

# This one weird trick could solve the $\sigma_8$ tension!

(You will never guess what!)

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Markus R. Mosbech

February 19, 2021

University of Sydney

1. Introduction

2. The Boltzmann Equation

3. Results

# Introduction

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## The full Boltzmann hierarchy for dark matter-massive neutrino interactions

Markus R. Mosbech,<sup>a</sup> Celine Boehm,<sup>a</sup> Steen Hannestad,<sup>b</sup> Olga Mena,<sup>c</sup> Julia Stadler,<sup>d</sup> and Yvonne Y. Y. Wong<sup>e</sup>

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The impact of dark matter-neutrino interactions on the measurement of the cosmological parameters has been investigated in the past in the context of massless neutrinos exclusively. Here we revisit the role of a neutrino-dark matter coupling in light of ongoing cosmological tensions by implementing the full Boltzmann hierarchy for three massive neutrinos. Our tightest 95% CL upper limit on the strength of the interactions, parameterized via  $u_k = \frac{\sigma_{\nu\chi}}{m_{\nu}} \left(\frac{m_{\nu}}{m_{\chi}}\right)^{-1}$ , is  $u_k \leq 3.34 \cdot 10^{-4}$ , arising from a combination of Planck TTTEEE data, Planck lensing data and SDSS BAO data. This upper bound is, as expected, slightly higher than previous results for interacting massless neutrinos, due to the correction factor associated with neutrino masses. We find that these interactions significantly relax the lower bounds on the value of  $\sigma_8$  that is inferred in the context of  $\Lambda$ CDM from the Planck data, leading to agreement within 1-2 $\sigma$  with weak lensing estimates of  $\sigma_8$ , as those from KiDS-1000. However, the presence of these interactions barely affects the value of the Hubble constant  $H_0$ .

- Boltzmann Hierarchy

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- Boltzmann Hierarchy
- Class

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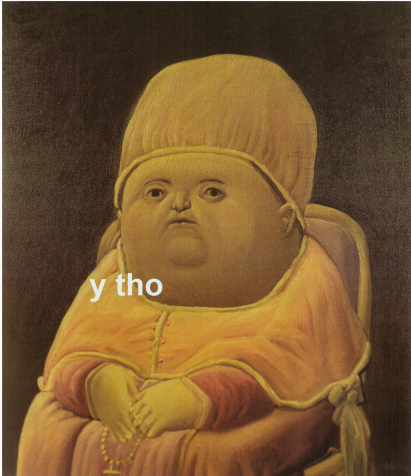
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- Boltzmann Hierarchy
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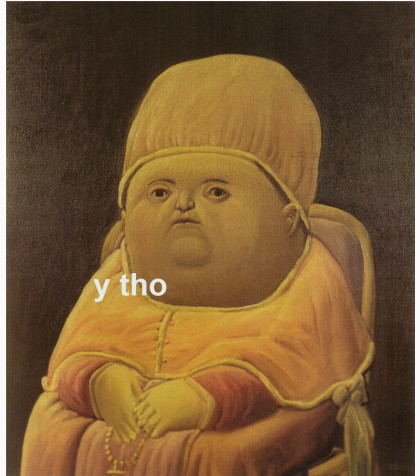
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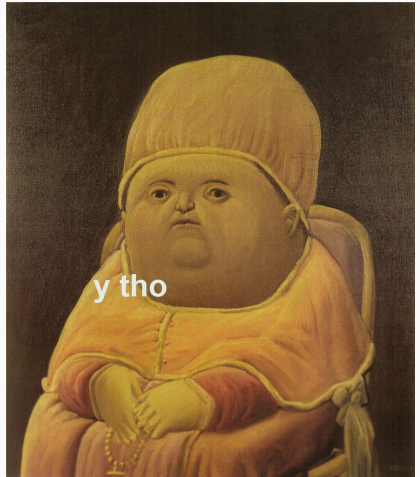
# Why this?

- Why not?



