

Quintessential Dark Energy crossing the

PHANTOM DIVIDE

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Based on [2508.19101]



Motivation

What is the evidence for
Dynamical Dark Energy?



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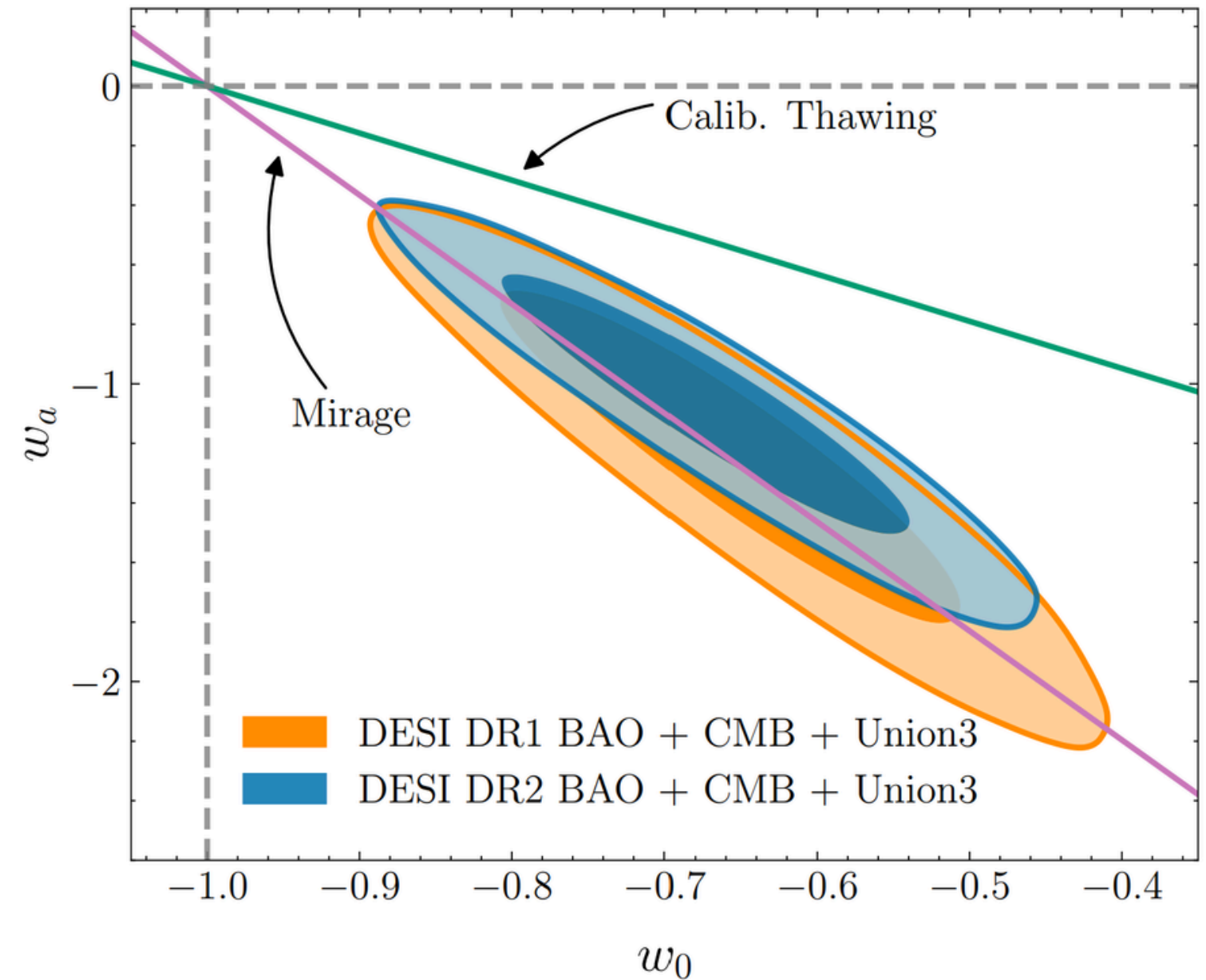


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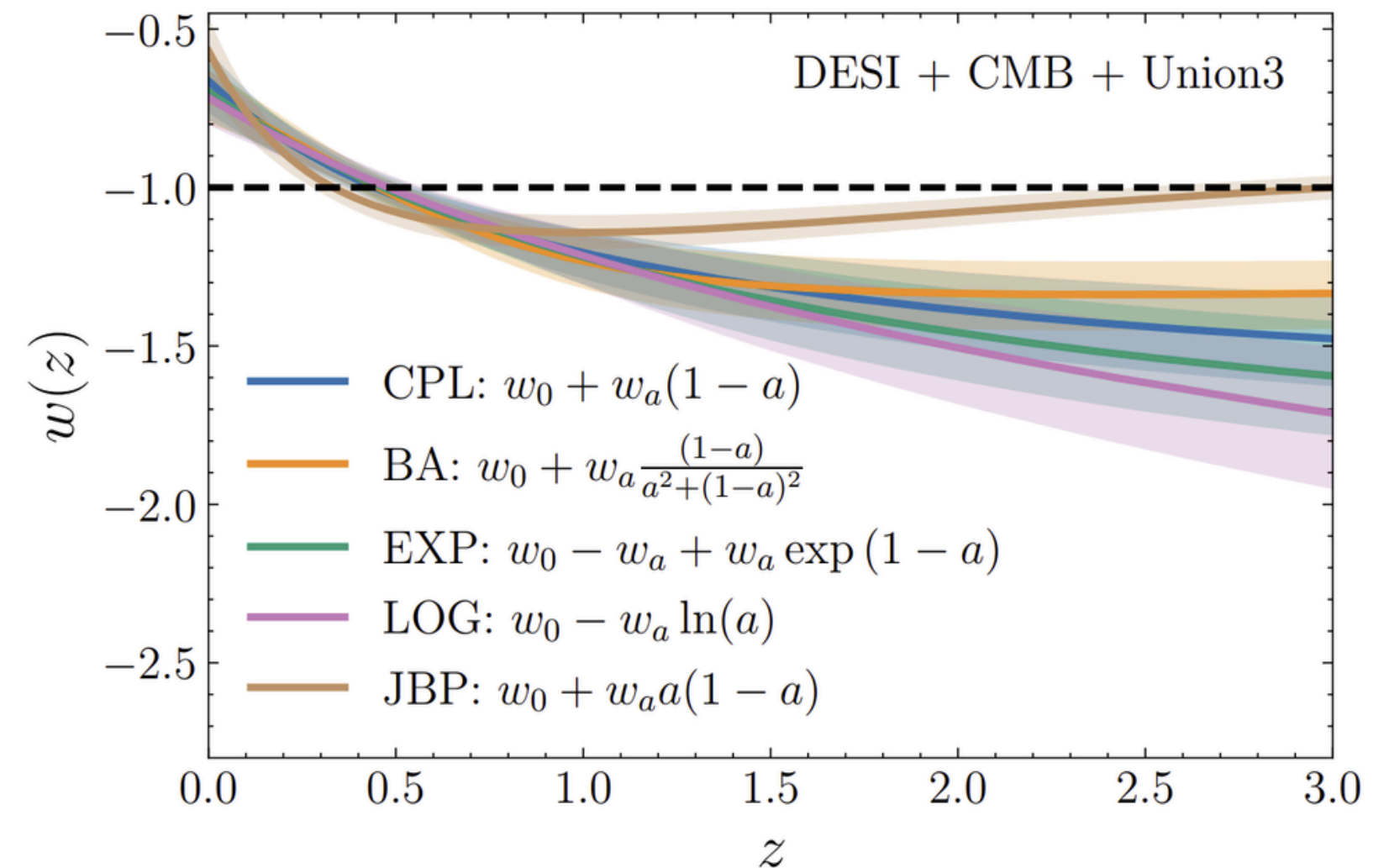


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Dark Energy is not only dynamical but also has an equation of state < -1 .

