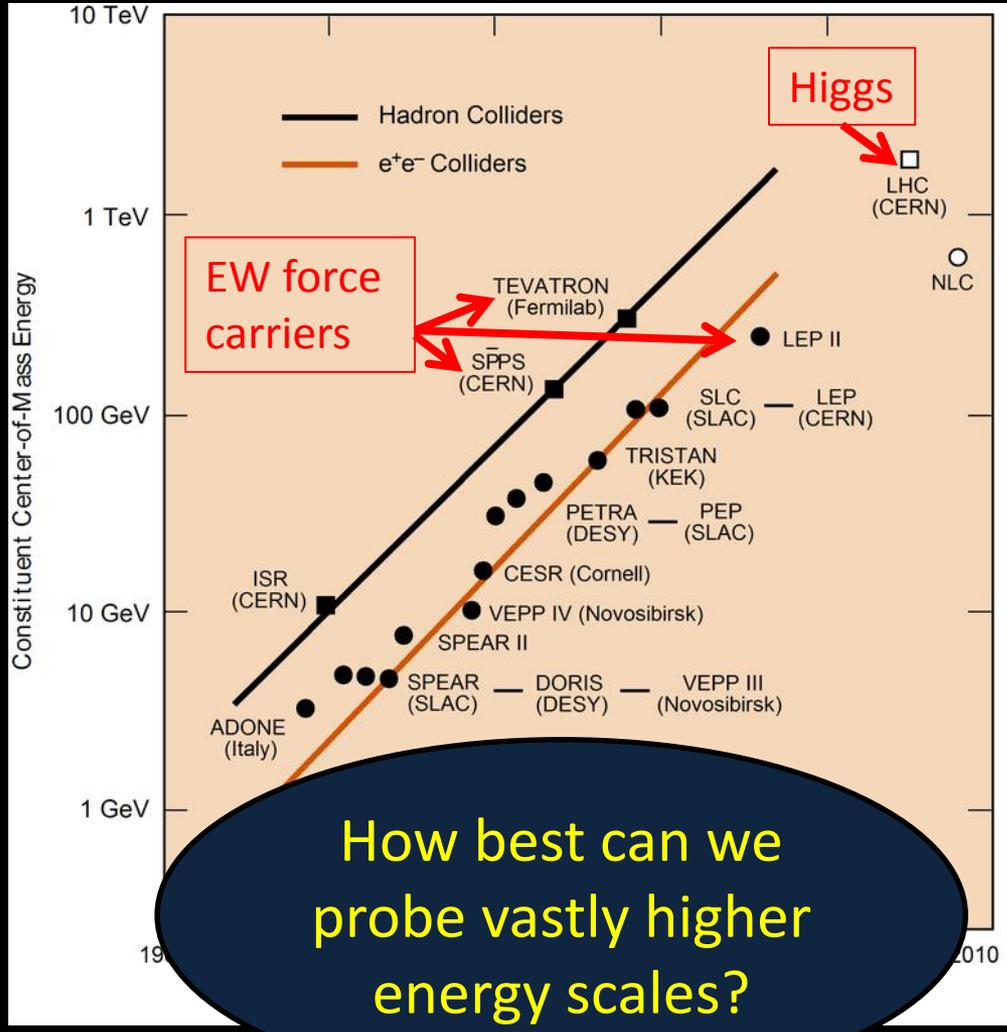


What Energy Accelerator is Needed to Probe New Physics?

We don't know!

There is no guarantee that a new accelerator 10X or 100X the LHC energy will see anything new: just confirm the Standard Model at a new energy scale.

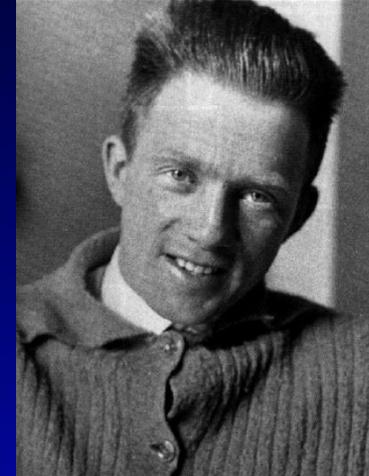
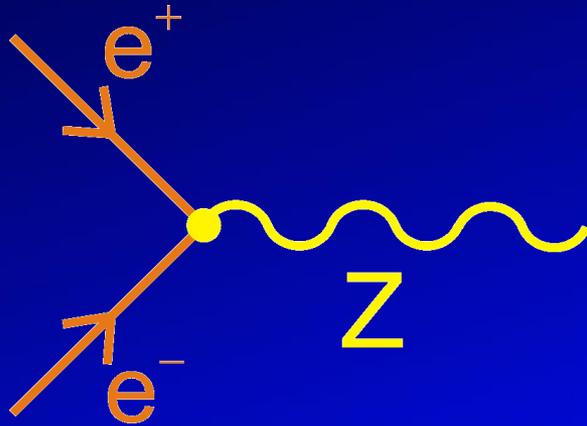
And accelerators are not cheap: a 10X the LHC energy accelerator would cost tens of billions of dollars.



Rather than exploiting Einstein's mass-energy relation, $E=mc^2$, exploit Heisenberg's uncertainty principle, $\Delta E\Delta t \geq \hbar/2$



Highest energy crucial



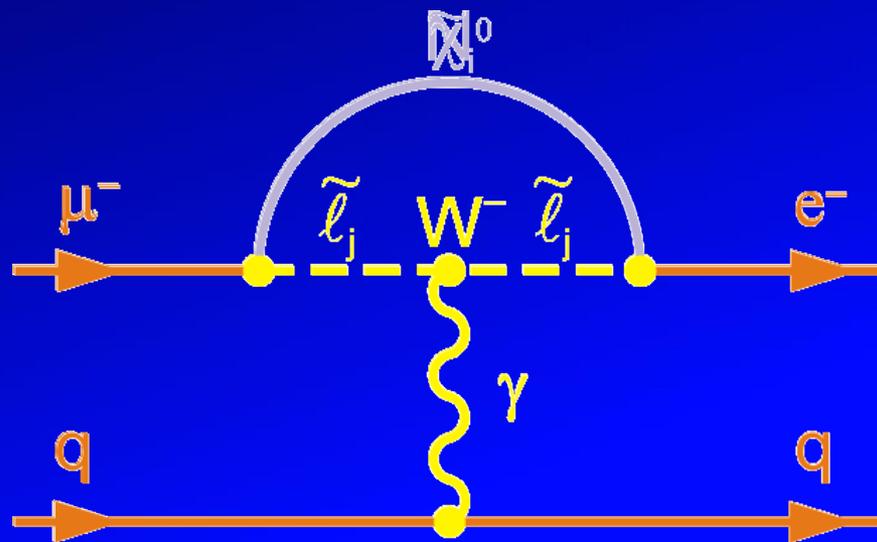
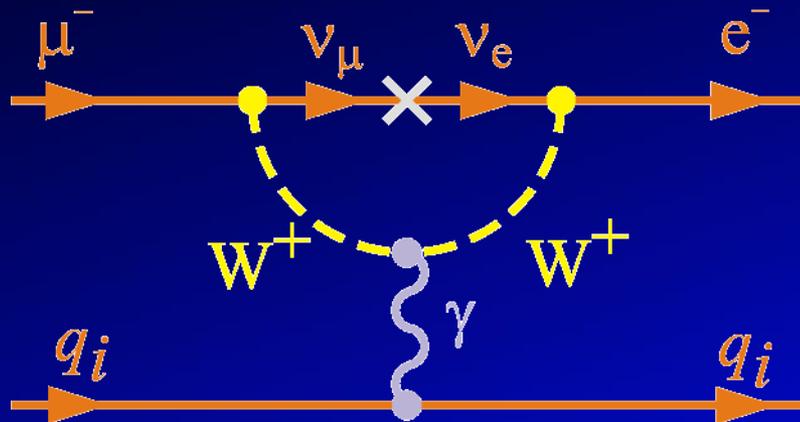
Highest intensities crucial



Mu2e: Probe New Physics Charged Lepton Flavor Violation

$$\mu^- N \rightarrow e^- N'$$

- In Standard Model not there \Rightarrow neutrino mass discovery implies an unobservable 10^{-52} rate
- Hence, any signal unambiguous evidence of new physics
- Exquisite sensitivities can be obtained experimentally
 - \Rightarrow sensitivities that allow favored beyond-the-standard-model theories to be tested



Supersymmetry



Why Muon-to-Electron Conversion?

Different SUSY and non-SUSY BSM models

★★★ Large effects

★★ Visible, but small

★ No sizable effect

	AC	RVV2	AKM	δ LL	FBMSSM	LHT	RS
$D^0 - \bar{D}^0$	★★★	★	★	★	★	★★★	?
ϵ_K	★	★★★	★★★	★	★	★★	★★★
$S_{\psi\phi}$	★★★	★★★	★★★	★	★	★★★	★★★
$S_{\phi K_S}$	★★★	★★	★	★★★	★★★	★	?
$A_{CP}(B \rightarrow X_s \gamma)$	★	★	★	★★★	★★★	★	?
$A_{7,8}(B \rightarrow K^* \mu^+ \mu^-)$	★	★	★	★★★	★★★	★★	?
$A_9(B \rightarrow K^* \mu^+ \mu^-)$	★	★	★	★	★	★	?
$B \rightarrow K^{(*)} \nu \bar{\nu}$	★	★	★	★	★	★	★
$B_s \rightarrow \mu^+ \mu^-$	★★★	★★★	★★★	★★★	★★★	★	★
$K^+ \rightarrow \pi^+ \nu \bar{\nu}$	★	★	★	★	★	★★★	★★★
$K_L \rightarrow \pi^0 \nu \bar{\nu}$	★	★	★	★	★	★★★	★★★
$\mu \rightarrow e \gamma$	★★★	★★★	★★★	★★★	★★★	★★★	★★★
$\tau \rightarrow \mu \gamma$	★★★	★★★	★	★★★	★★★	★★★	★★★
$\mu + N \rightarrow e + N$	★★★	★★★	★★★	★★★	★★★	★★★	★★★
d_n	★★★	★★★	★★★	★★	★★★	★	★★★
d_e	★★★	★★★	★★	★	★★★	★	★★★
$(g-2)_\mu$	★★★	★★★	★★	★★★	★★★	★	?

Table 8: “DNA” of flavour physics effects for the most interesting observables in a selection of SUSY and non-SUSY models ★★★ signals large effects, ★★ visible but small effects and ★ implies that the given model does not predict sizable effects in that observable.