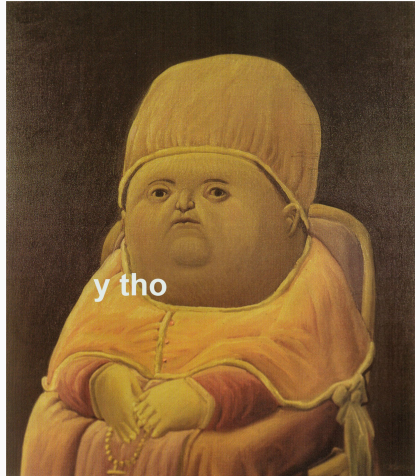
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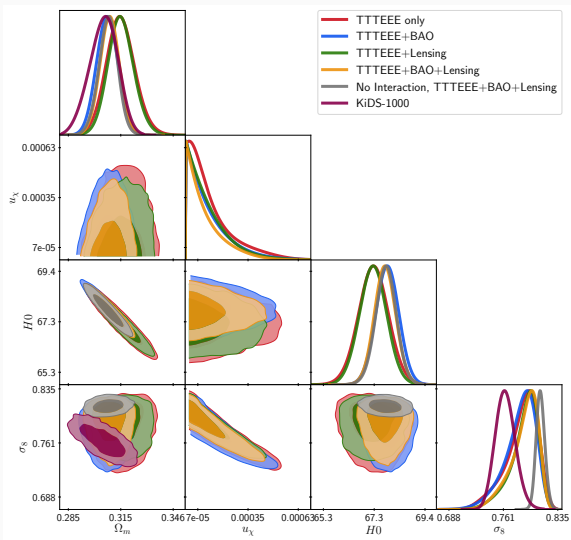


# Why this?

- Why not?
- Cosmo parameters -  $H_0$ ,  $\sigma_8$
- Neutrinos and DM



# Outcome



# The Boltzmann Equation

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# General Boltzmann Equation

$$P^\alpha \frac{\partial f}{\partial X^\alpha} - \Gamma_{\alpha\beta}^\gamma P^\alpha P^\beta \frac{\partial f}{\partial P^\gamma} = m \left( \frac{\partial f}{\partial \tau} \right)_C$$

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$$\frac{\partial \Psi}{\partial \tau} + i \frac{p}{E} (\mathbf{k} \cdot \hat{\mathbf{n}}) \Psi + \frac{d \ln f^{(0)}(p)}{d \ln p} \left[ \dot{\phi} - i \frac{E}{p} (\hat{\mathbf{k}} \cdot \hat{\mathbf{n}}) \psi \right] = \frac{1}{f_0} \left( \frac{\partial f}{\partial \tau} \right)_c$$



$$\dot{\delta}_{\text{cdm}} = -\theta_{\text{cdm}} + 3\dot{\phi}$$

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$$\dot{\mu}_{\chi}(\theta_{\nu} - \theta_{\chi}) = \frac{3}{4}k \frac{\int p^2 dp p f^{(0)}(p) C_{\chi}(p) \left( \frac{\theta_{\chi} E_{\nu}(p)}{3k f^{(0)}(p)} \frac{df^{(0)}(p)}{dp} + \Psi_1 \right)}{\int p^2 dp p f^{(0)}(p)}$$

$$\frac{\partial \Psi_0}{\partial \tau} = -\frac{pk}{E_\nu(p)} \Psi_1 - \phi \frac{d \ln f^{(0)}(p)}{d \ln p}$$

$$\frac{\partial \Psi_1}{\partial \tau} = \frac{1}{3} \frac{pk}{E_\nu(p)} (\Psi_0 - 2\Psi_2) - \frac{E_\nu(p) k}{3p} \psi \frac{d \ln f^{(0)}(p)}{d \ln p}$$

$$\frac{\partial \Psi_l}{\partial \tau} = \frac{1}{2l+1} \frac{pk}{E_\nu(p)} (l\Psi_{l-1} - (l+1)\Psi_{l+1}), \quad l \geq 2$$