Violates the **Null Energy Condition** (NEC)

$$\rho + p \geq 0.$$

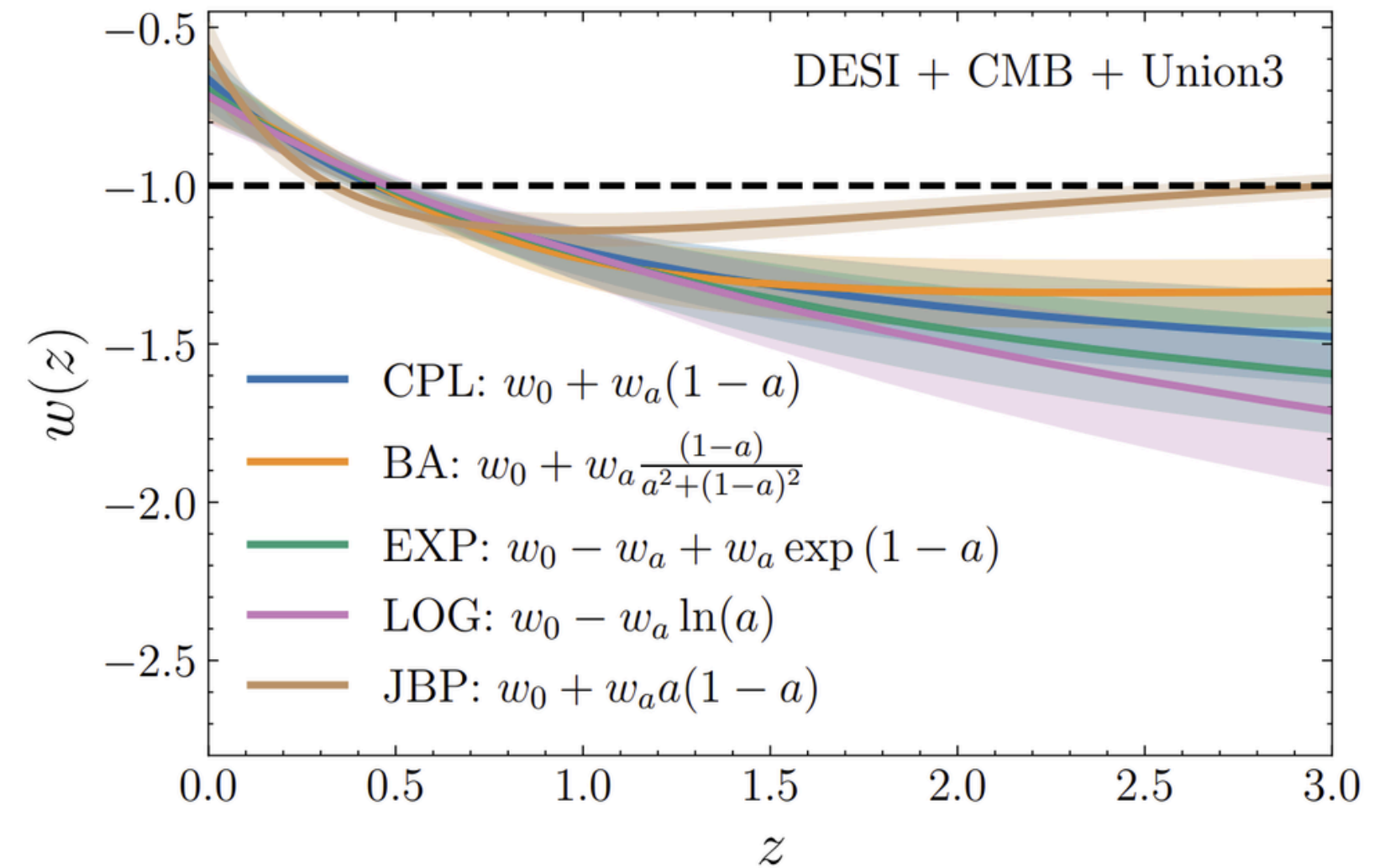


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Consider a scalar field rolling down some potential. The classical field equation can be written as

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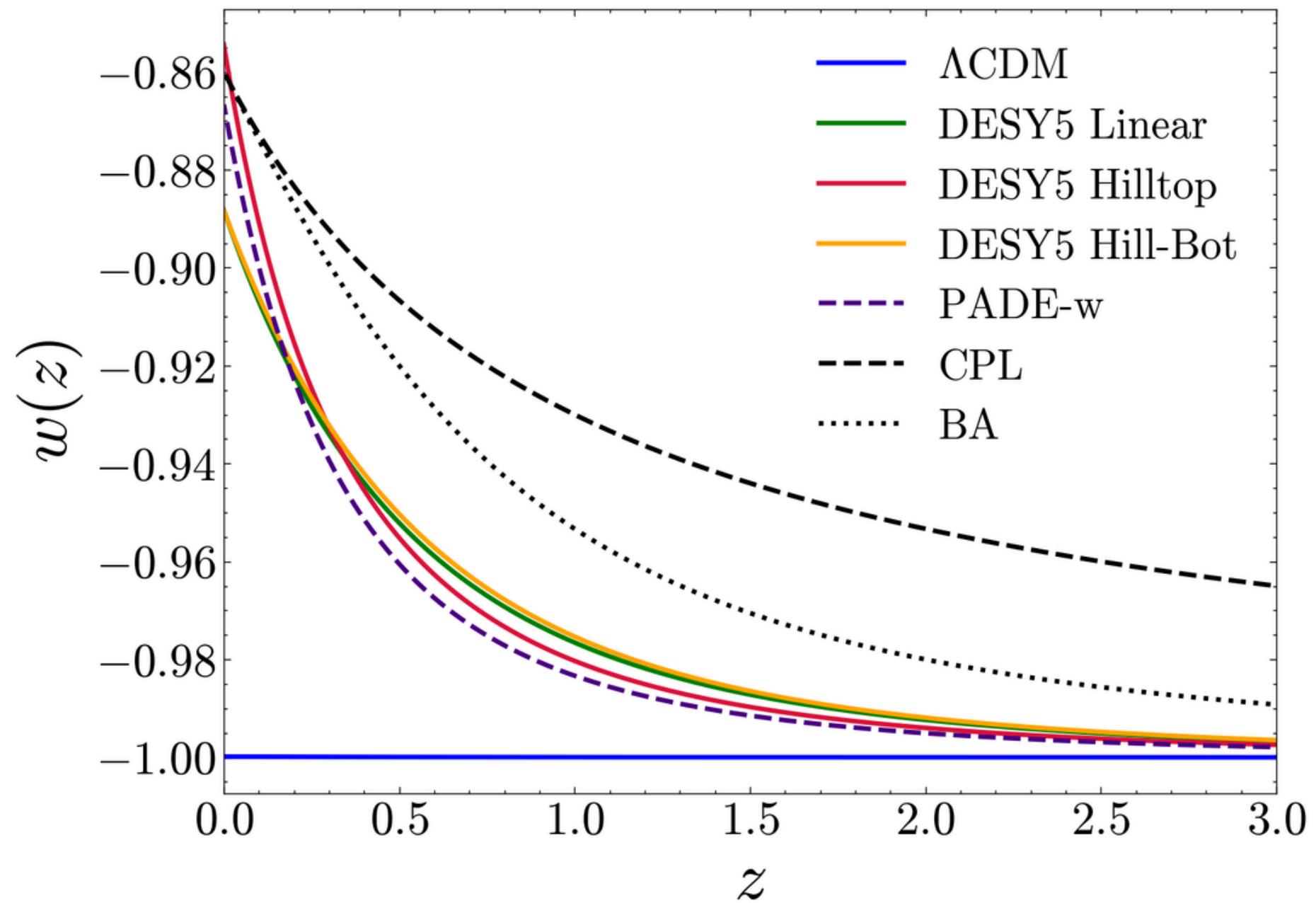
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Unphysical!