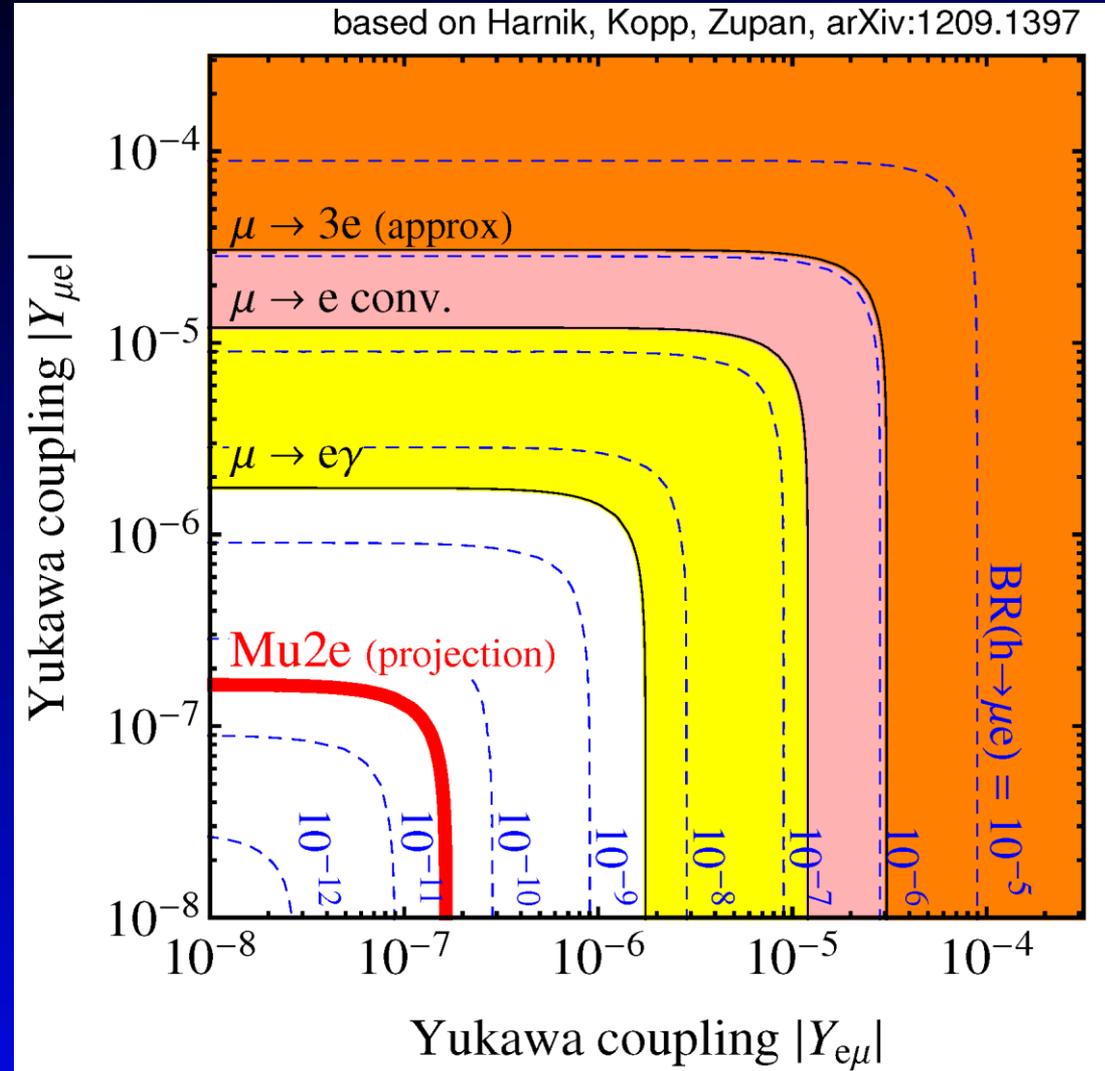
Altmannshofer et al.,
NPB 830, 17 (2010)

Example: Flavor-Violating Higgs Decays

Flavor violating Higgs couplings: $h \rightarrow \mu e$

Mu2e sensitivity:

$$\sqrt{|Y_{\mu e}|^2 + |Y_{e\mu}|^2} > \text{few} \times 10^{-7}$$

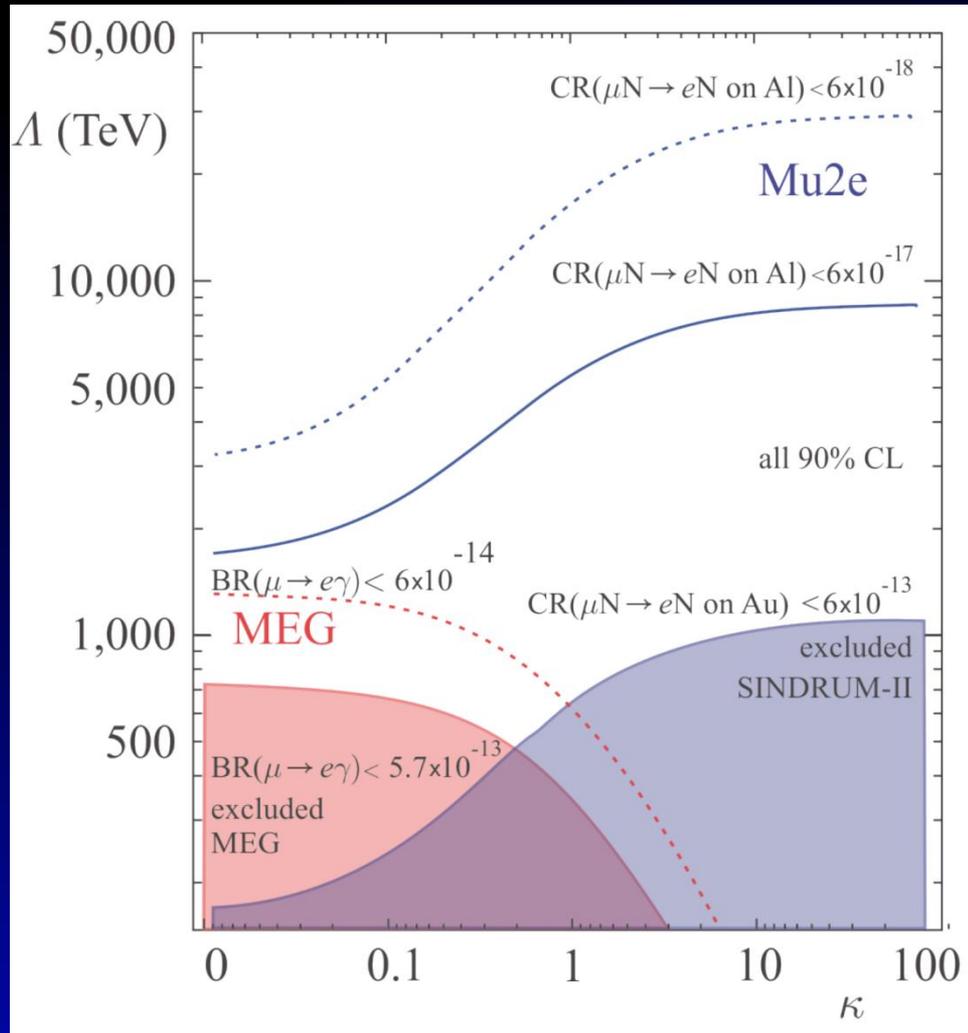


Mu2e Looking for an Exceedingly Rare Event

To achieve our proposed sensitivity of 3×10^{-17} , a factor of 10,000 better than present limit, the number of muons needed is $\sim 10^{19}$, roughly the number of grains of sand on Earth: $\sim 10^{19}$



Mu2e Will Indirectly Probe Extremely High Mass Scales



Up to 10,000 TeV
or about ~10,000 times
that probed directly by
the LHC

loop
dominated

contact
dominated

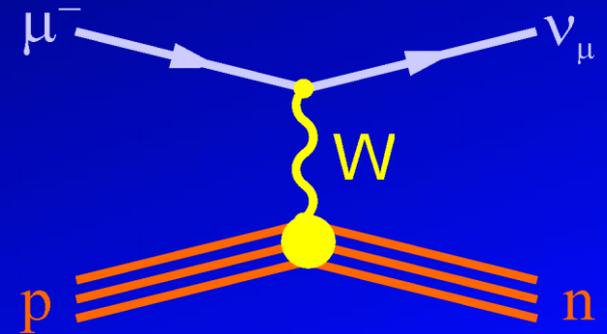
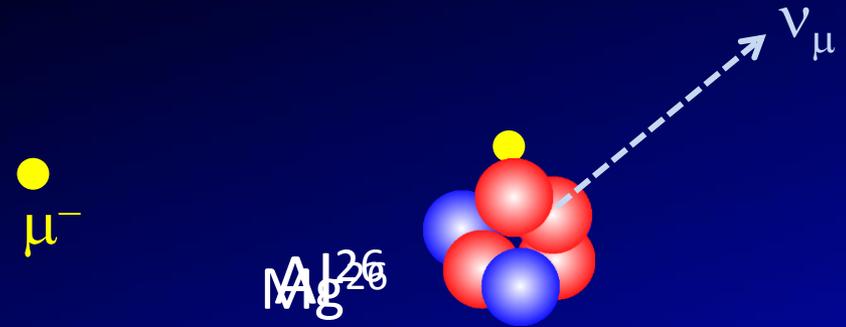
How to Search for $\mu^-N \rightarrow e^-N$

- Stop muon in atom
- Muon rapidly (10^{-16} s) cascades to 1S state
- Circles the nucleus for up to $\sim 2 \mu\text{s}$
- Two things most likely happen:

1. muon is captured by the nucleus:
 $\mu^- N_{A,Z} \rightarrow \nu_\mu N_{A,Z-1}$