$$\frac{\partial \Psi_1}{\partial \tau} = [\dots] - C_x \frac{v_x E_\nu(p)}{3f^{(0)}(p)} \frac{df^{(0)}(p)}{dp} - C_x \Psi_1,$$

$$\frac{\partial \Psi_2}{\partial \tau} = [\dots] - \frac{9}{10} C_x \Psi_2,$$

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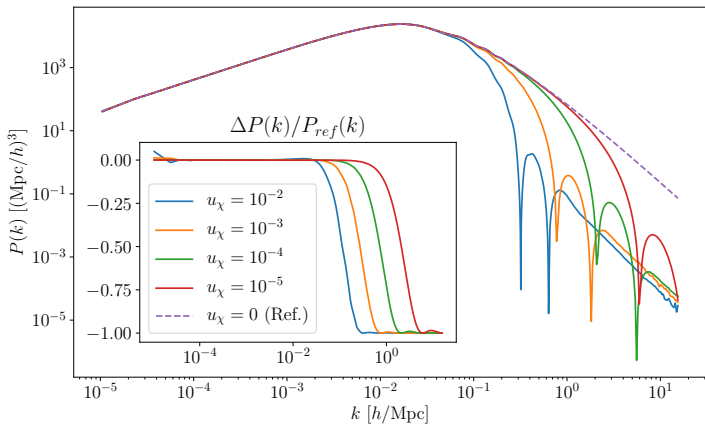
$$C_X = a u_{\nu X} \frac{\sigma_{\text{Th}} \rho_X}{100 \text{ GeV}} \frac{p^2}{E_\nu^2}, \quad u_{\nu X} = \frac{\sigma_0}{\sigma_{\text{Th}}} \left( \frac{m_X}{100 \text{ GeV}} \right)^{-1}$$

# Results

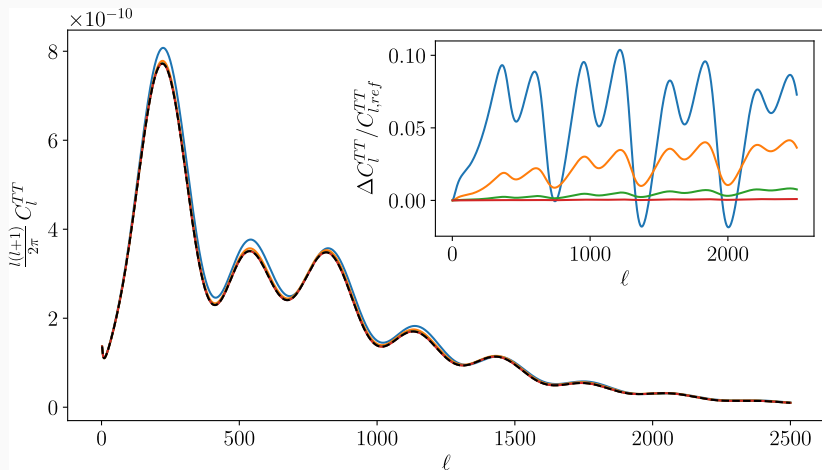
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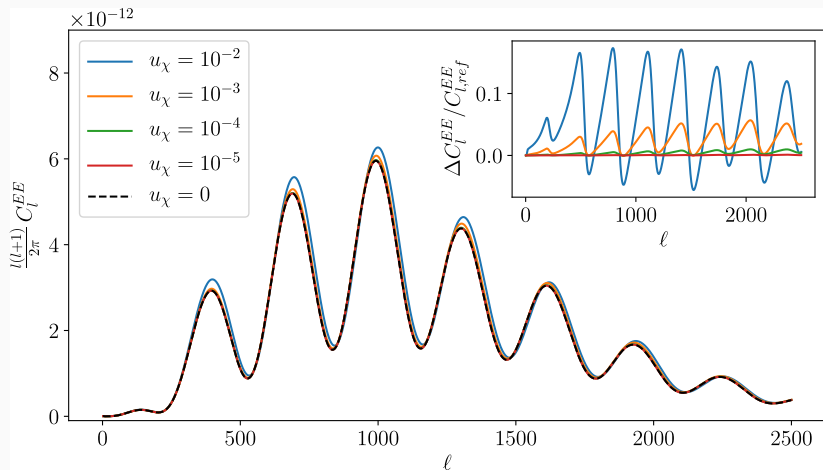
# Matter power spectrum