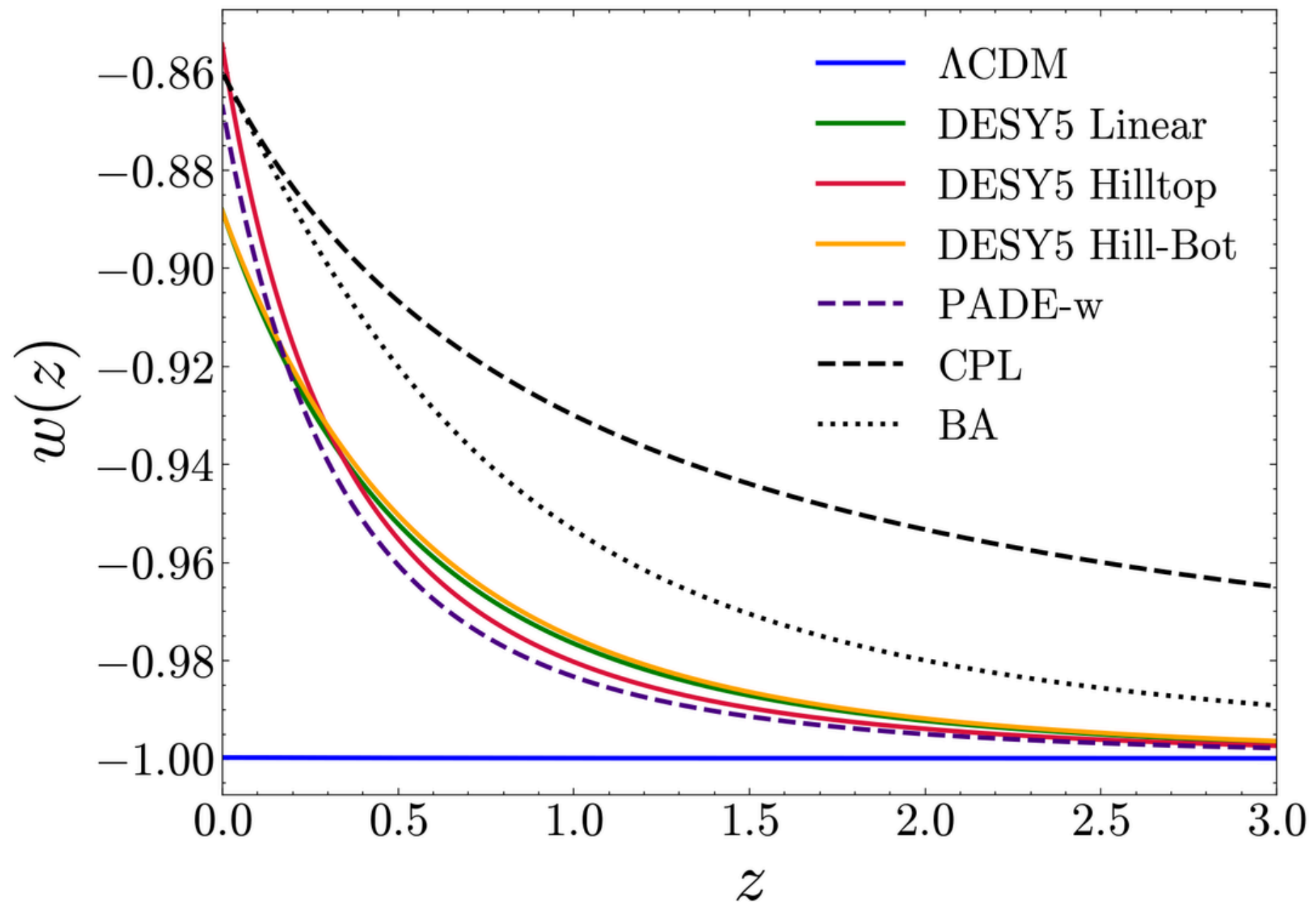
Quintessence Field

Figure from 2506.13047.



Equation of state vs redshift for quintessence models. The PADE-w parametrization with $\epsilon_0 = 0.2, \eta_0 = 3$ and CPL and BA with $w_0 = -0.86, w_a = -0.14$ are also shown.

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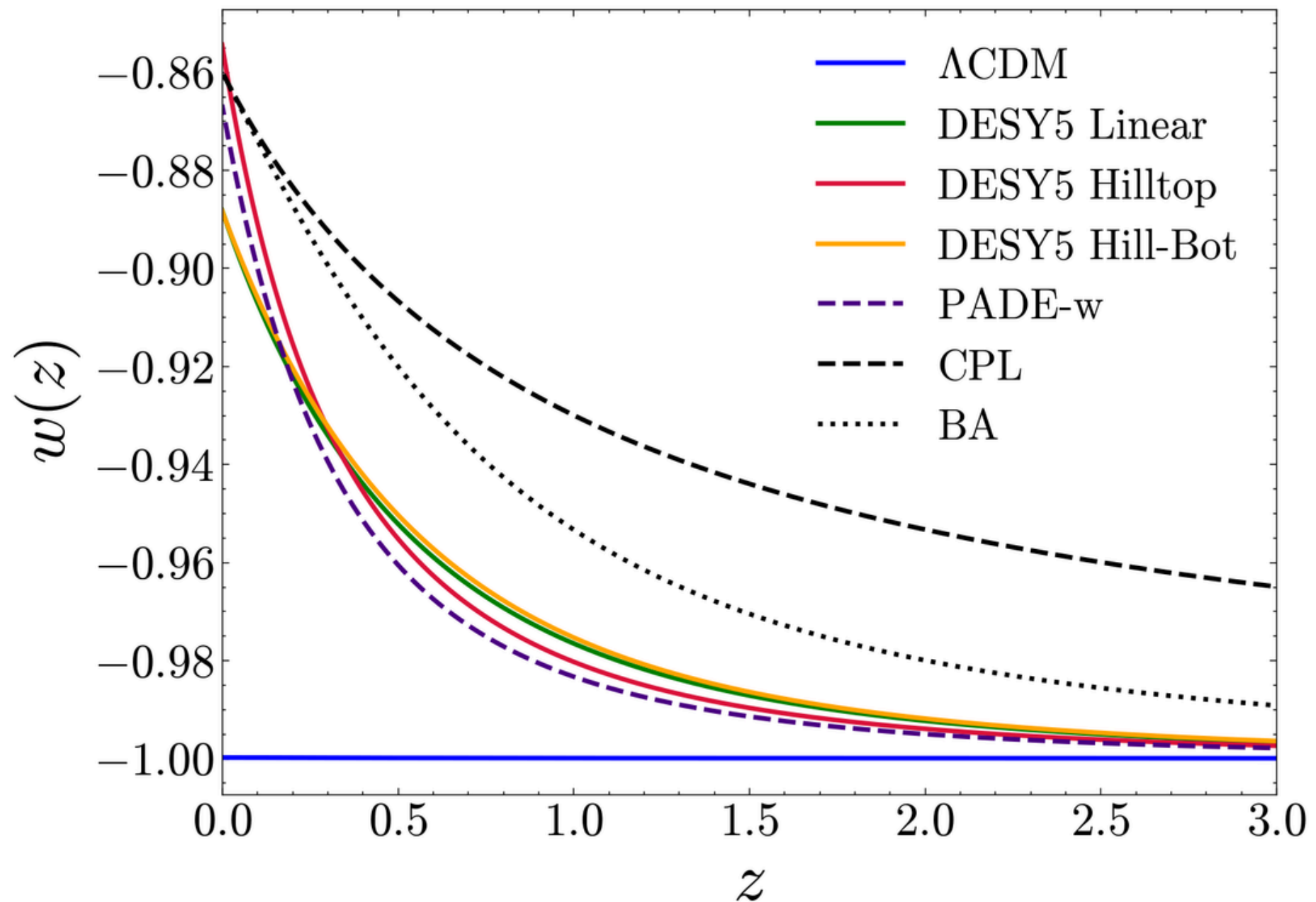
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Results show **minimal evidence** for Quintessence data. Only a modest improvement compared to a cosmological constant. [2506.13047]

