

5-MINUTE FLASH TALK

Dark Fluids & Black Holes

A Curvature-Dependent Equation of State in Strong-Field Regimes

Bisweswar Sen
(You can call me Sen 😊)

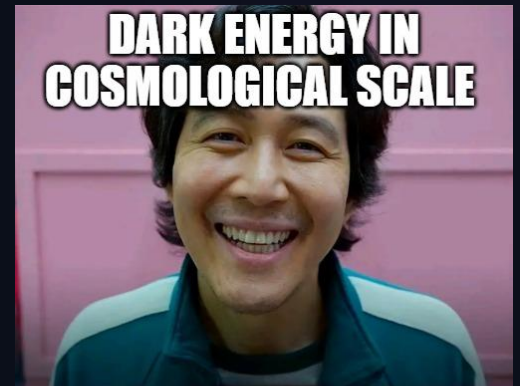
Supervisor- Dr. Nandan Roy
Mahidol University Nakhonsawan Campus, Thailand

Why Curvature-Dependent Dark Energy? (Or is it the other way round-Dark Energy influences Curvature?)

🌍 Cosmological Paradox

Standard Λ CDM presumes dark energy is perfectly static everywhere in the Universe. Normal EOS $P = w \rho$. $w = -1$

However, assuming identical DE behavior across both vast voids and extreme gravitational wells near black holes is physically restrictive.



🛡️ General Covariance

Setting a naive coordinate dependence violates General Relativity principles. $w(r)$

To preserve covariance while capturing strong-field response, the Equation of State must depend on **scalar curvature invariants**.

The Curvature Model & Dual Regimes

Choice of Scalar: Kretschmann Scalar K

$$R = 0$$

Outside Schwarzschild black holes, the simplest scalar - Ricci scalar vanishes. We adopt the second simple scalar - The Kretschmann scalar:

$$K = R_{\mu\nu\rho\sigma} R^{\mu\nu\rho\sigma} = \frac{48M^2}{r^6}$$

Non-zero in vacuum, non-singular at event horizon $r = 2M$, and diverges only at physical singularity $r = 0$



PROPOSED EQUATION OF STATE

$$w(K) = w_0 + w_1 \frac{K}{K + K^*}$$

- ♦ **Weak Curvature:** (Cosmological Λ CDM recovered)

$$w \rightarrow w_0 \approx -1 \quad K \ll K^*$$

- ♦ **Strong Curvature:** (Dynamical DE shift near BH)

$$w \rightarrow w_0 + w_1 \quad K \gg K^*$$

Accretion Dynamics & Flow

Relativistic Fluid Flow

Relativistic accretion rate scales as . $\dot{M} \propto \rho + P = \rho (1 + w)$

- ♦ **Standard Λ CDM:** . Zero dark energy accretion onto black hole.

$$w = -1 \Rightarrow \rho + P = 0$$

- ♦ **Modified Curvature Flow:** In strong gravity, if , then , allowing active DE inflow. As $\dot{M} > 0$ now.

$$w = -1 + \epsilon \quad \rho + P > 0$$

- ♦ **Sonic Point Shift:** Michel accretion critical radius shifts as local sound speed varies continuously. The sonic point changes its location

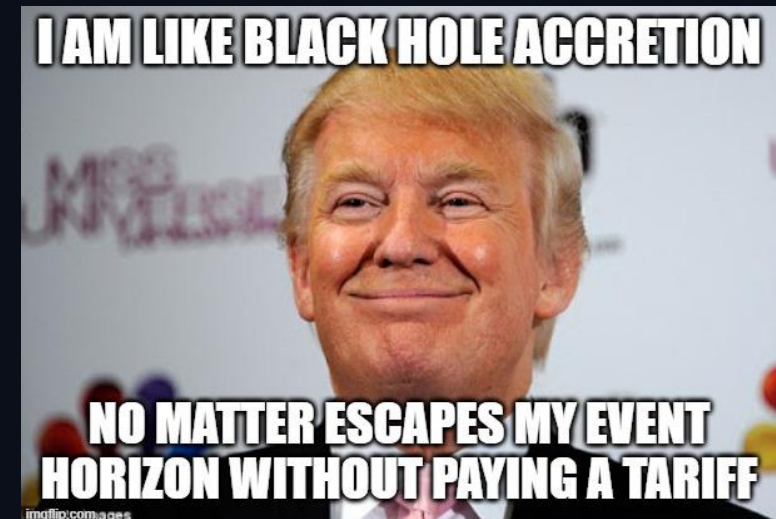
$$c_s^2 = w(K)$$

According to our model near a Black Hole:

$$M \propto \rho [1+w_0+w_1]$$

According to our model at cosmology scale:

$$M \propto \rho [1+ w_0]$$



Physics, Bounds & Extensions



Energy Conditions

Null, Weak, Dominant, and Strong Energy Conditions tightly constrain the parameter space, ruling out phantom instabilities & superluminal speeds.