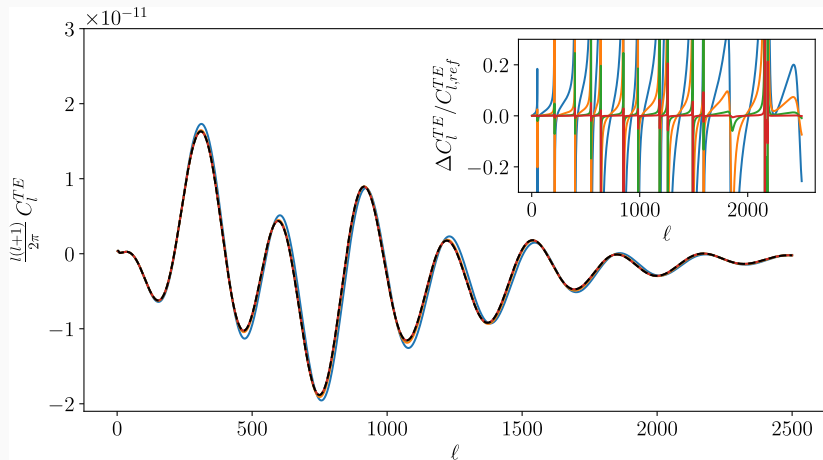
# CMB power spectrum



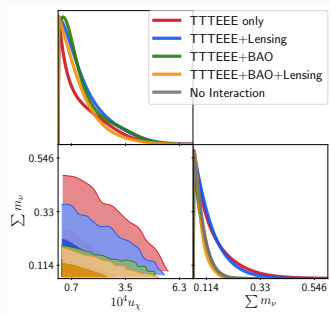
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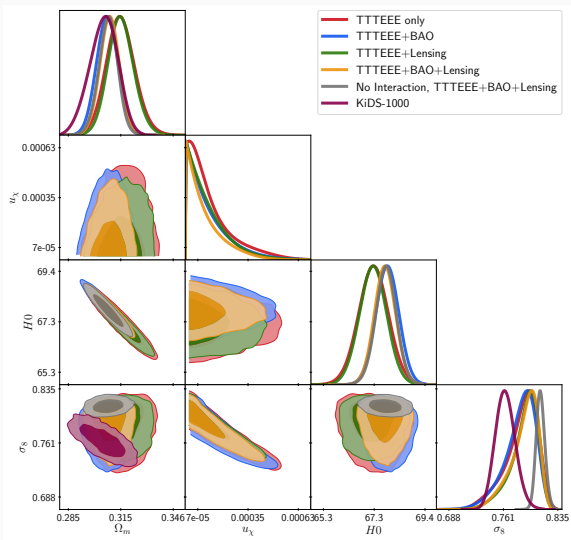
# CMB power spectrum



# Posteriors



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# Boring numbers

	Planck + Lensing + BAO
$100 \omega_b$	$2.24^{+0.03}_{-0.03}$
$\omega_{DM}$	$0.119^{+0.002}_{-0.002}$
$100 \theta_s$	$1.0419^{+0.0010}_{-0.0004}$
$\ln 10^{10} A_s$	$3.05^{+0.03}_{-0.03}$
$n_s$	$0.967^{+0.007}_{-0.010}$
$\tau_{reio}$	$0.057^{+0.017}_{-0.014}$
$u_\chi$	$3.34 \cdot 10^{-4}$
$\sum m_\nu$ [eV]	0.14
$H_0$ [km/s/Mpc]	$67.6^{+1.0}_{-1.0}$
$\sigma_8$	$0.81^{+0.01}_{-0.06}$

*fin*

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