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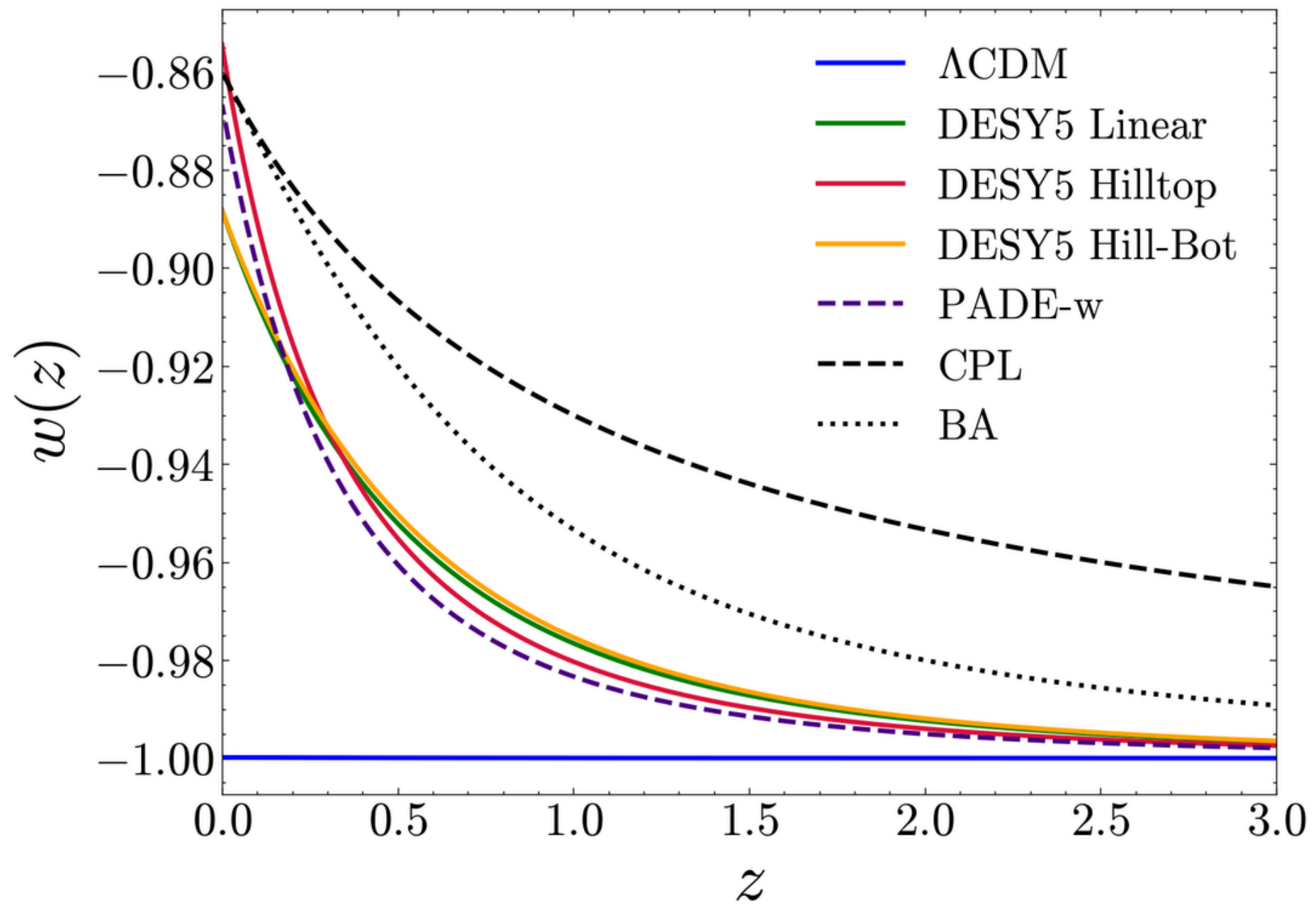
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Need phantom crossing to explain DESI results!

2

Motivation

How could one get
Phantom ($w < -1$) Dark Energy?



Interacting DM-DE

Suppose the dark matter mass varies with the field, which could occur through renormalizable operators such as $\phi^2 S^2$ for scalar DM or $\bar{\chi}\phi\chi$ for fermionic DM (Farrar and Peebles, 2003).

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Where,

$$V_{\text{eff}} = V(\phi) + \rho_{\text{dm}}^0 a^{-3} (g(\phi) - 1) = V(\phi) + \bar{g}(\phi) a^{-3}.$$

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Hence, the effective equation of state of dark energy becomes

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Thus, the dark energy “effectively” crosses into the phantom regime as the field evolves and the dark matter was **less massive** in the **past**.

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Analysis

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4 free parameters including the initial field velocity π_0 .

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Solve the differential equations (setting $8\pi G/3 = c = 1$)