2. muon decays in orbit:

$$\mu^- N_{A,Z} \rightarrow e^- \nu_\mu \nu_e N_{A,Z}$$



Muon Capture

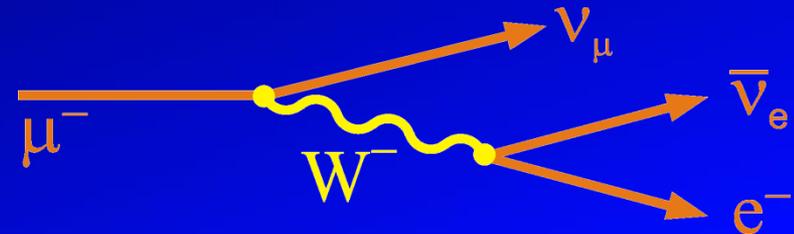
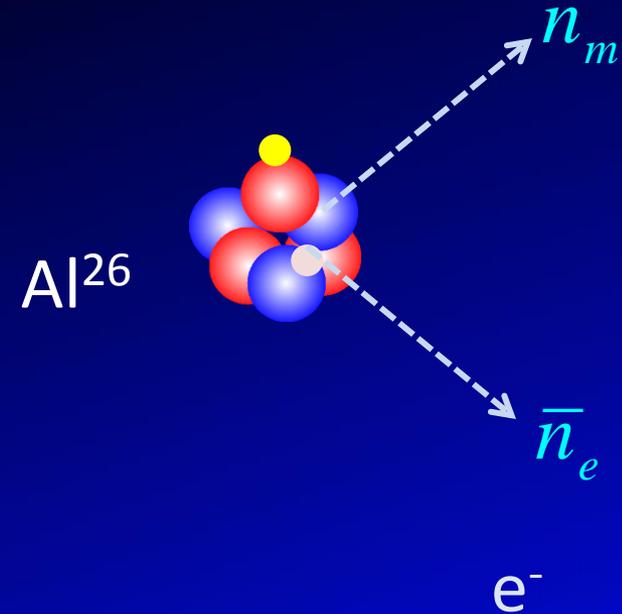
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2. muon decays in orbit:

$$\mu^- N_{A,Z} \rightarrow e^- \nu_\mu \bar{\nu}_e N_{A,Z}$$



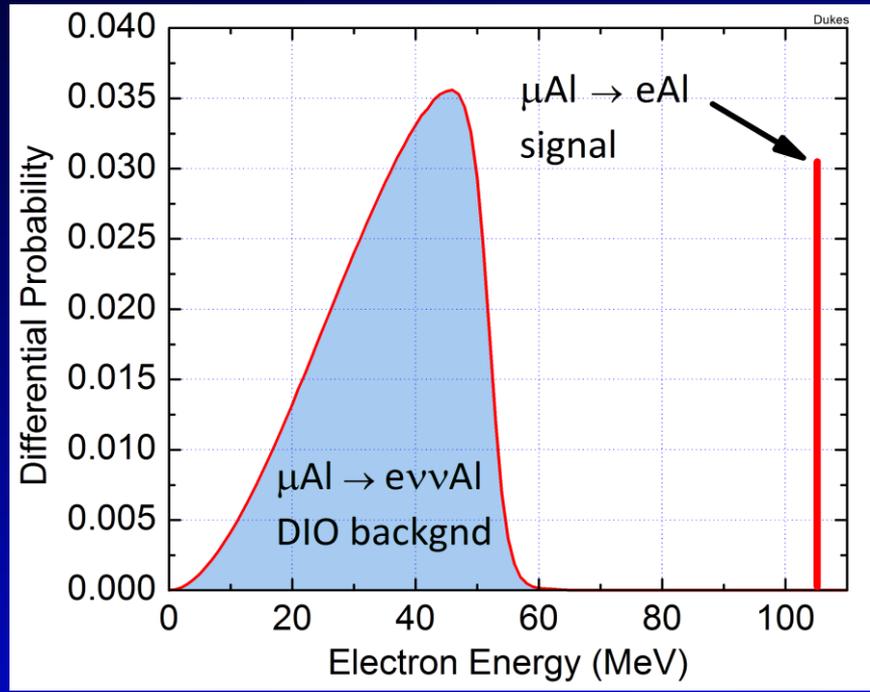
Muon Decay-in-Orbit

Mu2e Searching for a Third Process: $\mu^-N \rightarrow e^-N$

In $\mu^-N \rightarrow e^-N$ the muon coherently interacts with nucleus leaving it in ground state

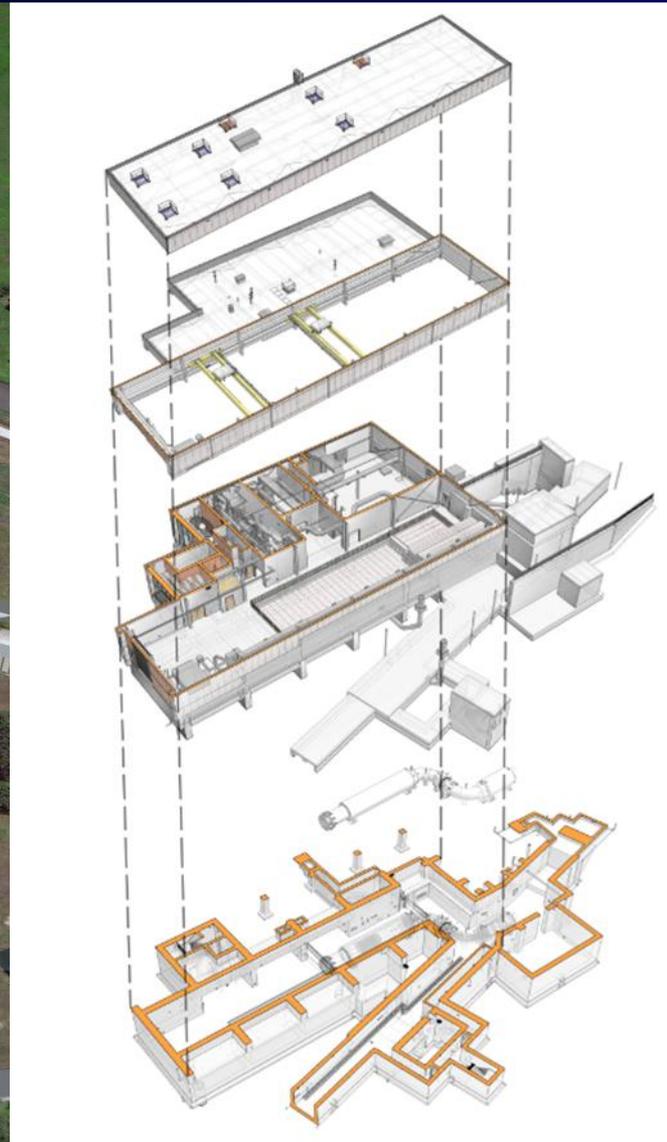
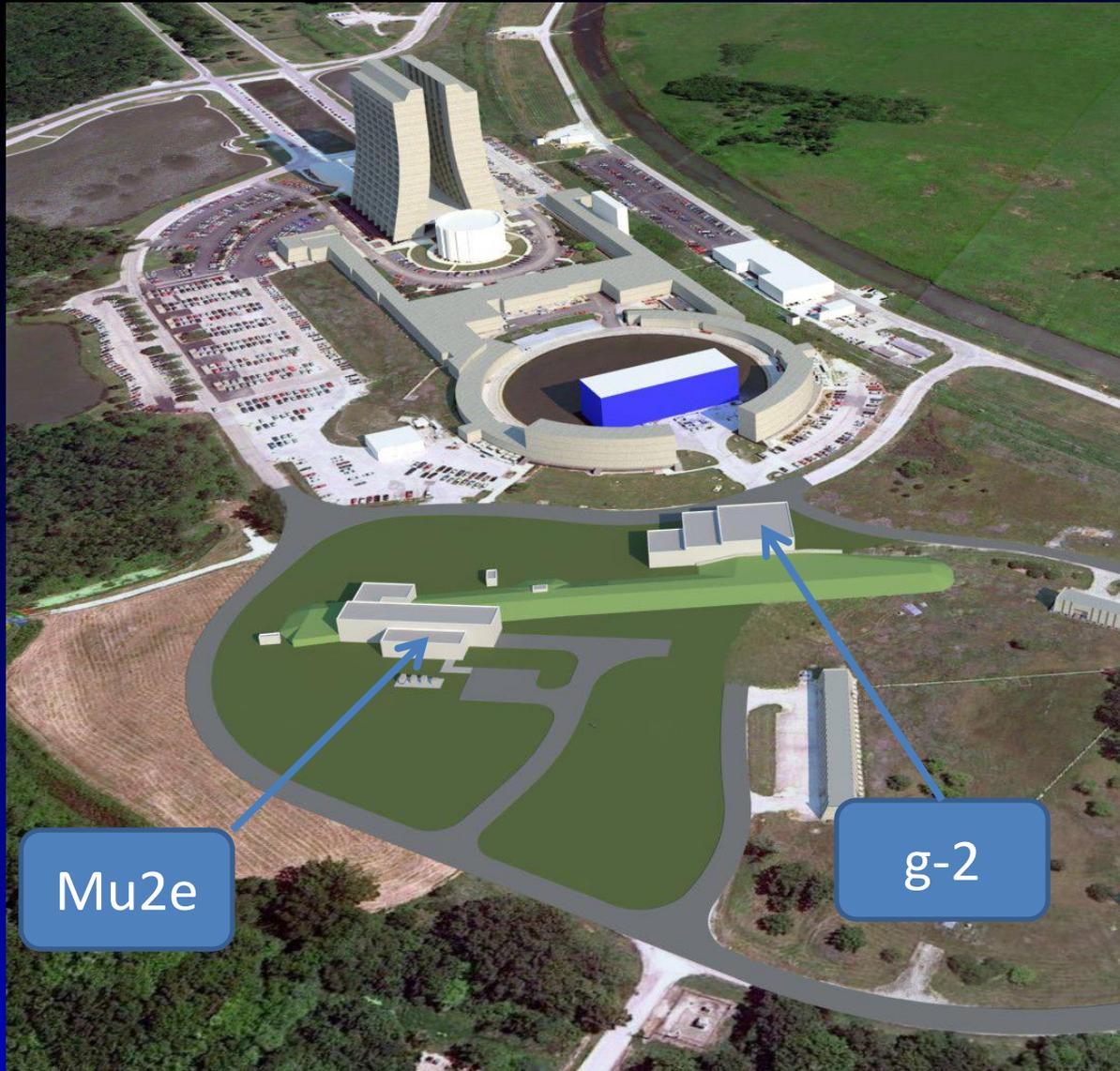
- signature delayed single isolated electron
- Electron energy is rest mass of the muon minus the nucleus recoil + binding energy:

$$E_e = m_\mu - E_{NR} - E_b \sim 104.97 \text{ MeV (Al)}$$

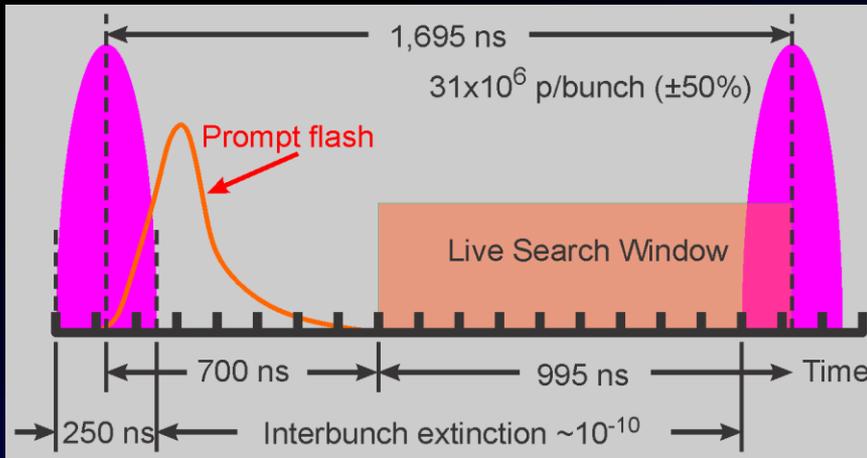


e⁻

New Beamline and Detector Hall to be Built



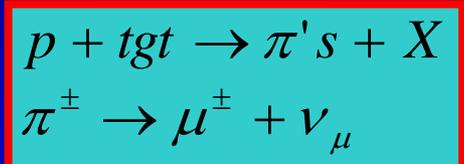
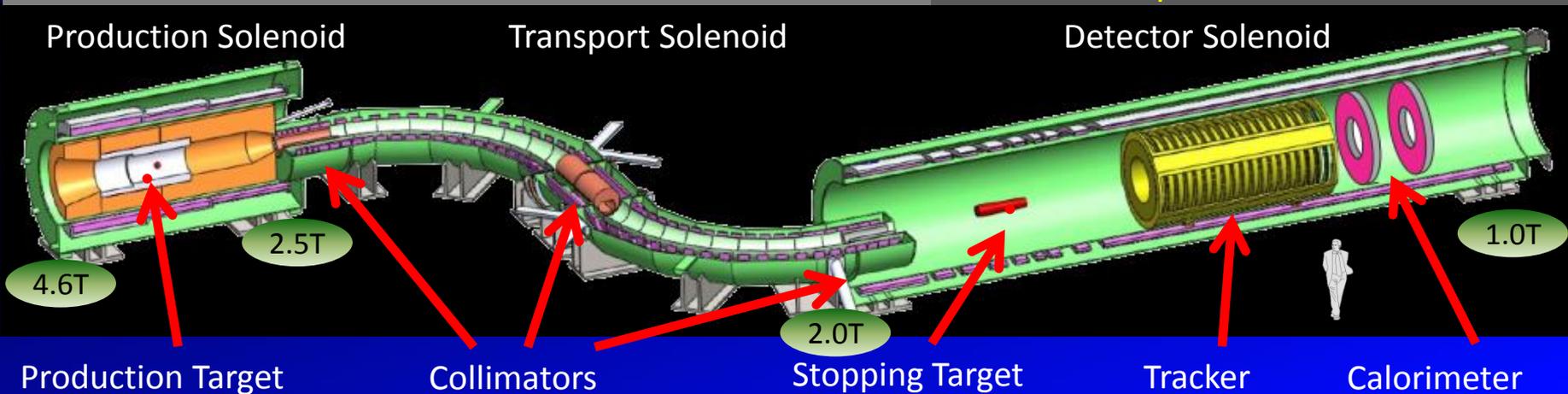
Mu2e Apparatus



- ### Salient Features
- Graded solenoidal field for pion capture
 - Muon transport in curved solenoid to eliminate neutral and positive particles
 - Pulsed beam to kill prompt backgrounds
 - No detector elements in muon beam

Muon Beam

Spectrometer

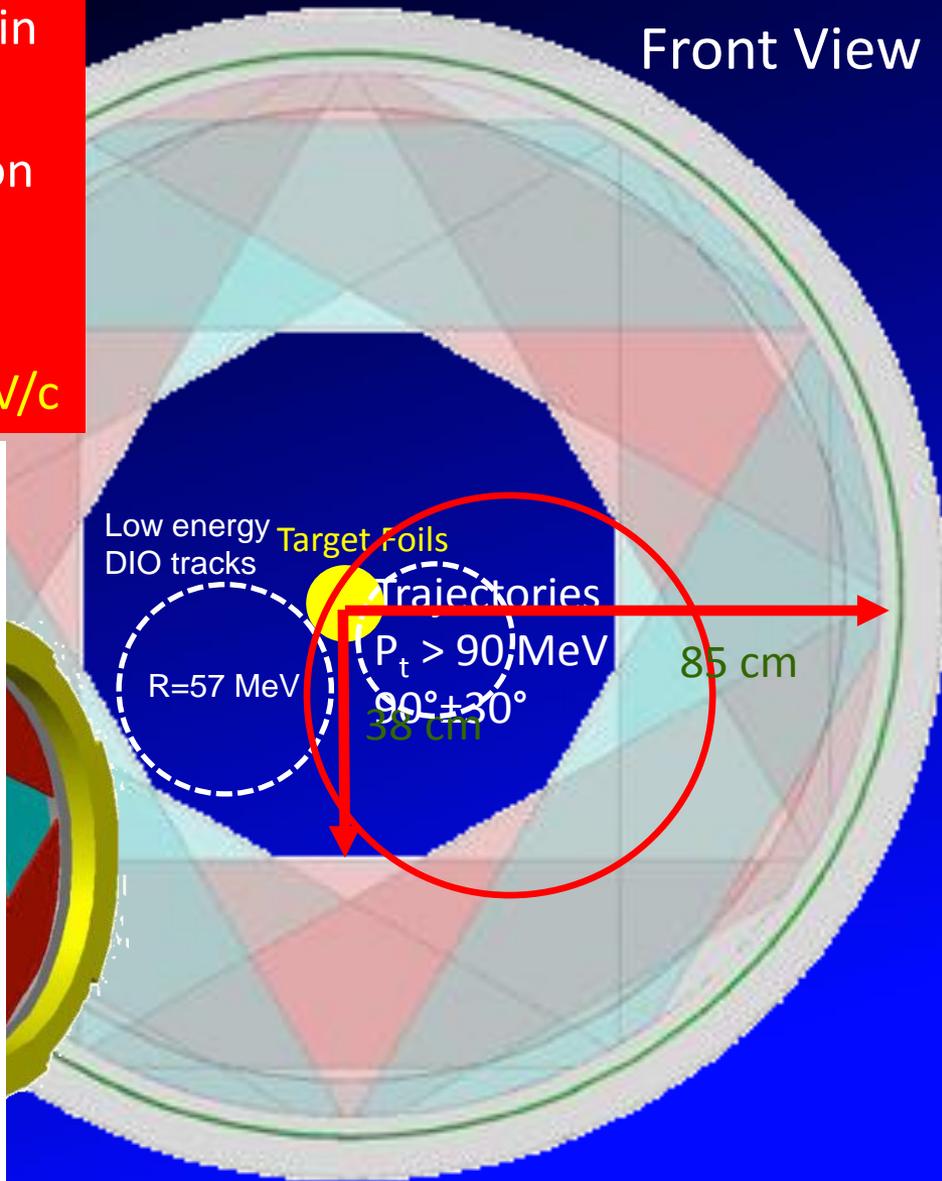
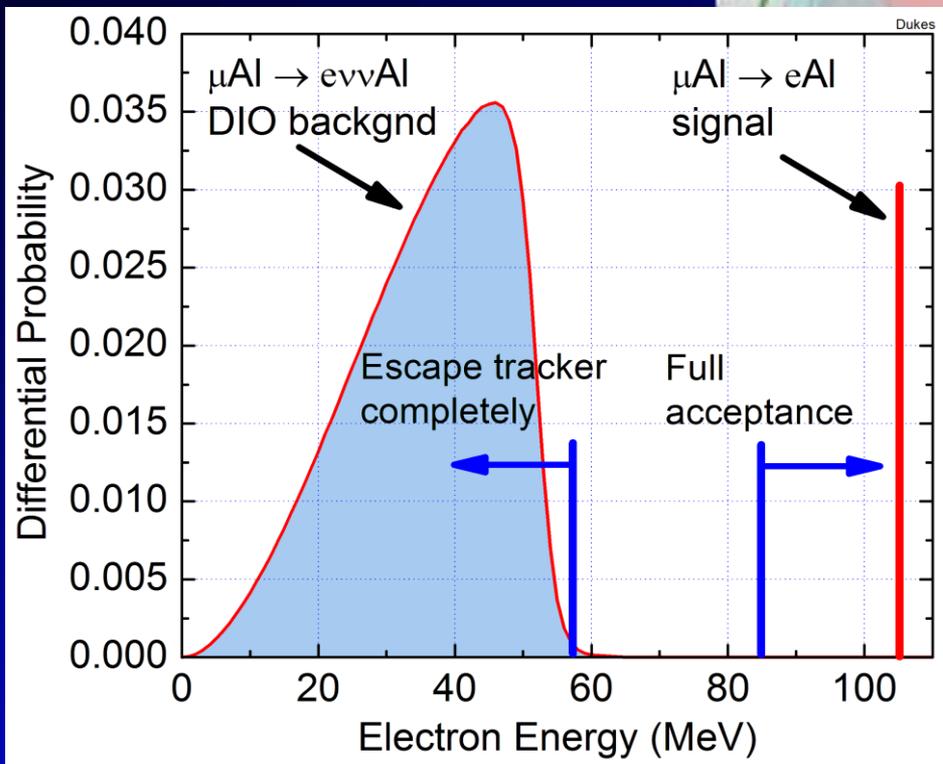


← 25 m →

50,000 μ^- stop every 1.7 μ s
 30 billion μ^- stops/spill-second

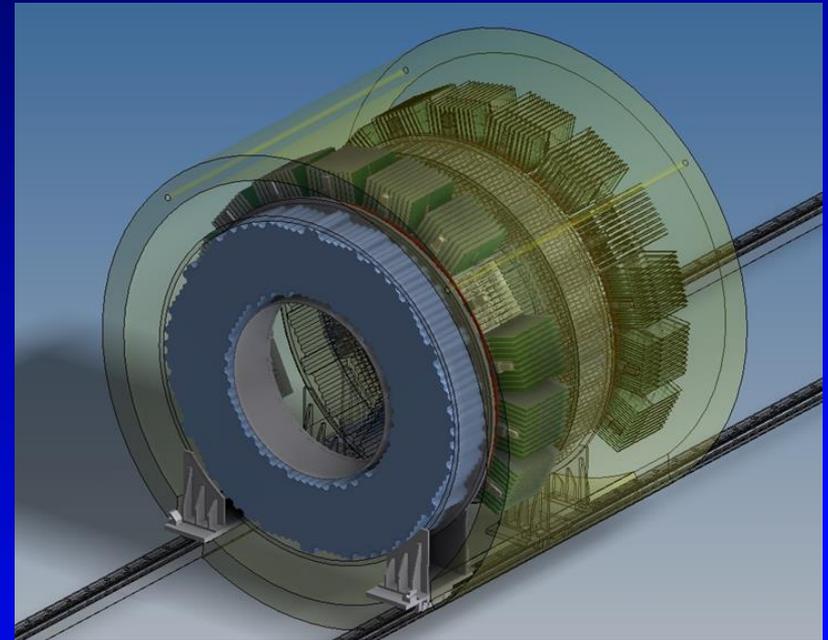
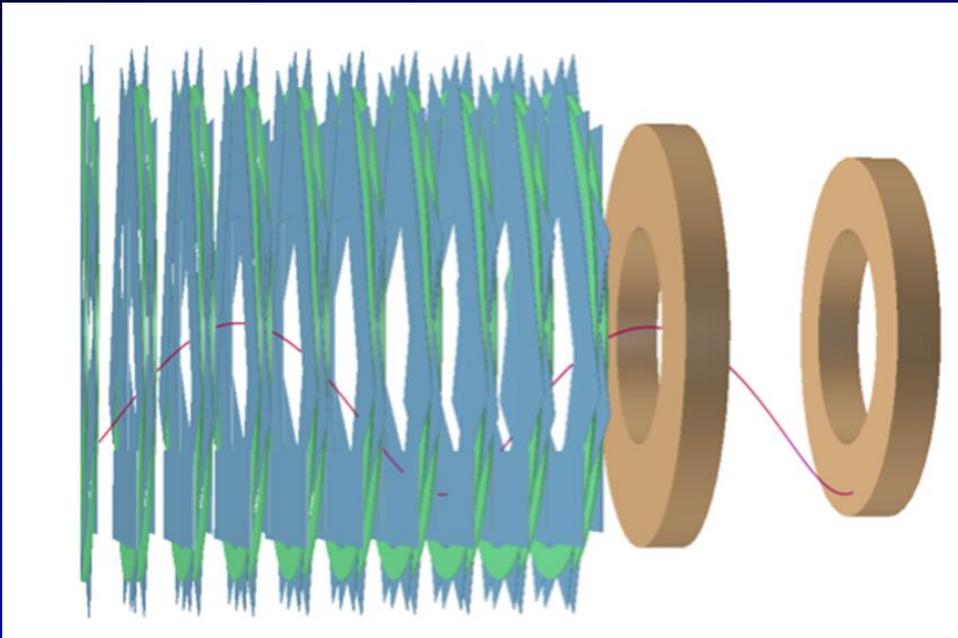
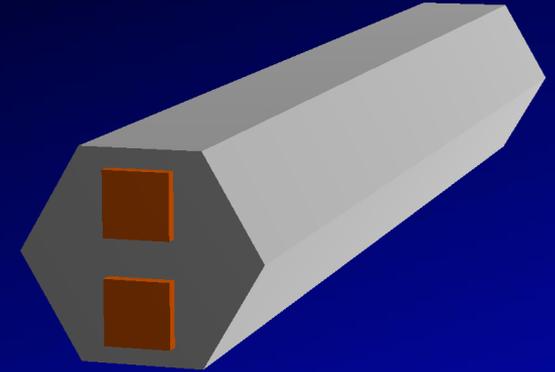
Heart of the Spectrometer: Straw Tracker