(w_0, w_1)



Scalar Field Origin

Can be Derived dynamically from effective field theory with scalar non-minimal coupling to curvature:

$$\alpha K \phi^2$$



Kerr BH Extension

Ψ_2

Using Newman-Penrose Weyl scalar, rotation introduces angular asymmetry, distorting the DE profile at the equator. $w(r, \theta)$

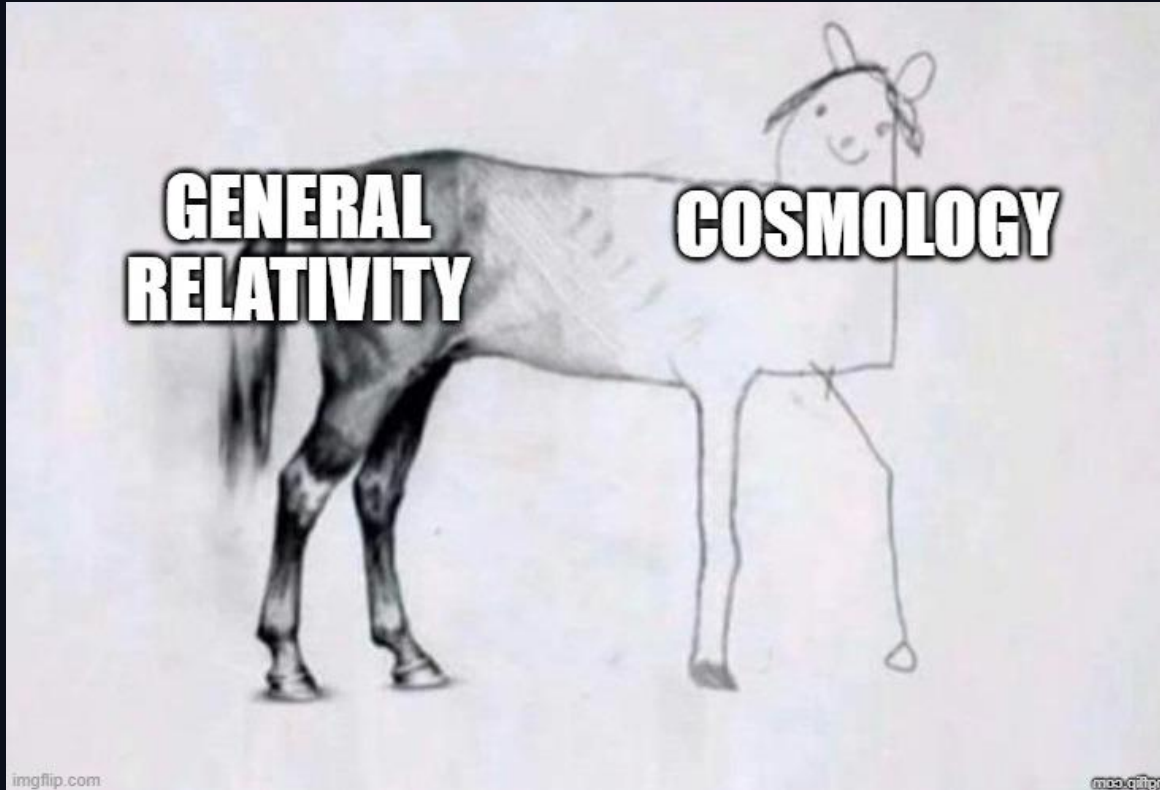
Motivation, Model , Summary , all in 1 Picture:

FLASH TALK SUMMARY



Core Results

Our Model EoS smoothly connects cosmological constants with strong-field modified dark energy near black hole horizons.



Observational Signature

Provides an analytical mechanism for DE accretion, opening pathways to test BH mass growth & gravitational wave ringdown shifts.

A fully covariant framework bridging cosmology and local strong gravity physics

Thank you for listening! Questions & Discussion welcome.