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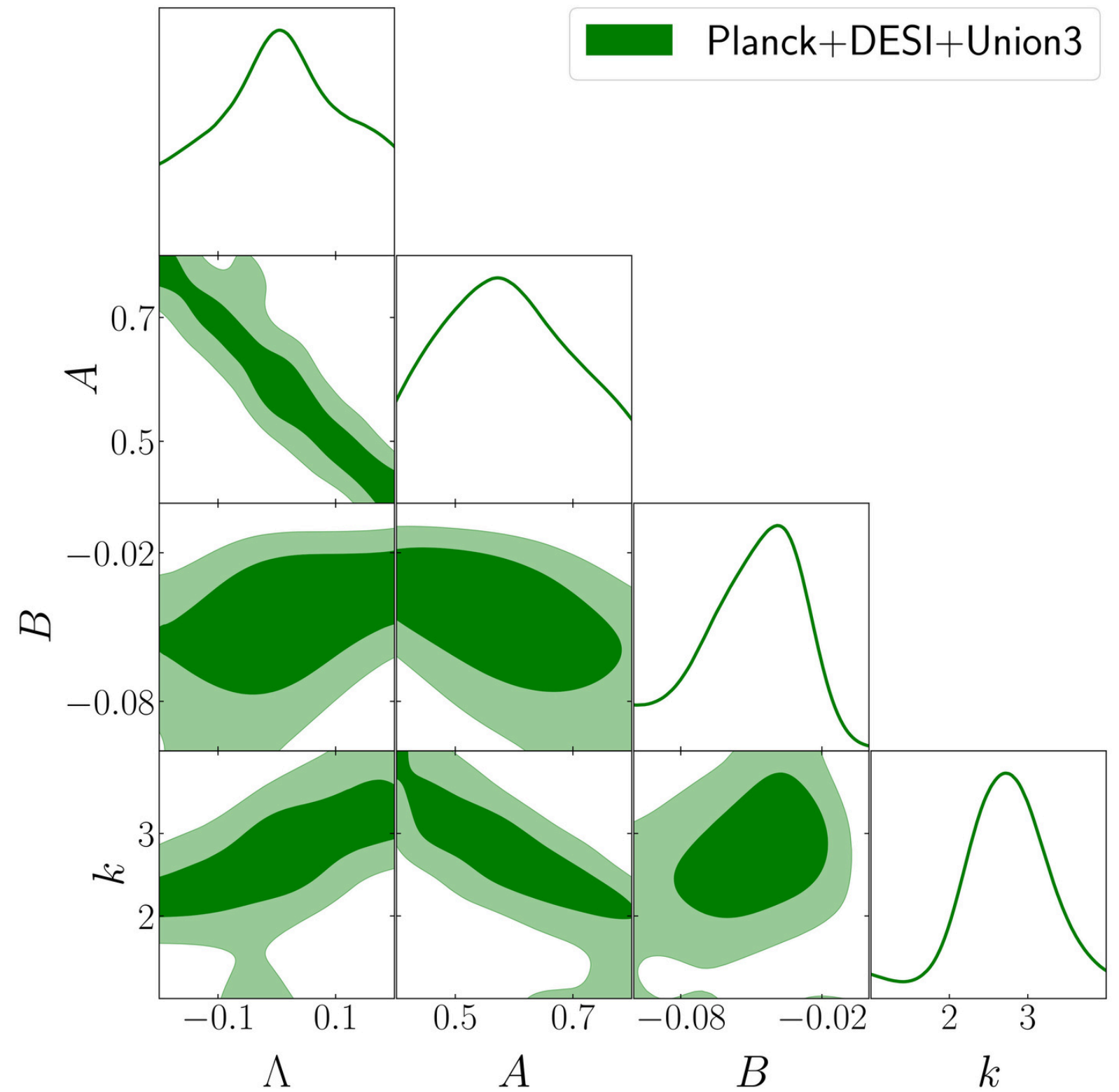
Datasets used

- Planck 2018 CMB + Lensing
- DESI 2024 BAO
- Union3 Type 1A SN.