- 23,000 5-mm diameter straws arranged in octagonal geometry
- Center beam region empty: intense muon beam passes freely through
- Signal/DIO acceptance: $>20\%/<10^{-3}$
- Excellent energy resolution: 115/175 KeV/c



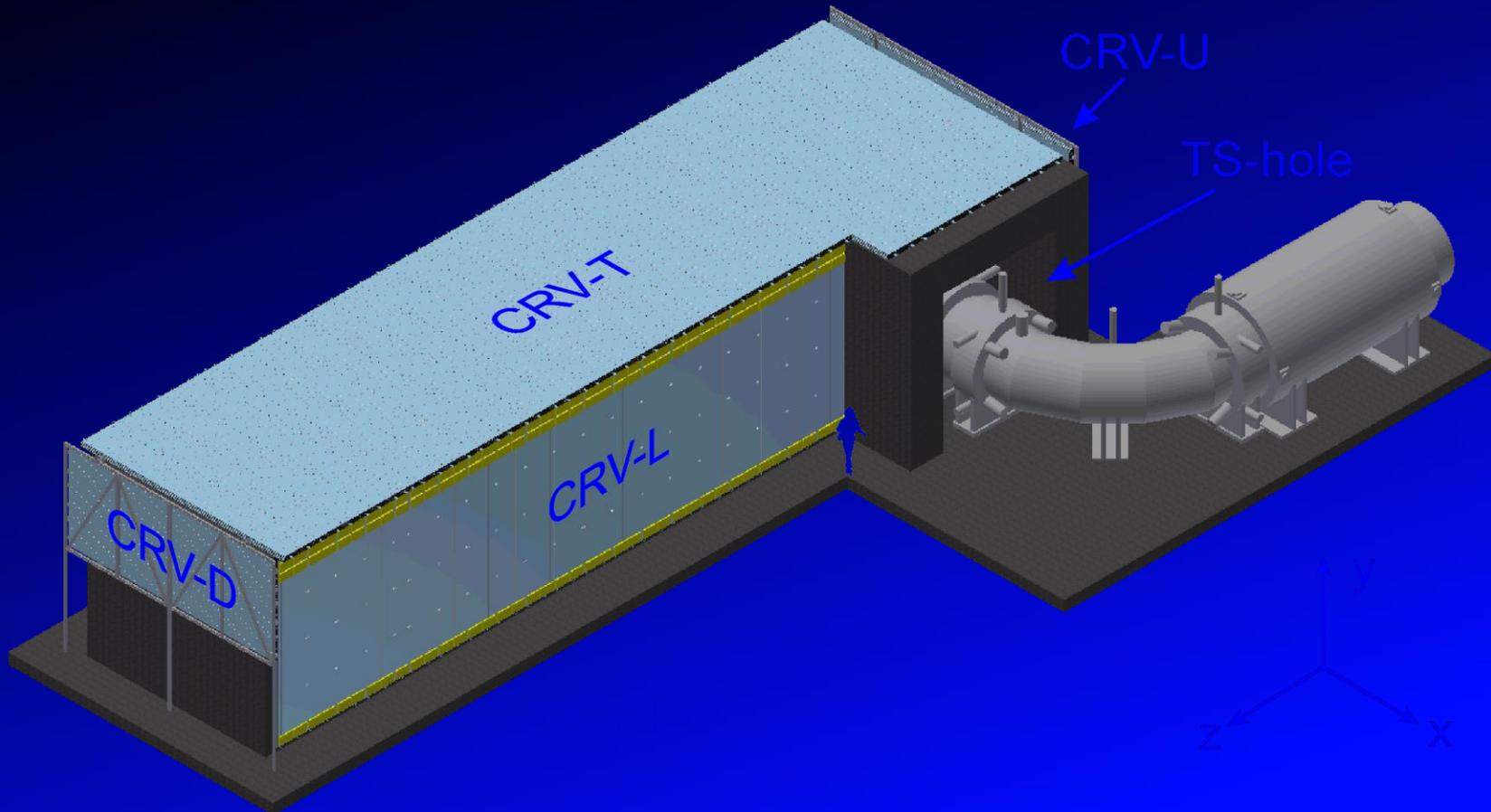
Electromagnetic Calorimeter

- Needed for:
 - Timing
 - Confirmation of tracker energy measurement
 - Particle ID: e vs μ
- 1500 BaF_2 crystals
- Dual APD readout



Cosmic Ray Veto

- One fake event/day from cosmic-ray muons
- Hence need 10^{-4} Inefficiency
- Surround Detector Solenoid by 4 layers of scintillator read out by waveshifting fibers, and silicon photomultipliers

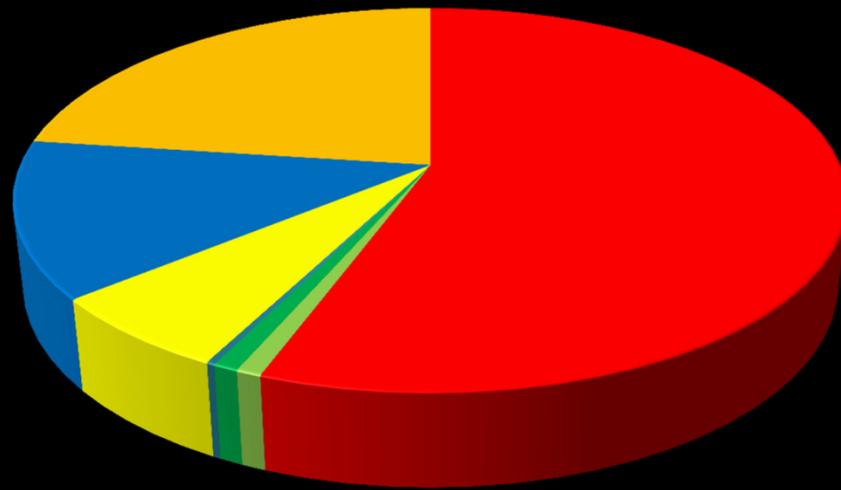


What we Get

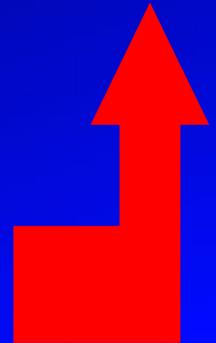
$$R_{\mu e} = \frac{\Gamma(\mu N \rightarrow e N)}{\Gamma(\mu N \rightarrow \nu_{\mu} N^*)}$$

10,000X improvement over present limit (SINDRUM II)

Proton flux	1.9x10 ¹³ p/s
Running time	6x10 ⁷ s
Total protons	3.6x10 ²⁰ p
μ ⁻ stops/incident proton	0.0019
μ ⁻ capture probability	0.609
Total acc. and selection eff.	0.085
Single event sensitivity	2.9x10⁻¹⁷
Detected events for R _{μe} = 10 ⁻¹⁶	3.4
Estimated background events	0.36



- Muon decay in orbit
- Radiative muon capture
- Beam electrons
- Muon decay in flight
- Pion decay in flight
- Radiative pion capture
- Antiproton induced
- Cosmic ray induced



Where Best to Find the New Physics?



“The sense of mystery has never been more acute in our field”
Ian Shipsey, Snowmass 2013