Results

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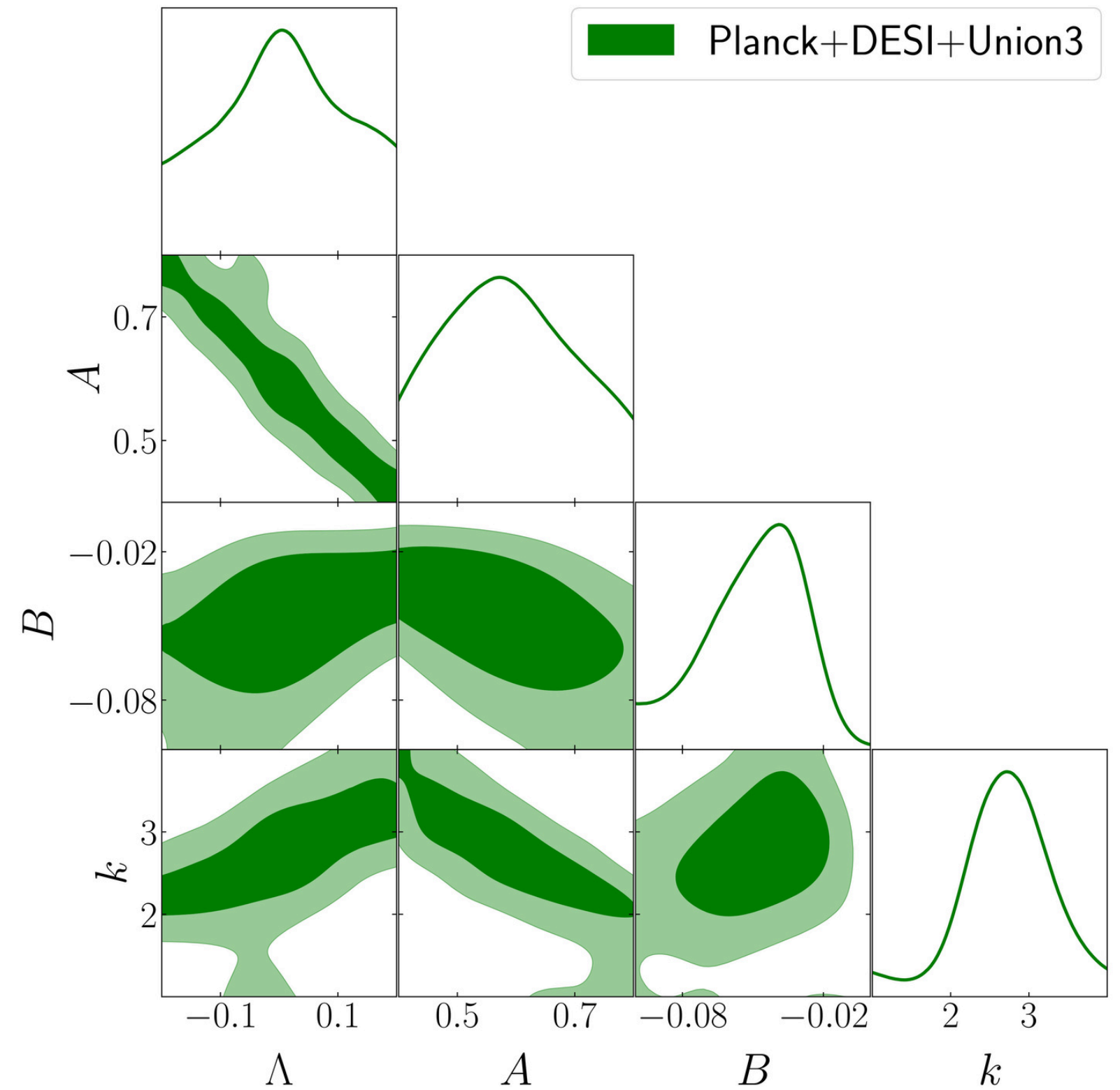
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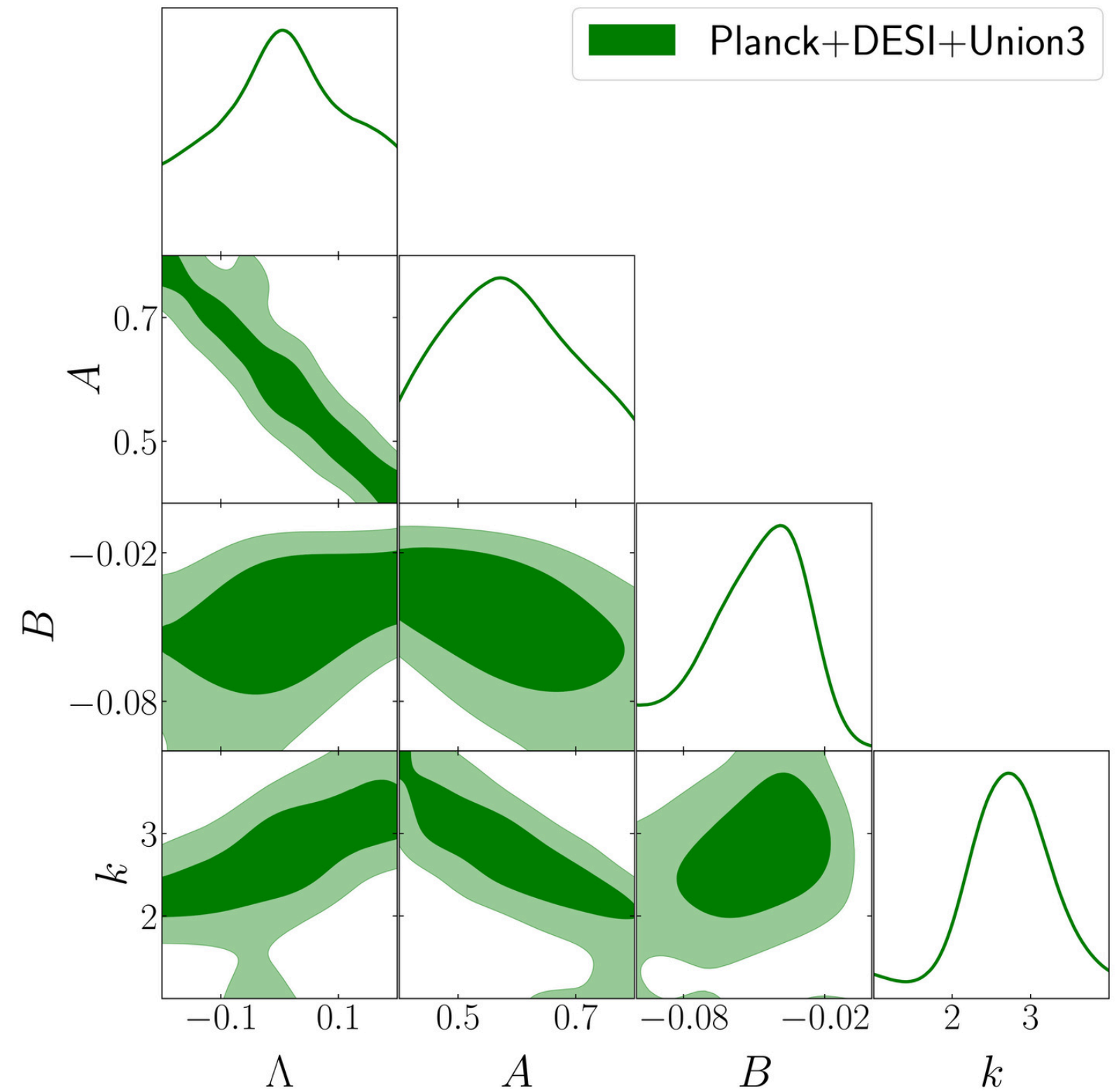
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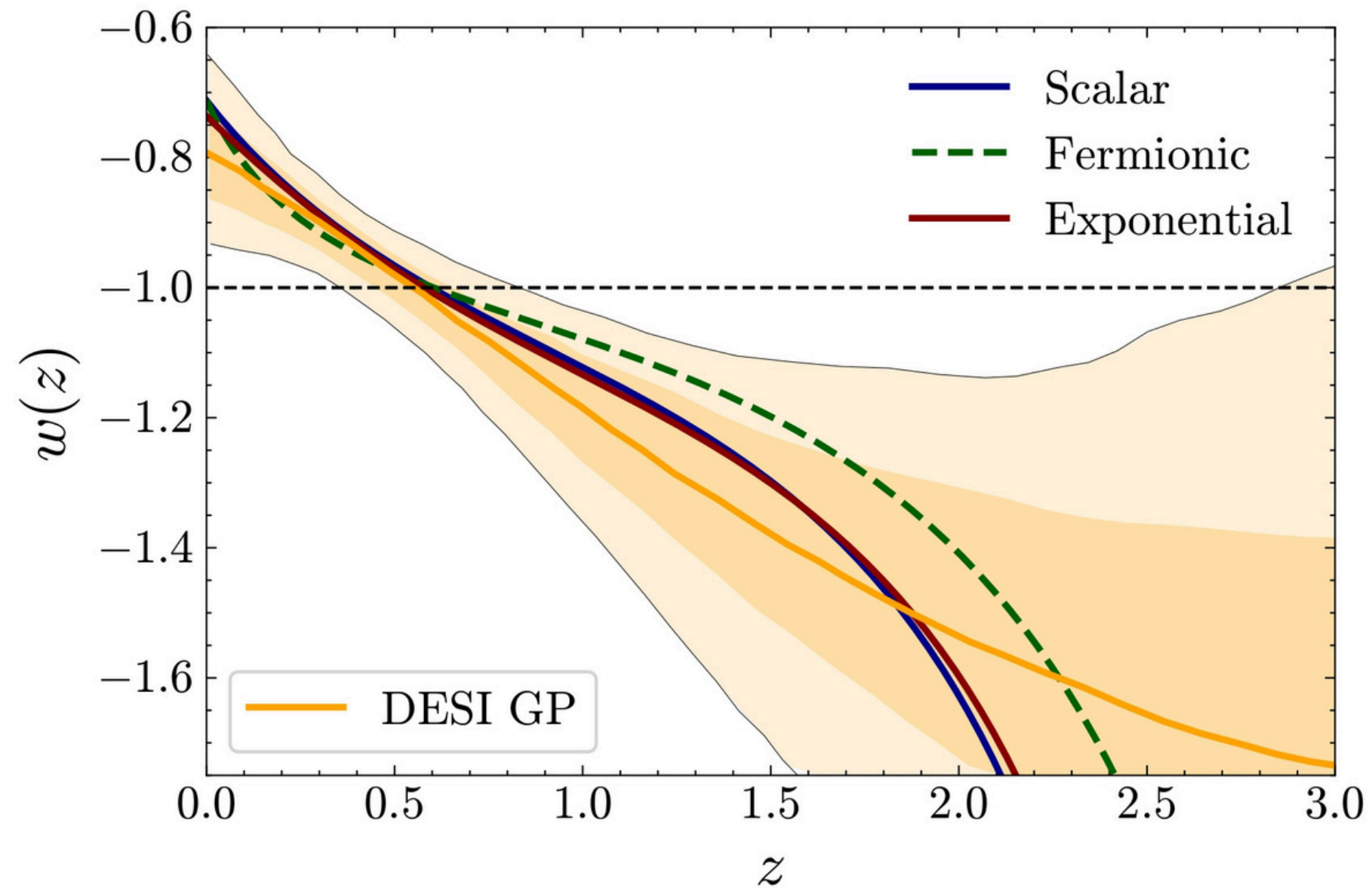
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But there's also another catch!

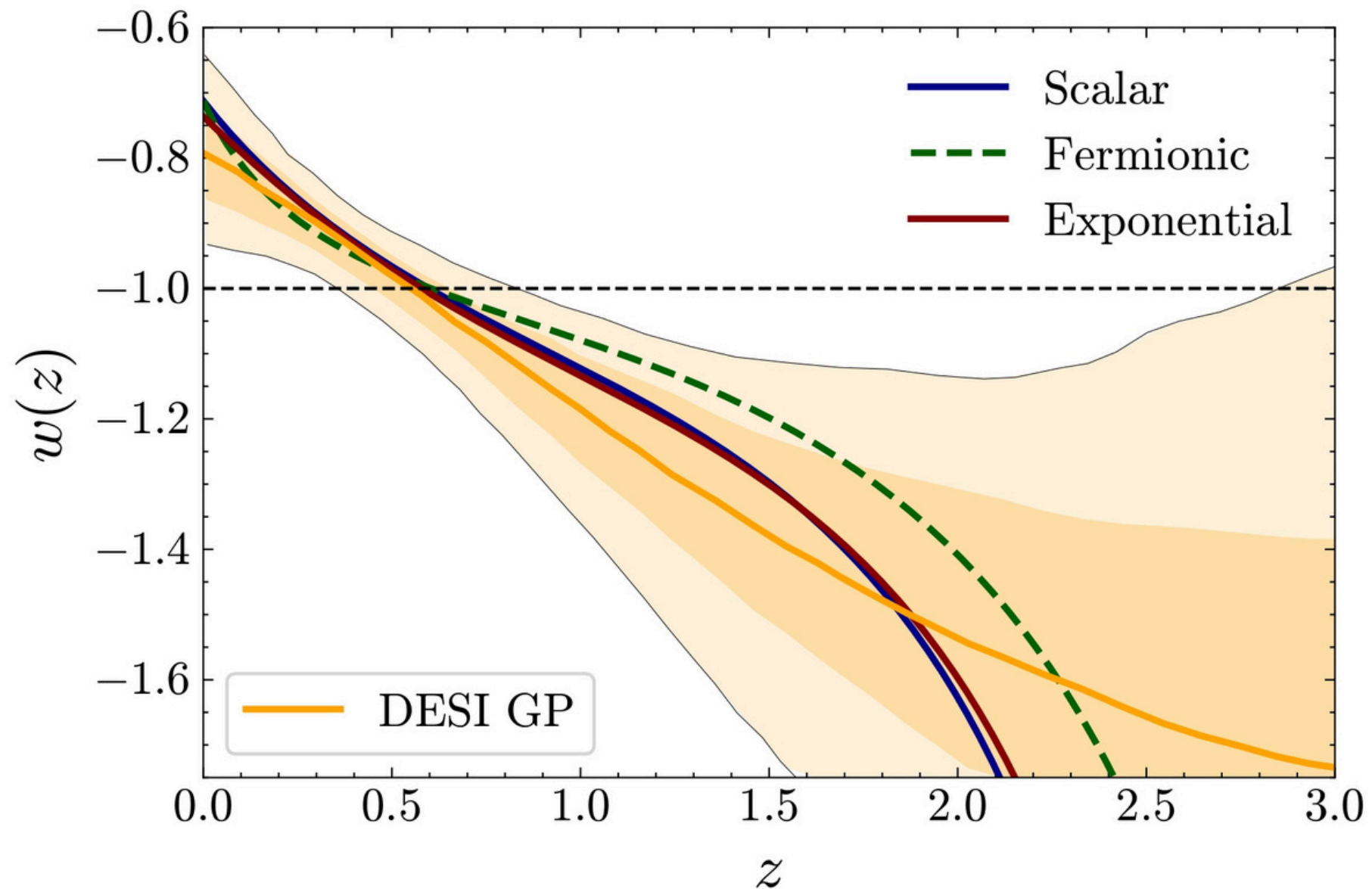
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Equation of state vs Redshift for the 3 models. The DESI Gaussian process (GP) reconstructed 1σ and 2σ contours are given in yellow.

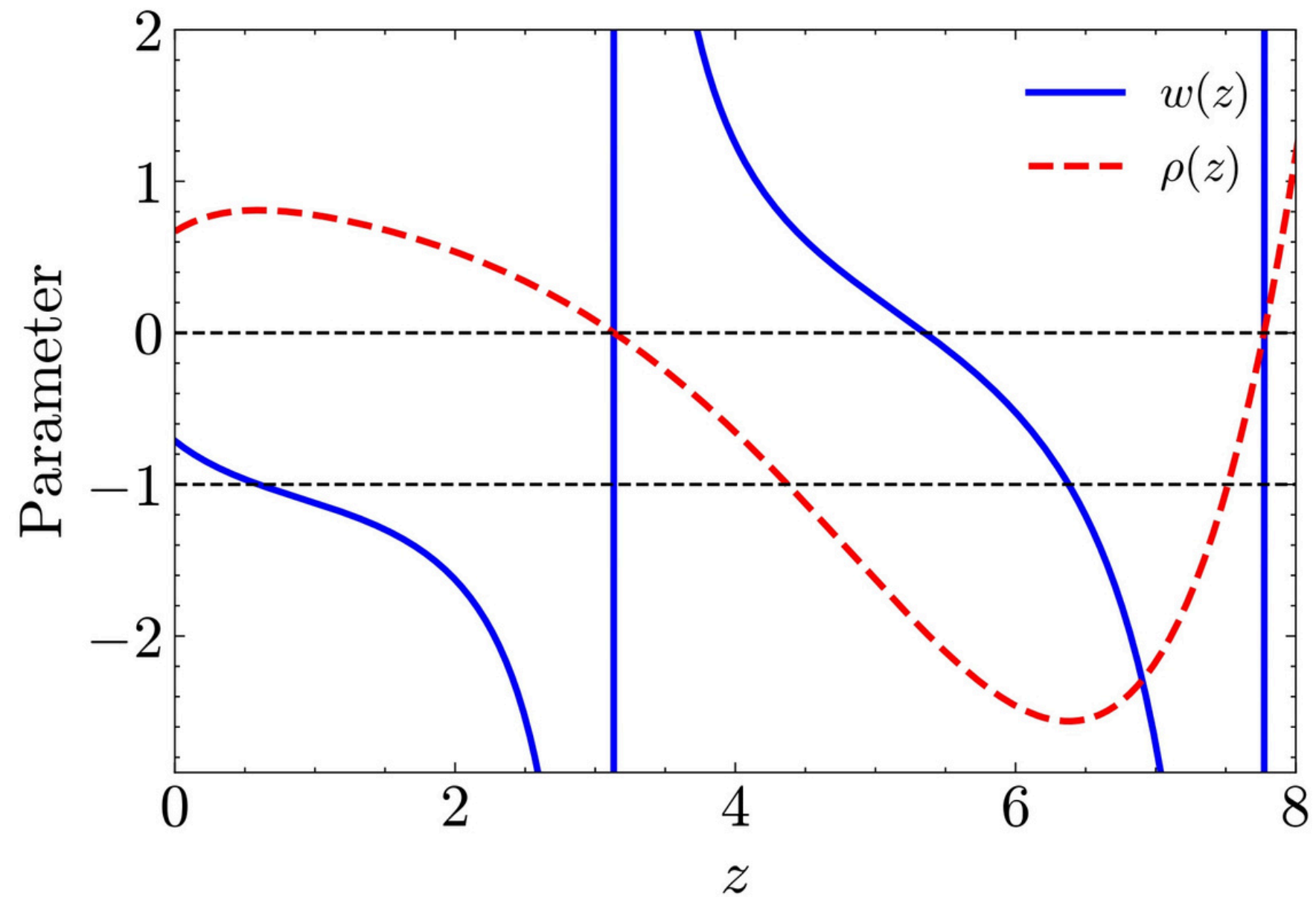
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The 3 models cross into the phantom regime around redshift 0.5 just like DESI and recreate the evolution history upto $z \approx 2.5$.