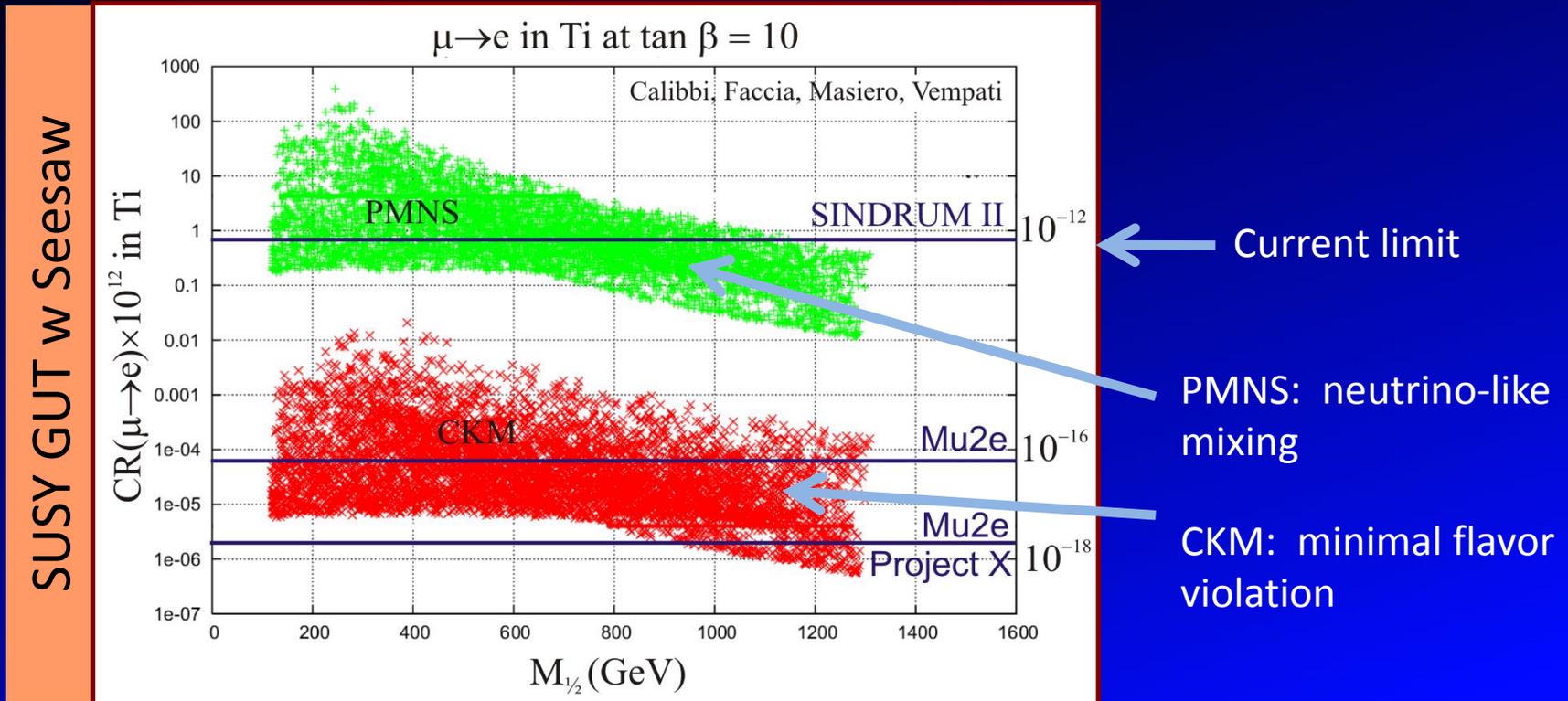
Backup Slides

What Sensitivity is Needed?

Present sensitivity already interesting and constraining!

$\sim 10^{-16}$ removes many models

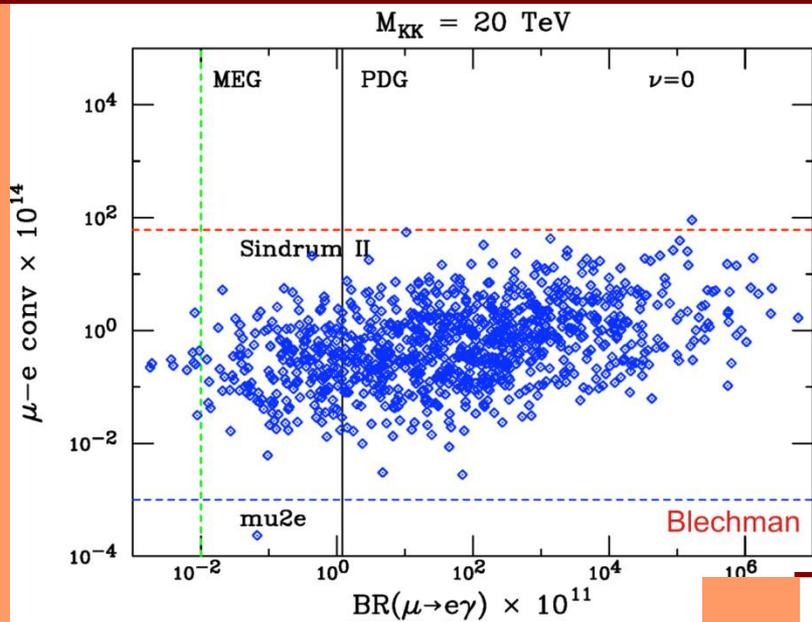
$\sim 10^{-18}$ extremely difficult for theorists to deal with



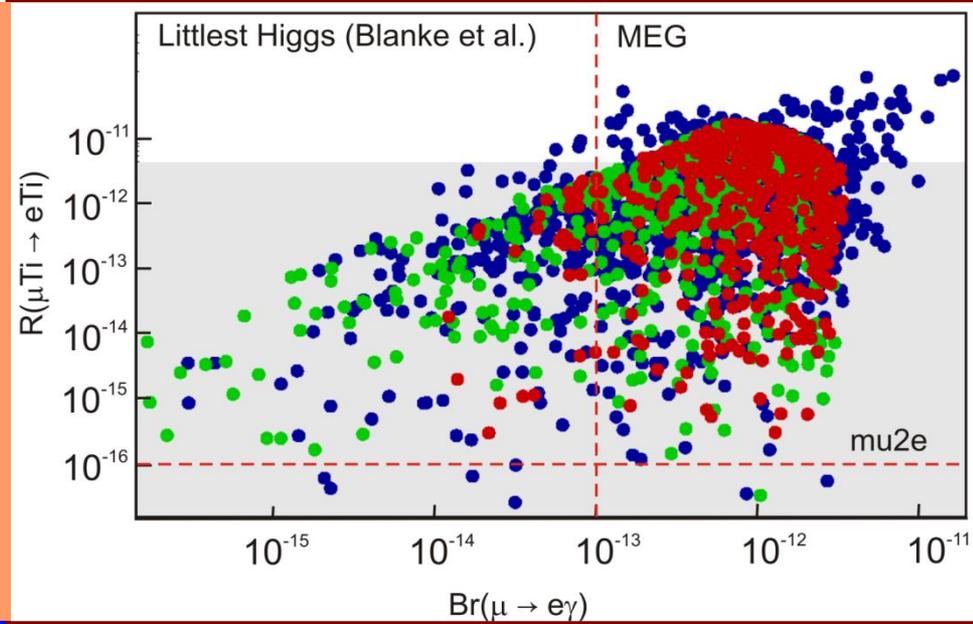
$M_{1/2}$ = gaugino mass

What Sensitivity is Needed?

Randall-Sundrum

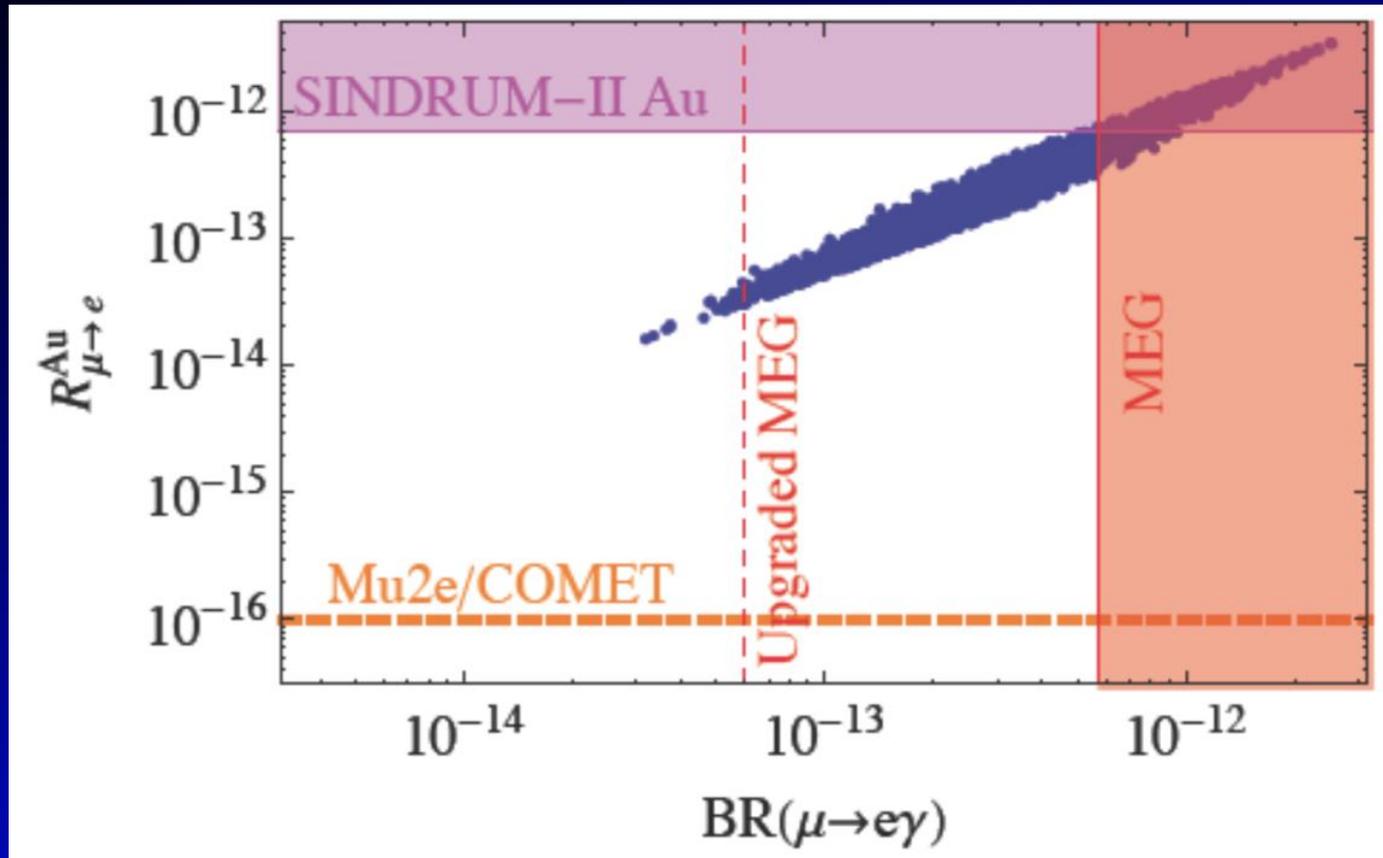


Littlest Higgs



What Sensitivity is Needed?

Left-right symmetric model with Type-I Seesaw



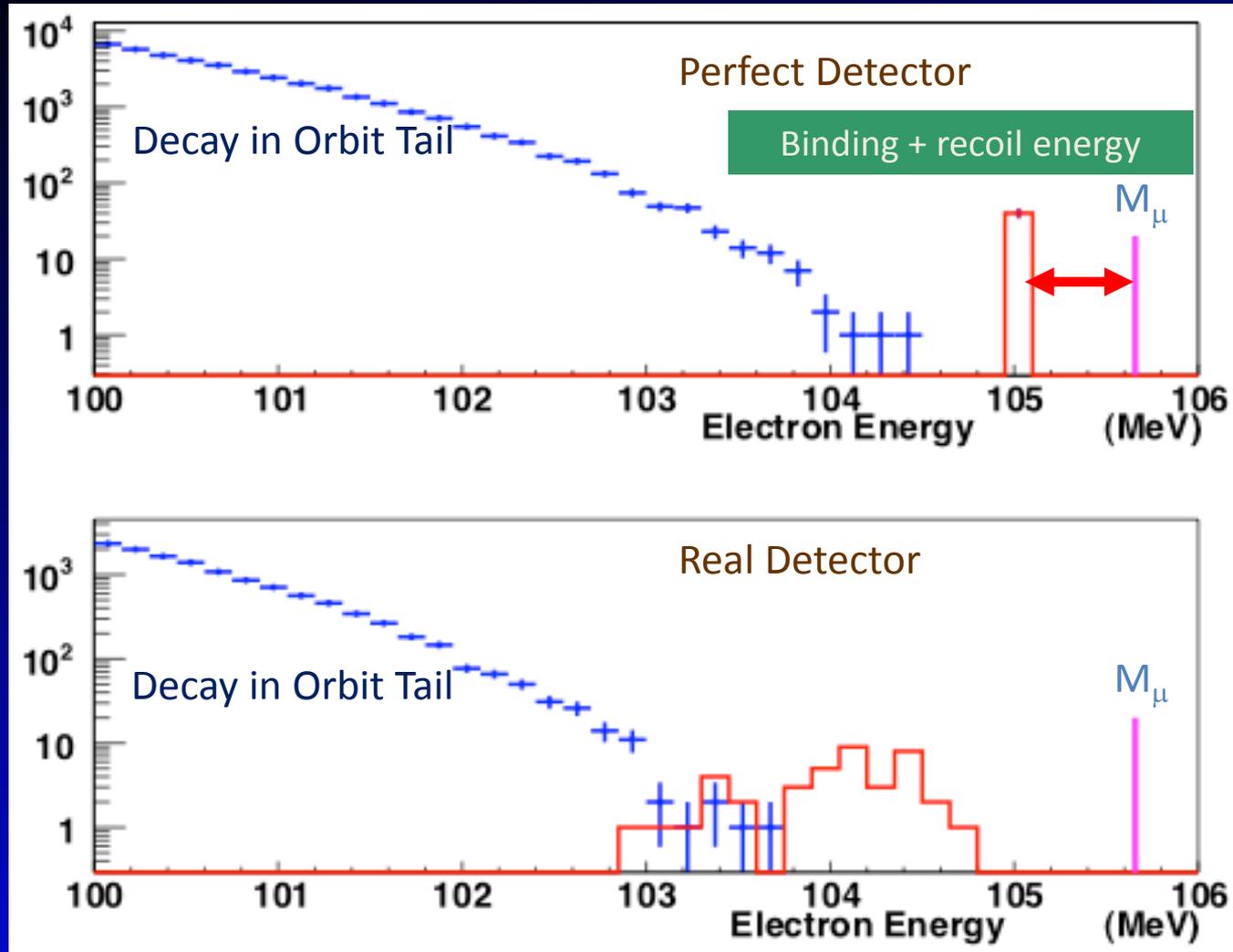
What we Get

Intrinsic tracker
energy resolution:

$$\sigma(E) \approx 150 \text{ keV}$$

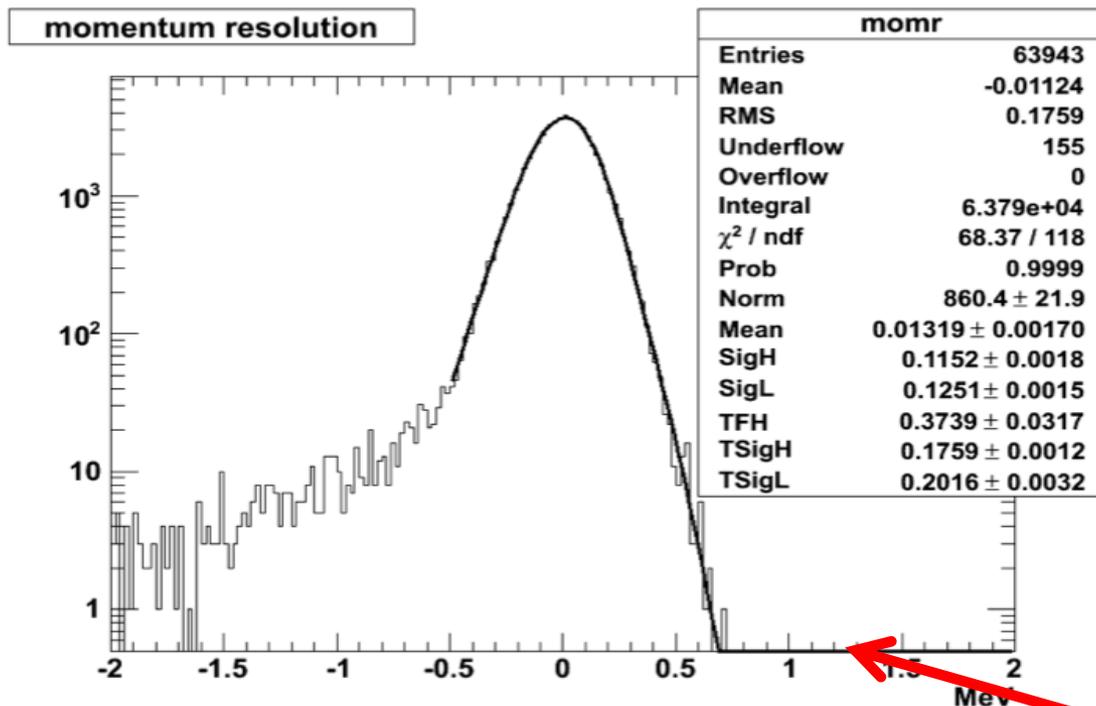
Average energy loss
due to spectrometer
material:

$$E(\text{shift}) \approx 1 \text{ MeV}$$



Tracker Performance

Tracker resolution for 105 MeV/c electrons



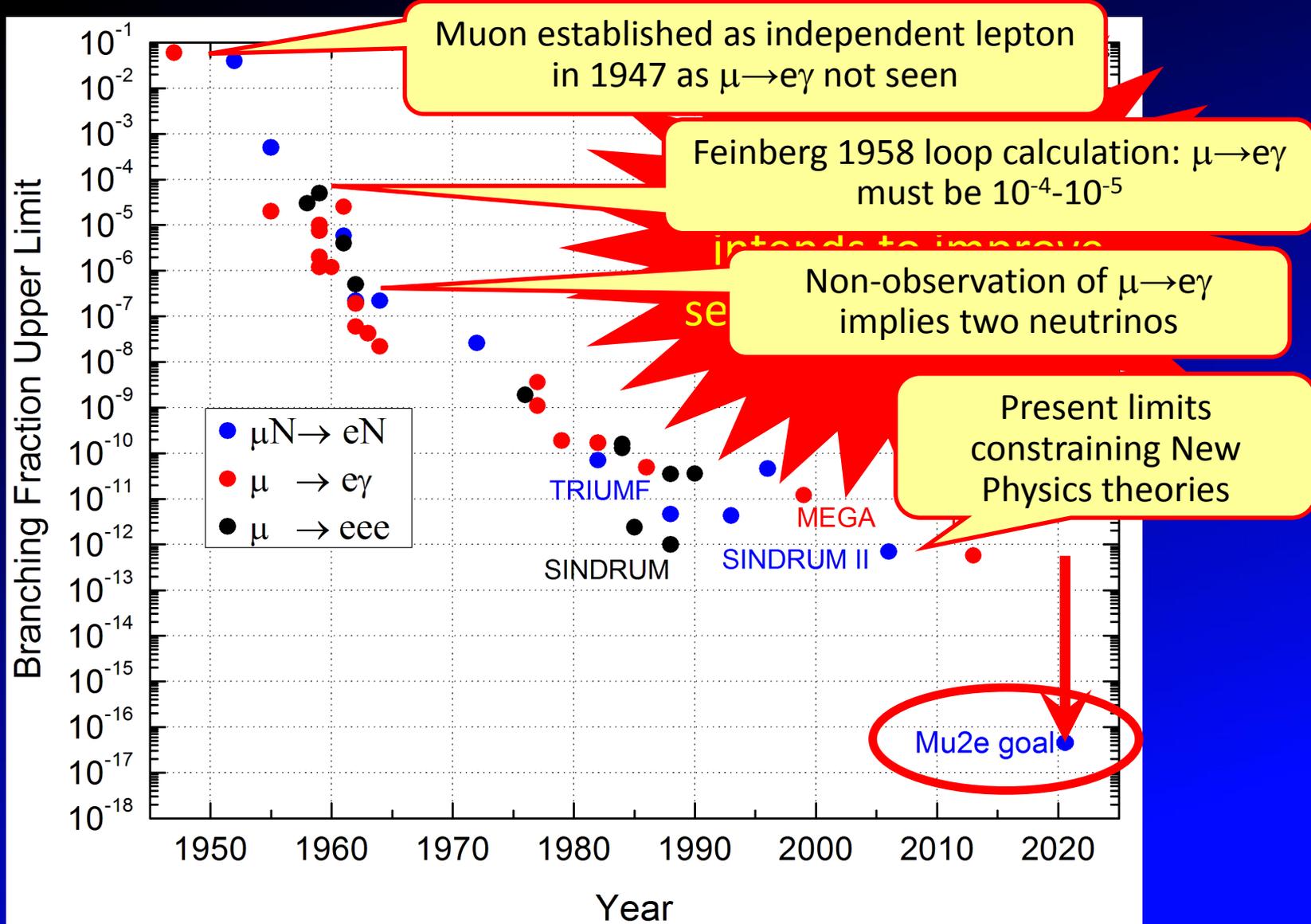
SigH	σ of core Gaussian, high side	115 KeV/c
SigL	σ of core Gaussian, low side	125 KeV/c
TFH	Fraction in the tail	37%
TSigH	σ of tail Gaussian, high side	176 KeV/c
TSigL	σ of tail Gaussian, low side	202 KeV/c

No high-side tail!