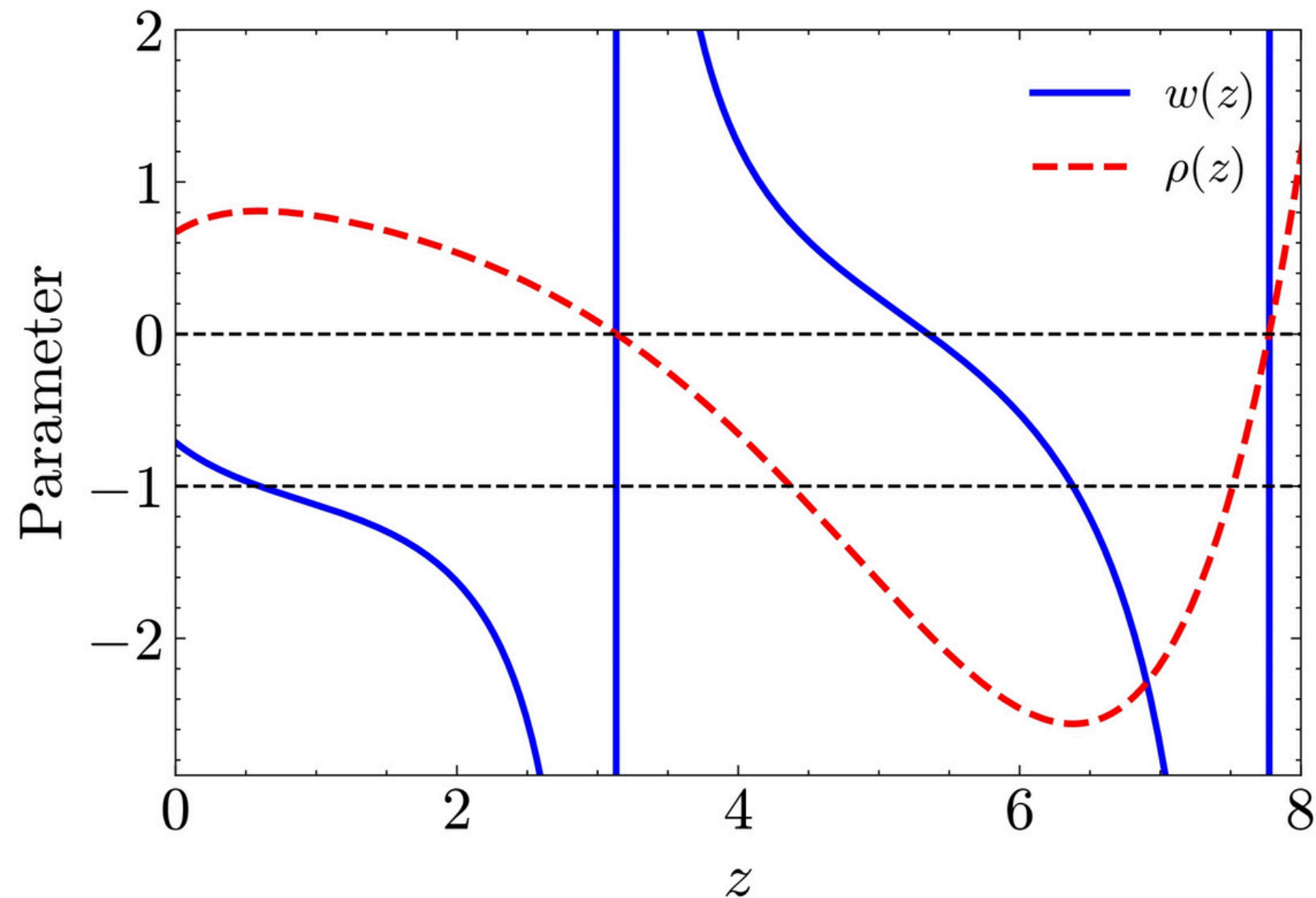
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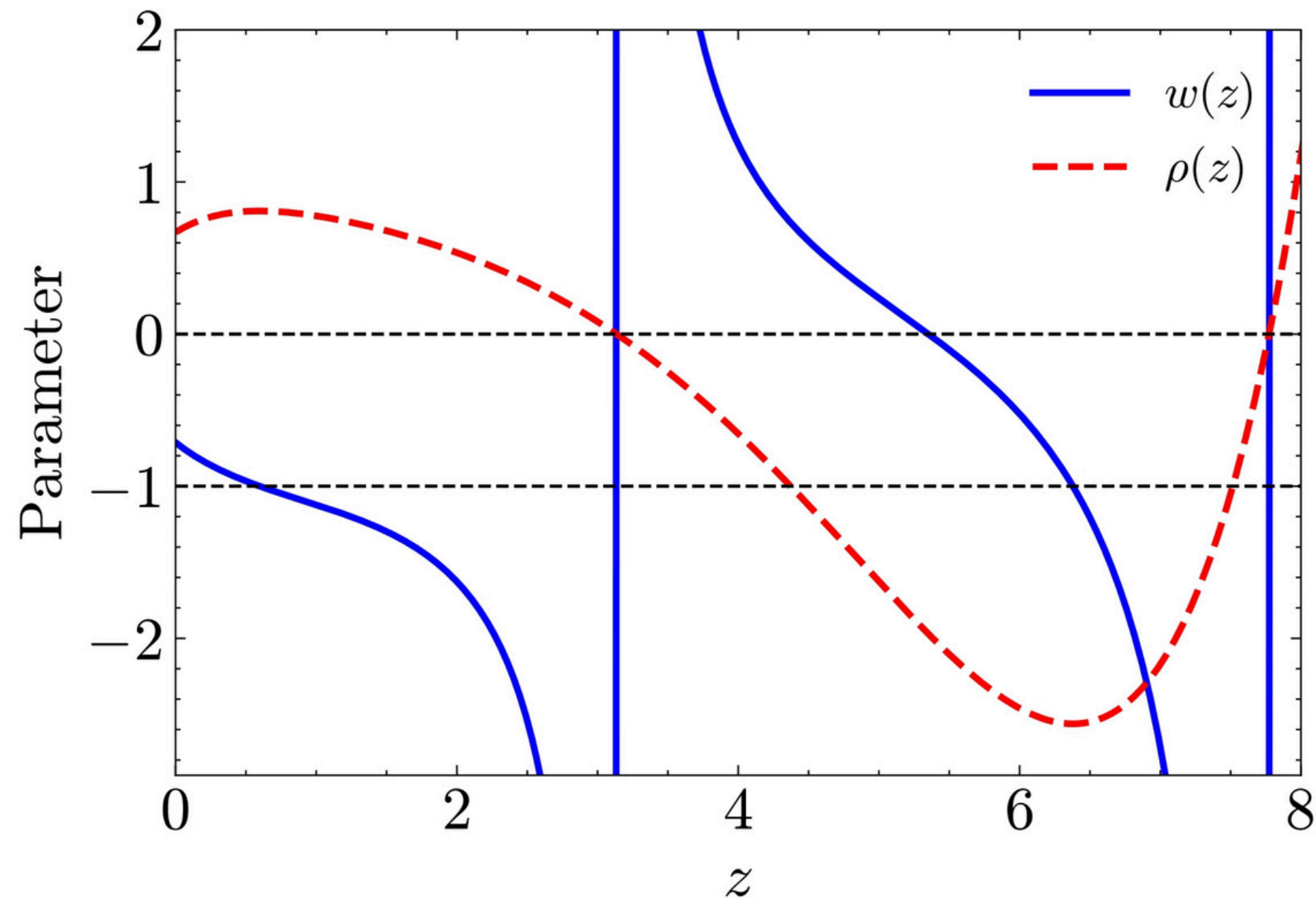
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But wait, there's more!

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Model	β
Scalar	$\gtrsim 0.6$
Fermionic	0.03
Exponential	0.006 – 0.01

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Discussion

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- The scalar interaction is ruled out by structure formation constraints and the fermionic one is at the verge. Exponential coupling small enough to escape them.
- Almost all phantom models make the **Hubble tension worse**. True even for parametrizations :(.

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 - Systematic Errors? [2504.16932]

Thank You!