Core $\sigma = 115$ KeV/c

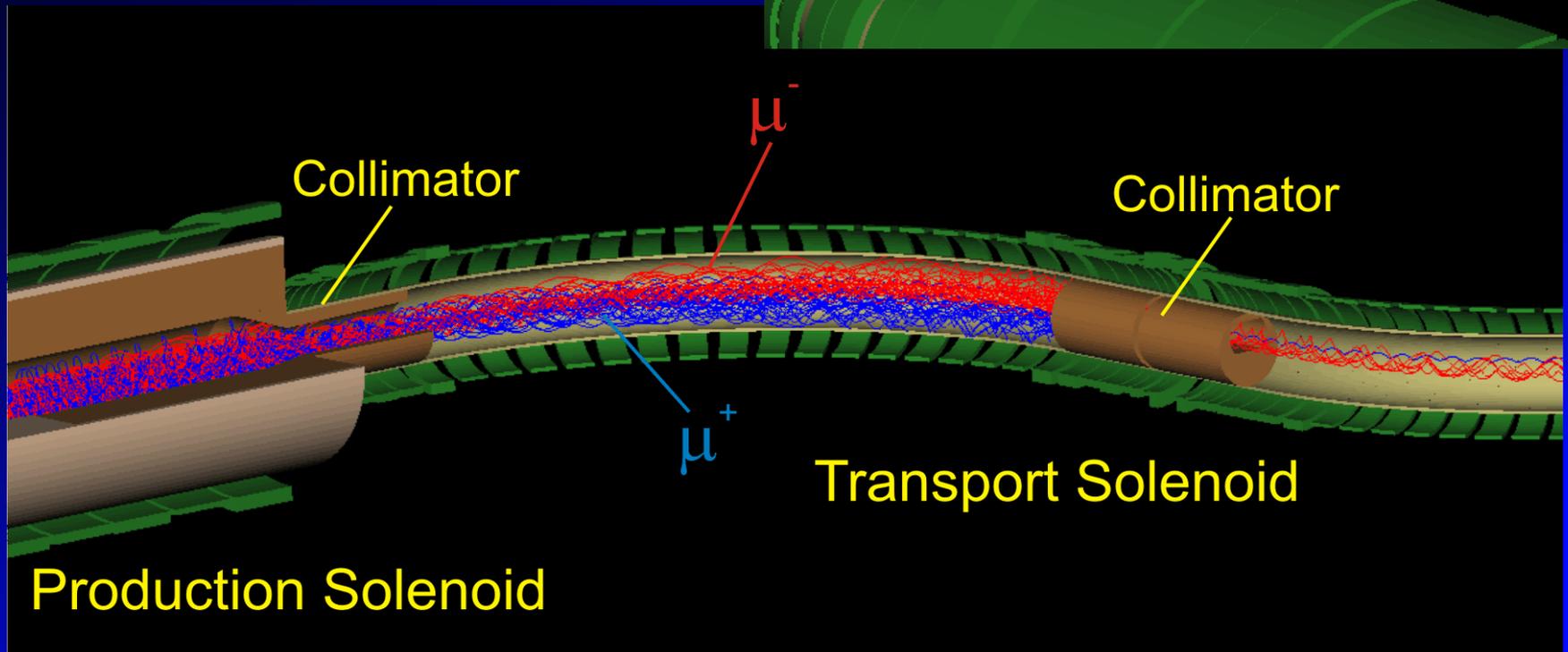
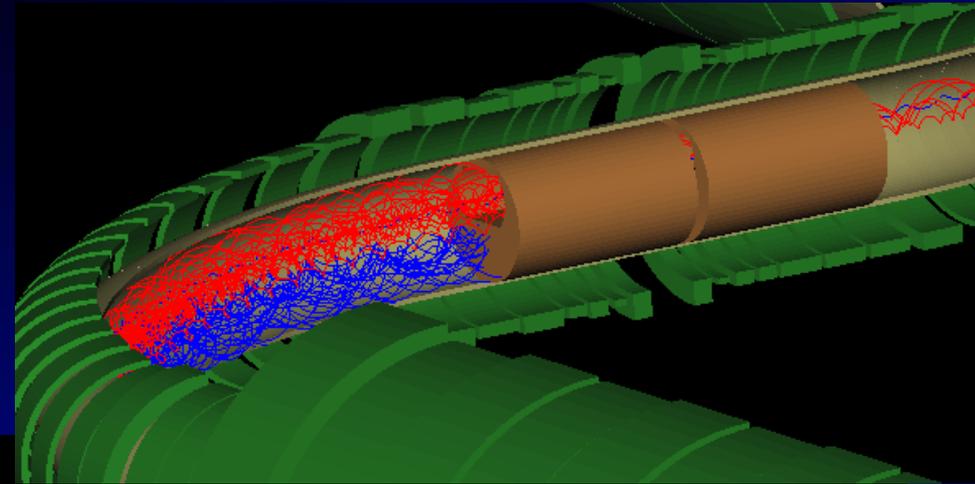
High-side tail $\sigma = 176$ KeV/c

History of Muon Lepton Flavor Violation Searches

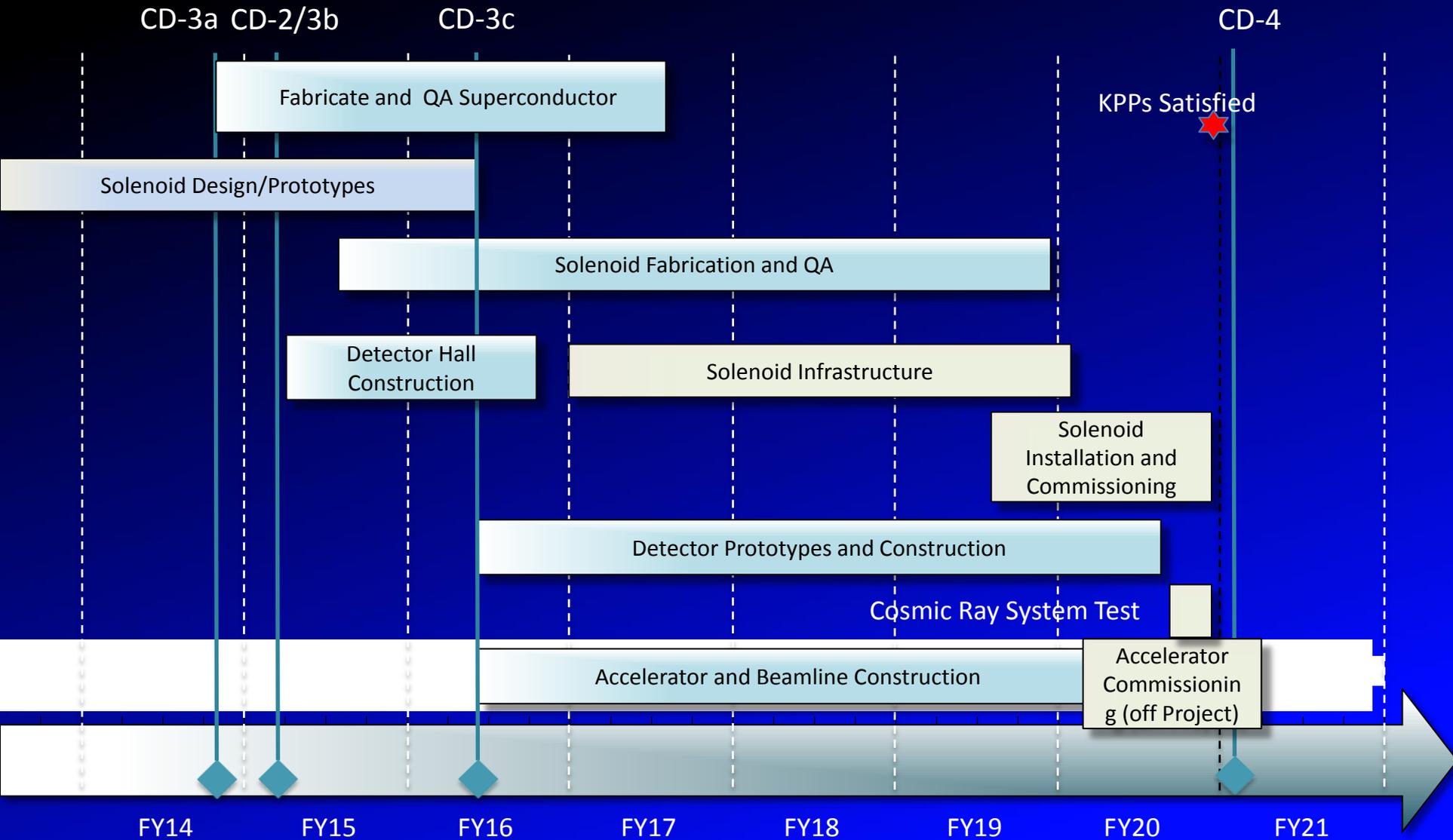


Transport Solenoidal Magnet

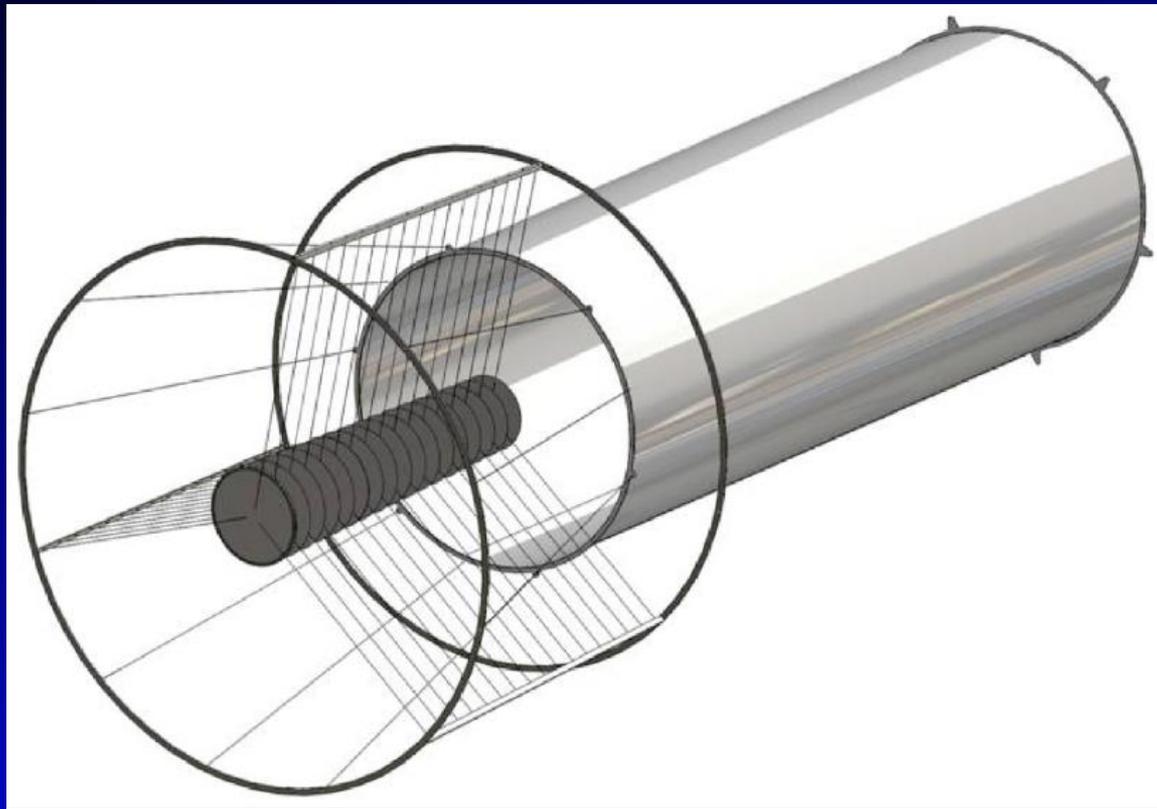
- Curved solenoid:
 1. separates charges by charge sign
 2. reduces line-of-sight transport of neutrals
- Collimators eliminate wrong-sign particles and particles with too large momentum



Schedule



Proton Absorber



Choice of Stopping Target Material

- **Large Z:**
 - rate $\propto Z |F_n|^2$ (F_n is the form factor)
 - can reveal nature of interaction
- **Small Z:**
 - longer lifetime
 - higher endpoint energy
- **Note:** Need $m_{Z-1} > m_Z$ to place max. energy of radiative capture muons below signal electrons

